
galpy Documentation

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galpy is a Python 2 and 3 package for galactic dynamics. It supports orbit integration in a variety of potentials, evaluating and sampling various distribution functions, and the calculation of action-angle coordinates for all static potentials. galpy is an [astropy affiliated package](#) and provides full support for astropy's [Quantity](#) framework for variables with units.

galpy is developed on GitHub. If you are looking to [report an issue](#) or for information on how to [contribute to the code](#), please head over to [galpy's GitHub page](#) for more information.

As a preview of the kinds of things you can do with galpy, here's an [animation](#) of the orbit of the Sun in galpy's MWPotential2014 potential over 7 Gyr:

1.1 Installation

1.1.1 With conda

The easiest way to install the latest released version of galpy is using conda:

```
conda install galpy -c conda-forge
```

or:

```
conda config --add channels conda-forge
conda install galpy
```

1.1.2 With pip

galpy can also be installed using pip. Some advanced features require the GNU Scientific Library (GSL; [see below](#)). If you want to use these, install the GSL first (or install it later and re-install using the upgrade command above). Then do:

```
pip install galpy
```

or to upgrade without upgrading the dependencies:

```
pip install -U --no-deps galpy
```

1.1.3 Latest version

The latest updates in galpy can be installed using:

```
pip install -U --no-deps git+git://github.com/jobovy/galpy.git#egg=galpy
```

or:

```
pip install -U --no-deps --install-option="--prefix=~/.local" git+git://github.com/
↪jobovy/galpy.git#egg=galpy
```

for a local installation. The latest updates can also be installed from the source code downloaded from github using the standard python `setup.py` installation:

```
python setup.py install
```

or:

```
python setup.py install --prefix=~/.local
```

for a local installation.

1.1.4 Installing from a branch

If you want to use a feature that is currently only available in a branch, do:

```
pip install -U --no-deps git+git://github.com/jobovy/galpy.git@dev#egg=galpy
```

to, for example, install the dev branch.

1.1.5 Installing from source on Windows

Versions >1.3 should be able to be compiled on Windows systems using the Microsoft Visual Studio C compiler (>= 2015). For this you need to first install the GNU Scientific Library (GSL), for example using Anaconda ([see below](#)). Similar to on a UNIX system, you need to set paths to the header and library files where the GSL is located. On Windows this is done as:

```
set INCLUDE=%CONDA_PREFIX%\Library\include;%INCLUDE%
set LIB=%CONDA_PREFIX%\Library\lib;%LIB%
set LIBPATH=%CONDA_PREFIX%\Library\lib;%LIBPATH%
```

where in this example `CONDA_PREFIX` is the path of your current conda environment (the path that ends in `\ENV_NAME`). If you have installed the GSL somewhere else, adjust these paths (but do not use `YOUR_PATH\include\gsl` or `YOUR_PATH\lib\gsl` as the paths, simply use `YOUR_PATH\include` and `YOUR_PATH\lib`).

To compile with OpenMP on Windows, you have to install Intel OpenMP via:

```
conda install -c anaconda intel-openmp
```

and then to compile the code:

```
python setup.py install
```

If you encounter any issue related to OpenMP during compilation, you can do:

```
python setup.py install --no-openmp
```


1.1.6 Installing the TorusMapper code

Since v1.2, `galpy` contains a basic interface to the TorusMapper code of [Binney & McMillan \(2016\)](#). This interface uses a stripped-down version of the TorusMapper code, that is not bundled with the `galpy` code, but kept in a fork of the original TorusMapper code. Installation of the TorusMapper interface is therefore only possible when installing from source after downloading or cloning the `galpy` code and using the `python setup.py install` method above.

To install the TorusMapper code, *before* running the installation of `galpy`, navigate to the top-level `galpy` directory (which contains the `setup.py` file) and do:

```
git clone https://github.com/jobovy/Torus.git galpy/actionAngle/actionAngleTorus_c_
→ext/torus
cd galpy/actionAngle/actionAngleTorus_c_ext/torus
git checkout galpy
cd -
```

Then proceed to install `galpy` using the `python setup.py install` technique or its variants as usual.

1.1.7 Installation FAQ

What is the required `numpy` version?

`galpy` should mostly work for any relatively recent version of `numpy`, but some advanced features, including calculating the normalization of certain distribution functions using Gauss-Legendre integration require `numpy` version 1.7.0 or higher.

I get warnings like “`galpyWarning: integrateFullOrbit_c extension module not loaded, because galpy_integrate_c.so image was not found`”

This typically means that the GNU Scientific Library ([GSL](#)) was unavailable during `galpy`’s installation, causing the C extensions not to be compiled. Most of the `galpy` code will still run, but slower because it will run in pure Python. The code requires `GSL` versions ≥ 1.14 . If you believe that the correct `GSL` version is installed for `galpy`, check that the library can be found during installation (see [below](#)).

I get the warning “`galpyWarning: actionAngleTorus_c extension module not loaded, because galpy_actionAngleTorus_c.so image was not found`”

This is typically because the TorusMapper code was not compiled, because it was unavailable during installation. This code is only necessary if you want to use `galpy.actionAngle.actionAngleTorus`. See [above](#) for instructions on how to install the TorusMapper code.

How do I install the `GSL`?

Certain advanced features require the GNU Scientific Library ([GSL](#)), with action calculations requiring version 1.14 or higher. The easiest way to install this is using its Anaconda build:

```
conda install -c conda-forge gsl
```

If you do not want to go that route, on a Mac, the next easiest way to install the `GSL` is using [Homebrew](#) as:

```
brew install gsl --universal
```

You should be able to check your version using (on Mac/Linux):

```
gsl-config --version
```

On Linux distributions with `apt-get`, the GSL can be installed using:

```
apt-get install libgsl0-dev
```

The galpy installation fails because of C compilation errors

galpy's installation can fail due to compilation errors, which look like:

```
error: command 'gcc' failed with exit status 1
```

or:

```
error: command 'clang' failed with exit status 1
```

or:

```
error: command 'cc' failed with exit status 1
```

This is typically because the compiler cannot locate the GSL header files or the GSL library. You can tell the installation about where you've installed the GSL library by defining (for example, when the GSL was installed under `/usr`; the `LD_LIBRARY_PATH` part of this may or may not be necessary depending on your system):

```
export CFLAGS=-I/usr/include
export LDFLAGS=-L/usr/lib
export LD_LIBRARY_PATH=-L/usr/lib
```

or:

```
setenv CFLAGS -I/usr/include
setenv LDFLAGS -L/usr/lib
setenv LD_LIBRARY_PATH -L/usr/lib
```

depending on your shell type (change the actual path to the include and lib directories that have the gsl directory). If you already have `CFLAGS`, `LDFLAGS`, and `LD_LIBRARY_PATH` defined you just have to add the `'-I/usr/include'` and `'-L/usr/lib'` to them.

If you are on a Mac or UNIX system (e.g., Linux), you can find the correct `CFLAGS` and `LDFLAGS/LD_LIBRARY_PATH` entries by doing:

```
gsl-config --cflags
gsl-config --libs
```

(don't include the `-lgsl` `lgslcblas` portion of the latter output.)

I have defined `CFLAGS`, `LDFLAGS`, and `LD_LIBRARY_PATH`, but the compiler does not seem to include these and still returns with errors

This typically happens if you install using `sudo`, but have defined the `CFLAGS` etc. environment variables without using `sudo`. Try using `sudo -E` instead, which propagates your own environment variables to the `sudo` user.

I'm having issues with OpenMP

galpy uses [OpenMP](#) to parallelize various of the computations done in C. galpy can be installed without OpenMP by specifying the option `--no-openmp` when running the `python setup.py` commands above:

```
python setup.py install --no-openmp
```

or when using pip as follows:

```
pip install -U --no-deps --install-option="--no-openmp" git+git://github.com/jobovy/  
↪galpy.git#egg=galpy
```

or:

```
pip install -U --no-deps --install-option="--prefix=~/.local" --install-option="--no-  
↪openmp" git+git://github.com/jobovy/galpy.git#egg=galpy
```

for a local installation. This might be useful if one is using the `clang` compiler, which is the new default on macs with OS X (≥ 10.8), but does not support OpenMP. `clang` might lead to errors in the installation of galpy such as:

```
ld: library not found for -lgomp  
  
clang: error: linker command failed with exit code 1 (use -v to see invocation)
```

If you get these errors, you can use the commands given above to install without OpenMP, or specify to use `gcc` by specifying the `CC` and `LD_SHARED` environment variables to use `gcc`. Note that `clang` does not seem to have this issue anymore in more recent versions, but it still does not support OpenMP.

1.1.8 Configuration file

Since v1.2, galpy uses a configuration file to set a small number of configuration variables. This configuration file is parsed using [ConfigParser/configparser](#). It is currently used:

- to set a default set of distance and velocity scales (`ro` and `vo` throughout galpy) for conversion between physical and internal galpy unit
- to decide whether to use seaborn plotting with galpy's defaults (which affects *all* plotting after importing `galpy.util.bovy_plot`),
- to specify whether output from functions or methods should be given as an [astropy Quantity](#) with units as much as possible or not, and whether or not to use astropy's [coordinate transformations](#) (these are typically somewhat slower than galpy's own coordinate transformations, but they are more accurate and more general)
- to set the level of verbosity of galpy's warning system (the default `verbose=False` turns off non-crucial warnings).

The current configuration file therefore looks like this:

```
[normalization]  
ro = 8.  
vo = 220.  
  
[plot]  
seaborn-bovy-defaults = False  
  
[astropy]  
astropy-units = False
```

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```
astropy-coords = True

[warnings]
verbose = False
```

where `ro` is the distance scale specified in kpc, `vo` the velocity scale in km/s, and the setting is to *not* return output as a Quantity. These are the current default settings.

A user-wide configuration file should be located at `$HOME/.galpyrc`. This user-wide file can be overridden by a `$PWD/.galpyrc` file in the current directory. If no configuration file is found, the code will automatically write the default configuration to `$HOME/.galpyrc`. Thus, after installing galpy, you can simply use some of its simplest functionality (e.g., integrate an orbit), and after this the default configuration file will be present at `$HOME/.galpyrc`. If you want to change any of the settings (for example, you want Quantity output), you can edit this file. The default configuration file can also be found [here](#).

1.2 What's new?

This page gives some of the key improvements in each galpy version. See the `HISTORY.txt` file in the galpy source for full details on what is new and different in each version.

1.2.1 v1.4

- Added dynamical friction as the `ChandrasekharDynamicalFrictionForce` class, an implementation of dynamical friction based on the classical Chandrasekhar formula (with recent tweaks from the literature to better represent the results from N-body simulations).
- A general `EllipsoidalPotential` superclass for implementing potentials with densities that are constant on ellipsoids (functions of $m^2 = x^2 + y^2/b^2 + z^2/c^2$). Also implemented in C. Implementing new types of ellipsoidal potentials now only requires three simple functions to be defined: the density as a function of `m`, its derivative with respect to `m`, and its integral with respect to `m^2`. Makes implementing any ellipsoidal potential a breeze. See examples in the new-potentials section below.
- New or improved potentials and *potential wrappers*:
 - `CorotatingRotationWrapperPotential`: wrapper to make a pattern (e.g., a `SpiralArmsPotential`) wind up over time such that it is always corotating (see [Hunt et al. \(2018\)](#) for an example of this).
 - `GaussianAmplitudeWrapperPotential`: wrapper to modulate the amplitude of a (list of) `Potential` (s) with a Gaussian.
 - `PerfectEllipsoidPotential`: Potential of a perfect triaxial ellipsoid ([de Zeeuw 1985](#)).
 - `SphericalShellPotential`: Potential of a thin, spherical shell.
 - `RingPotential`: Potential of a circular ring.
 - Re-implemented `TwoPowerTriaxialPotential`, `TriaxialHernquistPotential`, `TriaxialJaffePotential`, and `TriaxialNFWPotential` using the general `EllipsoidalPotential` class.
- New `Potential` methods and functions:
 - Use nested lists of `Potential` instances wherever lists of `Potential` instances can be used. Allows easy adding of components (e.g., a bar) to previously defined potentials (which may be lists themselves): `new_pot= [pot,bar_pot]`.

- `rtide` and `ttenor`: compute the tidal radius of an object and the full tidal tensor.
- `surfdens` method and `evaluateSurfaceDensities` function to evaluate the surface density up to a given z .
- `r2deriv` and `evaluator2derivs`: 2nd derivative wrt spherical radius.
- `evaluatephi2derivs`: second derivative wrt ϕ .
- `evaluateRphiderivs`: mixed (R, ϕ) derivative.
- New or improved `galpy.orbit.Orbit` functionality and methods:
 - `Orbit.from_name` to initialize an `Orbit` instance from an object's name. E.g., `orb= Orbit.from_name('LMC')`.
 - Orbit initialization without arguments is now the orbit of the Sun.
 - Orbits can be initialized with a `SkyCoord`.
 - Default `solarmotion=` parameter is now 'schoenrich' for the Solar motion of [Schoenrich et al. \(2010\)](#).
 - `rguiding`: Guiding-center radius.
 - `Lz`: vertical component of the angular momentum.
 - If `astropy` version > 3 , `Orbit.SkyCoord` method returns a `SkyCoord` object that includes the velocity information and the Galactocentric frame used by the `Orbit` instance.
- `galpy.df.jeans` module with tools for Jeans modeling. Currently only contains the functions `sigmar` and `sigmalos` to calculate the velocity dispersion in the radial or line-of-sight direction using the spherical Jeans equation in a given potential, density profile, and anisotropy profile (anisotropy can be radially varying).
- Support for compilation on Windows with MSVC.

1.2.2 v1.3

- A fast and precise method for approximating an orbit's eccentricity, peri- and apocenter radii, and maximum height above the midplane using the Staekel approximation (see [Mackereth & Bovy 2018](#)). Can determine these parameters to better than a few percent accuracy in as little as $10 \mu s$ per object, more than 1,000 times faster than through direct orbit integration. See [this section](#) of the documentation for more info.
- A general method for modifying `Potential` classes through potential wrappers—simple classes that wrap existing potentials to modify their behavior. See [this section](#) of the documentation for examples and [this section](#) for information on how to easily define new wrappers. Example wrappers include `SolidBodyRotationWrapperPotential` to allow *any* potential to rotate as a solid body and `DehnenSmoothWrapperPotential` to smoothly grow *any* potential. See [this section of the galpy.potential API page](#) for an up-to-date list of wrappers.
- New or improved potentials:
 - `DiskSCFPotential`: a general Poisson solver well suited for galactic disks
 - Bar potentials `SoftenedNeedleBarPotential` and `FerrersPotential` (latter only in Python for now)
 - 3D spiral arms model `SpiralArmsPotential`
 - Henon & Heiles (1964) potential `HenonHeilesPotential`
 - Triaxial version of `LogarithmicHaloPotential`
 - 3D version of `DehnenBarPotential`
 - Generalized version of `CosmphiDiskPotential`
- New or improved `galpy.orbit.Orbit` methods:

- Method to display an animation of an integrated orbit in jupyter notebooks: `Orbit.animate`. See [this section](#) of the documentation.
- Improved default method for fast calculation of eccentricity, `zmax`, `rperi`, `rap`, actions, frequencies, and angles by switching to the Staeckel approximation with automatically-estimated approximation parameters.
- Improved plotting functions: plotting of spherical radius and of arbitrary user-supplied functions of time in `Orbit.plot`, `Orbit.plot3d`, and `Orbit.animate`.
- `actionAngleStaeckel` upgrades:
 - `actionAngleStaeckel` methods now allow for different focal lengths `delta` for different phase-space points and for the order of the Gauss-Legendre integration to be specified (default: 10, which is good enough when using `actionAngleStaeckel` to compute approximate actions etc. for an axisymmetric potential).
 - Added an option to the `estimateDeltaStaeckel` function to facilitate the return of an estimated `delta` parameter at every phase space point passed, rather than returning a median of the estimate at each point.
- `galpy.df.schwarzschilddf`: the simple Schwarzschild distribution function for a razor-thin disk (useful for teaching).

1.2.3 v1.2

- Full support for providing inputs to all initializations, methods, and functions as `astropy Quantity` with `units` and for providing outputs as `astropy Quantities`.
- `galpy.potential.TwoPowerTriaxialPotential`, a set of triaxial potentials with iso-density contours that are arbitrary, similar, coaxial ellipsoids whose ‘radial’ density is a (different) power-law at small and large radii: $1/m^\alpha/(1+m)^\beta$ (the triaxial generalization of `TwoPowerSphericalPotential`, with flattening in the density rather than in the potential; includes triaxial Hernquist and NFW potentials).
- `galpy.potential.SCFPotential`, a class that implements general density/potential pairs through the basis expansion approach to solving the Poisson equation of Hernquist & Ostriker (1992). Also implemented functions to compute the coefficients for a given density function. See more explanation [here](#).
- `galpy.actionAngle.actionAngleTorus`: an experimental interface to Binney & McMillan’s `TorusMapper` code for computing positions and velocities for given actions and angles. See the installation instructions for how to properly install this. See [this section](#) and the `galpy.actionAngle` API page for documentation.
- `galpy.actionAngle.actionAngleIsochroneApprox` (Bovy 2014) now implemented for the general case of a time-independent potential.
- `galpy.df.streamgapdf`, a module for modeling the effect of a dark-matter subhalo on a tidal stream. See [Sanders et al. \(2016\)](#). Also includes the fast methods for computing the density along the stream and the stream track for a perturbed stream from [Bovy et al. \(2016\)](#).
- `Orbit.flip` can now flip the velocities of an orbit in-place by specifying `inplace=True`. This allows correct velocities to be easily obtained for backwards-integrated orbits.
- `galpy.potential.PseudoIsothermalPotential`, a standard pseudo-isothermal-sphere potential. `galpy.potential.KuzminDiskPotential`, a razor-thin disk potential.
- Internal transformations between equatorial and Galactic coordinates are now performed by default using `astropy’s coordinates` module. Transformation of (ra,dec) to Galactic coordinates for general epochs.

1.2.4 v1.1

- Full support for Python 3.
- `galpy.potential.SnapshotRZPotential`, a potential class that can be used to get a frozen snapshot of the potential of an N-body simulation.
- Various other potentials: `PlummerPotential`, a standard Plummer potential; `MN3ExponentialDiskPotential`, an approximation to an exponential disk using three Miyamoto-Nagai potentials (Smith et al. 2015); `KuzminKutuzovStaeckelPotential`, a Staeckel potential that can be used to approximate the potential of a disk galaxy (Batsleer & Dejonghe 1994).
- Support for converting potential parameters to NEMO format and units.
- Orbit fitting in custom sky coordinates.

1.3 Introduction

The most basic features of galpy are its ability to display rotation curves and perform orbit integration for arbitrary combinations of potentials. This section introduces the most basic features of `galpy.potential` and `galpy.orbit`.

1.3.1 Rotation curves

The following code example shows how to initialize a Miyamoto-Nagai disk potential and plot its rotation curve

```
>>> from galpy.potential import MiyamotoNagaiPotential
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,normalize=1.)
>>> mp.plotRotcurve(Range=[0.01,10.],grid=1001)
```

The `normalize=1.` option normalizes the potential such that the radial force is a fraction `normalize=1.` of the radial force necessary to make the circular velocity 1 at $R=1$. Starting in v1.2 you can also initialize potentials with amplitudes and other parameters in physical units; see below and other parts of this documentation.

Tip: You can copy all of the code examples in this documentation to your clipboard by clicking the button in the top, right corner of each example. This can be directly pasted into a Python interpreter (including the `>>>`).

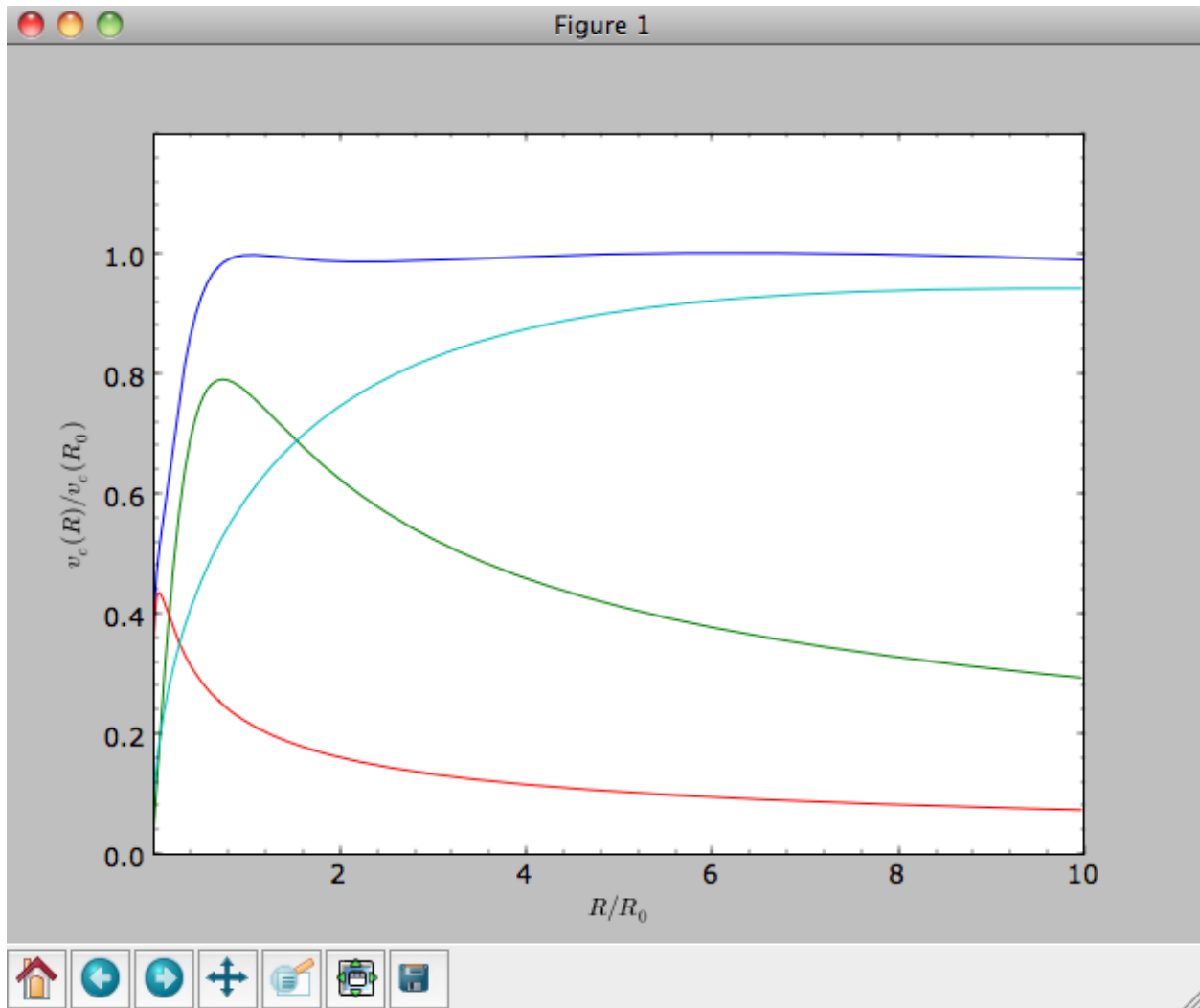
Similarly we can initialize other potentials and plot the combined rotation curve

```
>>> from galpy.potential import NFWPotential, HernquistPotential
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,normalize=.6)
>>> np= NFWPotential(a=4.5,normalize=.35)
>>> hp= HernquistPotential(a=0.6/8,normalize=0.05)
>>> from galpy.potential import plotRotcurve
>>> plotRotcurve([hp,mp,np],Range=[0.01,10.],grid=1001,yrange=[0.,1.2])
```

Note that the `normalize` values add up to 1. such that the circular velocity will be 1 at $R=1$. The resulting rotation curve is approximately flat. To show the rotation curves of the three components do

```
>>> mp.plotRotcurve(Range=[0.01,10.],grid=1001,overplot=True)
>>> hp.plotRotcurve(Range=[0.01,10.],grid=1001,overplot=True)
>>> np.plotRotcurve(Range=[0.01,10.],grid=1001,overplot=True)
```

You'll see the following



As a shortcut the [hp, mp, np] Milky-Way-like potential is defined as

```
>>> from galpy.potential import MWPotential
```

This is *not* the recommended Milky-Way-like potential in galpy. The (currently) recommended Milky-Way-like potential is MWPotential2014:

```
>>> from galpy.potential import MWPotential2014
```

MWPotential2014 has a more realistic bulge model and is actually fit to various dynamical constraints on the Milky Way (see [here](#) and the galpy paper).

1.3.2 Units in galpy

Internal (natural) units

Above we normalized the potentials such that they give a circular velocity of 1 at $R=1$. These are the standard galpy units (sometimes referred to as *natural units* in the documentation). galpy will work most robustly when using these

natural units. When using galpy to model a real galaxy with, say, a circular velocity of 220 km/s at $R=8$ kpc, all of the velocities should be scaled as $v = V/[220 \text{ km/s}]$ and all of the positions should be scaled as $x = X/[8 \text{ kpc}]$ when using galpy's natural units.

For convenience, a utility module `bovy_conversion` is included in galpy that helps in converting between physical units and natural units for various quantities. Alternatively, you can use the `astropy.units` module to specify inputs in physical units and get outputs with units (see the [next subsection](#) below). For example, in natural units the orbital time of a circular orbit at $R=1$ is 2π ; in physical units this corresponds to

```
>>> from galpy.util import bovy_conversion
>>> print(2.*numpy.pi*bovy_conversion.time_in_Gyr(220.,8.))
# 0.223405444283
```

or about 223 Myr. We can also express forces in various physical units. For example, for the Milky-Way-like potential defined in galpy, we have that the vertical force at 1.1 kpc is

```
>>> from galpy.potential import MWPotential2014, evaluatezforces
>>> -evaluatezforces(MWPotential2014, 1.,1.1/8.)*bovy_conversion.force_in_pcMyr2(220.,
↪8.)
# 2.0259181908629933
```

which we can also express as an equivalent surface-density by dividing by $2\pi G$

```
>>> -evaluatezforces(MWPotential2014, 1.,1.1/8.)*bovy_conversion.force_in_
↪2piGmsolpc2(220.,8.)
# 71.658016957792356
```

Because the vertical force at the solar circle in the Milky Way at 1.1 kpc above the plane is approximately $70 (2\pi G M_\odot \text{ pc}^{-2})$ (e.g., [2013arXiv1309.0809B](#)), this shows that our Milky-Way-like potential has a realistic disk (at least in this respect).

`bovy_conversion` further has functions to convert densities, masses, surface densities, and frequencies to physical units (actions are considered to be too obvious to be included); see [here](#) for a full list. As a final example, the local dark matter density in the Milky-Way-like potential is given by

```
>>> MWPotential2014[2].dens(1.,0.)*bovy_conversion.dens_in_msolpc3(220.,8.)
# 0.0075419566970079373
```

or

```
>>> MWPotential2014[2].dens(1.,0.)*bovy_conversion.dens_in_gevcc(220.,8.)
# 0.28643101789044584
```

or about $0.0075 M_\odot \text{ pc}^{-3} \approx 0.3 \text{ GeV cm}^{-3}$, in line with current measurements (e.g., [2012ApJ...756...89B](#)).

When galpy Potentials, Orbits, actionAngles, or DFs are initialized using a distance scale `ro=` and a velocity scale `vo=` output quantities returned and plotted in physical coordinates. Specifically, positions are returned in the units in the table below. If `astropy-units = True` in the [configuration file](#), then an `astropy.Quantity` which includes the units is returned instead (see below).

Quantity	Default unit
position	kpc
velocity	km/s
energy	(km/s)^2
Jacobi integral	(km/s)^2
angular momentum	km/s x kpc
actions	km/s x kpc
frequencies	1/Gyr
time	Gyr
period	Gyr
potential	(km/s)^2
force	km/s/Myr
force derivative	1/Gyr^2
density	Msun/pc^3
number density	1/pc^3
surface density	Msun/pc^2
mass	Msun
angle	rad
proper motion	mas/yr
phase-space density	1/(kpc x km/s)^3

Physical units

Tip: With `apy-units = True` in the configuration file and specifying all inputs using `astropy Quantity` with units, `galpy` will return outputs in convenient, unambiguous units.

Full support for unitful quantities using `astropy Quantity` was added in v1.2. Thus, *any* input to a `galpy Potential`, `Orbit`, `actionAngle`, or `DF` instantiation, method, or function can now be specified in physical units as a `Quantity`. For example, we can set up a Miyamoto-Nagai disk potential with a mass of $5 \times 10^{10} M_{\odot}$, a scale length of 3 kpc, and a scale height of 300 pc as follows

```
>>> from galpy.potential import MiyamotoNagaiPotential
>>> from astropy import units
>>> mp= MiyamotoNagaiPotential(amp=5*10**10*units.Msun,a=3.*units.kpc,b=300.*units.pc)
```

Internally, `galpy` uses a set of normalized units, where positions are divided by a scale `ro` and velocities are divided by a scale `vo`. If these are not specified, the default set from the [configuration file](#) is used. However, they can also be specified on an instance-by-instance manner for all `Potential`, `Orbit`, `actionAngle`, and `DF` instances. For example

```
>>> mp= MiyamotoNagaiPotential(amp=5*10**10*units.Msun,a=3.*units.kpc,b=300.*units.pc,
↪ro=9*units.kpc,vo=230.*units.km/units.s)
```

uses differently normalized internal units. When you specify the parameters of a `Potential`, `Orbit`, etc. in physical units (e.g., the Miyamoto-Nagai setup above), the internal set of units is unimportant as long as you receive output in physical units (see below) and it is unnecessary to change the values of `ro` and `vo`, unless you are modeling a system with very different distance and velocity scales from the default set (for example, if you are looking at internal globular cluster dynamics rather than galaxy dynamics). If you find an input to any `galpy` function that does not take a `Quantity` as an input (or that does it wrong), please report an [Issue](#).

Warning: If you combine potentials in a list, galpy uses the `ro` and `vo` scales from the first potential in the list for physical <-> internal unit conversion. galpy does **not** always check whether the unit systems of various objects are consistent when they are combined (but does check this for many common cases, e.g., integrating an Orbit in a Potential).

galpy can also return values with units as an astropy Quantity. Whether or not this is done is specified by the `apy-units` option in the *configuration file*. If you want to get return values as a Quantity, set `apy-units = True` in the configuration file. Then you can do for the Miyamoto-Nagai potential above

```
>>> mp.vcirc(10.*units.kpc)
# <Quantity 135.72399857308042 km / s>
```

Note that if you do not specify the argument as a Quantity with units, galpy will assume that it is given in natural units, viz.

```
>>> mp.vcirc(10.)
# <Quantity 51.78776595740726 km / s>
```

because this input is considered equal to 10 times the distance scale (this is for the case using the default `ro` and `vo`, the first Miyamoto-Nagai instantiation of this subsection)

```
>>> mp.vcirc(10.*8.*units.kpc)
# <Quantity 51.78776595740726 km / s>
```

Warning: If you do not specify arguments of methods and functions using a Quantity with units, galpy assumes that the argument has internal (natural) units.

If you do not use astropy Quantities (`apy-units = False` in the configuration file), you can still get output in physical units when you have specified `ro=` and `vo=` during instantiation of the Potential, Orbit, etc. For example, for the Miyamoto-Nagai potential above in a session with `apy-units = False`

```
>>> mp= MiyamotoNagaiPotential(amp=5*10**10*units.Msun,a=3.*units.kpc,b=300.*units.pc)
>>> mp.vcirc(10.*units.kpc)
# 135.72399857308042
```

This return value is in km/s (see the *table* at the end of the previous section for default units for different quantities). Note that as long as astropy is installed, we can still provide arguments as a Quantity, but the return value will not be a Quantity when `apy-units = False`. If you setup a Potential, Orbit, actionAngle, or DF object with parameters specified as a Quantity, the default is to return any output in physical units. This is why `mp.vcirc` returns the velocity in km/s above. Potential and Orbit instances (or lists of Potentials) also support the functions `turn_physical_off` and `turn_physical_on` to turn physical output off or on. For example, if we do

```
>>> mp.turn_physical_off()
```

outputs will be in internal units

```
>>> mp.vcirc(10.*units.kpc)
# 0.61692726624127459
```

If you setup a Potential, Orbit, etc. object without specifying the parameters as a Quantity, the default is to return output in natural units, except when `ro=` and `vo=` scales are specified. `ro=` and `vo=` can always be given as a Quantity themselves. `ro=` and `vo=` can always also be specified on a method-by-method basis, overwriting an object's default. For example

```
>>> mp.vcirc(10.*units.kpc,ro=12.*units.kpc)
# 0.69273212489609337
```

Physical output can also be turned off on a method-by-method or function-by-function basis, for example

```
>>> mp.turn_physical_on() # turn overall physical output on
>>> mp.vcirc(10.*units.kpc)
135.72399857308042 # km/s
>>> mp.vcirc(10.*units.kpc,use_physical=False)
# 0.61692726624127459 # in natural units
```

Further examples of specifying inputs with units will be given throughout the documentation.

1.3.3 Orbit integration

Warning: galpy uses a left-handed coordinate frame, as is common in studies of the kinematics of the Milky Way. This means that in particular cross-products, like the angular momentum $\vec{L} = \vec{r} \times \vec{p}$, behave differently than in a right-handed coordinate frame.

We can also integrate orbits in all galpy potentials. Going back to a simple Miyamoto-Nagai potential, we initialize an orbit as follows

```
>>> from galpy.orbit import Orbit
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,amp=1.,normalize=1.)
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.,0.1])
```

Since we gave `Orbit()` a five-dimensional initial condition $[R, vR, vT, z, vz]$, we assume we are dealing with a three-dimensional axisymmetric potential in which we do not wish to track the azimuth. We then integrate the orbit for a set of times `ts`

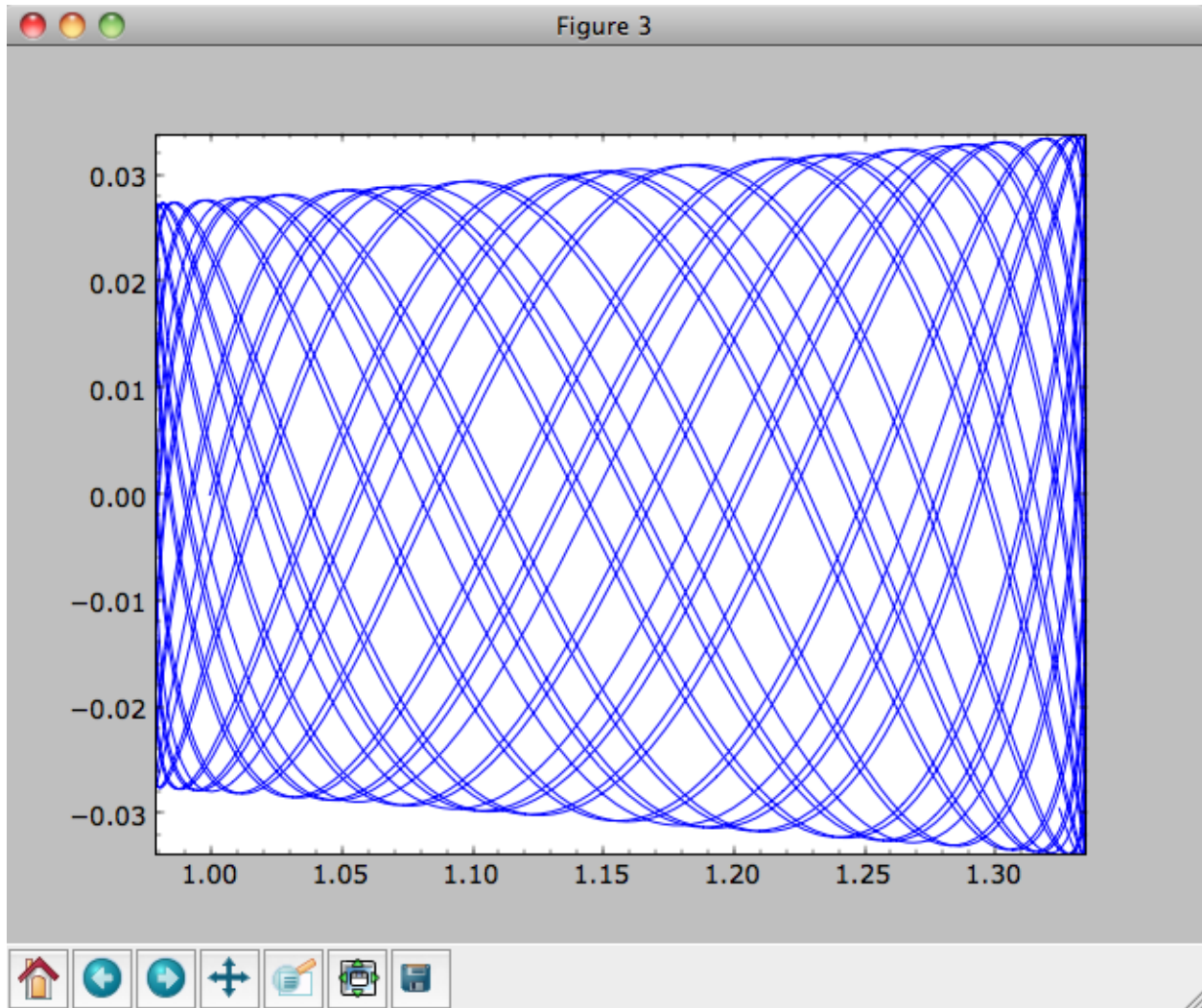
```
>>> import numpy
>>> ts= numpy.linspace(0,100,10000)
>>> o.integrate(ts,mp,method='odeint')
```

Tip: Like for the Miyamoto-Nagai example in the section above, the `Orbit` and integration times can also be specified in physical units, e.g., `o= Orbit(vxvv=[8.*units.kpc,22.*units.km/units.s,242.*units.km/units.s,0.*units.pc,20.*units.km/s])` and `ts= numpy.linspace(0.,10.,10000)*units.Gyr`

Now we plot the resulting orbit as

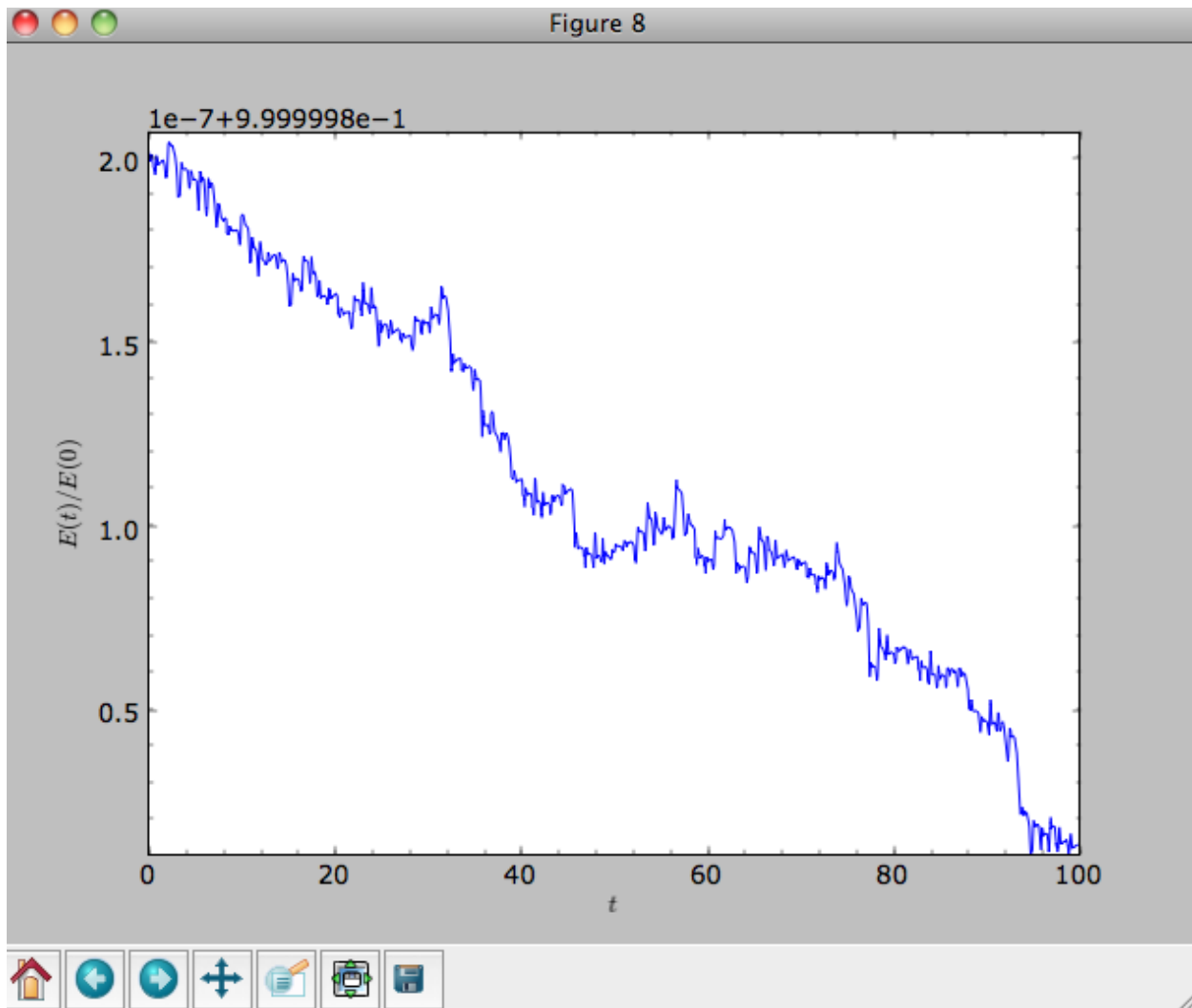
```
>>> o.plot()
```

Which gives



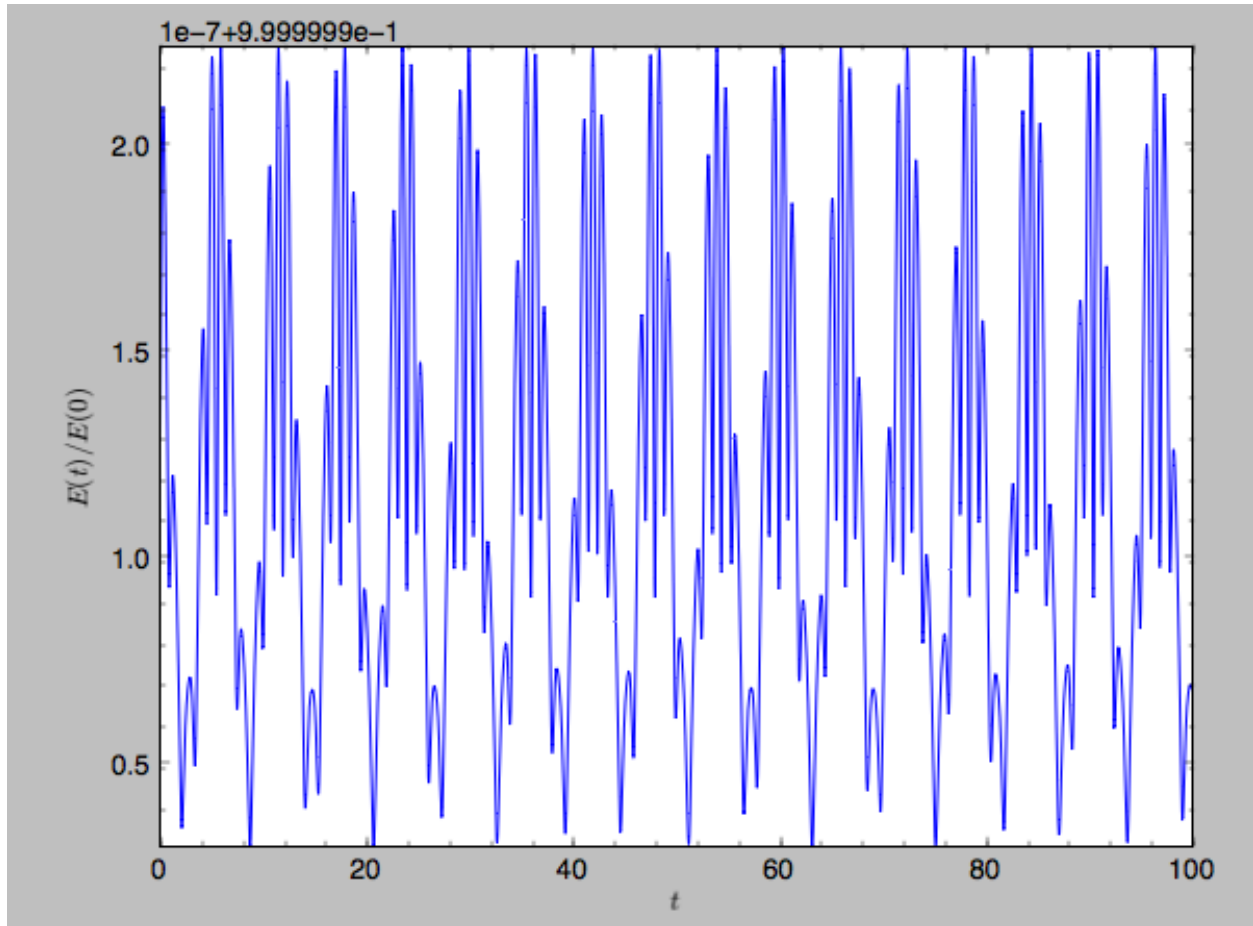
The integrator used is not symplectic, so the energy error grows with time, but is small nonetheless

```
>>> o.plotE(normed=True)
```



When we use a symplectic leapfrog integrator, we see that the energy error remains constant

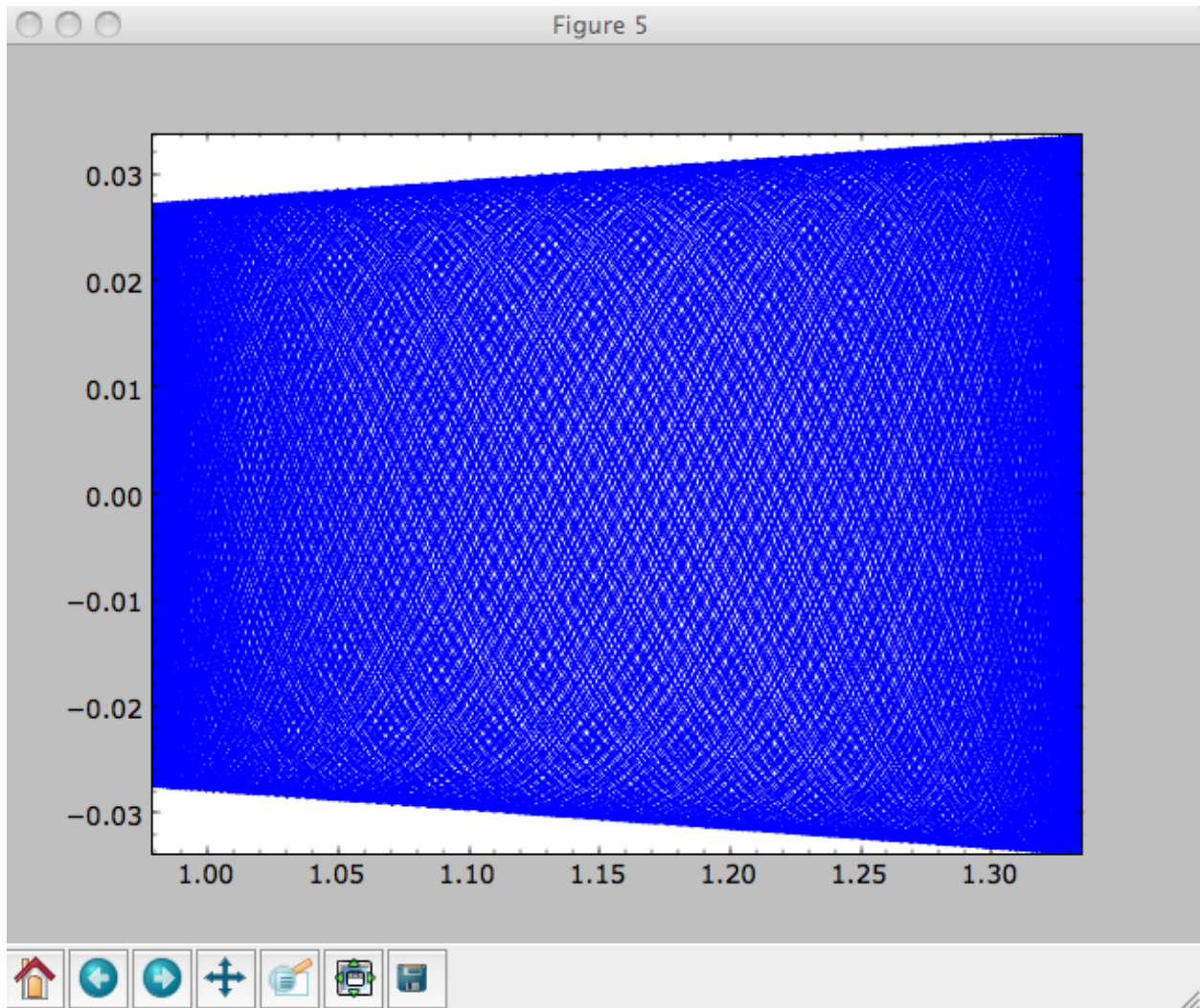
```
>>> o.integrate(ts,mp,method='leapfrog')
>>> o.plotE(xlabel=r'$t$',ylabel=r'$E(t)/E(0)$')
```



Because stars have typically only orbited the center of their galaxy tens of times, using symplectic integrators is mostly unnecessary (compared to planetary systems which orbits millions or billions of times). `galpy` contains fast integrators written in C, which can be accessed through the `method=` keyword (e.g., `integrate(..., method='dopr54_c')` is a fast high-order Dormand-Prince method).

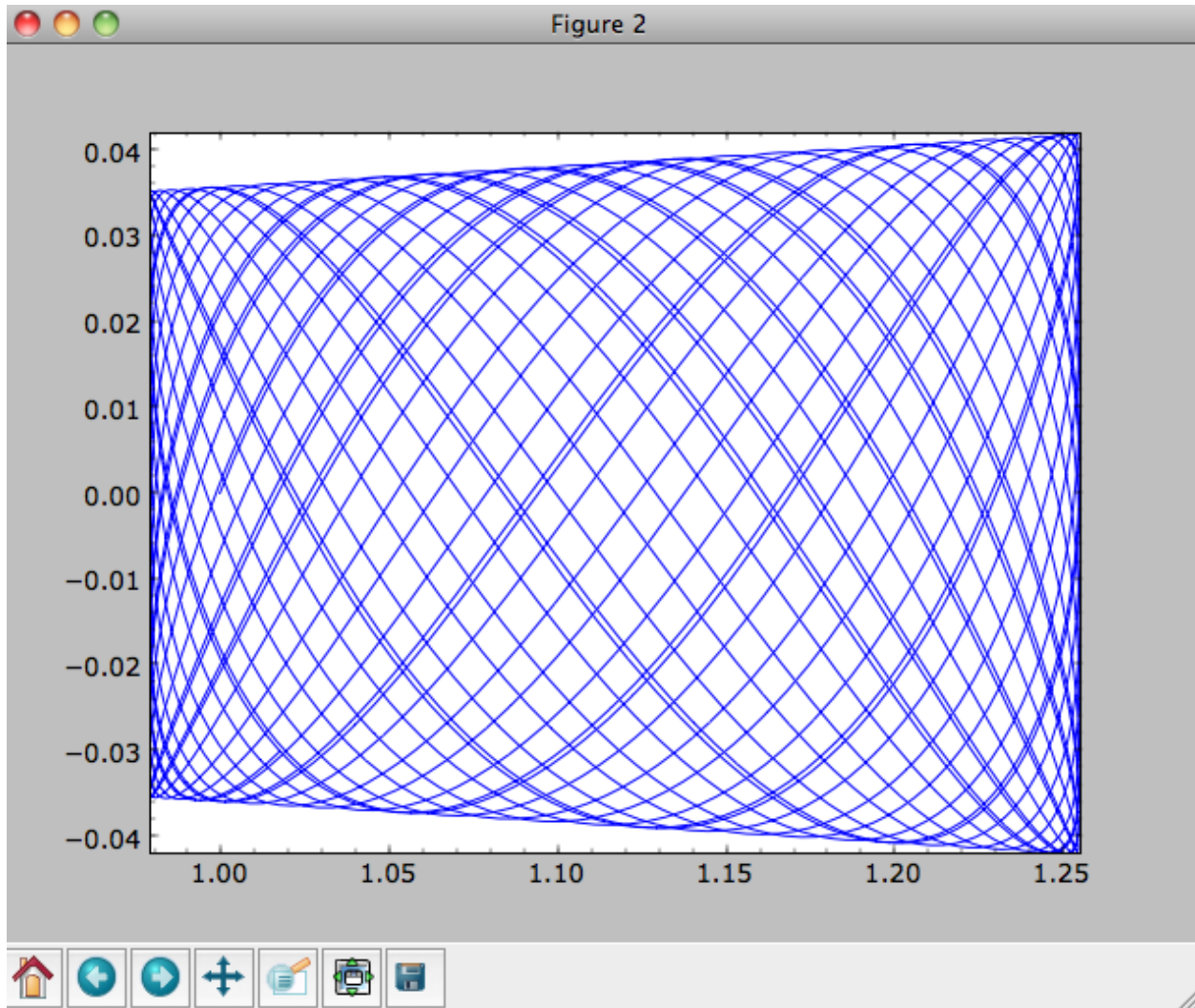
When we integrate for much longer we see how the orbit fills up a torus (this could take a minute)

```
>>> ts= numpy.linspace(0,1000,10000)
>>> o.integrate(ts,mp,method='odeint')
>>> o.plot()
```

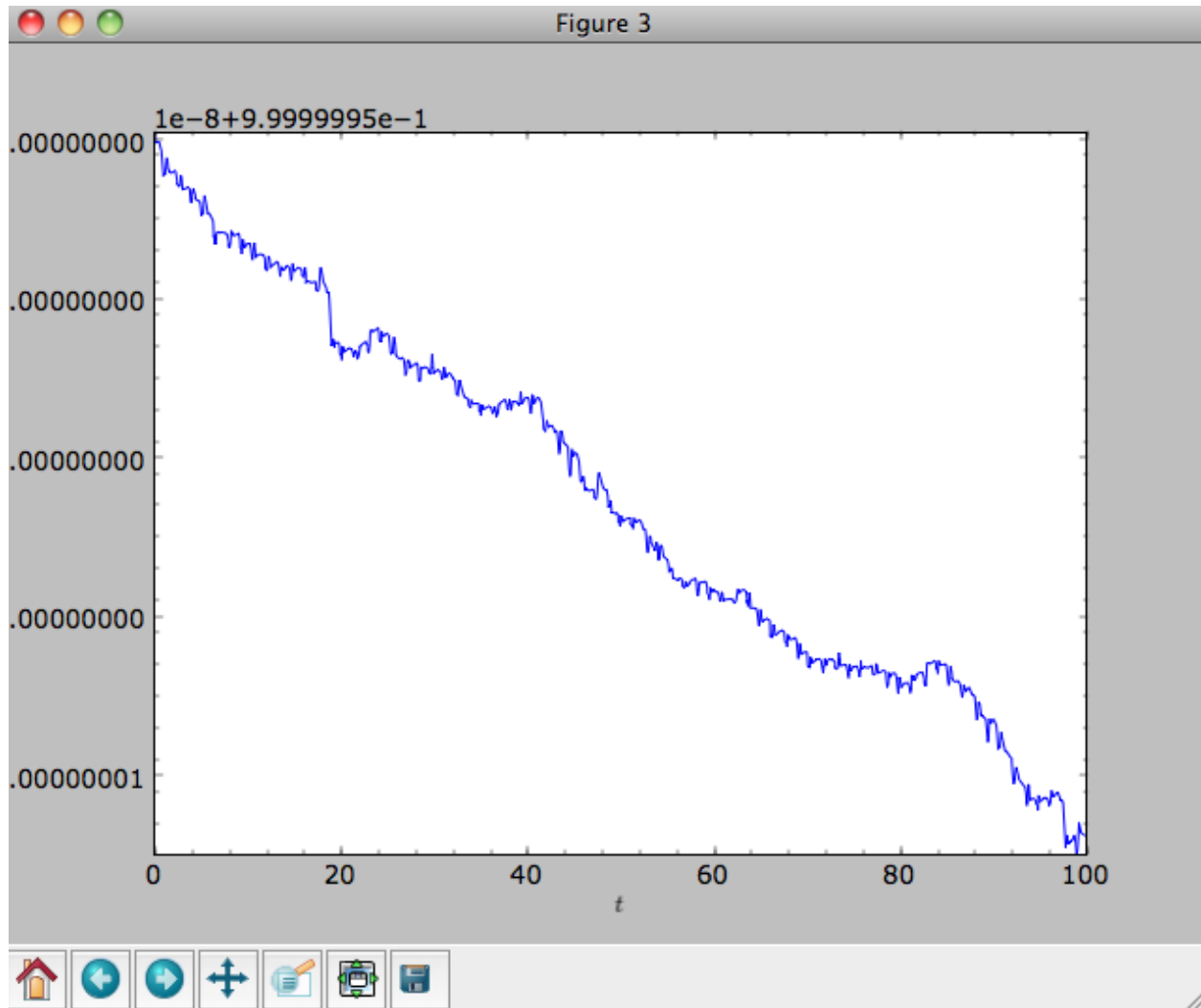
As before, we can also integrate orbits in combinations of potentials. Assuming `mp`, `np`, and `hp` were defined as above, we can

```
>>> ts= numpy.linspace(0,100,10000)
>>> o.integrate(ts,[mp,hp,np])
>>> o.plot()
```

Energy is again approximately conserved

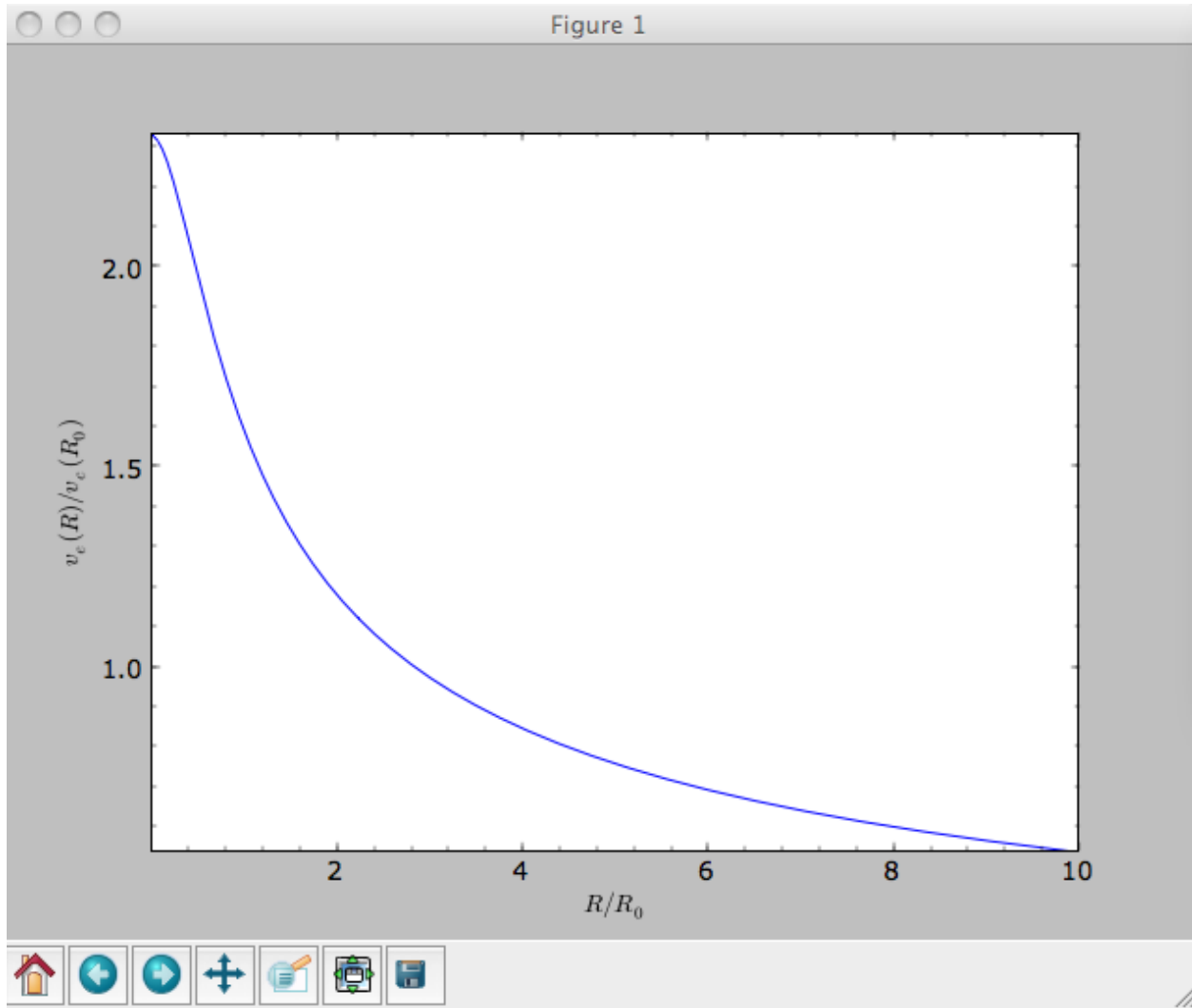
```
>>> o.plotE(xlabel=r'$t$', ylabel=r'$E(t)/E(0)$')
```



1.3.4 Escape velocity curves

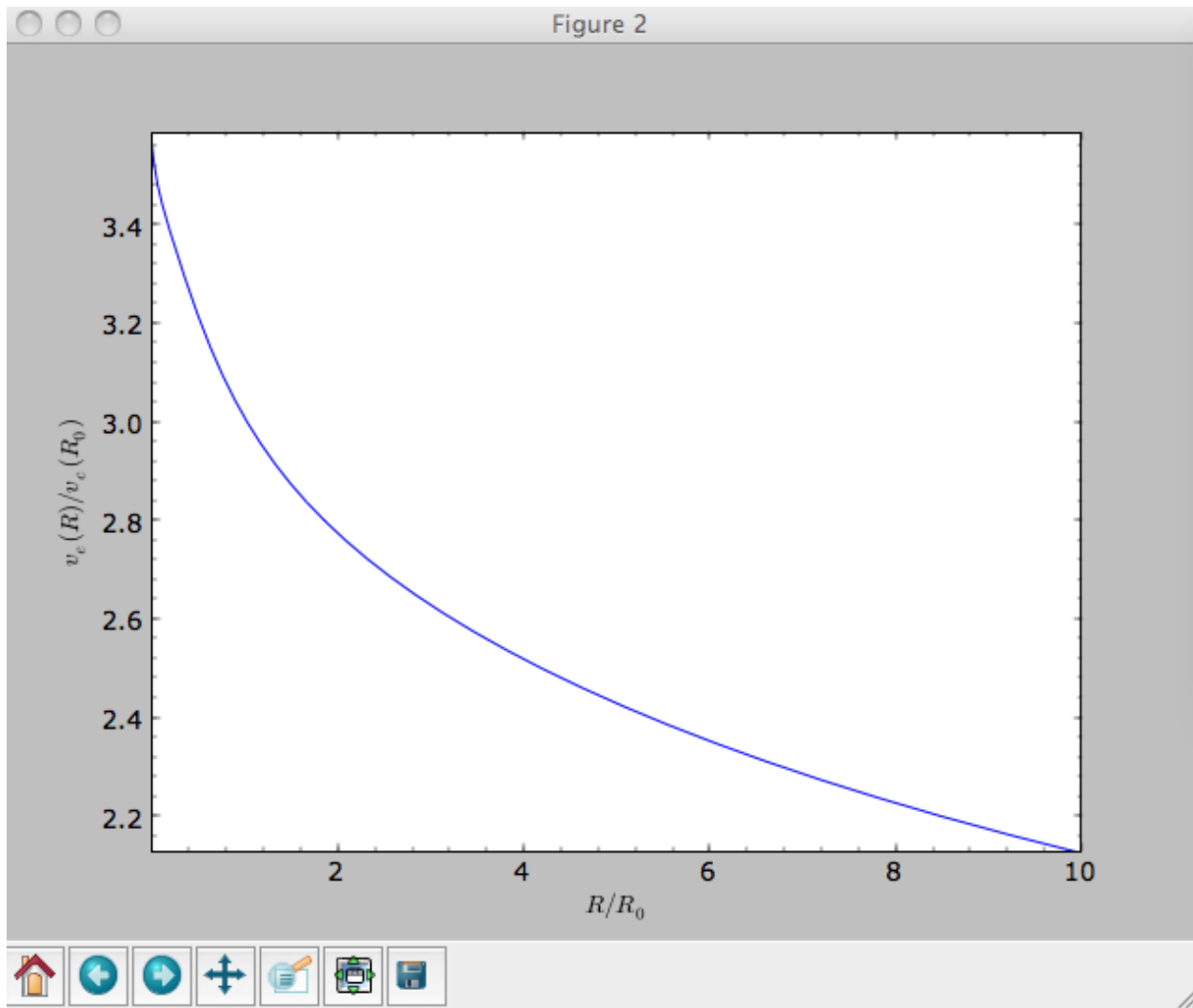
Just like we can plot the rotation curve for a potential or a combination of potentials, we can plot the escape velocity curve. For example, the escape velocity curve for the Miyamoto-Nagai disk defined above

```
>>> mp.plotEscapecurve(Rrange=[0.01,10.],grid=1001)
```



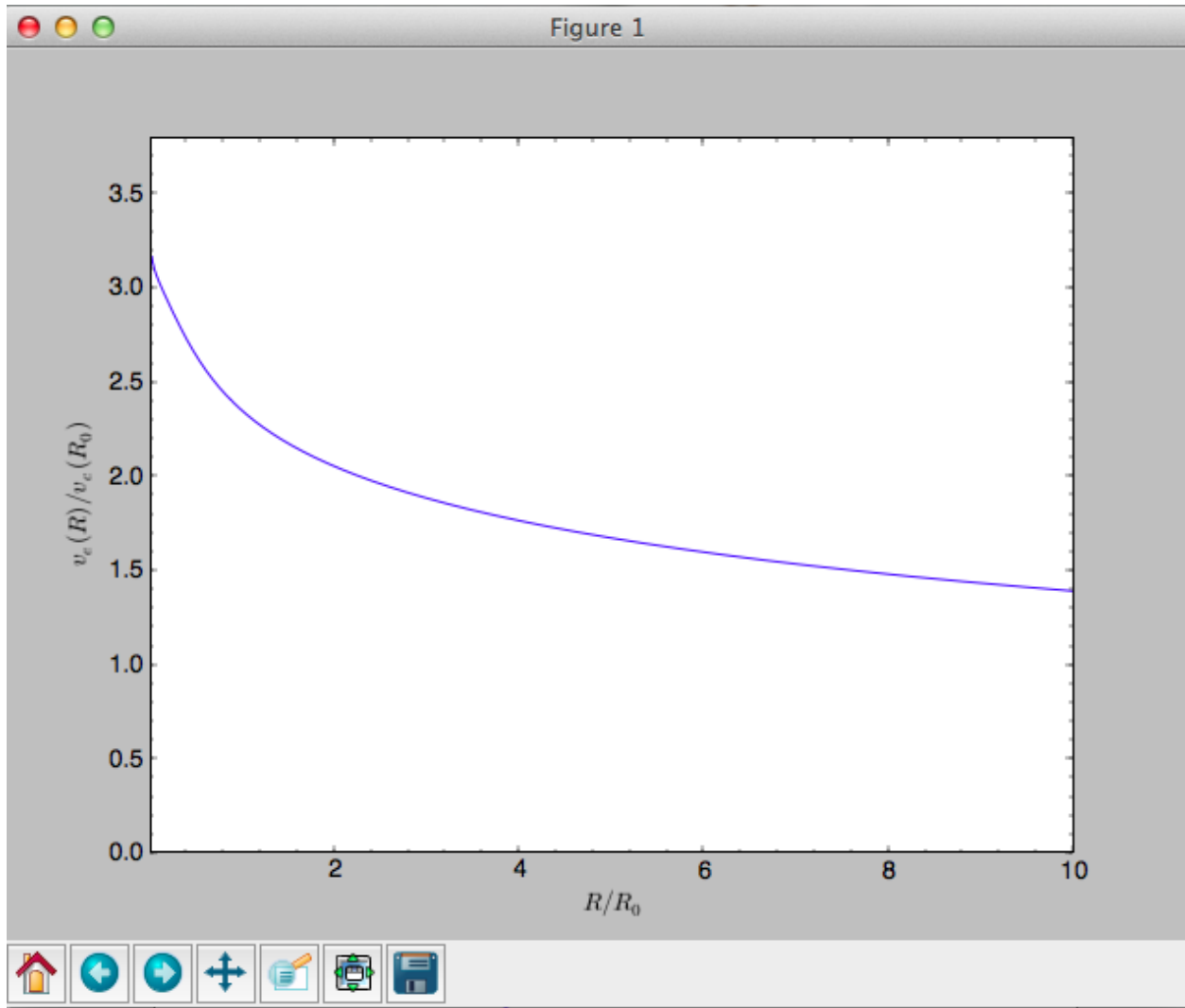
or of the combination of potentials defined above

```
>>> from galpy.potential import plotEscapecurve
>>> plotEscapecurve([mp, hp, np], Rrange=[0.01, 10.], grid=1001)
```



For the Milky-Way-like potential `MWPotential2014`, the escape-velocity curve is

```
>>> plotEscapecurve(MWPotential2014, Range=[0.01, 10.], grid=1001)
```



At the solar radius, the escape velocity is

```
>>> from galpy.potential import vesc
>>> vesc(MWPotential2014, 1.)
2.3316389848832784
```

Or, for a local circular velocity of 220 km/s

```
>>> vesc(MWPotential2014, 1.) * 220.
# 512.96057667432126
```

similar to direct measurements of this (e.g., [2007MNRAS.379..755S](#) and [2014A%26A...562A..91P](#)).

1.4 Potentials in galpy

galpy contains a large variety of potentials in `galpy.potential` that can be used for orbit integration, the calculation of action-angle coordinates, as part of steady-state distribution functions, and to study the properties of gravitational potentials. This section introduces some of these features.

1.4.1 Potentials and forces

Various 3D and 2D potentials are contained in galpy, list in the *API page*. Another way to list the latest overview of potentials included with galpy is to run

```
>>> import galpy.potential
>>> print([p for p in dir(galpy.potential) if 'Potential' in p])
# ['CosmphiDiskPotential',
#  'DehnenBarPotential',
#  'DoubleExponentialDiskPotential',
#  'EllipticalDiskPotential',
#  'FlattenedPowerPotential',
#  'HernquistPotential',
#  ...]
```

(list cut here for brevity). Section *Rotation curves* explains how to initialize potentials and how to display the rotation curve of single Potential instances or of combinations of such instances. Similarly, we can evaluate a Potential instance

```
>>> from galpy.potential import MiyamotoNagaiPotential
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,normalize=1.)
>>> mp(1.,0.)
# -1.2889062500000001
```

Most member functions of Potential instances have corresponding functions in the galpy.potential module that allow them to be evaluated for lists of multiple Potential instances (and in versions ≥ 1.4 even for nested lists of Potential instances). galpy.potential.MWPotential2014 is such a list of three Potential instances

```
>>> from galpy.potential import MWPotential2014
>>> print(MWPotential2014)
# [<galpy.potential.PowerSphericalPotentialwCutoff.PowerSphericalPotentialwCutoff_
↪instance at 0x1089b23b0>, <galpy.potential.MiyamotoNagaiPotential.
↪MiyamotoNagaiPotential instance at 0x1089b2320>, <galpy.potential.
↪TwoPowerSphericalPotential.NFWPotential instance at 0x1089b2248>]
```

and we can evaluate the potential by using the evaluatePotentials function

```
>>> from galpy.potential import evaluatePotentials
>>> evaluatePotentials(MWPotential2014,1.,0.)
# -1.3733506513947895
```

Tip: Lists of Potential instances can be nested, allowing you to easily add components to existing gravitational-potential models. For example, to add a DehnenBarPotential to MWPotential2014, you can do: `pot=[MWPotential2014,DehnenBarPotential()]` and then use this `pot` everywhere where you can use a list of Potential instances.

Warning: galpy potentials do *not* necessarily approach zero at infinity. To compute, for example, the escape velocity or whether or not an orbit is unbound, you need to take into account the value of the potential at infinity. E.g., $v_{\text{esc}}(r) = \sqrt{2[\Phi(\infty) - \Phi(r)]}$.

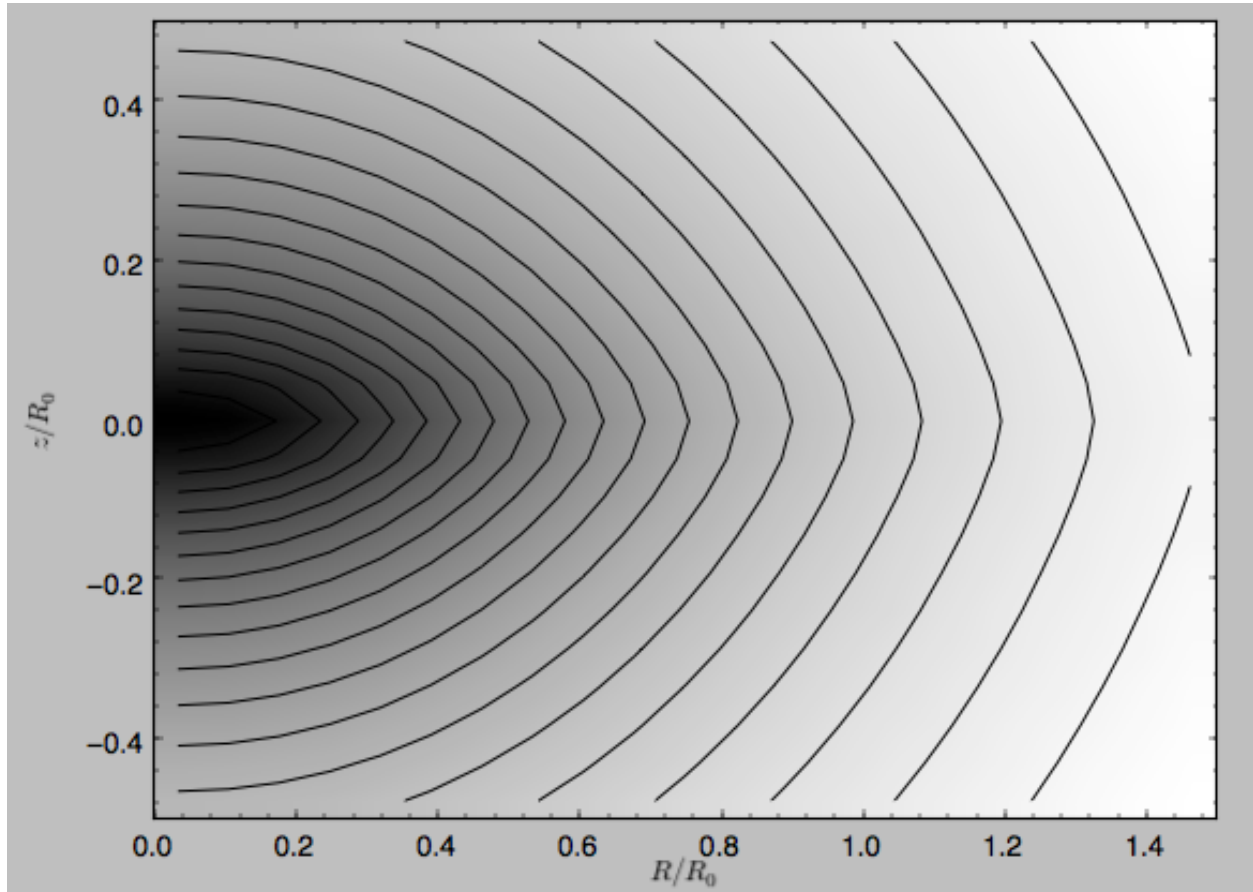
Tip: As discussed in the section on *physical units*, potentials can be initialized and evaluated with arguments specified as a astropy Quantity with units. Use the configuration parameter `apy-units = True` to get output values as a

Quantity. See also the subsection on *Initializing potentials with parameters with units* below.

We can plot the potential of axisymmetric potentials (or of non-axisymmetric potentials at $\phi=0$) using the `plot` member function

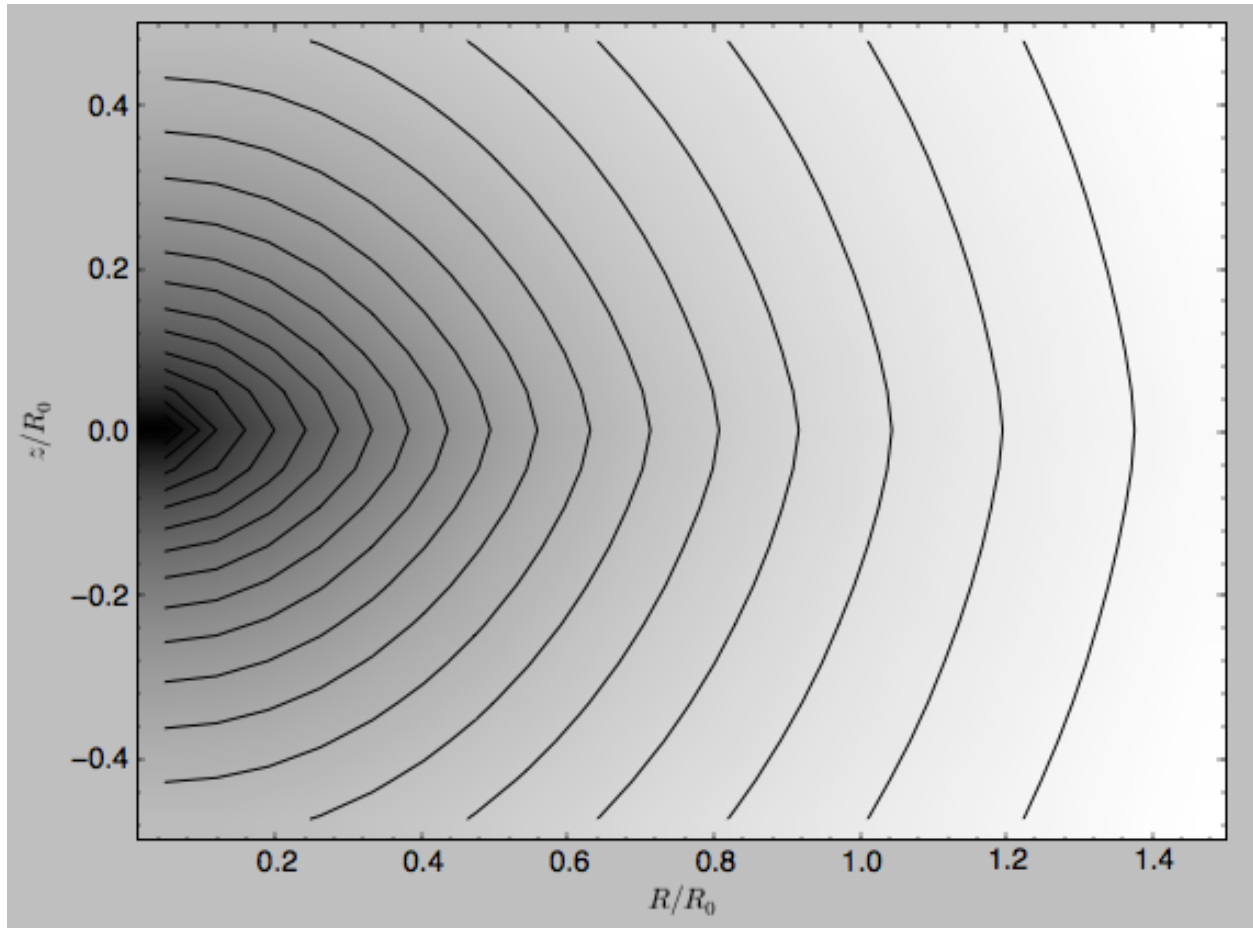
```
>>> mp.plot()
```

which produces the following plot



Similarly, we can plot combinations of Potentials using `plotPotentials`, e.g.,

```
>>> from galpy.potential import plotPotentials
>>> plotPotentials(MWPotential2014, rmin=0.01)
```



These functions have arguments that can provide custom R and z ranges for the plot, the number of grid points, the number of contours, and many other parameters determining the appearance of these figures.

galpy also allows the forces corresponding to a gravitational potential to be calculated. Again for the Miyamoto-Nagai Potential instance from above

```
>>> mp.Rforce(1., 0.)
# -1.0
```

This value of -1.0 is due to the normalization of the potential such that the circular velocity is 1. at $R=1$. Similarly, the vertical force is zero in the mid-plane

```
>>> mp.zforce(1., 0.)
# -0.0
```

but not further from the mid-plane

```
>>> mp.zforce(1., 0.125)
# -0.53488743705310848
```

As explained in [Units in galpy](#), these forces are in standard galpy units, and we can convert them to physical units using methods in the `galpy.util.bovy_conversion` module. For example, assuming a physical circular velocity of 220 km/s at $R=8$ kpc


```
>>> from galpy.util import bovy_conversion
>>> mp.zforce(1.,0.125)*bovy_conversion.force_in_kmsMyr(220.,8.)
# -3.3095671288657584 #km/s/Myr
>>> mp.zforce(1.,0.125)*bovy_conversion.force_in_2piGmsolpc2(220.,8.)
# -119.72021771473301 #2 \pi G Msol / pc^2
```

Again, there are functions in `galpy.potential` that allow for the evaluation of the forces for lists of Potential instances, such that

```
>>> from galpy.potential import evaluateRforces
>>> evaluateRforces(MWPotential2014,1.,0.)
# -1.0
>>> from galpy.potential import evaluatezforces
>>> evaluatezforces(MWPotential2014,1.,0.125)*bovy_conversion.force_in_
  ↪ 2piGmsolpc2(220.,8.)
>>> -69.680720137571114 #2 \pi G Msol / pc^2
```

We can evaluate the flattening of the potential as $\sqrt{|z F_R / R F_Z|}$ for a Potential instance as well as for a list of such instances

```
>>> mp.flattening(1.,0.125)
# 0.4549542914935209
>>> from galpy.potential import flattening
>>> flattening(MWPotential2014,1.,0.125)
# 0.61231675305658628
```

1.4.2 Densities

galpy can also calculate the densities corresponding to gravitational potentials. For many potentials, the densities are explicitly implemented, but if they are not, the density is calculated using the Poisson equation (second derivatives of the potential have to be implemented for this). For example, for the Miyamoto-Nagai potential, the density is explicitly implemented

```
>>> mp.dens(1.,0.)
# 1.1145444383277576
```

and we can also calculate this using the Poisson equation

```
>>> mp.dens(1.,0.,forcepoisson=True)
# 1.1145444383277574
```

which are the same to machine precision

```
>>> mp.dens(1.,0.,forcepoisson=True)-mp.dens(1.,0.)
# -2.2204460492503131e-16
```

Similarly, all of the potentials in `galpy.potential.MWPotential2014` have explicitly-implemented densities, so we can do

```
>>> from galpy.potential import evaluateDensities
>>> evaluateDensities(MWPotential2014,1.,0.)
# 0.57508603122264867
```

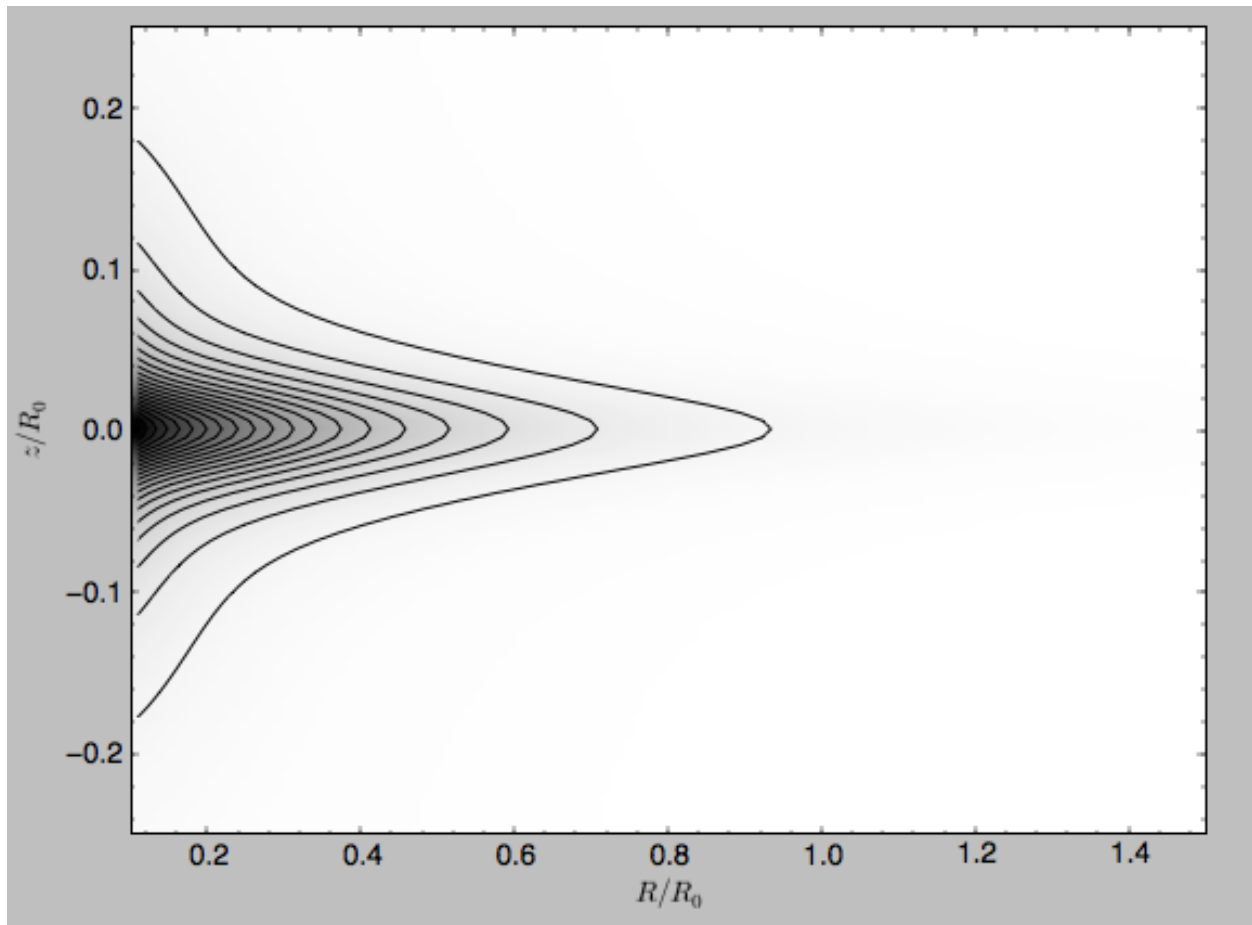
In physical coordinates, this becomes

```
>>> evaluateDensities(MWPotential2014,1.,0.)*bovy_conversion.dens_in_msolpc3(220.,8.)
# 0.1010945632524705 #Msol / pc^3
```

We can also plot densities

```
>>> from galpy.potential import plotDensities
>>> plotDensities(MWPotential2014,rmin=0.1,zmax=0.25,zmin=-0.25,nrs=101,nzs=101)
```

which gives



Another example of this is for an exponential disk potential

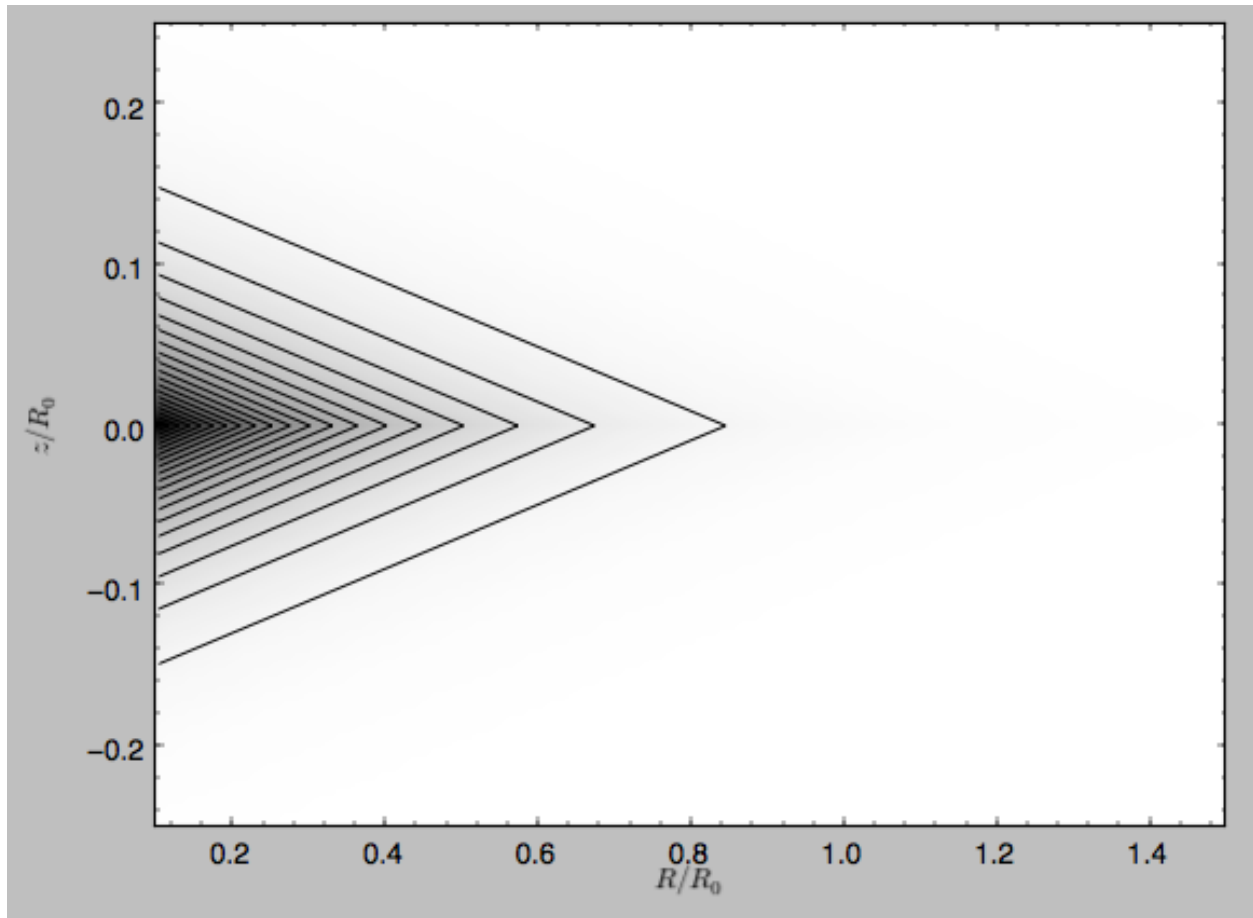
```
>>> from galpy.potential import DoubleExponentialDiskPotential
>>> dp= DoubleExponentialDiskPotential(hr=1./4.,hz=1./20.,normalize=1.)
```

The density computed using the Poisson equation now requires multiple numerical integrations, so the agreement between the analytical density and that computed using the Poisson equation is slightly less good, but still better than a percent

```
>>> (dp.dens(1.,0.,forcepoisson=True)-dp.dens(1.,0.))/dp.dens(1.,0.)
# 0.0032522956769123019
```

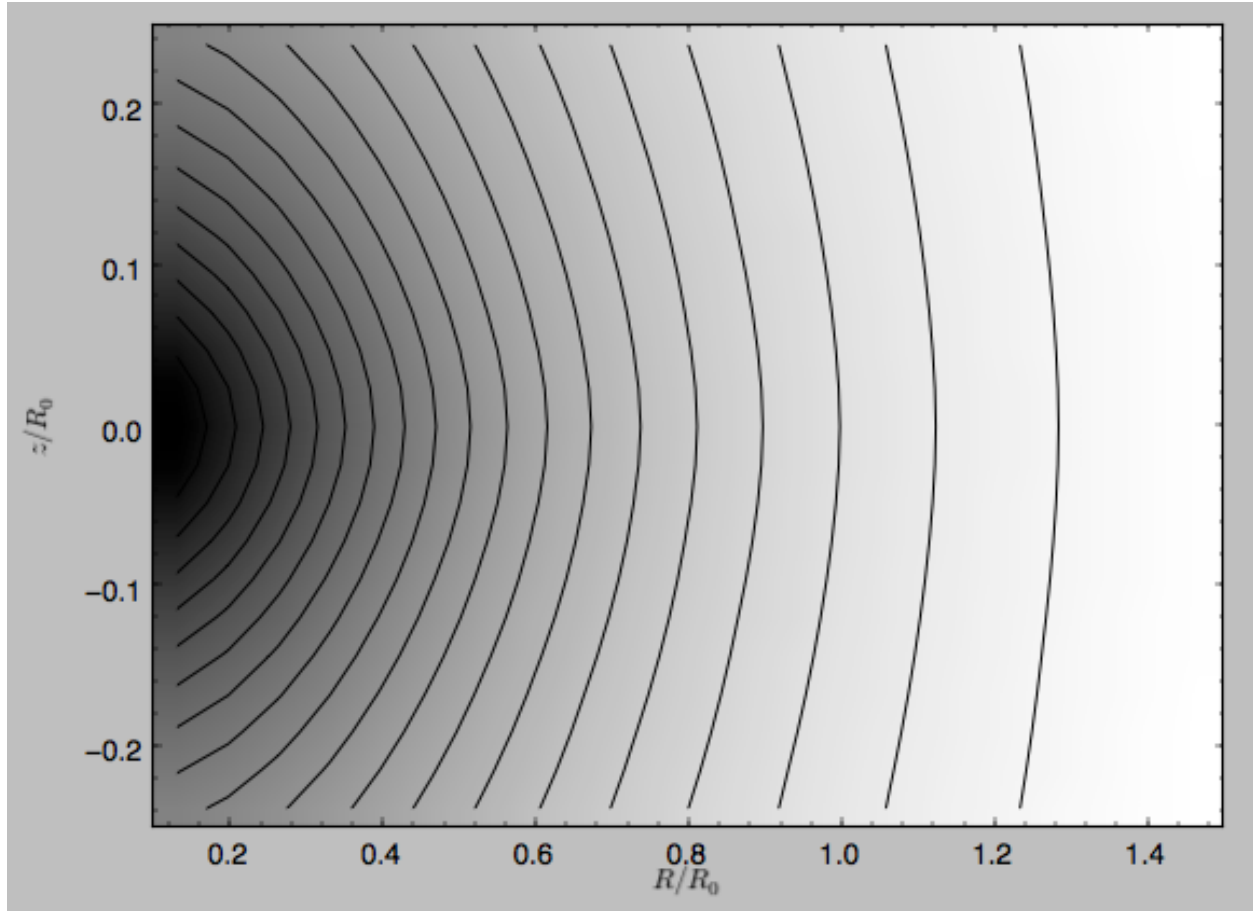
The density is

```
>>> dp.plotDensity(rmin=0.1,zmax=0.25,zmin=-0.25,nrs=101,nzs=101)
```



and the potential is

```
>>> dp.plot(rmin=0.1, zmin=-0.25, zmax=0.25)
```



Clearly, the potential is much less flattened than the density.

1.4.3 Modifying potential instances using wrappers

Potentials implemented in galpy can be modified using different kinds of wrappers. These wrappers modify potentials to, for example, change their amplitude as a function of time (e.g., to grow or decay the bar contribution to a potential) or to make a potential rotate. Specific kinds of wrappers are listed on the [Potential wrapper API page](#). These wrappers can be applied to instances of *any* potential implemented in galpy (including other wrappers). An example is to grow a bar using the polynomial smoothing of [Dehnen \(2000\)](#). We first setup an instance of a `DehnenBarPotential` that is essentially fully grown already

```
>>> from galpy.potential import DehnenBarPotential
>>> dpn= DehnenBarPotential(tform=-100.,tsteady=0.) # DehnenBarPotential has a custom_
↳ implementation of growth that we ignore by setting tform to -100
```

and then wrap it

```
>>> from galpy.potential import DehnenSmoothWrapperPotential
>>> dswp= DehnenSmoothWrapperPotential(pot=dpn,tform=-4.*2.*numpy.pi/dpn.OmegaP(),
↳ tsteady=2.*2.*numpy.pi/dpn.OmegaP())
```

This grows the `DehnenBarPotential` starting at 4 bar periods before $t=0$ over a period of 2 bar periods. `DehnenBarPotential` has an older, custom implementation of the same smoothing and the $(tform, tsteady)$ pair used here corresponds to the default setting for `DehnenBarPotential`. Thus we can compare the two

```
>>> dp= DehnenBarPotential()
>>> print(dp(0.9,0.3,phi=3.,t=-2.)-dswp(0.9,0.3,phi=3.,t=-2.))
# 0.0
>>> print(dp.Rforce(0.9,0.3,phi=3.,t=-2.)-dswp.Rforce(0.9,0.3,phi=3.,t=-2.))
# 0.0
```

The wrapper `SolidBodyRotationWrapperPotential` allows one to make any potential rotate around the z axis. This can be used, for example, to make general bar-shaped potentials, which one could construct from a basis-function expansion with `SCFPotential`, rotate without having to implement the rotation directly. As an example consider this `SoftenedNeedleBarPotential` (which has a potential-specific implementation of rotation)

```
>>> sp= SoftenedNeedleBarPotential(normalize=1.,omegab=1.8,pa=0.)
```

The same potential can be obtained from a non-rotating `SoftenedNeedleBarPotential` run through the `SolidBodyRotationWrapperPotential` to add rotation

```
>>> sp_still= SoftenedNeedleBarPotential(omegab=0.,pa=0.,normalize=1.)
>>> swp= SolidBodyRotationWrapperPotential(pot=sp_still,omega=1.8,pa=0.)
```

Compare for example

```
>>> print(sp(0.8,0.2,phi=0.2,t=3.)-swp(0.8,0.2,phi=0.2,t=3.))
# 0.0
>>> print(sp.Rforce(0.8,0.2,phi=0.2,t=3.)-swp.Rforce(0.8,0.2,phi=0.2,t=3.))
# 8.881784197e-16
```

Wrapper potentials can be used anywhere in galpy where general potentials can be used. They can be part of lists of Potential instances. They can also be used in C for orbit integration provided that both the wrapper and the potentials that it wraps are implemented in C. For example, a static `LogarithmicHaloPotential` with a bar potential grown as above would be

```
>>> from galpy.potential import LogarithmicHaloPotential, evaluateRforces
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> pot= [lp,dswp]
>>> print(evaluateRforces(pot,0.9,0.3,phi=3.,t=-2.))
# -1.00965326579
```

1.4.4 Close-to-circular orbits and orbital frequencies

We can also compute the properties of close-to-circular orbits. First of all, we can calculate the circular velocity and its derivative

```
>>> mp.vcirc(1.)
# 1.0
>>> mp.dvcircdR(1.)
# -0.163777427566978
```

or, for lists of Potential instances

```
>>> from galpy.potential import vcirc
>>> vcirc(MWPotential2014,1.)
# 1.0
>>> from galpy.potential import dvcircdR
>>> dvcircdR(MWPotential2014,1.)
# -0.10091361254334696
```

We can also calculate the various frequencies for close-to-circular orbits. For example, the rotational frequency

```
>>> mp.omegac(0.8)
# 1.2784598203204887
>>> from galpy.potential import omegac
>>> omegac(MWPotential2014,0.8)
# 1.2733514576122869
```

and the epicycle frequency

```
>>> mp.epifreq(0.8)
# 1.7774973530267848
>>> from galpy.potential import epifreq
>>> epifreq(MWPotential2014,0.8)
# 1.7452189766287691
```

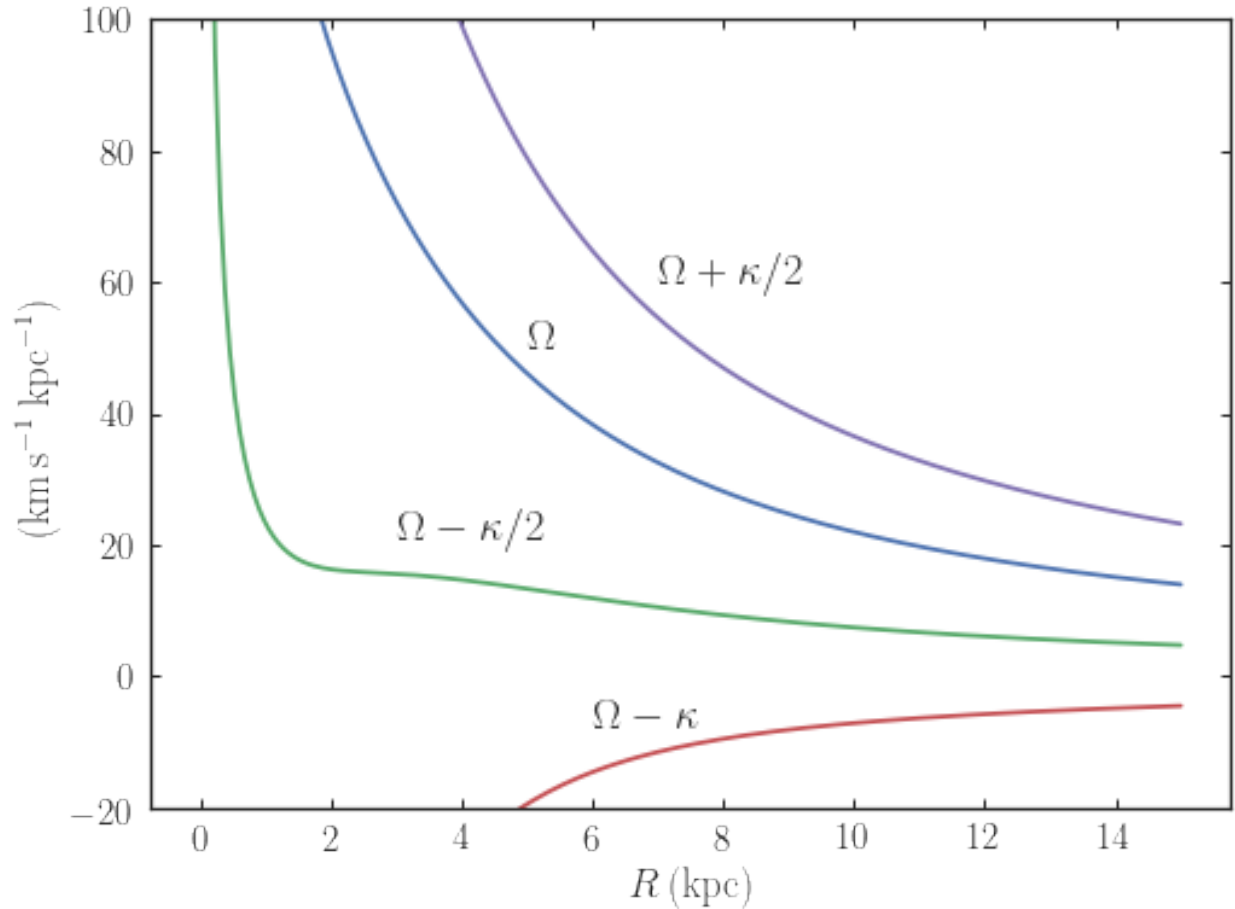
as well as the vertical frequency

```
>>> mp.verticalfreq(1.0)
# 3.7859388972001828
>>> from galpy.potential import verticalfreq
>>> verticalfreq(MWPotential2014,1.)
# 2.7255405754769875
```

We can also for example easily make the diagram of $\Omega - n\kappa/m$ that is important for understanding kinematic spiral density waves. For example, for MWPotential2014

```
>>> def OmegaMinusKappa(pot,Rs,n,m,ro=8.,vo=220.): # ro,vo for physical units
    return omegac(pot,Rs,ro=ro,vo=vo)-n/m*epifreq(pot,Rs,ro=ro,vo=vo)
>>> plot(Rs,OmegaMinusKappa(MWPotential2014,Rs,0,1))
>>> plot(Rs,OmegaMinusKappa(MWPotential2014,Rs,1,2))
>>> plot(Rs,OmegaMinusKappa(MWPotential2014,Rs,1,1))
>>> plot(Rs,OmegaMinusKappa(MWPotential2014,Rs,1,-2))
>>> ylim(-20.,100.)
>>> xlabel(r'$R\, (\mathrm{kpc})$')
>>> ylabel(r'$\mathrm{km\,s}^{-1}\, \mathrm{kpc}^{-1}$')
>>> text(3.,21.,r'$\Omega-\kappa/2$',size=18.)
>>> text(5.,50.,r'$\Omega$',size=18.)
>>> text(7.,60.,r'$\Omega+\kappa/2$',size=18.)
>>> text(6.,-7.,r'$\Omega-\kappa$',size=18.)
```

which gives



For close-to-circular orbits, we can also compute the radii of the Lindblad resonances. For example, for a frequency similar to that of the Milky Way's bar

```
>>> mp.lindbladR(5./3.,m='corotation') #args are pattern speed and m of pattern
# 0.6027911166042229 #~ 5kpc
>>> print(mp.lindbladR(5./3.,m=2))
# None
>>> mp.lindbladR(5./3.,m=-2)
# 0.9906190683480501
```

The None here means that there is no inner Lindblad resonance, the $m=-2$ resonance is in the Solar neighborhood (see the section on the *Hercules stream* in this documentation).

1.4.5 Using interpolations of potentials

galpy contains a general Potential class `interpRZPotential` that can be used to generate interpolations of potentials that can be used in their stead to speed up calculations when the calculation of the original potential is computationally expensive (for example, for the `DoubleExponentialDiskPotential`). Full details on how to set this up are given [here](#). Interpolated potentials can be used anywhere that general three-dimensional galpy potentials can be used. Some care must be taken with outside-the-interpolation-grid evaluations for functions that use C to speed up computations.

1.4.6 Initializing potentials with parameters with units

As already discussed in the section on *physical units*, potentials in galpy can be specified with parameters with units since v1.2. For most inputs to the initialization it is straightforward to know what type of units the input Quantity needs to have. For example, the scale length parameter `a=` of a Miyamoto-Nagai disk needs to have units of distance.

The amplitude of a potential is specified through the `amp=` initialization parameter. The units of this parameter vary from potential to potential. For example, for a logarithmic potential the units are velocity squared, while for a Miyamoto-Nagai potential they are units of mass. Check the documentation of each potential on the *API page* for the units of the `amp=` parameter of the potential that you are trying to initialize and please report an [Issue](#) if you find any problems with this.

1.4.7 General density/potential pairs with basis-function expansions

galpy allows for the potential and forces of general, time-independent density functions to be computed by expanding the potential and density in terms of basis functions. This is supported for ellipsoidal-ish as well as for disk-y density distributions, in both cases using the basis-function expansion of the self-consistent-field (SCF) method of [Hernquist & Ostriker \(1992\)](#). On its own, the SCF technique works well for ellipsoidal-ish density distributions, but using a trick due to [Kuijken & Dubinski \(1995\)](#) it can also be made to work well for disk-y potentials. We first describe the basic SCF implementation and then discuss how to use it for disk-y potentials.

The basis-function approach in the SCF method is implemented in the *SCFPotential* class, which is also implemented in C for fast orbit integration. The coefficients of the basis-function expansion can be computed using the *scf_compute_coeffs_spherical* (for spherically-symmetric density distribution), *scf_compute_coeffs_axi* (for axisymmetric densities), and *scf_compute_coeffs* (for the general case). The coefficients obtained from these functions can be directly fed into the *SCFPotential* initialization. The basis-function expansion has a free scale parameter `a`, which can be specified for the *scf_compute_coeffs_XX* functions and for the *SCFPotential* itself. Make sure that you use the same `a`! Note that the general functions are quite slow.

The simplest example is that of the Hernquist potential, which is the lowest-order basis function. When we compute the first ten radial coefficients for this density we obtain that only the lowest-order coefficient is non-zero

```
>>> from galpy.potential import HernquistPotential
>>> from galpy.potential import scf_compute_coeffs_spherical
>>> hp= HernquistPotential(amp=1.,a=2.)
>>> Acos, Asin= scf_compute_coeffs_spherical(hp.dens,10,a=2.)
>>> print(Acos)
# array([[ 1.00000000e+00]],
#        [[ -2.83370393e-17]],
#        [[  3.31150709e-19]],
#        [[ -6.66748299e-18]],
#        [[  8.19285777e-18]],
#        [[ -4.26730651e-19]],
#        [[ -7.16849567e-19]],
#        [[  1.52355608e-18]],
#        [[ -2.24030288e-18]],
#        [[ -5.24936820e-19]])
```

As a more complicated example, consider a prolate NFW potential

```
>>> from galpy.potential import TriaxialNFWPotential
>>> np= TriaxialNFWPotential(normalize=1.,c=1.4,a=1.)
```

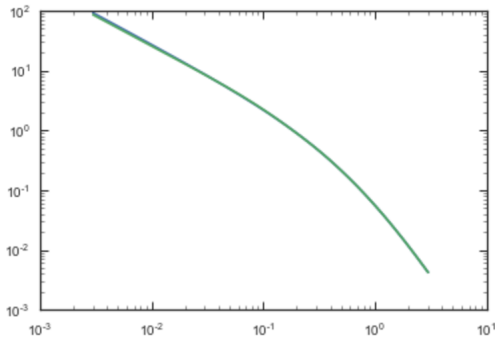
and we compute the coefficients using the axisymmetric *scf_compute_coeffs_axi*


```
>>> a_SCF= 50. # much larger a than true scale radius works well for NFW
>>> Acos, Asin= scf_compute_coeffs_axi(np.dens,80,40,a=a_SCF)
>>> sp= SCFPotential(Acos=Acos,Asin=Asin,a=a_SCF)
```

If we compare the densities along the $R=Z$ line as

```
>>> xs= numpy.linspace(0.,3.,1001)
>>> loglog(xs,np.dens(xs,xs))
>>> loglog(xs,sp.dens(xs,xs))
```

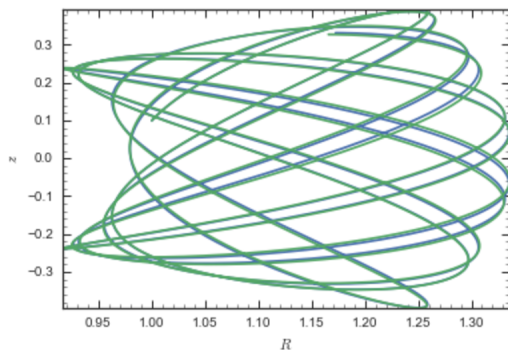
we get



If we then integrate an orbit, we also get good agreement

```
>>> from galpy.orbit import Orbit
>>> o= Orbit([1.,0.1,1.1,0.1,0.3,0.])
>>> ts= numpy.linspace(0.,100.,10001)
>>> o.integrate(ts,hp)
>>> o.plot()
>>> o.integrate(ts,sp)
>>> o.plot(overplot=True)
```

which gives



Near the end of the orbit integration, the slight differences between the original potential and the basis-expansion version cause the two orbits to deviate from each other.

To use the SCF method for disk potentials, we use the trick from [Kuijken & Dubinski \(1995\)](#). This trick works by approximating the disk density as $\rho_{\text{disk}}(R, \phi, z) \approx \sum_i \Sigma_i(R) h_i(z)$, with $h_i(z) = d^2 H(z)/dz^2$ and searching for solutions of the form

$$\Phi(R, \phi, z) = \Phi_{\text{ME}}(R, \phi, z) + 4\pi G \sum_i \Sigma_i(r) H_i(z),$$

where r is the spherical radius $r^2 = R^2 + z^2$. The density which gives rise to $\Phi_{\text{ME}}(R, \phi, z)$ is not strongly confined to a plane when $\rho_{\text{disk}}(R, \phi, z) \approx \sum_i \Sigma_i(R) h_i(z)$ and can be obtained using the SCF basis-function-expansion technique discussed above. See the documentation of the [DiskSCFPotential](#) class for more details on this procedure.

As an example, consider a double-exponential disk, which we can compare to the `DoubleExponentialDiskPotential` implementation

```
>>> from galpy import potential
>>> dp= potential.DoubleExponentialDiskPotential(amp=13.5, hr=1./3., hz=1./27.)
```

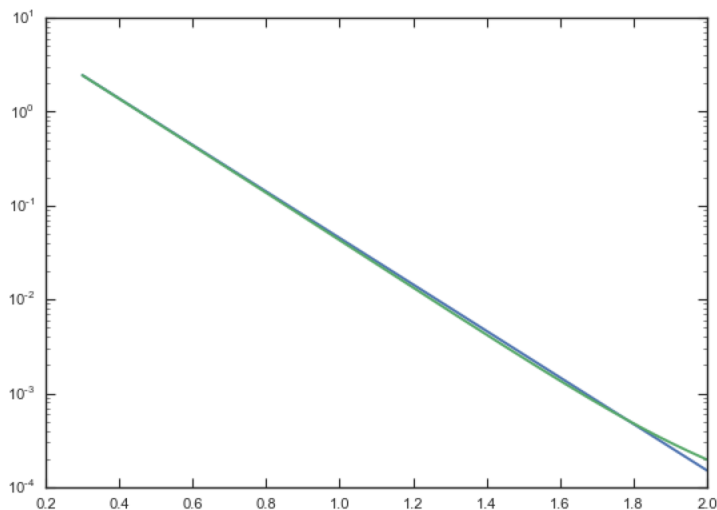
and then setup the `DiskSCFPotential` approximation to this as

```
>>> dscfp= potential.DiskSCFPotential(dens=lambda R, z: dp.dens(R, z),
                                     Sigma={'type': 'exp', 'h': 1./3., 'amp': 1.},
                                     hz={'type': 'exp', 'h': 1./27.},
                                     a=1., N=10, L=10)
```

The `dens=` keyword specifies the target density, while the `Sigma=` and `hz=` inputs specify the approximation functions $\Sigma_i(R)$ and $h_i(z)$. These are specified as dictionaries here for a few pre-defined approximation functions, but general functions are supported as well. Care should be taken that the `dens=` input density and the approximation functions have the same normalization. We can compare the density along the $R=10 \ z$ line as

```
>>> xs= numpy.linspace(0.3, 2., 1001)
>>> semilogy(xs, dp.dens(xs, xs/10.))
>>> semilogy(xs, dscfp.dens(xs, xs/10.))
```

which gives



The agreement is good out to 5 scale lengths and scale heights and then starts to degrade. We can also integrate orbits and compare them

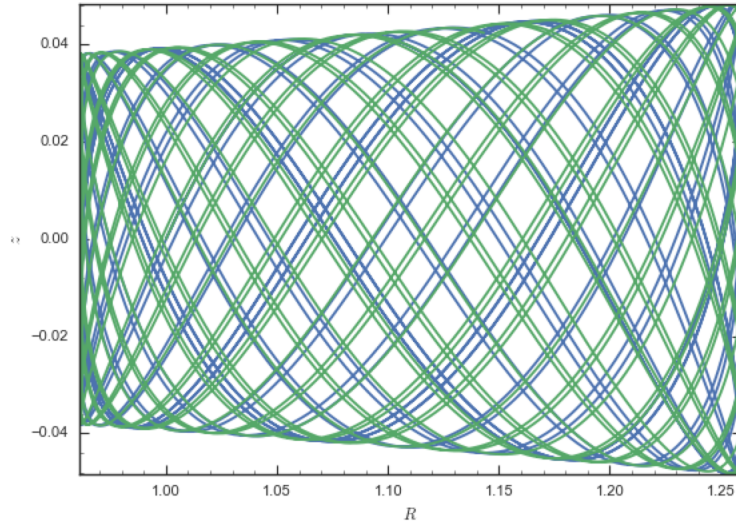
```
>>> from galpy.orbit import Orbit
>>> o= Orbit([1., 0.1, 0.9, 0., 0.1, 0.])
>>> ts= numpy.linspace(0., 100., 10001)
```

(continues on next page)

(continued from previous page)

```
>>> o.integrate(ts, dp)
>>> o.plot()
>>> o.integrate(ts, dscfp)
>>> o.plot(overplot=True)
```

which gives



The orbits diverge slightly because the potentials are not quite the same, but have very similar properties otherwise (peri- and apogalacticons, eccentricity, ...). By increasing the order of the SCF approximation, the potential can be gotten closer to the target density. Note that orbit integration in the `DiskSCFPotential` is much faster than that of the `DoubleExponentialDisk` potential

```
>>> timeit(o.integrate(ts, dp))
# 1 loops, best of 3: 5.83 s per loop
>>> timeit(o.integrate(ts, dscfp))
# 1 loops, best of 3: 286 ms per loop
```

The *SCFPotential* and *DiskSCFPotential* can be used wherever general potentials can be used in galpy.

1.4.8 The potential of N-body simulations

galpy can setup and work with the frozen potential of an N-body simulation. This allows us to study the properties of such potentials in the same way as other potentials in galpy. We can also investigate the properties of orbits in these potentials and calculate action-angle coordinates, using the galpy framework. Currently, this functionality is limited to axisymmetrized versions of the N-body snapshots, although this capability could be somewhat straightforwardly expanded to full triaxial potentials. The use of this functionality requires `pynbody` to be installed; the potential of any snapshot that can be loaded with `pynbody` can be used within galpy.

As a first, simple example of this we look at the potential of a single simulation particle, which should correspond to galpy's `KeplerPotential`. We can create such a single-particle snapshot using `pynbody` by doing

```
>>> import pynbody
>>> s= pynbody.new(star=1)
>>> s['mass']= 1.
>>> s['eps']= 0.
```

and we get the potential of this snapshot in galpy by doing

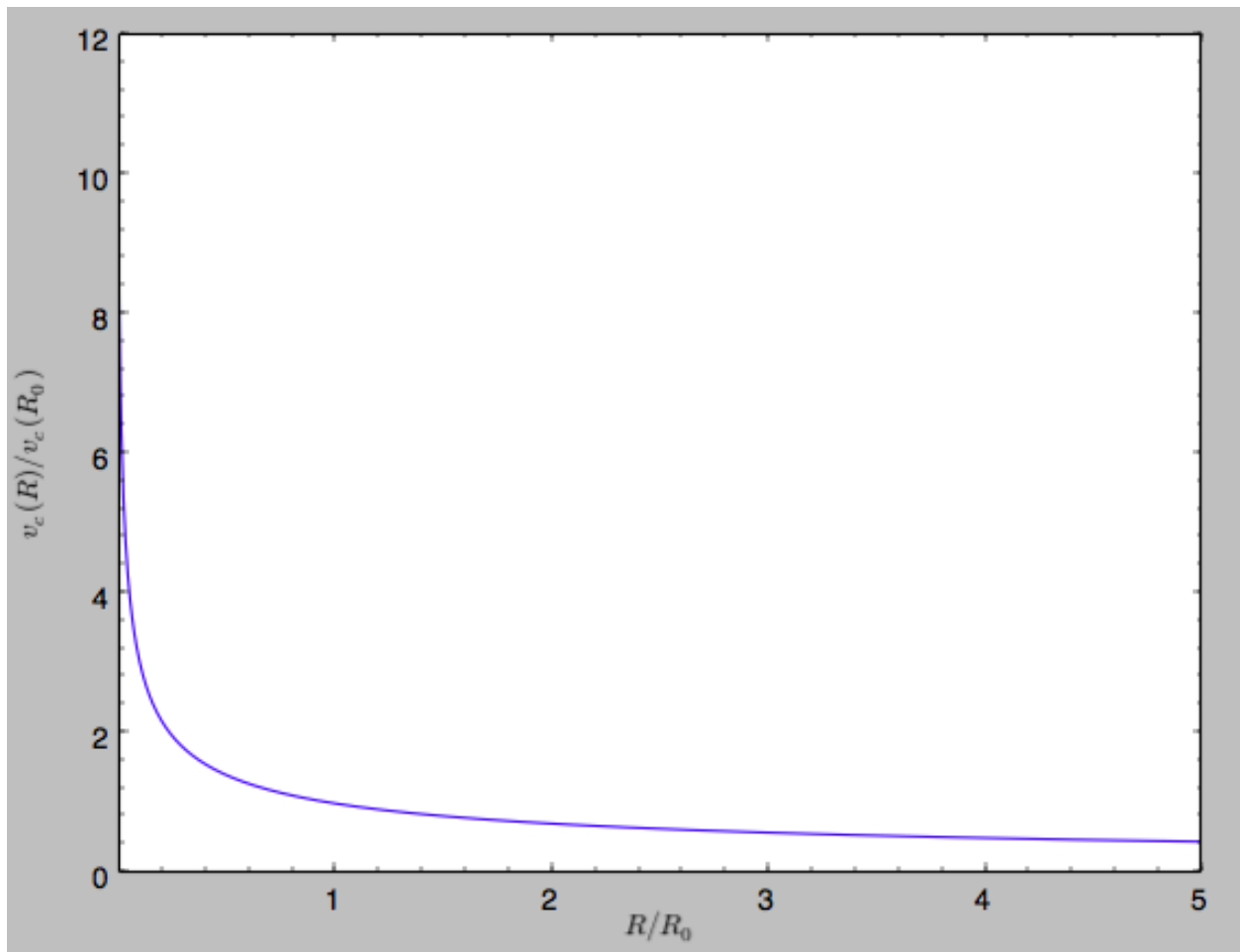
```
>>> from galpy.potential import SnapshotRZPotential
>>> sp= SnapshotRZPotential(s,num_threads=1)
```

With these definitions, this snapshot potential should be the same as `KeplerPotential` with an amplitude of one, which we can test as follows

```
>>> from galpy.potential import KeplerPotential
>>> kp= KeplerPotential(amp=1.)
>>> print(sp(1.1,0.),kp(1.1,0.),sp(1.1,0.)-kp(1.1,0.))
# (-0.90909090909090906, -0.9090909090909091, 0.0)
>>> print(sp.Rforce(1.1,0.),kp.Rforce(1.1,0.),sp.Rforce(1.1,0.)-kp.Rforce(1.1,0.))
# (-0.82644628099173545, -0.8264462809917353, -1.1102230246251565e-16)
```

`SnapshotRZPotential` instances can be used wherever other `galpy` potentials can be used (note that the second derivatives have not been implemented, such that functions depending on those will not work). For example, we can plot the rotation curve

```
>>> sp.plotRotcurve()
```



Because evaluating the potential and forces of a snapshot is computationally expensive, most useful applications of frozen N-body potentials employ interpolated versions of the snapshot potential. These can be setup in `galpy` using an `InterpSnapshotRZPotential` class that is a subclass of the `interpRZPotential` described above and that can be used in the same manner. To illustrate its use we will make use of one of `pynbody`'s example snapshots, `g15784`. This snapshot is used [here](#) to illustrate `pynbody`'s use. Please follow the instructions there on how to

download this snapshot.

Once you have downloaded the pynbody testdata, we can load this snapshot using

```
>>> s = pynbody.load('testdata/g15784.1r.01024.gz')
```

(please adjust the path according to where you downloaded the pynbody testdata). We get the main galaxy in this snapshot, center the simulation on it, and align the galaxy face-on using

```
>>> h = s.halos()
>>> h1 = h[1]
>>> pynbody.analysis.halo.center(h1,mode='hyb')
>>> pynbody.analysis.angmom.faceon(h1, cen=(0,0,0),mode='ssc')
```

we also convert the simulation to physical units, but set $G=1$ by doing the following

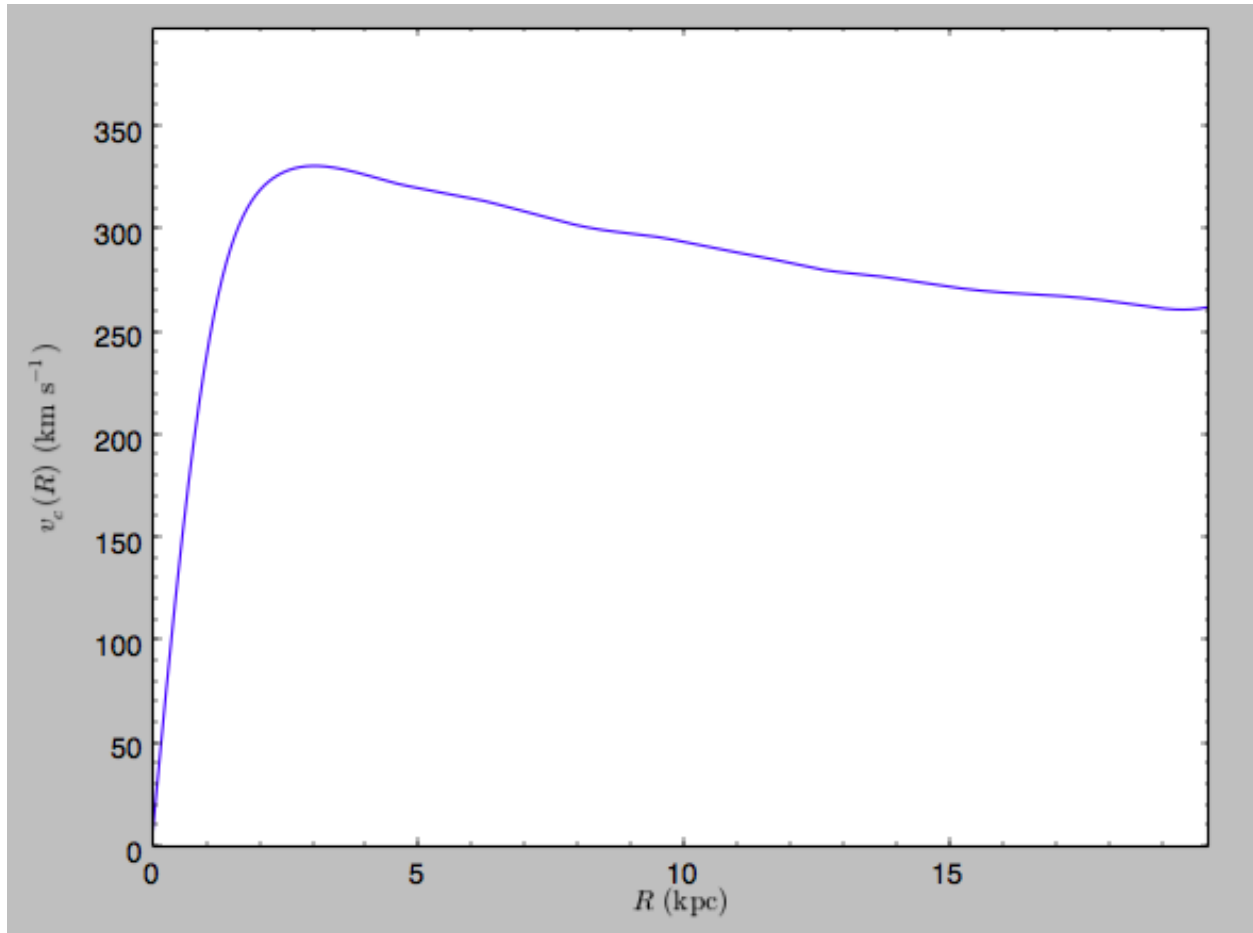
```
>>> s.physical_units()
>>> from galpy.util.bovy_conversion import _G
>>> g= pynbody.array.SimArray(_G/1000.)
>>> g.units= 'kpc Msol**-1 km**2 s**-2 G**-1'
>>> s._arrays['mass']= s._arrays['mass']*g
```

We can now load an interpolated version of this snapshot's potential into galpy using

```
>>> from galpy.potential import InterpSnapshotRZPotential
>>> spi= InterpSnapshotRZPotential(h1,rgrid=(numpy.log(0.01),numpy.log(20.)),101),
↳ logR=True,zgrid=(0.,10.,101),interpPot=True,zsym=True)
```

where we further assume that the potential is symmetric around the mid-plane ($z=0$). This instantiation will take about ten to fifteen minutes. This potential instance has *physical* units (and thus the `rgrid=` and `zgrid=` inputs are given in kpc if the simulation's distance unit is kpc). For example, if we ask for the rotation curve, we get the following:

```
>>> spi.plotRotcurve(Rrange=[0.01,19.9],xlabel=r'$R\,(\mathrm{kpc})$',ylabel=r'$v_{_c(R)\,(\mathrm{km}\,s^{-1})}$')
```



This can be compared to the rotation curve calculated by `pynbody`, see [here](#).

Because `galpy` works best in a system of *natural units* as explained in [Units in galpy](#), we will convert this instance to natural units using the circular velocity at $R=10$ kpc, which is

```
>>> spi.vcirc(10.)  
# 294.62723076942245
```

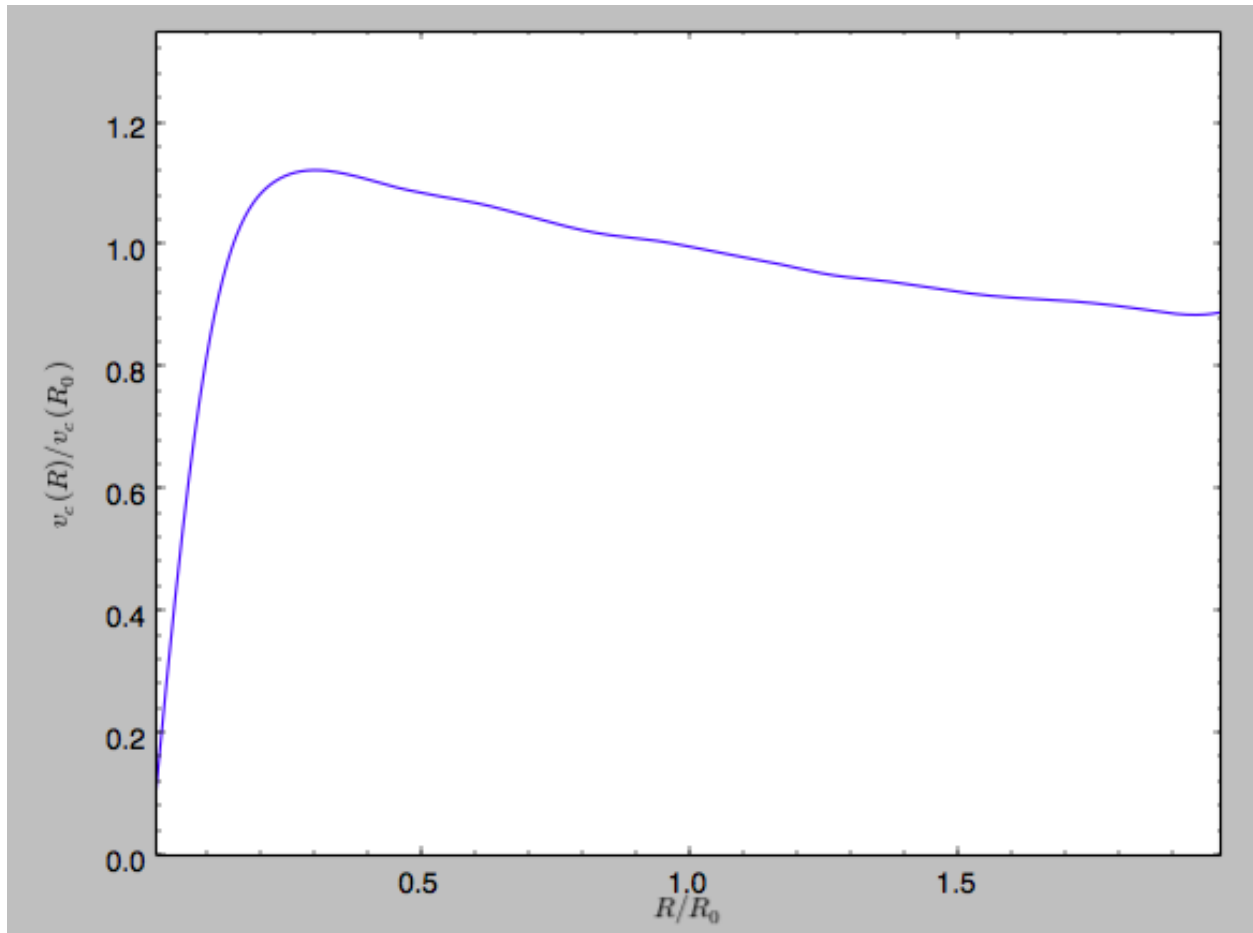
To convert to *natural units* we do

```
>>> spi.normalize(R0=10.)
```

We can then again plot the rotation curve, keeping in mind that the distance unit is now R_0

```
>>> spi.plotRotcurve(Rrange=[0.01,1.99])
```

which gives

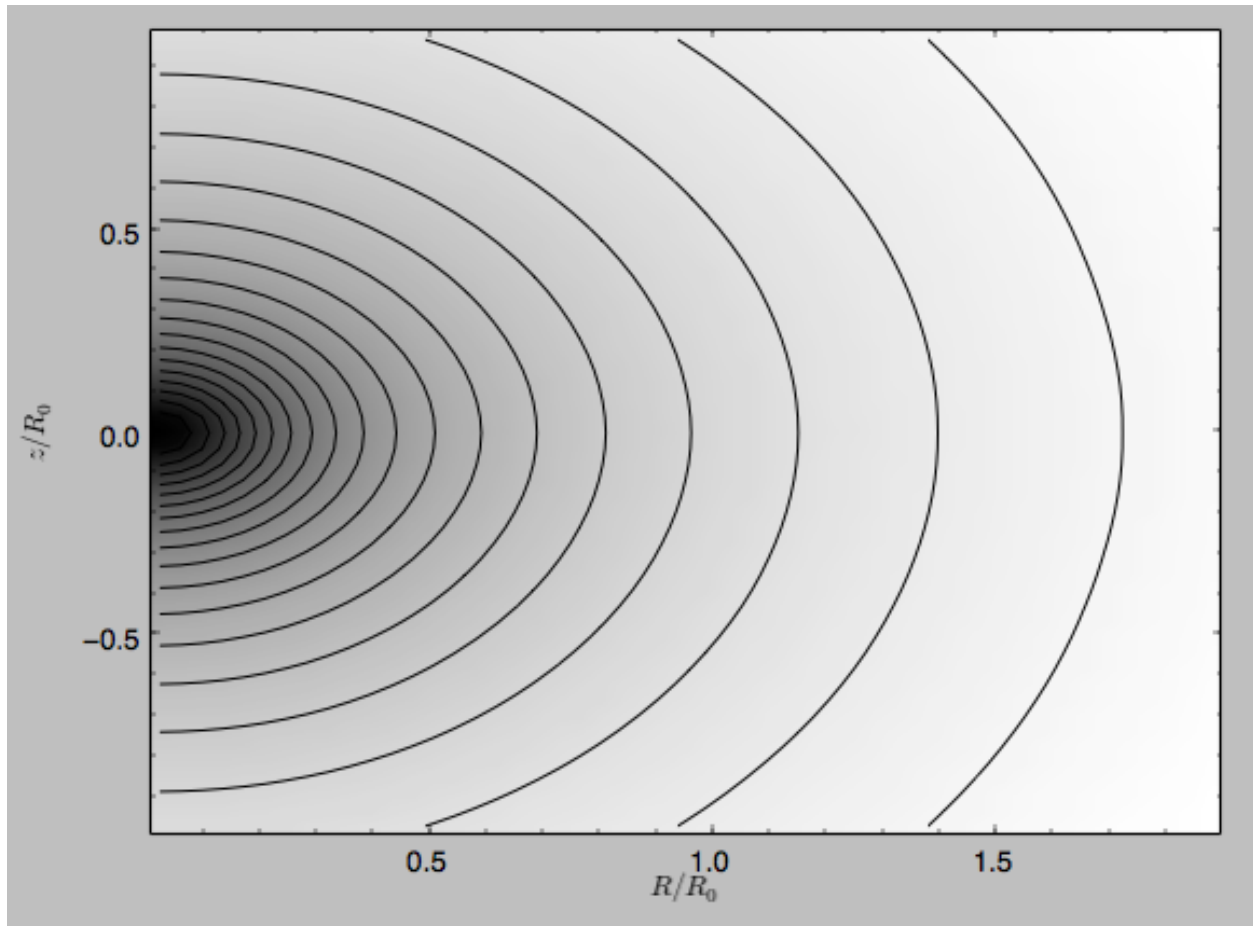


in particular

```
>>> spi.vcirc(1.)  
# 1.00000000000000002
```

We can also plot the potential

```
>>> spi.plot(rmin=0.01, rmax=1.9, nrs=51, zmin=-0.99, zmax=0.99, nzs=51)
```



Clearly, this simulation's potential is quite spherical, which is confirmed by looking at the flattening

```
>>> spi.flattening(1.,0.1)
# 0.86675711023391921
>>> spi.flattening(1.5,0.1)
# 0.94442750306256895
```

The epicycle and vertical frequencies can also be interpolated by setting the `interpepifreq=True` or `interpverticalfreq=True` keywords when instantiating the `InterpSnapshotRZPotential` object.

1.4.9 Conversion to NEMO potentials

NEMO is a set of tools for studying stellar dynamics. Some of its functionality overlaps with that of `galpy`, but many of its programs are very complementary to `galpy`. In particular, it has the ability to perform N-body simulations with a variety of poisson solvers, which is currently not supported by `galpy` (and likely will never be directly supported). To encourage interaction between `galpy` and **NEMO** it is quite useful to be able to convert potentials between these two frameworks, which is not completely trivial. In particular, **NEMO** contains Walter Dehnen's fast collisionless `gyrfalcon` code (see [2000ApJ...536L..39D](#) and [2002JCoPh.179...27D](#)) and the discussion here focuses on how to run N-body simulations using external potentials defined in `galpy`.

Some `galpy` potential instances support the functions `nemo_accname` and `nemo_accpars` that return the name of the **NEMO** potential corresponding to this `galpy` Potential and its parameters in **NEMO** units. These functions assume that you use **NEMO** with `WD_units`, that is, positions are specified in kpc, velocities in kpc/Gyr, times in Gyr, and $G=1$. For the Miyamoto-Nagai potential above, you can get its name in the **NEMO** framework as


```
>>> mp.nemo_accname()
# 'MiyamotoNagai'
```

and its parameters as

```
>>> mp.nemo_accpars(220.,8.)
# '0,592617.11132,4.0,0.3'
```

assuming that we scale velocities by $v_0=220$ km/s and positions by $r_0=8$ kpc in galpy. These two strings can then be given to the `gyrfalcON` `accname=` and `accpars=` keywords.

We can do the same for lists of potentials. For example, for `MWPotential2014` we do

```
>>> from galpy.potential import nemo_accname, nemo_accpars
>>> nemo_accname(MWPotential2014)
# 'PowSphwCut+MiyamotoNagai+NFW'
>>> nemo_accpars(MWPotential2014,220.,8.)
# '0,1001.79126907,1.8,1.9#0,306770.418682,3.0,0.28#0,16.0,162.958241887'
```

Therefore, these are the `accname=` and `accpars=` that one needs to provide to `gyrfalcON` to run a simulation in `MWPotential2014`.

Note that the NEMO potential `PowSphwCut` is *not* a standard NEMO potential. This potential can be found in the `nemo/` directory of the `galpy` source code; this directory also contains a `Makefile` that can be used to compile the extra NEMO potential and install it in the correct NEMO directory (this requires one to have NEMO running, i.e., having sourced `nemo_start`).

You can use the `PowSphwCut.cc` file in the `nemo/` directory as a template for adding additional potentials in `galpy` to the NEMO framework. To figure out how to convert the normalized `galpy` potential to an amplitude when scaling to physical coordinates (like kpc and kpc/Gyr), one needs to look at the scaling of the radial force with R . For example, from the definition of `MiyamotoNagaiPotential`, we see that the radial force scales as R^{-2} . For a general scaling $R^{-\alpha}$, the amplitude will scale as $V_0^2 R_0^{\alpha-1}$ with the velocity V_0 and position R_0 of the $v=1$ at $R=1$ normalization. Therefore, for the `MiyamotoNagaiPotential`, the physical amplitude scales as $V_0^2 R_0$. For the `LogarithmicHaloPotential`, the radial force scales as R^{-1} , so the amplitude scales as V_0^2 .

Currently, only the `MiyamotoNagaiPotential`, `NFWPotential`, `PowerSphericalPotentialwCutoff`, `PlummerPotential`, `MN3ExponentialDiskPotential`, and the `LogarithmicHaloPotential` have this NEMO support. Combinations of the first three are also supported (e.g., `MWPotential2014`); they can also be combined with spherical `LogarithmicHaloPotentials`. Because of the definition of the logarithmic potential in NEMO, it cannot be flattened in z , so to use a flattened logarithmic potential, one has to flip y and z between `galpy` and NEMO (one can flatten in y).

1.4.10 NEW in v1.4: : Dissipative forces

While almost all of the forces that you can use in `galpy` derive from a potential (that is, the force is the gradient of a scalar function, the potential, meaning that the forces are *conservative*), `galpy` also supports dissipative forces. Dissipative forces all inherit from the `DissipativeForce` class and they are required to take the velocity $\mathbf{v}=[v_R, v_T, v_Z]$ in cylindrical coordinates as an argument to the force in addition to the standard $(R, z, \phi=0, t=0)$. The set of functions `evaluateXforces` (with $X=R, z, r, \phi$, etc.) will evaluate the force due to `Potential` instances, `DissipativeForce` instances, or lists of combinations of these two.

Currently, the only dissipative force implemented in `galpy` is `ChandrasekharDynamicalFrictionForce`, an implementation of the classic Chandrasekhar dynamical-friction formula, with recent tweaks to better represent the results from N -body simulations.

Note that there is currently no support for implementing dissipative forces in C. Thus, only Python-based integration methods are available for any dissipative forces.

Warning: Dissipative forces can currently only be used for 3D orbits in `galpy`. The code should throw an error when they are used for 2D orbits.

1.4.11 Adding potentials to the galpy framework

Potentials in `galpy` can be used in many places such as orbit integration, distribution functions, or the calculation of action-angle variables, and in most cases any instance of a potential class that inherits from the general `Potential` class (or a list of such instances) can be given. For example, all orbit integration routines work with any list of instances of the general `Potential` class. Adding new potentials to `galpy` therefore allows them to be used everywhere in `galpy` where general `Potential` instances can be used. Adding a new class of potentials to `galpy` consists of the following series of steps (for steps to add a new wrapper potential, also see [the next section](#)):

1. Implement the new potential in a class that inherits from `galpy.potential.Potential`. The new class should have an `__init__` method that sets up the necessary parameters for the class. An amplitude parameter `amp=` and two units parameters `ro=` and `vo=` should be taken as an argument for this class and before performing any other setup, the `galpy.potential.Potential.__init__(self, amp=amp, ro=ro, vo=vo, amp_units=)` method should be called to setup the amplitude and the system of units; the `amp_units=` keyword specifies the physical units of the amplitude parameter (e.g., `amp_units='velocity2'` when the units of the amplitude are velocity-squared) To add support for normalizing the potential to standard `galpy` units, one can call the `galpy.potential.Potential.normalize` function at the end of the `__init__` function.

The new potential class should implement some of the following functions:

- `_evaluate(self, R, z, phi=0., t=0)` which evaluates the potential itself (*without* the `amp` factor, which is added in the `__call__` method of the general `Potential` class).
- `_Rforce(self, R, z, phi=0., t=0.)` which evaluates the radial force in cylindrical coordinates ($-d \text{ potential} / d R$).
- `_zforce(self, R, z, phi=0., t=0.)` which evaluates the vertical force in cylindrical coordinates ($-d \text{ potential} / d z$).
- `_R2deriv(self, R, z, phi=0., t=0.)` which evaluates the second (cylindrical) radial derivative of the potential ($d^2 \text{ potential} / d R^2$).
- `_z2deriv(self, R, z, phi=0., t=0.)` which evaluates the second (cylindrical) vertical derivative of the potential ($d^2 \text{ potential} / d z^2$).
- `_Rzderiv(self, R, z, phi=0., t=0.)` which evaluates the mixed (cylindrical) radial and vertical derivative of the potential ($d^2 \text{ potential} / d R d z$).
- `_dens(self, R, z, phi=0., t=0.)` which evaluates the density. If not given, the density is computed using the Poisson equation from the first and second derivatives of the potential (if all are implemented).
- `_mass(self, R, z=0., t=0.)` which evaluates the mass. For spherical potentials this should give the mass enclosed within the spherical radius; for axisymmetric potentials this should return the mass up to R and between $-Z$ and Z . If not given, the mass is computed by integrating the density (if it is implemented or can be calculated from the Poisson equation).
- `_phiforce(self, R, z, phi=0., t=0.)`: the azimuthal force in cylindrical coordinates (assumed zero if not implemented).
- `_phi2deriv(self, R, z, phi=0., t=0.)`: the second azimuthal derivative of the potential in cylindrical coordinates ($d^2 \text{ potential} / d \phi^2$; assumed zero if not given).

- `_Rphideriv(self, R, z, phi=0., t=0.)`: the mixed radial and azimuthal derivative of the potential in cylindrical coordinates ($d^2 \text{ potential} / dR dz$; assumed zero if not given).
- `OmegaP(self)`: returns the pattern speed for potentials with a pattern speed (used to compute the Jacobi integral for orbits).

If you want to be able to calculate the concentration for a potential, you also have to set `self._scale` to a scale parameter for your potential.

The code for `galpy.potential.MiyamotoNagaiPotential` gives a good template to follow for 3D axisymmetric potentials. Similarly, the code for `galpy.potential.CosmphiDiskPotential` provides a good template for 2D, non-axisymmetric potentials.

After this step, the new potential will work in any part of galpy that uses pure python potentials. To get the potential to work with the C implementations of orbit integration or action-angle calculations, the potential also has to be implemented in C and the potential has to be passed from python to C.

The `__init__` method should be written in such a way that a relevant object can be initialized using `Classname()` (i.e., there have to be reasonable defaults given for all parameters, including the amplitude); doing this allows the nose tests for potentials to automatically check that your Potential's potential function, force functions, second derivatives, and density (through the Poisson equation) are correctly implemented (if they are implemented). The continuous-integration platform that builds the galpy codebase upon code pushes will then automatically test all of this, streamlining push requests of new potentials.

A few attributes need to be set depending on the potential: `hasC=True` for potentials for which the forces and potential are implemented in C (see below); `self.hasC_dxdv=True` for potentials for which the (planar) second derivatives are implemented in C; `self.isNonAxi=True` for non-axisymmetric potentials.

2. To add a C implementation of the potential, implement it in a .c file under `potential/potential_c_ext`. Look at `potential/potential_c_ext/LogarithmicHaloPotential.c` for the right format for 3D, axisymmetric potentials, or at `potential/potential_c_ext/LopsidedDiskPotential.c` for 2D, non-axisymmetric potentials.

For orbit integration, the functions such as:

- `double LogarithmicHaloPotentialRforce(double R, double Z, double phi, double t, struct potentialArg * potentialArgs)`
- `double LogarithmicHaloPotentialzforce(double R, double Z, double phi, double t, struct potentialArg * potentialArgs)`

are most important. For some of the action-angle calculations

- `double LogarithmicHaloPotentialEval(double R, double Z, double phi, double t, struct potentialArg * potentialArgs)`

is most important (i.e., for those algorithms that evaluate the potential). The arguments of the potential are passed in a `potentialArgs` structure that contains `args`, which are the arguments that should be unpacked. Again, looking at some example code will make this clear. The `potentialArgs` structure is defined in `potential/potential_c_ext/galpy_potentials.h`.

3. Add the potential's function declarations to `potential/potential_c_ext/galpy_potentials.h`
4. (4. and 5. for planar orbit integration) Edit the code under `orbit/orbit_c_ext/integratePlanarOrbit.c` to set up your new potential (in the `parse_leapFuncArgs` function).
5. Edit the code in `orbit/integratePlanarOrbit.py` to set up your new potential (in the `_parse_pot` function).
6. Edit the code under `orbit/orbit_c_ext/integrateFullOrbit.c` to set up your new potential (in the `parse_leapFuncArgs_Full` function).

7. Edit the code in `orbit/integrateFullOrbit.py` to set up your new potential (in the `_parse_pot` function).
8. (for using the `actionAngleStaeckel` methods in C) Edit the code in `actionAngle/actionAngle_c_ext/actionAngle.c` to parse the new potential (in the `parse_actionAngleArgs` function).
9. Finally, add `self.hasC= True` to the initialization of the potential in question (after the initialization of the super class, or otherwise it will be undone). If you have implemented the necessary second derivatives for integrating phase-space volumes, also add `self.hasC_dxdv=True`.

After following the relevant steps, the new potential class can be used in any galpy context in which C is used to speed up computations.

1.4.12 Adding wrapper potentials to the galpy framework

Wrappers all inherit from the general `WrapperPotential` or `planarWrapperPotential` classes (which themselves inherit from the `Potential` and `planarPotential` classes and therefore all wrappers are `Potentials` or `planarPotentials`). Depending on the complexity of the wrapper, wrappers can be implemented much more economically in Python than new `Potential` instances as described [above](#).

To add a Python implementation of a new wrapper, classes need to inherit from `parentWrapperPotential`, take the potentials to be wrapped as a `pot=` (a `Potential`, `planarPotential`, or a list thereof; automatically assigned to `self._pot`) input to `__init__`, and implement the `_wrap(self, attribute, *args, **kwargs)` function. This function modifies the `Potential` functions `_evaluate`, `_Rforce`, etc. (all of those listed [above](#)), with `attribute` the function that is being modified. Inheriting from `parentWrapperPotential` gives the class access to the `self._wrap_pot_func(attribute)` function which returns the relevant function for each attribute. For example, `self._wrap_pot_func('_evaluate')` returns the `evaluatePotentials` function that can then be called as `self._wrap_pot_func('_evaluate')(self._pot, R, Z, phi=phi, t=t)` to evaluate the potentials being wrapped. By making use of `self._wrap_pot_func`, wrapper potentials can be implemented in just a few lines. Your `__init__` function should *only* initialize things in your wrapper; there is no need to manually assign `self._pot` or to call the superclass' `__init__` (all automatically done for you!).

To correctly work with both 3D and 2D potentials, inputs to `_wrap` need to be specified as `*args, **kwargs`: grab the values you need for `R, z, phi, t` from these as `R=args[0]`, `z=0` if `len(args) == 1` else `args[1]`, `phi=kwargs.get('phi', 0.)`, `t=kwargs.get('t', 0.)`, where the complicated expression for `z` is to correctly deal with both 3D and 2D potentials (of course, if your wrapper depends on `z`, it probably doesn't make much sense to apply it to a 2D `planarPotential`; you could check the dimensionality of `self._pot` in your wrapper's `__init__` function with `from galpy.potential.Potential._dim` and raise an error if it is not 3 in this case). Wrapping a 2D potential automatically results in a wrapper that is a subclass of `planarPotential` rather than `Potential`; this is done by the setup in `parentWrapperPotential` and hidden from the user. For wrappers of `planarPotentials`, `self._wrap_pot_func(attribute)` will return the `evaluateplanarPotentials` etc. functions instead, but this is again hidden from the user if you implement the `_wrap` function as explained above.

As an example, for the `DehnenSmoothWrapperPotential`, the `_wrap` function is

```
def _wrap(self, attribute, *args, **kwargs):
    return self._smooth(kwargs.get('t', 0.)) \
        *self._wrap_pot_func(attribute)(self._pot, *args, **kwargs)
```

where `smooth(t)` returns the smoothing function of the amplitude. When any of the basic `Potential` functions are called (`_evaluate`, `_Rforce`, etc.), `_wrap` gets called by the superclass `WrapperPotential`, and the `_wrap` function returns the corresponding function for the wrapped potentials with the amplitude modified by `smooth(t)`. Therefore, one does not need to implement each of the `_evaluate`, `_Rforce`, etc. functions like for regular potential. The rest of the `DehnenSmoothWrapperPotential` is essentially (slightly simplified in non-crucial aspects)

```

def __init__(self, amp=1., pot=None, tform=-4., tsteady=None, ro=None, vo=None) :
    # Note: (i) don't assign self._pot and (ii) don't run super.__init__
    self._tform= tform
    if tsteady is None:
        self._tsteady= self._tform/2.
    else:
        self._tsteady= self._tform+tsteady
    self.hasC= True
    self.hasC_dxdv= True

def _smooth(self,t):
    #Calculate relevant time
    if t < self._tform:
        smooth= 0.
    elif t < self._tsteady:
        deltat= t-self._tform
        xi= 2.*deltat/(self._tsteady-self._tform)-1.
        smooth= (3./16.*xi**5.-5./8*xi**3.+15./16.*xi+.5)
    else: #bar is fully on
        smooth= 1.
    return smooth

```

The source code for `DehnenSmoothWrapperPotential` potential may act as a guide to implementing new wrappers.

C implementations of potential wrappers can also be added in a similar way as C implementations of regular potentials (all of the steps listed in the [previous section](#) for adding a potential to C need to be followed). All of the necessary functions (`...Rforce`, `...zforce`, `...phiforce`, etc.) need to be implemented separately, but by including `galpy_potentials.h` calling the relevant functions of the wrapped potentials is easy. Look at `DehnenSmoothWrapperPotential.c` for an example that can be straightforwardly edited for other wrappers.

The glue between Python and C for wrapper potentials needs to glue both the wrapper and the wrapped potentials. This can be easily achieved by recursively calling the `_parse_pot` glue functions in Python (see the previous section; this needs to be done separately for each potential currently) and the `parse_leapFuncArgs` and `parse_leapFuncArgs_Full` functions in C (done automatically for all wrappers). Again, following the example of `DehnenSmoothWrapperPotential.py` should allow for a straightforward implementation of the glue for any new wrappers. Wrapper potentials should be given negative potential types in the glue to distinguish them from regular potentials.

1.4.13 NEW in v1.4: Adding dissipative forces to the galpy framework

Dissipative forces are implemented in much the same way as forces that derive from potentials. Rather than inheriting from `galpy.potential.Potential`, dissipative forces inherit from `galpy.potential.DissipativeForce`. The procedure for implementing a new class of dissipative force is therefore very similar to that for [implementing a new potential](#). The main differences are that (a) you only need to implement the forces and (b) the forces are required to take an extra keyword argument `v=` that gives the velocity in cylindrical coordinates (because dissipative forces will in general depend on the current velocity). Thus, the steps are:

1. Implement the new dissipative force in a class that inherits from `galpy.potential.DissipativeForce`. The new class should have an `__init__` method that sets up the necessary parameters for the class. An amplitude parameter `amp=` and two units parameters `ro=` and `vo=` should be taken as an argument for this class and before performing any other setup, the `galpy.potential.DissipativeForce.__init__(self, amp=amp, ro=ro, vo=vo, amp_units=)` method should be called to setup the amplitude and the system of units; the `amp_units=` keyword specifies the physical units of the amplitude parameter (e.g., `amp_units='mass'` when the units of the amplitude are mass)

The new dissipative-force class should implement the following functions:

- `_Rforce(self, R, z, phi=0., t=0., v=None)` which evaluates the radial force in cylindrical coordinates
- `_phiforce(self, R, z, phi=0., t=0., v=None)` which evaluates the azimuthal force in cylindrical coordinates
- `_zforce(self, R, z, phi=0., t=0., v=None)` which evaluates the vertical force in cylindrical coordinates

The code for `galpy.potential.ChandrasekharDynamicalFrictionForce` gives a good template to follow.

2. That's it, as for now there is no support for implementing a C version of dissipative forces.

1.5 A closer look at orbit integration

1.5.1 UPDATED in v1.4: Orbit initialization

Standard initialization

`Orbits` can be initialized in various coordinate frames. The simplest initialization gives the initial conditions directly in the Galactocentric cylindrical coordinate frame (or in the rectangular coordinate frame in one dimension). `Orbit()` automatically figures out the dimensionality of the space from the initial conditions in this case. In three dimensions initial conditions are given either as `vxvv=[R, vR, vT, z, vz, phi]` or one can choose not to specify the azimuth of the orbit and initialize with `vxvv=[R, vR, vT, z, vz]`. Since potentials in galpy are easily initialized to have a circular velocity of one at a radius equal to one, initial coordinates are best given as a fraction of the radius at which one specifies the circular velocity, and initial velocities are best expressed as fractions of this circular velocity. For example,

```
>>> from galpy.orbit import Orbit
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.,0.1,0.] )
```

initializes a fully three-dimensional orbit, while

```
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.,0.1])
```

initializes an orbit in which the azimuth is not tracked, as might be useful for axisymmetric potentials.

In two dimensions, we can similarly specify fully two-dimensional orbits `o=Orbit(vxvv=[R, vR, vT, phi])` or choose not to track the azimuth and initialize with `o= Orbit(vxvv=[R, vR, vT])`.

In one dimension we simply initialize with `o= Orbit(vxvv=[x, vx])`.

Initialization with physical units

Orbits are normally used in galpy's *natural coordinates*. When `Orbits` are initialized using a distance scale `ro=` and a velocity scale `vo=`, then many `Orbit` methods return quantities in physical coordinates. Specifically, physical distance and velocity scales are specified as

```
>>> op= Orbit(vxvv=[1.,0.1,1.1,0.,0.1,0.],ro=8.,vo=220.)
```

All output quantities will then be automatically be specified in physical units: kpc for positions, km/s for velocities, (km/s)² for energies and the Jacobi integral, km/s kpc for the angular momentum `o.L()` and actions, 1/Gyr for frequencies, and Gyr for times and periods. See below for examples of this.

The actual initial condition can also be specified in physical units. For example, the Orbit above can be initialized as

```
>>> from astropy import units
>>> op= Orbit(vxvv=[8.*units.kpc,22.*units.km/units.s,242*units.km/units.s,0.*units.
↳kpc,22.*units.km/units.s,0.*units.deg])
```

In this case, it is unnecessary to specify the `ro=` and `vo=` scales; when they are not specified, `ro` and `vo` are set to the default values from the *configuration file*. However, if they are specified, then those values rather than the ones from the configuration file are used.

Tip: If you do input and output in physical units, the internal unit conversion specified by `ro=` and `vo=` does not matter!

Inputs to any Orbit method can also be specified with units as an astropy Quantity. galpy's natural units are still used under the hood, as explained in the section on *physical units in galpy*. For example, integration times can be specified in Gyr if you want to integrate for a specific time period.

If for any output you do *not* want the output in physical units, you can specify this by supplying the keyword argument `use_physical=False`.

Initialization from observed coordinates

For orbit integration and characterization of observed stars or clusters, initial conditions can also be specified directly as observed quantities when `radec=True` is set (see further down in this section on how to use an astropy `SkyCoord` instead). In this case a full three-dimensional orbit is initialized as `o= Orbit(vxvv=[RA,Dec,distance,pmRA,pmDec,Vlos],radec=True)` where RA and Dec are expressed in degrees, the distance is expressed in kpc, proper motions are expressed in mas/yr (`pmra = pmra' * cos[Dec]`), and `Vlos` is the heliocentric line-of-sight velocity given in km/s. The observed epoch is currently assumed to be J2000.00. These observed coordinates are translated to the Galactocentric cylindrical coordinate frame by assuming a Solar motion that can be specified as either `solarmotion=hogg (2005ApJ...629..268H)`, `solarmotion=dehnen (1998MNRAS.298..387D)` or `solarmotion=schoenrich` (default; `2010MNRAS.403.1829S`). A circular velocity can be specified as `vo=220` in km/s and a value for the distance between the Galactic center and the Sun can be given as `ro=8.0` in kpc (e.g., `2012ApJ...759..131B`). While the inputs are given in physical units, the orbit is initialized assuming a circular velocity of one at the distance of the Sun (that is, the orbit's position and velocity is scaled to galpy's *natural* units after converting to the Galactocentric coordinate frame, using the specified `ro=` and `vo=`). The parameters of the coordinate transformations are stored internally, such that they are automatically used for relevant outputs (for example, when the RA of an orbit is requested). An example of all of this is:

```
>>> o= Orbit(vxvv=[20.,30.,2.,-10.,20.,50.],radec=True,ro=8.,vo=220.)
```

However, the internally stored position/velocity vector is

```
>>> print(o._orb.vxvv)
# [1.1480792664061401, 0.1994859759019009, 1.8306295160508093, -0.13064400474040533,
↳0.58167185623715167, 0.14066246212987227]
```

and is therefore in *natural* units.

Tip: Initialization using observed coordinates can also use units. So, for example, proper motions can be specified as `2*units.mas/units.yr`.

Similarly, one can also initialize orbits from Galactic coordinates using `o= Orbit(vxvv=[glon,glat,distance,pmll,pmbb,Vlos],lb=True)`, where `glon` and `glat` are Galactic longitude and latitude expressed

in degrees, and the proper motions are again given in mas/yr ($\text{pmll} = \text{pmll}' * \cos[\text{glat}]$):

```
>>> o= Orbit(vxvv=[20., 30., 2., -10., 20., 50.], lb=True, ro=8., vo=220.)
>>> print(o._orb.vxvv)
# [0.79959714332811838, 0.073287283885367677, 0.5286278286083651, 0.12748861331872263,
  ↪ 0.89074407199364924, 0.0927414387396788]
```

When `radec=True` or `lb=True` is set, velocities can also be specified in Galactic coordinates if `UVW=True` is set. The input is then `vxvv=[RA, Dec, distance, U, V, W]`, where the velocities are expressed in km/s. U is, as usual, defined as $-v_R$ (minus v_R).

Finally, orbits can also be initialized using an `astropy.coordinates.SkyCoord` object. For example, the (ra,dec) example from above can also be initialized as:

```
>>> from astropy.coordinates import SkyCoord
>>> import astropy.units as u
>>> c= SkyCoord(ra=20.*u.deg, dec=30.*u.deg, distance=2.*u.kpc,
               pm_ra_cosdec=-10.*u.mas/u.yr, pm_dec=20.*u.mas/u.yr,
               radial_velocity=50.*u.km/u.s)
>>> o= Orbit(c)
```

In this case, you can still specify the properties of the transformation to Galactocentric coordinates using the standard `ro`, `vo`, `zo`, and `solar_motion` keywords, or you can use the `SkyCoord` [Galactocentric frame specification](#) and these are propagated to the `Orbit` instance. For example,

```
>>> from astropy.coordinates import CartesianDifferential
>>> c= SkyCoord(ra=20.*u.deg, dec=30.*u.deg, distance=2.*u.kpc,
               pm_ra_cosdec=-10.*u.mas/u.yr, pm_dec=20.*u.mas/u.yr,
               radial_velocity=50.*u.km/u.s,
               galcen_distance=8.*u.kpc, z_sun=15.*u.pc,
               galcen_v_sun=CartesianDifferential([10.0, 235., 7.]*u.km/u.s))
>>> o= Orbit(c)
```

A subtlety here is that the `galcen_distance` and `ro` keywords are not interchangeable, because the former is the distance between the Sun and the Galactic center and `ro` is the projection of this distance onto the Galactic midplane. Another subtlety is that the `astropy` Galactocentric frame is a right-handed frame, while `galpy` normally uses a left-handed frame, so the sign of the x component of `galcen_v_sun` is the opposite of what it would be in `solar_motion`. Because the Galactocentric frame in `astropy` does not specify the circular velocity, but only the Sun's velocity, you still need to specify `vo` to use a non-default circular velocity.

When orbits are initialized using `radec=True`, `lb=True`, or using a `SkyCoord` physical scales `ro=` and `vo=` are automatically specified (because they have defaults of `ro=8` and `vo=220`). Therefore, all output quantities will be specified in physical units (see above). If you do want to get outputs in `galpy`'s natural coordinates, you can turn this behavior off by doing

```
>>> o.turn_physical_off()
```

All outputs will then be specified in `galpy`'s natural coordinates.

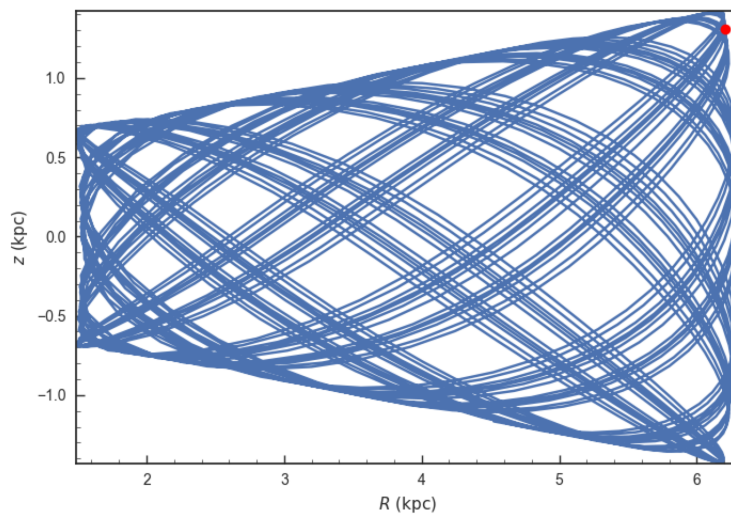
Initialization from an object's name

A convenience method, `Orbit.from_name`, is also available to initialize orbits from the name of an object. For example, for the star [Lacaille 8760](#):


```
>>> o= Orbit.from_name('Lacaille 8760', ro=8., vo=220.)
>>> [o.ra(), o.dec(), o.dist(), o.pmra(), o.pmdec(), o.vlos()]
# [319.31362023999276, -38.86736390000036, 0.003970940656277758, -3258.5529999996584,
↪ -1145.3959999996205, 20.560000000006063]
```

but this also works for some globular clusters, e.g., to obtain *Omega Cen*'s orbit and current location in the Milky Way do:

```
>>> o= Orbit.from_name('Omega Cen')
>>> from galpy.potential import MWPotential2014
>>> ts= numpy.linspace(0.,100.,2001)
>>> o.integrate(ts,MWPotential2014)
>>> o.plot()
>>> plot([o.R()], [o.z()], 'ro')
```



We see that *Omega Cen* is currently close to its maximum distance from both the Galactic center and from the Galactic midplane.

Similarly, you can do:

```
>>> o= Orbit.from_name('LMC')
>>> [o.ra(), o.dec(), o.dist(), o.pmra(), o.pmdec(), o.vlos()]
# [80.894200000000055, -69.756099999999847, 49.99999999999993, 1.909999999999999, 0.
↪ 22900000000000037, 262.19999999999993]
```

The `Orbit.from_name` method attempts to resolve the name of the object in SIMBAD, and then use the observed coordinates found there to generate an `Orbit` instance. In order to query SIMBAD, `Orbit.from_name` requires the `astroquery` package to be installed.

Tip: Setting up an `Orbit` instance *without* arguments will return an `Orbit` instance representing the Sun: `o=Orbit()`. This instance has physical units *turned on by default*, so methods will return outputs in physical units unless you `o.turn_physical_off()`.

Warning: Orbits initialized using `Orbit.from_name` have physical output *turned on by default*, so methods will return outputs in physical units unless you `o.turn_physical_off()`.

1.5.2 Orbit integration

After an orbit is initialized, we can integrate it for a set of times `ts`, given as a numpy array. For example, in a simple logarithmic potential we can do the following

```
>>> from galpy.potential import LogarithmicHaloPotential
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.,0.1,0.])
>>> import numpy
>>> ts= numpy.linspace(0,100,10000)
>>> o.integrate(ts,lp)
```

to integrate the orbit from $t=0$ to $t=100$, saving the orbit at 10000 instances. In physical units, we can integrate for 10 Gyr as follows

```
>>> from astropy import units
>>> ts= numpy.linspace(0,10.,10000)*units.Gyr
>>> o.integrate(ts,lp)
```

If we initialize the Orbit using a distance scale `ro=` and a velocity scale `vo=`, then Orbit plots and outputs will use physical coordinates (currently, times, positions, and velocities)

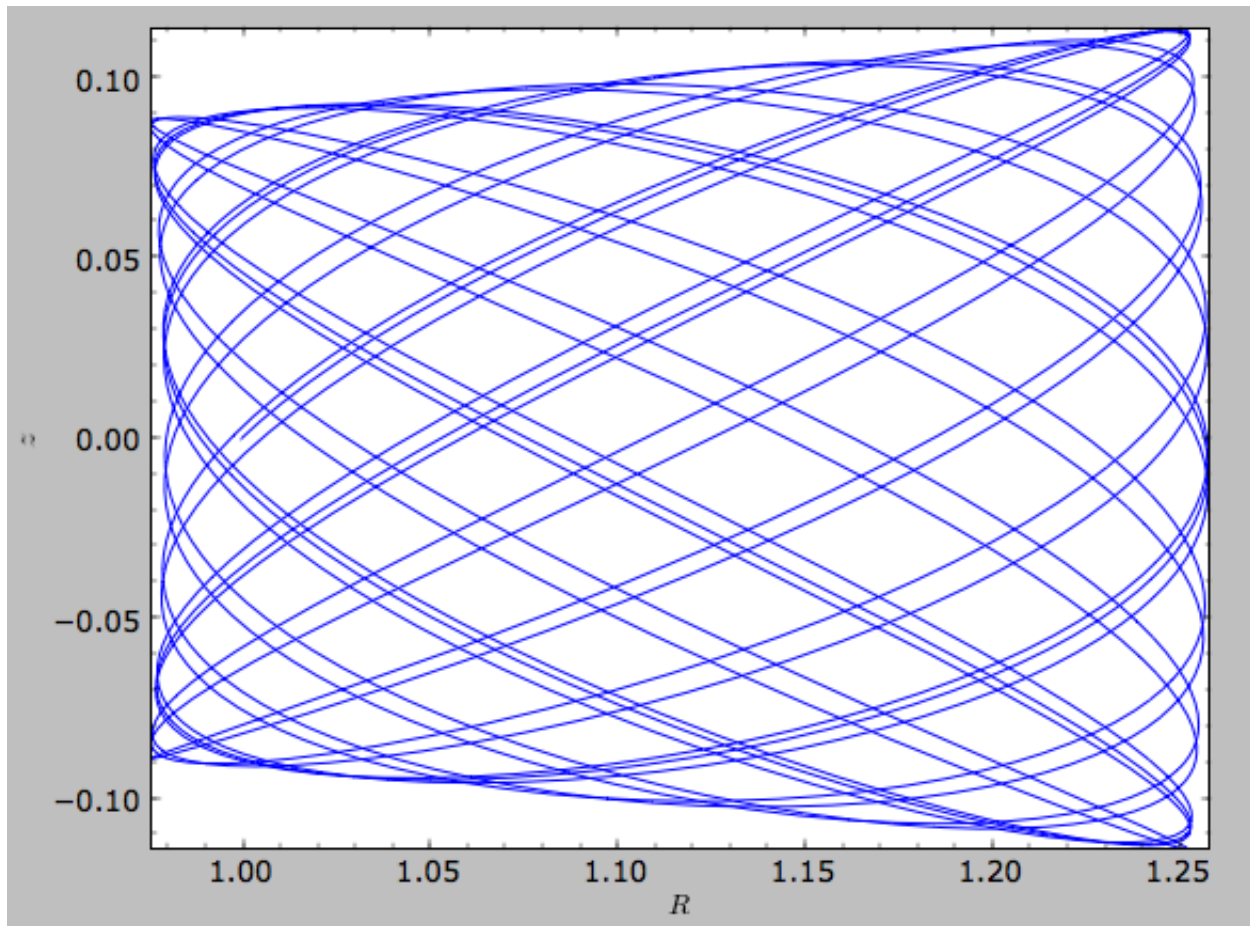
```
>>> op= Orbit(vxvv=[1.,0.1,1.1,0.,0.1,0.],ro=8.,vo=220.) #Use Vc=220 km/s at R= 8 kpc
↪as the normalization
>>> op.integrate(ts,lp)
```

1.5.3 Displaying the orbit

After integrating the orbit, it can be displayed by using the `plot()` function. The quantities that are plotted when `plot()` is called depend on the dimensionality of the orbit: in 3D the (R,z) projection of the orbit is shown; in 2D either (X,Y) is plotted if the azimuth is tracked and (R,vR) is shown otherwise; in 1D (x,vx) is shown. E.g., for the example given above,

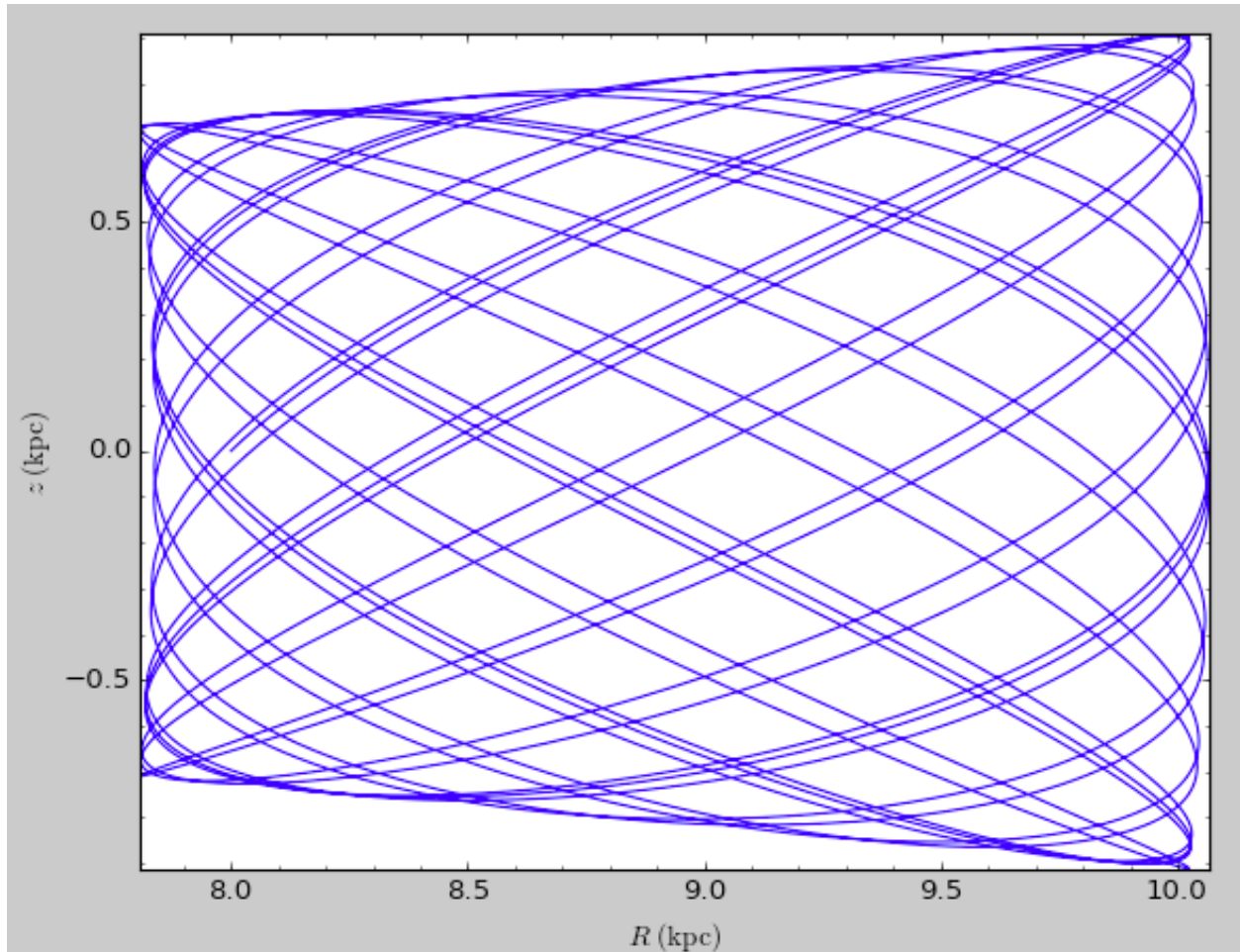
```
>>> o.plot()
```

gives



If we do the same for the Orbit that has physical distance and velocity scales associated with it, we get the following

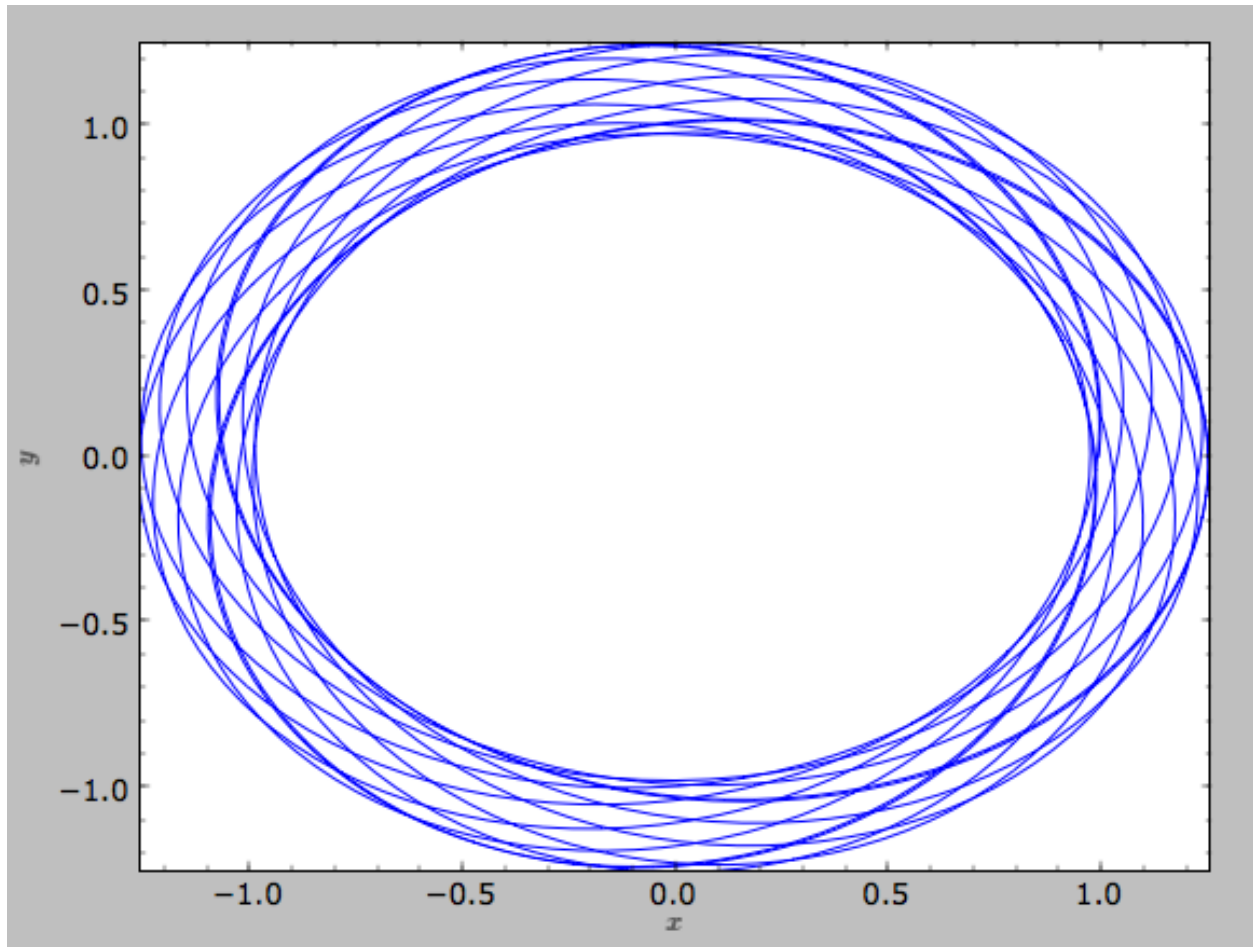
```
>>> op.plot()
```



If we call `op.plot(use_physical=False)`, the quantities will be displayed in natural galpy coordinates. Other projections of the orbit can be displayed by specifying the quantities to plot. E.g.,

```
>>> o.plot(d1='x', d2='y')
```

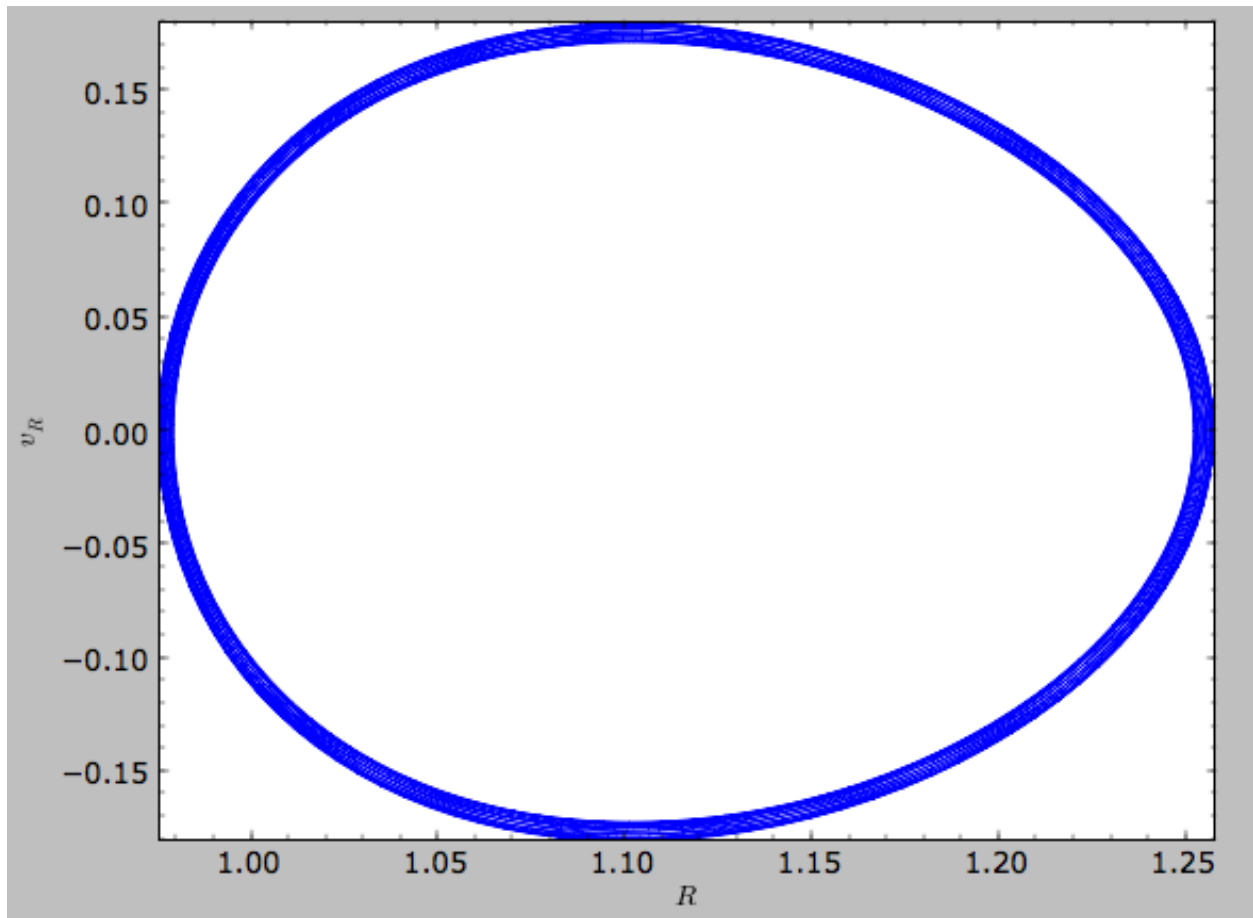
gives the projection onto the plane of the orbit:



while

```
>>> o.plot(d1='R', d2='vR')
```

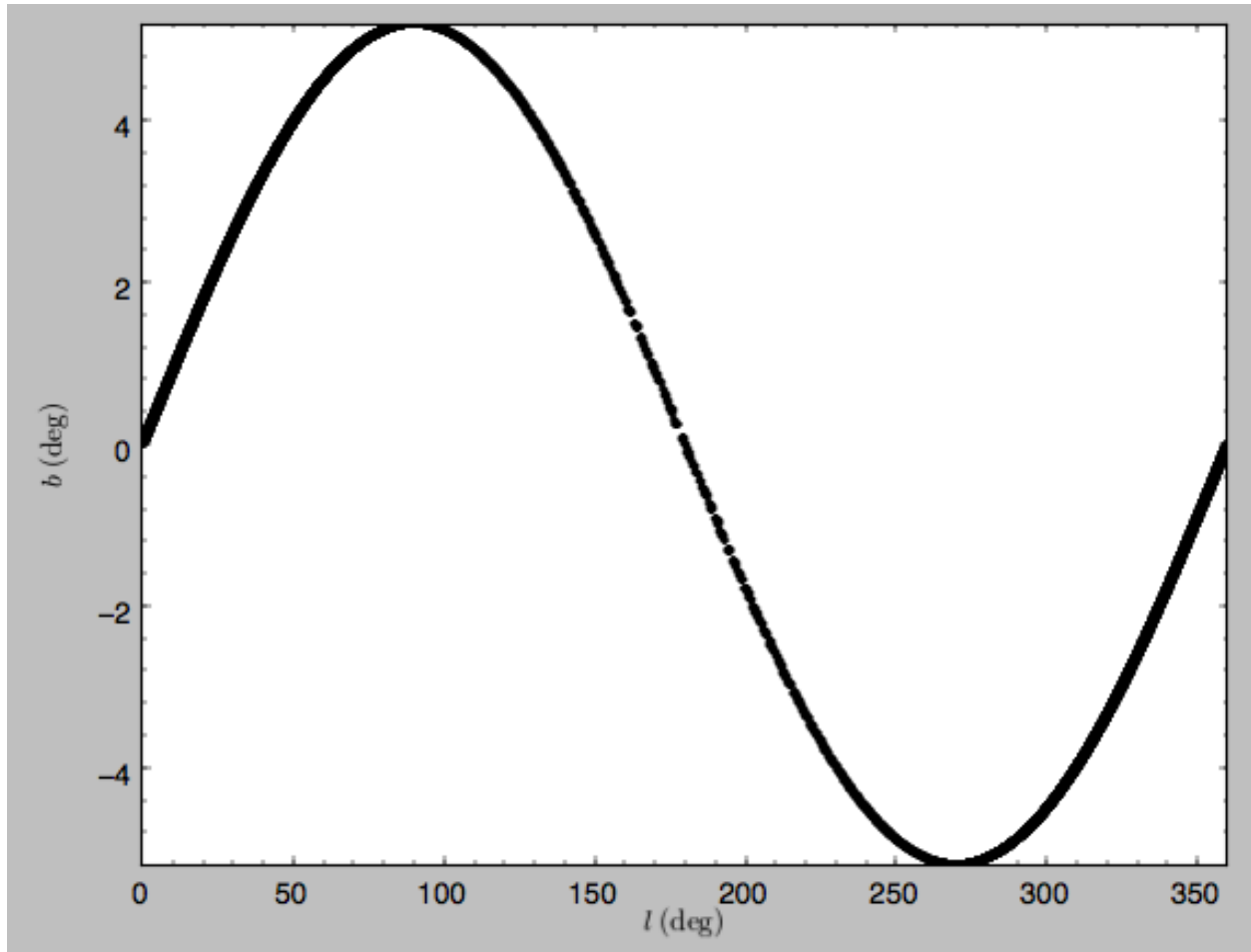
gives the projection onto (R, vR) :



We can also plot the orbit in other coordinate systems such as Galactic longitude and latitude

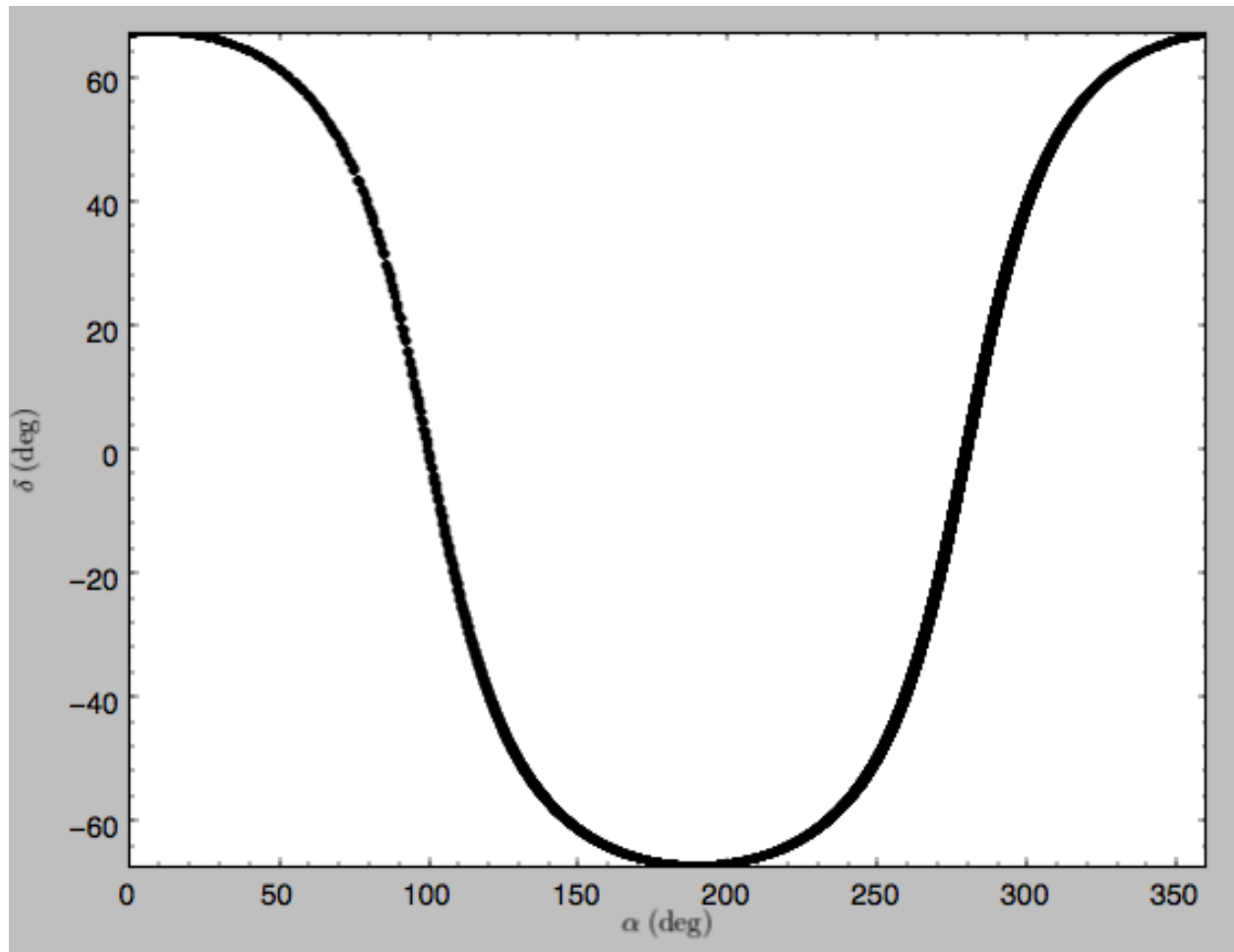
```
>>> o.plot('k.', d1='l1', d2='bb')
```

which shows



or RA and Dec

```
>>> o.plot('k.', d1='ra', d2='dec')
```

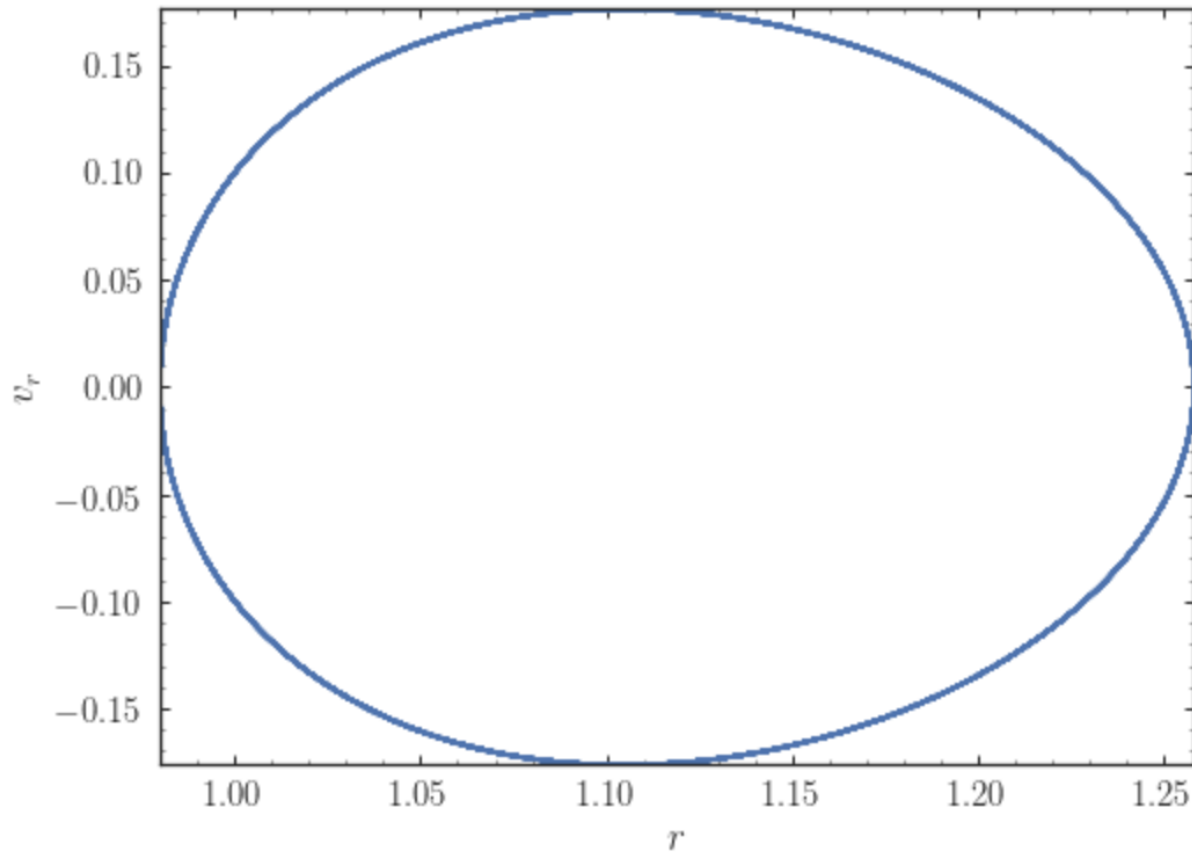


See the documentation of the `o.plot` function and the `o.ra()`, `o.ll()`, etc. functions on how to provide the necessary parameters for the coordinate transformations.

Finally, it is also possible to plot arbitrary functions of time with `Orbit.plot`, by specifying `d1=` or `d2=` as a function. This is for example useful if you want to display the orbit in a different coordinate system. For example, to display the orbital velocity in the spherical radial direction (which is currently not a pre-defined option), you can do the following

```
>>> o.plot(d1='r',
           d2=lambda t: o.vR(t)*o.R(t)/o.r(t)+o.vz(t)*o.z(t)/o.r(t),
           ylabel='v_r')
```

where `d2=` converts the velocity to spherical coordinates. This gives the following orbit (which is closed in this projection, because we are using a spherical potential):



1.5.4 Animating the orbit

Warning: Animating orbits is a new, experimental feature at this time that may be changed in later versions. It has only been tested in a limited fashion. If you are having problems with it, please open an [Issue](#) and list all relevant details about your setup (python version, jupyter version, browser, any error message in full). It may also be helpful to check the javascript console for any errors.

In a [jupyter notebook](#) or in [jupyterlab](#) (jupyterlab versions ≥ 0.33) you can also create an animation of an orbit *after* you have integrated it. For example, to do this for the `op` orbit from above (but only integrated for 2 Gyr to create a shorter animation as an example here), do

```
>>> op.animate()
```

This will create the following animation

Tip: There is currently no option to save the animation within `galpy`, but you could use screen capture software (for example, QuickTime's [Screen Recording](#) feature) to record your screen while the animation is running and save it as a video.

`animate` has options to specify the width and height of the resulting animation, and it can also animate up to three projections of an orbit at the same time. For example, we can look at the orbit in both (x,y) and (R,z) at the same time with

```
>>> op.animate(d1=['x', 'R'], d2=['y', 'z'], width=800)
```

which gives

If you want to embed the animation in a webpage, you can obtain the necessary HTML using the `_repr_html_()` function of the `IPython.core.display.HTML` object returned by `animate`. By default, the HTML includes the entire orbit's data, but `animate` also has an option to store the orbit in a separate JSON file that will then be loaded by the output HTML code.

1.5.5 Orbit characterization

The properties of the orbit can also be found using `galpy`. For example, we can calculate the peri- and apocenter radii of an orbit, its eccentricity, and the maximal height above the plane of the orbit

```
>>> o.rap(), o.rperi(), o.e(), o.zmax()
# (1.2581455175173673, 0.97981663263371377, 0.12436710999105324, 0.11388132751079502)
```

These four quantities can also be computed using analytical means (exact or approximations depending on the potential) by specifying `analytic=True`

```
>>> o.rap(analytic=True), o.rperi(analytic=True), o.e(analytic=True), o.
↪zmax(analytic=True)
# (1.2581448917376636, 0.97981640959995842, 0.12436697719989584, 0.11390708640305315)
```

We can also calculate the energy of the orbit, either in the potential that the orbit was integrated in, or in another potential:

```
>>> o.E(), o.E(pot=mp)
# (0.6150000000000001, -0.67390625000000015)
```

where `mp` is the Miyamoto-Nagai potential of [Introduction: Rotation curves](#).

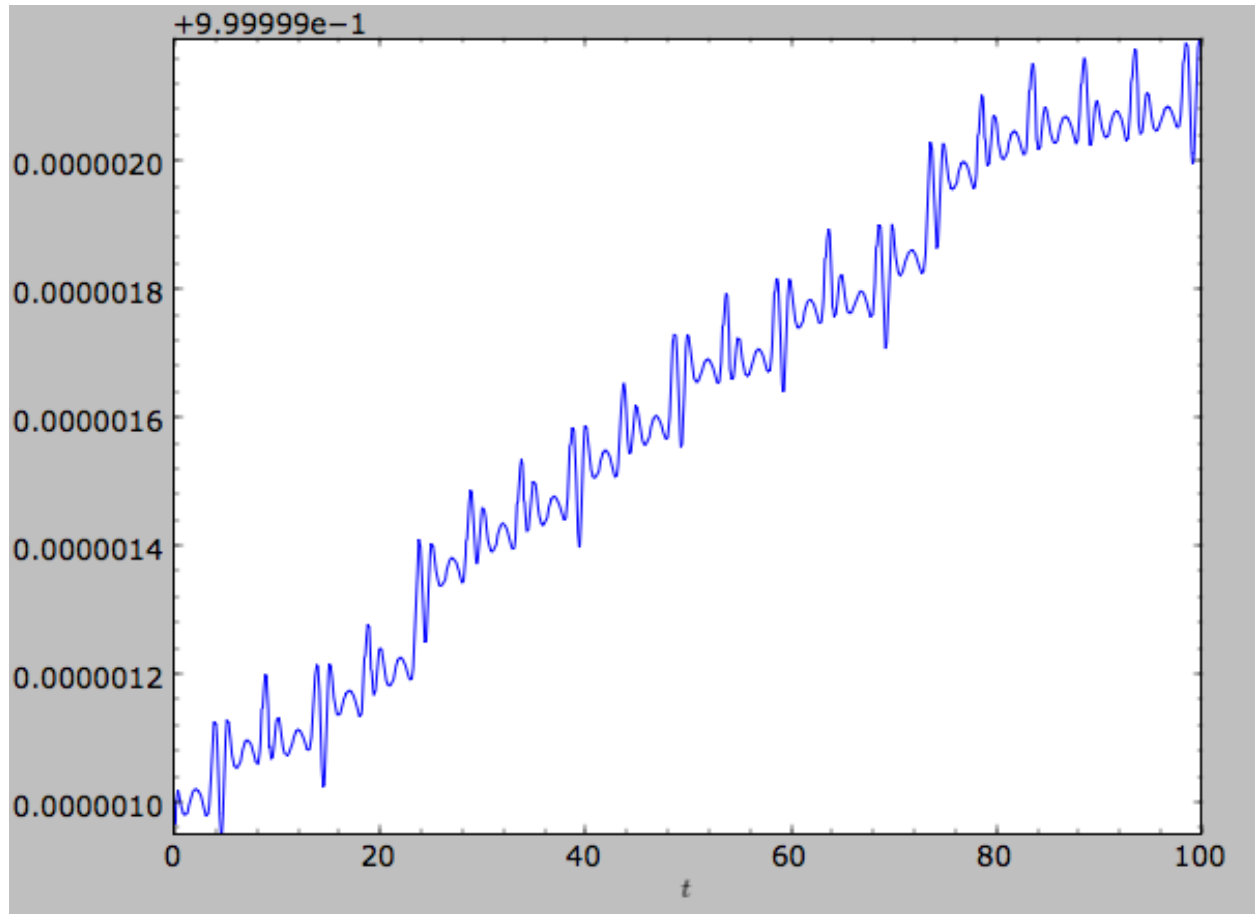
For the Orbit `op` that was initialized above with a distance scale `ro=` and a velocity scale `vo=`, these outputs are all in physical units

```
>>> op.rap(), op.rperi(), op.e(), op.zmax()
# (10.065158988860341, 7.8385312810643057, 0.12436696983841462, 0.91105035688072711) #kpc
>>> op.E(), op.E(pot=mp)
# (29766.000000000004, -32617.062500000007) # (km/s)^2
```

We can also show the energy as a function of time (to check energy conservation)

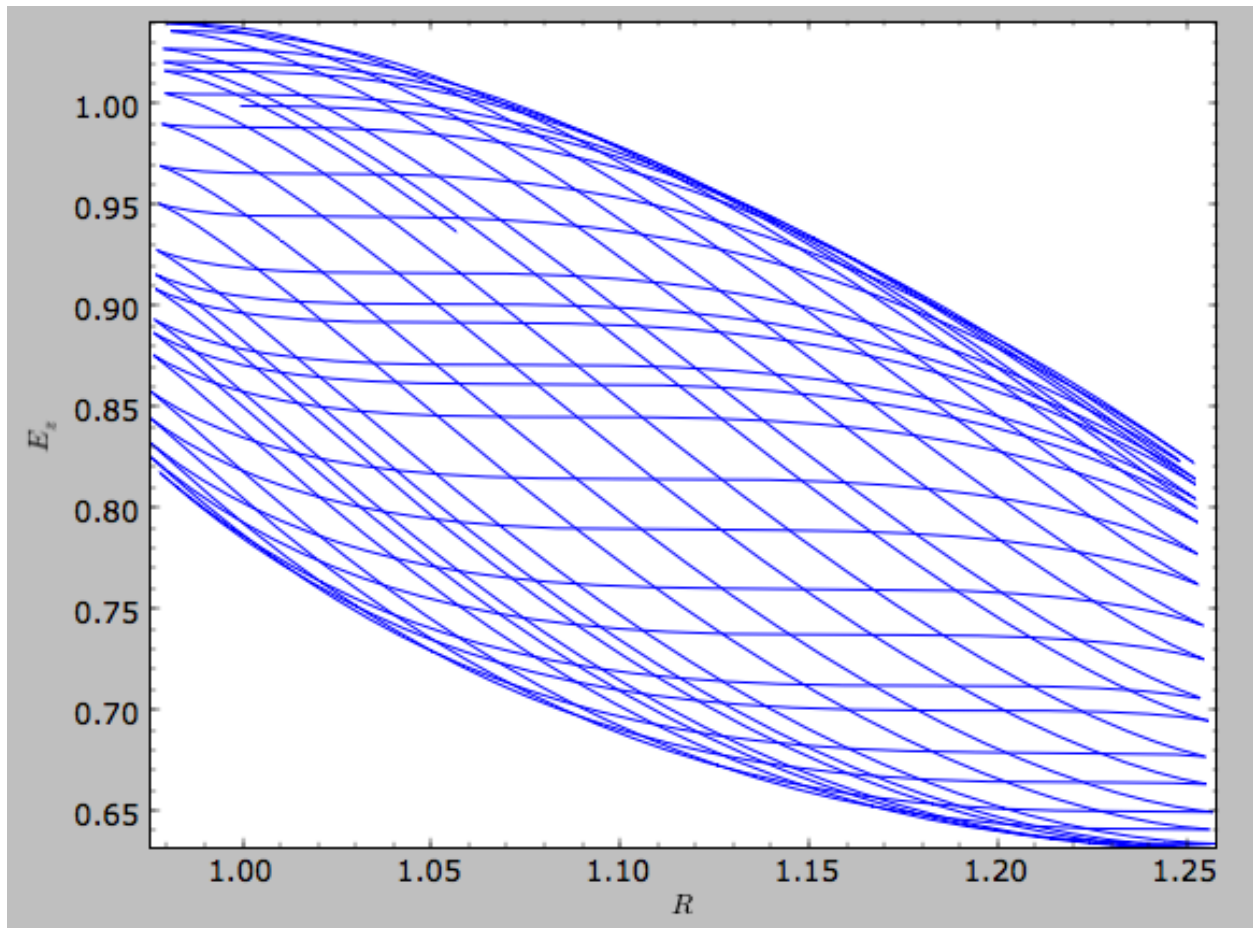
```
>>> o.plotE(normed=True)
```

gives



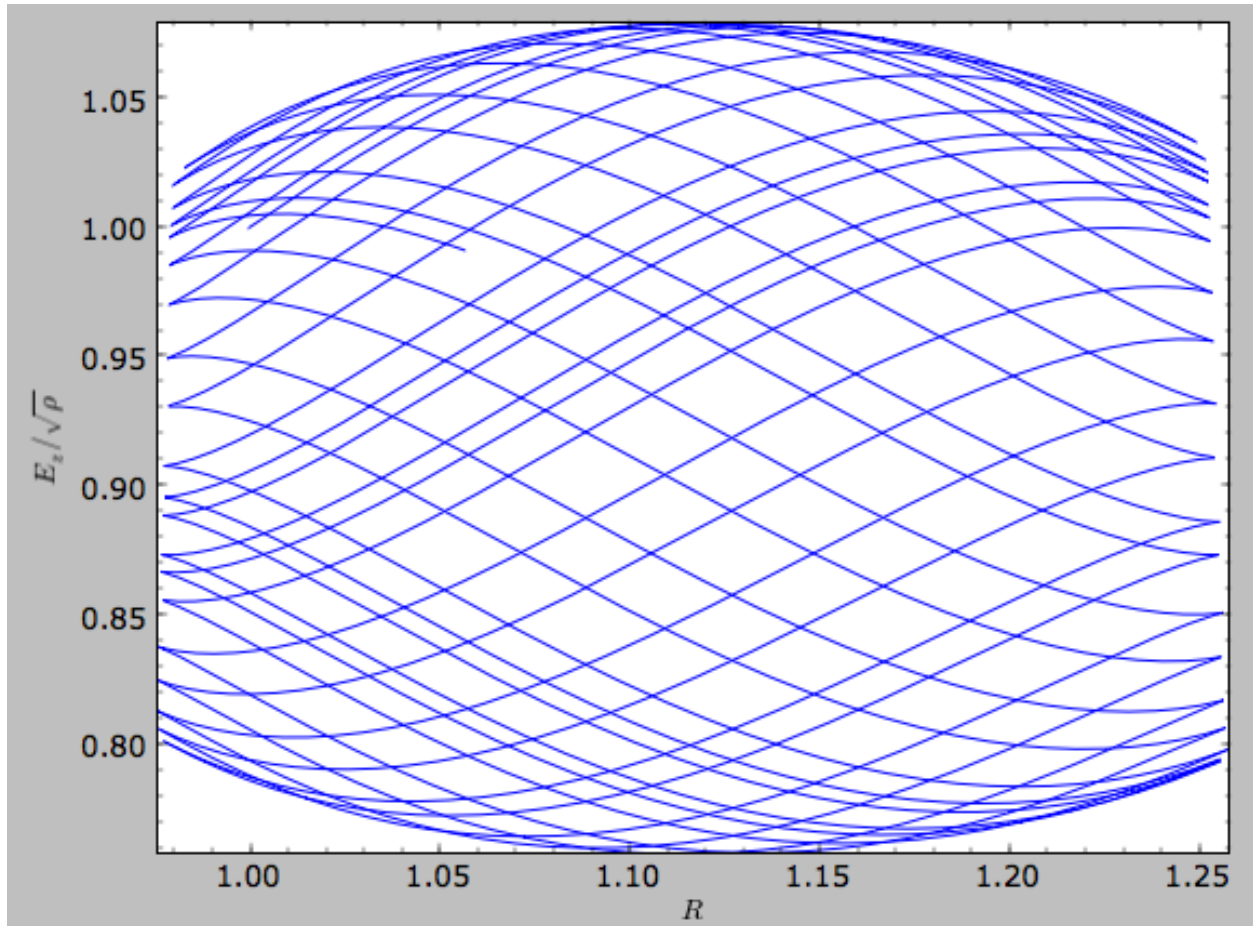
We can specify another quantity to plot the energy against by specifying `d1=`. We can also show the vertical energy, for example, as a function of R

```
>>> o.plotEz(d1='R', normed=True)
```



Often, a better approximation to an integral of the motion is given by $E_z/\sqrt{\text{density}[R]}$. We refer to this quantity as E_zJ_z and we can plot its behavior

```
>>> o.plotEzJz(d1='R', normed=True)
```



1.5.6 Fast orbit characterization

It is also possible to use galpy for the fast estimation of orbit parameters as demonstrated in Mackereth & Bovy (2018, in prep.) via the Staeckel approximation (originally used by Binney (2012) for the approximation of actions in axisymmetric potentials), without performing any orbit integration. The method uses the geometry of the orbit tori to estimate the orbit parameters. After initialising an `Orbit` instance, the method is applied by specifying `analytic=True` and selecting `type='staeckel'`.

```
>>> o.e(analytic=True, type='staeckel')
```

if running the above without integrating the orbit, the potential should also be specified in the usual way

```
>>> o.e(analytic=True, type='staeckel', pot=mp)
```

This interface automatically estimates the necessary delta parameter based on the initial condition of the `Orbit` object.

While this is useful and fast for individual `Orbit` objects, it is likely that users will want to rapidly evaluate the orbit parameters of large numbers of objects. It is possible to perform the orbital parameter estimation above through the `actionAngle` interface. To do this, we need arrays of the phase-space points `R`, `vR`, `vT`, `z`, `vz`, and `phi` for the objects. The orbit parameters are then calculated by first specifying an `actionAngleStaeckel` instance (this requires a single `delta` focal-length parameter, see [the documentation of the `actionAngleStaeckel` class](#)), then using the `EccZmaxRperiRap` method with the data points:

```
>>> aAS = actionAngleStaeckel(pot=mp, delta=0.4)
>>> e, Zmax, rperi, rap = aAS.EccZmaxRperiRap(R, vR, vT, z, vz, phi)
```

Alternatively, you can specify an array for delta when calling `aAS.EccZmaxRperiRap`, for example by first estimating good delta parameters as follows:

```
>>> from galpy.actionAngle import estimateDeltaStaeckel
>>> delta = estimateDeltaStaeckel(mp, R, z, no_median=True)
```

where `no_median=True` specifies that the function return the delta parameter at each given point rather than the median of the calculated deltas (which is the default option). Then one can compute the eccentricity etc. using individual delta values as:

```
>>> e, Zmax, rperi, rap = aAS.EccZmaxRperiRap(R, vR, vT, z, vz, phi, delta=delta)
```

The `EccZmaxRperiRap` method also exists for the `actionAngleIsochrone`, `actionAngleSpherical`, and `actionAngleAdiabatic` modules.

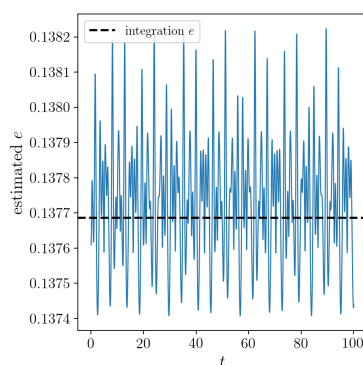
We can test the speed of this method in iPython by finding the parameters at 100000 steps along an orbit in MWPotential2014, like this

```
>>> o = Orbit(vxvv=[1.,0.1,1.1,0.,0.1,0.])
>>> ts = numpy.linspace(0,100,100000)
>>> o.integrate(ts,MWPotential2014)
>>> aAS = actionAngleStaeckel(pot=MWPotential2014,delta=0.3)
>>> R, vR, vT, z, vz, phi = o.getOrbit().T
>>> delta = estimateDeltaStaeckel(MWPotential2014, R, z, no_median=True)
>>> %timeit -n 10 es, zms, rps, ras = aAS.EccZmaxRperiRap(R,vR,vT,z,vz,phi,
↳ delta=delta)
#10 loops, best of 3: 899 ms per loop
```

you can see that in this potential, each phase space point is calculated in roughly 9 μ s. further speed-ups can be gained by using the `actionAngleStaeckelGrid` module, which first calculates the parameters using a grid-based interpolation

```
>>> from galpy.actionAngle import actionAngleStaeckelGrid
>>> aASG = actionAngleStaeckelGrid(pot=mp,delta=0.4,nE=51,npsi=51,nLz=61,c=True,
↳ interpecc=True)
>>> %timeit -n 10 es, zms, rps, ras = aASG.EccZmaxRperiRap(R,vR,vT,z,vz,phi)
#10 loops, best of 3: 587 ms per loop
```

where `interpecc=True` is required to perform the interpolation of the orbit parameter grid. Looking at how the eccentricity estimation varies along the orbit, and comparing to the calculation using the orbit integration, we see that the estimation good job



1.5.7 Accessing the raw orbit

The value of `R`, `vR`, `vT`, `z`, `vz`, `x`, `vx`, `y`, `vy`, `phi`, and `vphi` at any time can be obtained by calling the corresponding function with as argument the time (the same holds for other coordinates `ra`, `dec`, `pmra`, `pmdec`, `vra`, `vdec`, `ll`, `bb`, `pmll`, `pmll`, `vll`, `vbb`, `vlos`, `dist`, `helioX`, `helioY`, `helioZ`, `U`, `V`, and `W`). If no time is given the initial condition is returned, and if a time is requested at which the orbit was not saved spline interpolation is used to return the value. Examples include

```
>>> o.R(1.)
# 1.1545076874679474
>>> o.phi(99.)
# 88.105603035901169
>>> o.ra(2.,obs=[8.,0.,0.],ro=8.)
# array([ 285.76403985])
>>> o.helioX(5.)
# array([ 1.24888927])
>>> o.pmll(10.,obs=[8.,0.,0.,0.,245.,0.],ro=8.,vo=230.)
# array([-6.45263888])
```

For the Orbit `op` that was initialized above with a distance scale `ro=` and a velocity scale `vo=`, the first of these would be

```
>>> op.R(1.)
# 9.2360614837829225 #kpc
```

which we can also access in natural coordinates as

```
>>> op.R(1.,use_physical=False)
# 1.1545076854728653
```

We can also specify a different distance or velocity scale on the fly, e.g.,

```
>>> op.R(1.,ro=4.) #different velocity scale would be vo=
# 4.6180307418914612
```

We can also initialize an `Orbit` instance using the phase-space position of another `Orbit` instance evaluated at time `t`. For example,

```
>>> newOrbit= o(10.)
```

will initialize a new `Orbit` instance with as initial condition the phase-space position of orbit `o` at `time=10..`

The whole orbit can also be obtained using the function `getOrbit`

```
>>> o.getOrbit()
```

which returns a matrix of phase-space points with dimensions `[ntimes,ndim]`.

1.5.8 Fast orbit integration

The standard orbit integration is done purely in python using standard `scipy` integrators. When fast orbit integration is needed for batch integration of a large number of orbits, a set of orbit integration routines are written in C that can be accessed for most potentials, as long as they have C implementations, which can be checked by using the attribute `hasC`

```
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,amp=1.,normalize=1.)
>>> mp.hasC
# True
```

Fast C integrators can be accessed through the `method=` keyword of the `orbit.integrate` method. Currently available integrators are

- `rk4_c`
- `rk6_c`
- `dopr54_c`

which are Runge-Kutta and Dormand-Prince methods. There are also a number of symplectic integrators available

- `leapfrog_c`
- `symplec4_c`
- `symplec6_c`

The higher order symplectic integrators are described in [Yoshida \(1993\)](#).

For most applications I recommend `symplec4_c`, which is speedy and reliable. For example, compare

```
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.,0.1])
>>> timeit(o.integrate(ts,mp,method='leapfrog'))
# 1.34 s ± 41.8 ms per loop (mean ± std. dev. of 7 runs, 1 loop each)
>>> timeit(o.integrate(ts,mp,method='leapfrog_c'))
# galpyWarning: Using C implementation to integrate orbits
# 91 ms ± 2.42 ms per loop (mean ± std. dev. of 7 runs, 10 loops each)
>>> timeit(o.integrate(ts,mp,method='symplec4_c'))
# galpyWarning: Using C implementation to integrate orbits
# 9.67 ms ± 48.3 µs per loop (mean ± std. dev. of 7 runs, 100 loops each)
```

As this example shows, galpy will issue a warning that C is being used.

1.5.9 Integration of the phase-space volume

galpy further supports the integration of the phase-space volume through the method `integrate_dxdv`, although this is currently only implemented for two-dimensional orbits (`planarOrbit`). As an example, we can check Liouville's theorem explicitly. We initialize the orbit

```
>>> o= Orbit(vxvv=[1.,0.1,1.1,0.])
```

and then integrate small deviations in each of the four phase-space directions

```
>>> ts= numpy.linspace(0.,28.,1001) #~1 Gyr at the Solar circle
>>> o.integrate_dxdv([1.,0.,0.,0.],ts,mp,method='dopr54_c',rectIn=True,rectOut=True)
>>> dx= o.getOrbit_dxdv()[-1,:] # evolution of dxdv[0] along the orbit
>>> o.integrate_dxdv([0.,1.,0.,0.],ts,mp,method='dopr54_c',rectIn=True,rectOut=True)
>>> dy= o.getOrbit_dxdv()[-1,:]
>>> o.integrate_dxdv([0.,0.,1.,0.],ts,mp,method='dopr54_c',rectIn=True,rectOut=True)
>>> dvx= o.getOrbit_dxdv()[-1,:]
>>> o.integrate_dxdv([0.,0.,0.,1.],ts,mp,method='dopr54_c',rectIn=True,rectOut=True)
>>> dvx= o.getOrbit_dxdv()[-1,:]
```

We can then compute the determinant of the Jacobian of the mapping defined by the orbit integration from time zero to the final time


```
>>> tjac= numpy.linalg.det(numpy.array([dx,dy,dvx,dvy]))
```

This determinant should be equal to one

```
>>> print(tjac)
# 0.999999991189
>>> numpy.fabs(tjac-1.) < 10.**-8.
# True
```

The calls to `integrate_dxdv` above set the keywords `rectIn=` and `rectOut=` to `True`, as the default input and output uses phase-space volumes defined as $(dR, dvR, dvT, d\phi)$ in cylindrical coordinates. When `rectIn` or `rectOut` is set, the in- or output is in rectangular coordinates (x, y, vx, vy) in two dimensions).

Implementing the phase-space integration for three-dimensional `FullOrbit` instances is straightforward and is part of the longer term development plan for `galpy`. Let the main developer know if you would like this functionality, or better yet, implement it yourself in a fork of the code and send a pull request!

1.5.10 Example: The eccentricity distribution of the Milky Way's thick disk

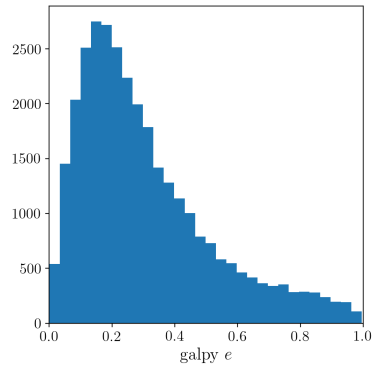
A straightforward application of `galpy`'s orbit initialization and integration capabilities is to derive the eccentricity distribution of a set of thick disk stars. We start by downloading the sample of SDSS SEGUE (2009AJ...137.4377Y) thick disk stars compiled by Dierickx et al. (2010arXiv1009.1616D) from CDS at [this link](#). Downloading the table and the ReadMe will allow you to read in the data using `astropy.io.ascii` like so

```
>>> from astropy.io import ascii
>>> dierickx = ascii.read('table2.dat', readme='ReadMe')
>>> vxvv = numpy.dstack([dierickx['RAdeg'], dierickx['DEdeg'], dierickx['Dist']/1e3,
↳ dierickx['pmRA'], dierickx['pmDE'], dierickx['HRV']])[0]
```

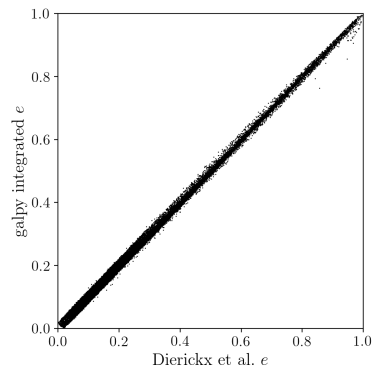
After reading in the data (RA, Dec, distance, pmRA, pmDec, vlos; see above) as a vector `vxvv` with dimensions $[6, n_{\text{data}}]$ we (a) define the potential in which we want to integrate the orbits, and (b) integrate each orbit and save its eccentricity as calculated analytically following the *Staeckel approximation method* and by orbit integration (running this for all 30,000-ish stars will take about half an hour)

```
>>> from galpy.actionAngle import UnboundError
>>> ts= np.linspace(0.,20.,10000)
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> e_ana = numpy.zeros(len(vxvv))
>>> e_int = numpy.zeros(len(vxvv))
>>> for i in range(len(vxvv)):
...     #calculate analytic e estimate, catch any 'unbound' orbits
...     try:
...         orbit = Orbit(vxvv[i], radec=True, vo=220., ro=8.)
...         e_ana[i] = orbit.e(analytic=True, pot=lp, c=True)
...     except UnboundError:
...         #parameters cannot be estimated analytically
...         e_ana[i] = np.nan
...     #integrate the orbit and return the numerical e value
...     orbit.integrate(ts, lp)
...     e_int[i] = orbit.e(analytic=False)
```

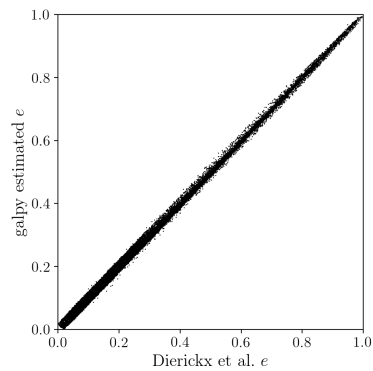
We then find the following eccentricity distribution (from the numerical eccentricities)



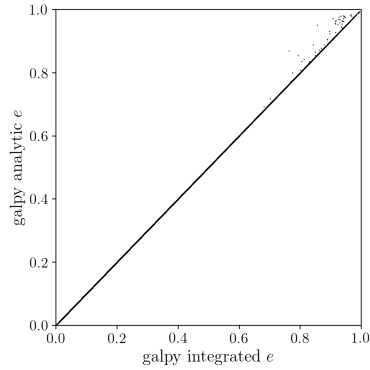
The eccentricity calculated by integration in galpy compare well with those calculated by Dierickx et al., except for a few objects



and the analytical estimates are equally as good:



In comparing the analytic and integrated eccentricity estimates - one can see that in this case the estimation is almost exact, due to the spherical symmetry of the chosen potential:



A script that calculates and plots everything can be downloaded [here](#). To generate the plots just run:

```
python dierickx_eccentricities.py ../path/to/folder
```

specifying the location you want to put the plots and data.

Alternatively - one can transform the observed coordinates into spherical coordinates and perform the estimations in one batch using the `actionAngle` interface, which takes considerably less time:

```
>>> from galpy import actionAngle
>>> deltas = actionAngle.estimateDeltaStaeckel(lp, Rphiz[:,0], Rphiz[:,2], no_
↳ median=True)
>>> aAS = actionAngleStaeckel(pot=lp, delta=0.)
>>> par = aAS.EccZmaxRperiRap(Rphiz[:,0], vRvTvz[:,0], vRvTvz[:,1], Rphiz[:,2],
↳ vRvTvz[:,2], Rphiz[:,1], delta=deltas)
```

The above code calculates the parameters in roughly 100ms on a single core.

1.5.11 NEW in v1.4 Example: The orbit of the Large Magellanic Cloud in the presence of dynamical friction

As a further example of what you can do with galpy, we investigate the Large Magellanic Cloud's (LMC) past and future orbit. Because the LMC is a massive satellite of the Milky Way, its orbit is affected by dynamical friction, a frictional force of gravitational origin that occurs when a massive object travels through a sea of low-mass objects (halo stars and dark matter in this case). First we import all the necessary packages:

```
>>> from astropy import units
>>> from galpy.potential import MWPotential2014, ChandrasekharDynamicalFrictionForce
>>> from galpy.orbit import Orbit
```

(also do `%pylab inline` if running this in a jupyter notebook or turn on the `pylab` option in `ipython` for plotting). We can load the current phase-space coordinates for the LMC using the `Orbit.from_name` function described [above](#):

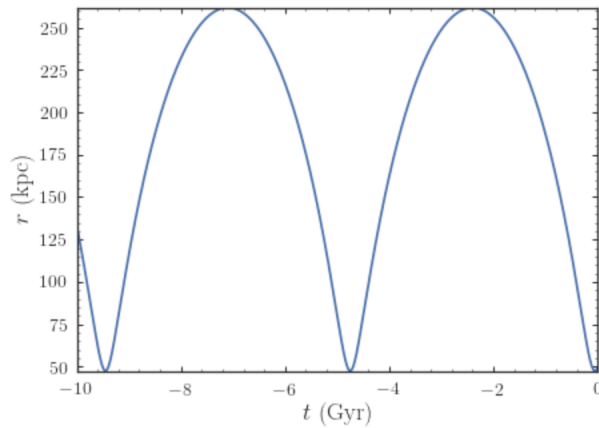
```
>>> o= Orbit.from_name('LMC')
```

We will use `MWPotential2014` as our Milky-Way potential model. Because the LMC is in fact unbound in `MWPotential2014`, we increase the halo mass by 50% to make it bound (this corresponds to a Milky-Way halo mass of $\approx 1.2 \times 10^{12} M_{\odot}$, a not unreasonable value). We can hack this together as

```
>>> MWPotential2014[2]._amp*= 1.5
```

(Note that this is *not* a generally recommended route for changing the mass of an object, since it relies on editing a private attribute). Let us now integrate the orbit backwards in time for 10 Gyr and plot it:

```
>>> ts= numpy.linspace(0.,-10.,1001)*units.Gyr
>>> o.integrate(ts,MWPotential2014)
>>> o.plot(d1='t',d2='r')
```



We see that the LMC is indeed bound, with an apocenter just over 250 kpc. Now let's add dynamical friction for the LMC, assuming that its mass is $5 \times 10^{10} M_{\odot}$. We setup the dynamical-friction object:

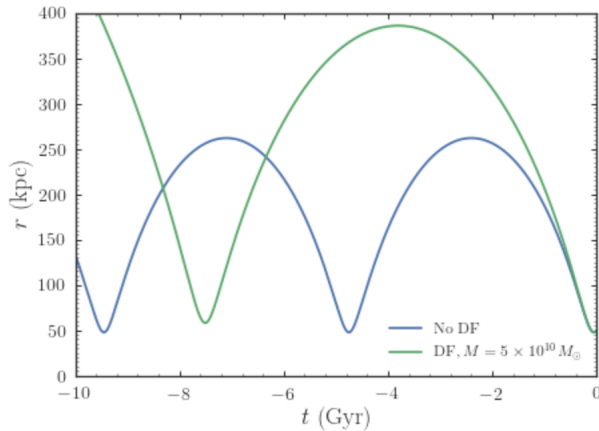
```
>>> cdf= ChandrasekharDynamicalFrictionForce(GMs=5.*10.**10.*units.Msun, rhm=5.*units.
↪ kpc,
                                           dens=MWPotential2014)
```

Dynamical friction depends on the velocity distribution of the halo, which is assumed to be an isotropic Gaussian distribution with a radially-dependent velocity dispersion. If the velocity dispersion is not given (like in the example above), it is computed from the spherical Jeans equation. We have set the half-mass radius to 5 kpc for definiteness. We now make a copy of the orbit instance above and integrate it in the potential that includes dynamical friction:

```
>>> odf= o()
>>> odf.integrate(ts,[MWPotential2014,cdf])
```

(Note that specifying the forces as the list `[MWPotential2014,cdf]` works even though `MWPotential2014` is itself a list of potentials, because we can use nested lists of potentials or forces wherever a list is allowed in `galpy`). Overlaying the orbits, we can see the difference in the evolution:

```
>>> o.plot(d1='t',d2='r',label=r'\mathrm{No\ DF}')
>>> odf.plot(d1='t',d2='r',overplot=True,label=r'\mathrm{DF}, M=5\times10^{10}\,M_
↪ \odot')
>>> ylim(0.,400.)
>>> legend()
```



We see that dynamical friction removes energy from the LMC's orbit, such that its past apocenter is now around 400 kpc rather than 250 kpc! The period of the orbit is therefore also much longer. Clearly, dynamical friction has a big impact on the orbit of the LMC.

Recent observations have suggested that the LMC may be even more massive than what we have assumed so far, with masses over $10^{11} M_{\odot}$ seeming in good agreement with various observations. Let's see how a mass of $10^{11} M_{\odot}$ changes the past orbit of the LMC. We can change the mass of the LMC used in the dynamical-friction calculation as

```
>>> cdf.GMs= 10.**11.*units.Msun
```

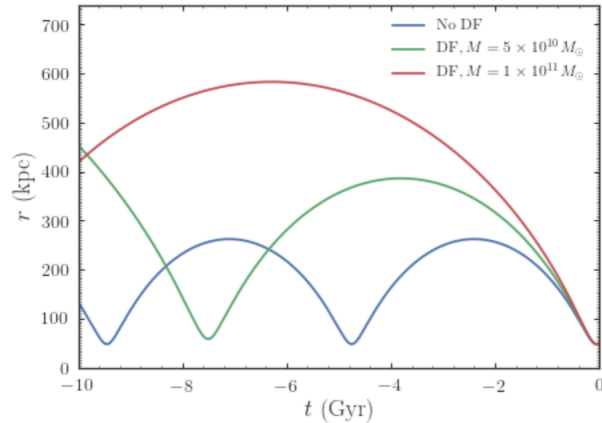
This way of changing the mass is preferred over re-initializing the `ChandrasekharDynamicalFrictionForce` object, because it avoids having to solve the Jeans equation again to obtain the velocity dispersion. Then we integrate the orbit and overplot it on the previous results:

```
>>> odf2= o()
>>> odf2.integrate(ts, [MWPotential2014, cdf])
```

and

```
>>> o.plot(d1='t', d2='r', label=r'$\mathrm{No\ DF}$')
>>> odf.plot(d1='t', d2='r', overplot=True, label=r'$\mathrm{DF}$', M=5\times10^{10}\,M_{\odot}$')
>>> odf2.plot(d1='t', d2='r', overplot=True, label=r'$\mathrm{DF}$', M=1\times10^{11}\,M_{\odot}$')
>>> ylim(0., 740.)
>>> legend()
```

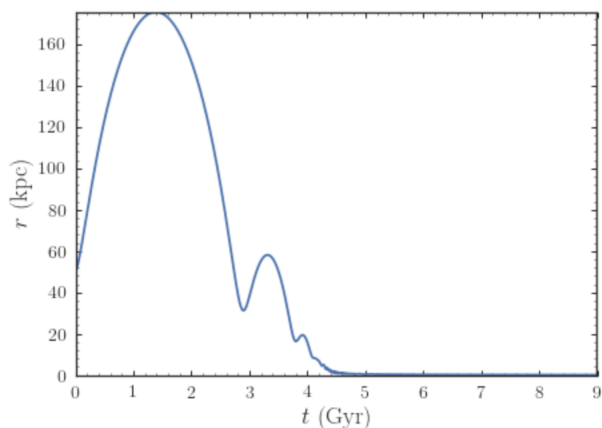
which gives



Now the apocenter increases to about 600 kpc and the LMC doesn't perform a full orbit over the last 10 Gyr.

Finally, let's see what will happen in the future if the LMC is as massive as $10^{11} M_{\odot}$. We simply flip the sign of the integration times to get the future trajectory:

```
>>> odf2.integrate(-ts[-ts < 9*units.Gyr], [MWPotential2014, cdf])
>>> odf2.plot(d1='t', d2='r')
```



Because of the large effect of dynamical friction, the LMC will merge with the Milky-Way in about 4 Gyr after a few more pericenter passages. Note that we have not taken any mass-loss into account. Because mass-loss would lead to a smaller dynamical-friction force, this would somewhat increase the merging timescale, but dynamical friction will inevitably lead to the merger of the LMC with the Milky Way.

Warning: When using dynamical friction, if the radius gets very small, the integration sometimes becomes very erroneous, which can lead to a big, unphysical kick (even though we turn off friction at very small radii); this is the reason why we have limited the future integration to 9 Gyr in the example above. When using dynamical friction, inspect the full orbit to make sure to catch whether a merger has happened.

1.6 Two-dimensional disk distribution functions

galpy contains various disk distribution functions, both in two and three dimensions. This section introduces the two-dimensional distribution functions, useful for studying the dynamics of stars that stay relatively close to the mid-plane

of a galaxy. The vertical motions of these stars may be approximated as being entirely decoupled from the motion in the plane.

1.6.1 Types of disk distribution functions

galpy contains the following distribution functions for razor-thin disks: `galpy.df.dehnendf`, `galpy.df.shudf`, and `galpy.df.schwarzschilddf`. These are the distribution functions of Dehnen (1999AJ...118.1201D), Shu (1969ApJ...158..505S), and Schwarzschild (the Shu DF in the epicycle approximation, see Binney & Tremaine 2008). Everything shown below for `dehnendf` can also be done for `shudf` and `schwarzschilddf`. The Schwarzschild DF is primarily included as an educational tool; it is *not* a true steady-state DF, because it uses the approximate energy from the epicycle approximation rather than the true energy, and is fully superseded by the Shu DF, which *is* a good steady-state DF.

These disk distribution functions are functions of the energy and the angular momentum alone. They can be evaluated for orbits, or for a given energy and angular momentum. At this point, only power-law rotation curves are supported. A `dehnendf` instance is initialized as follows

```
>>> from galpy.df import dehnendf
>>> dfc= dehnendf(beta=0.)
```

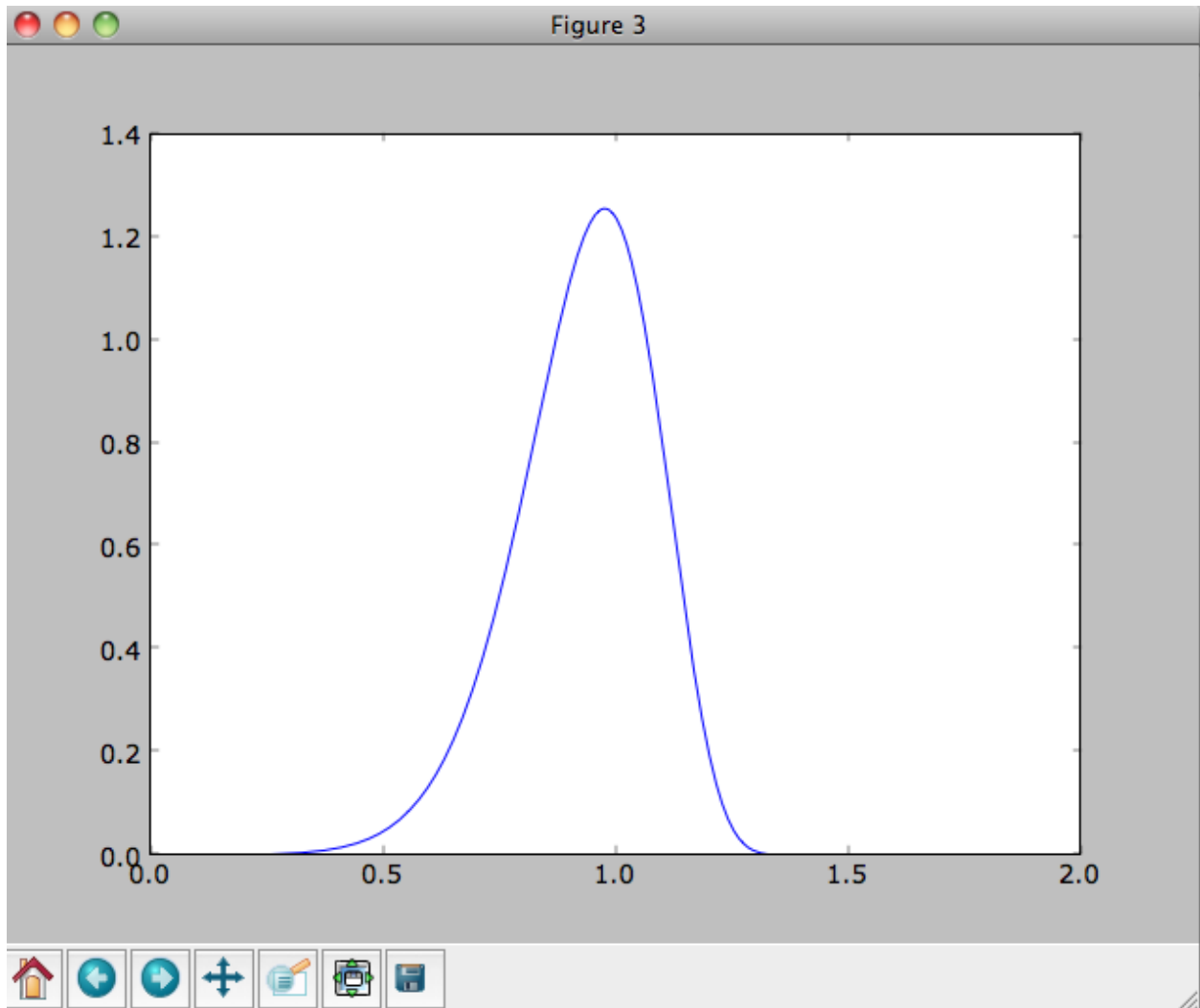
This initializes a `dehnendf` instance based on an exponential surface-mass profile with scale-length 1/3 and an exponential radial-velocity-dispersion profile with scale-length 1 and a value of 0.2 at $R=1$. Different parameters for these profiles can be provided as an initialization keyword. For example,

```
>>> dfc= dehnendf(beta=0.,profileParams=(1./4.,1.,0.2))
```

initializes the distribution function with a radial scale length of 1/4 instead.

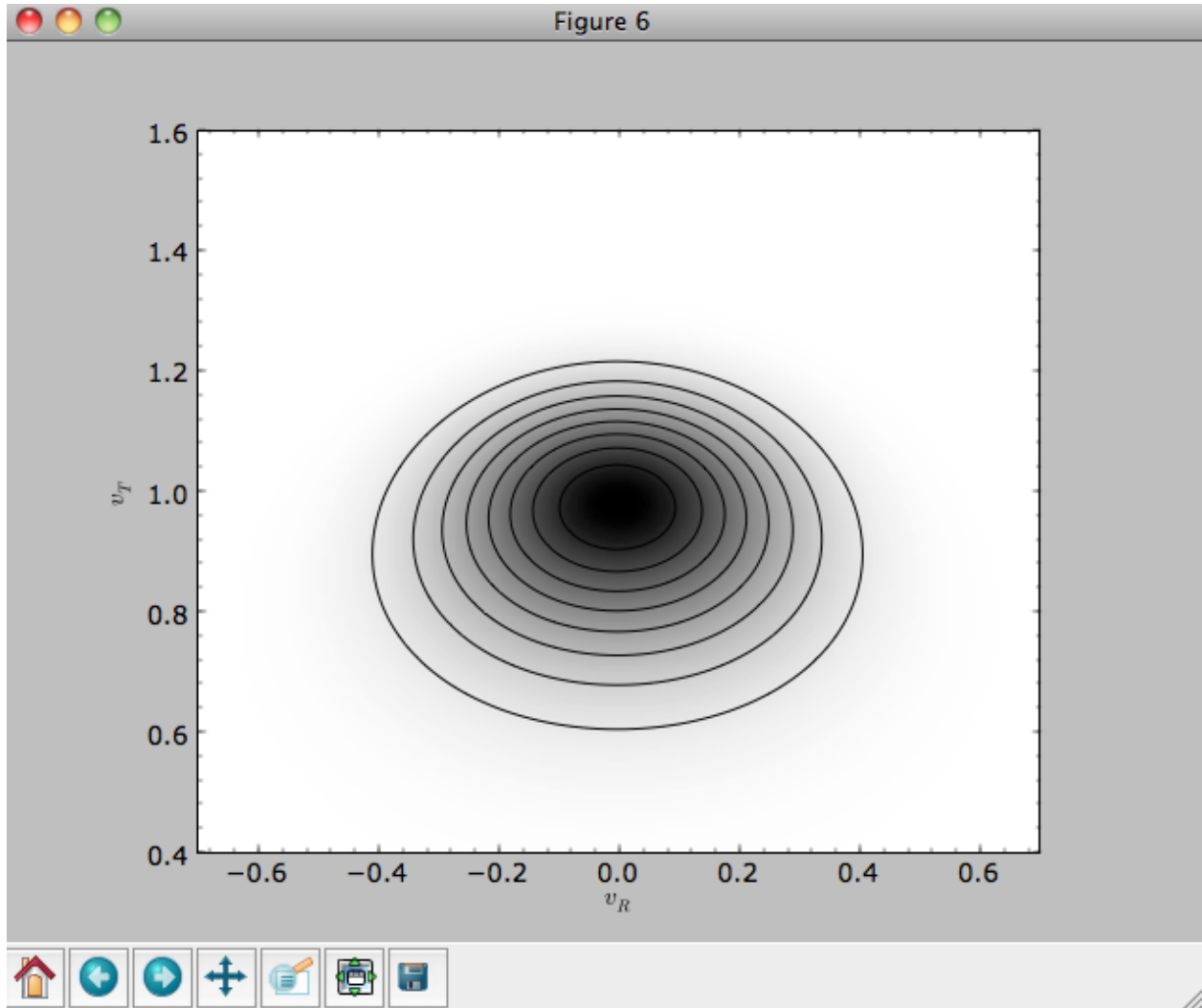
We can show that these distribution functions have an asymmetric drift built-in by evaluating the DF at $R=1$. We first create a set of orbit-instances and then evaluate the DF at them

```
>>> from galpy.orbit import Orbit
>>> os= [Orbit([1.,0.,1.+0.9+1.8/1000*ii]) for ii in range(1001)]
>>> dfro= [dfc(o) for o in os]
>>> plot([1.+0.9+1.8/1000*ii for ii in range(1001)],dfro)
```



Or we can plot the two-dimensional density at $R=1$.

```
>>> dfro= [[dfc(Orbit([1.,-0.7+1.4/200*jj,1.-0.6+1.2/200*ii])) for jj in_
↳range(201)]for ii in range(201)]
>>> dfro= numpy.array(dfro)
>>> from galpy.util.bovy_plot import bovy_dens2d
>>> bovy_dens2d(dfro,origin='lower',cmap='gist_yarg',contours=True,xrange=[-0.7,0.7],
↳yrange=[0.4,1.6],xlabel=r'$v_R$',ylabel=r'$v_T$')
```

1.6.2 Evaluating moments of the DF

galpy can evaluate various moments of the disk distribution functions. For example, we can calculate the mean velocities (for the DF with a scale length of 1/3 above)

```
>>> dfc.meanvT(1.)
# 0.91715276979447324
>>> dfc.meanvR(1.)
# 0.0
```

and the velocity dispersions

```
>>> numpy.sqrt(dfc.sigmaR2(1.))
# 0.19321086259083936
>>> numpy.sqrt(dfc.sigmaT2(1.))
# 0.15084122011271159
```

and their ratio

```
>>> dfc.sigmaR2(1.)/dfc.sigmaT2(1.)  
# 1.6406766813028968
```

In the limit of zero velocity dispersion (the epicycle approximation) this ratio should be equal to 2, which we can check as follows

```
>>> dfccold= dehnendf(beta=0.,profileParams=(1./3.,1.,0.02))  
>>> dfccold.sigmaR2(1.)/dfccold.sigmaT2(1.)  
# 1.9947493895454664
```

We can also calculate higher order moments

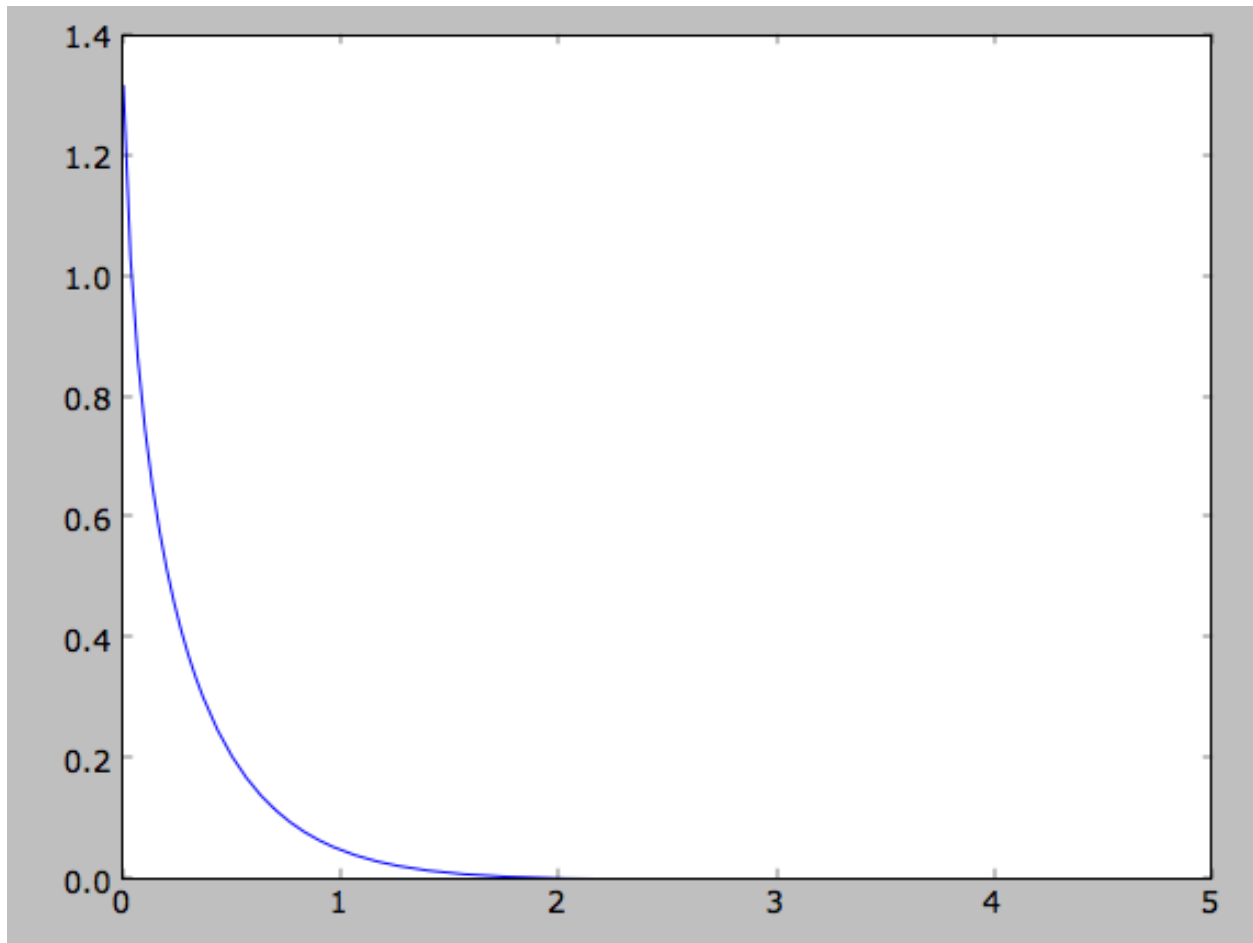
```
>>> dfc.skewvT(1.)  
# -0.48617143862047763  
>>> dfc.kurtosisvT(1.)  
# 0.13338978593181494  
>>> dfc.kurtosisvR(1.)  
# -0.12159407676394096
```

We already saw above that the velocity dispersion at $R=1$ is not exactly equal to the input velocity dispersion (0.19321086259083936 vs. 0.2). Similarly, the whole surface-density and velocity-dispersion profiles are not quite equal to the exponential input profiles. We can calculate the resulting surface-mass density profile using `surfacemass`, `sigmaR2`, and `sigma2surfacemass`. The latter calculates the product of the velocity dispersion squared and the surface-mass density. E.g.,

```
>>> dfc.surfacemass(1.)  
# 0.050820867101511534
```

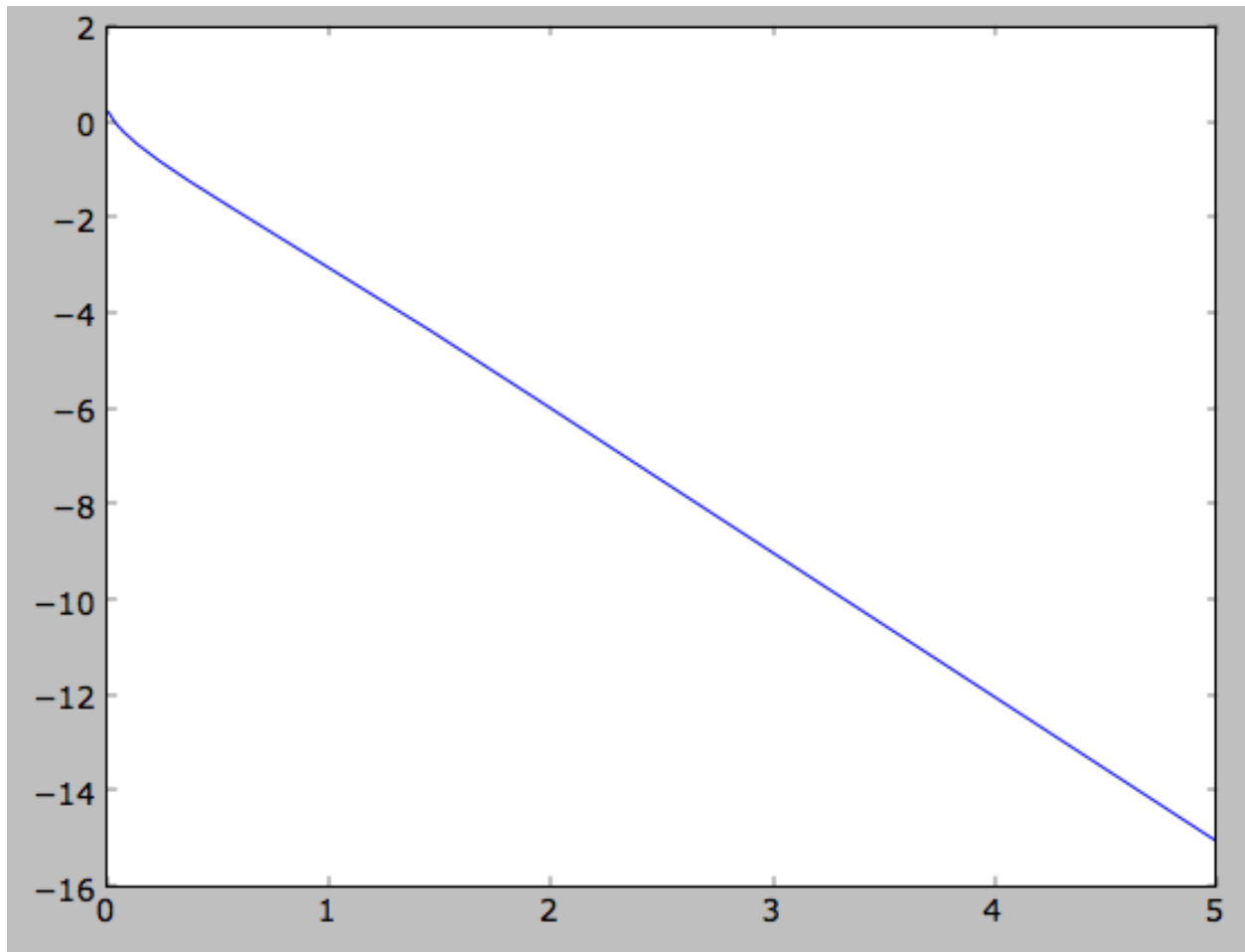
We can plot the surface-mass density as follows

```
>>> Rs= numpy.linspace(0.01,5.,151)  
>>> out= [dfc.surfacemass(r) for r in Rs]  
>>> plot(Rs, out)
```



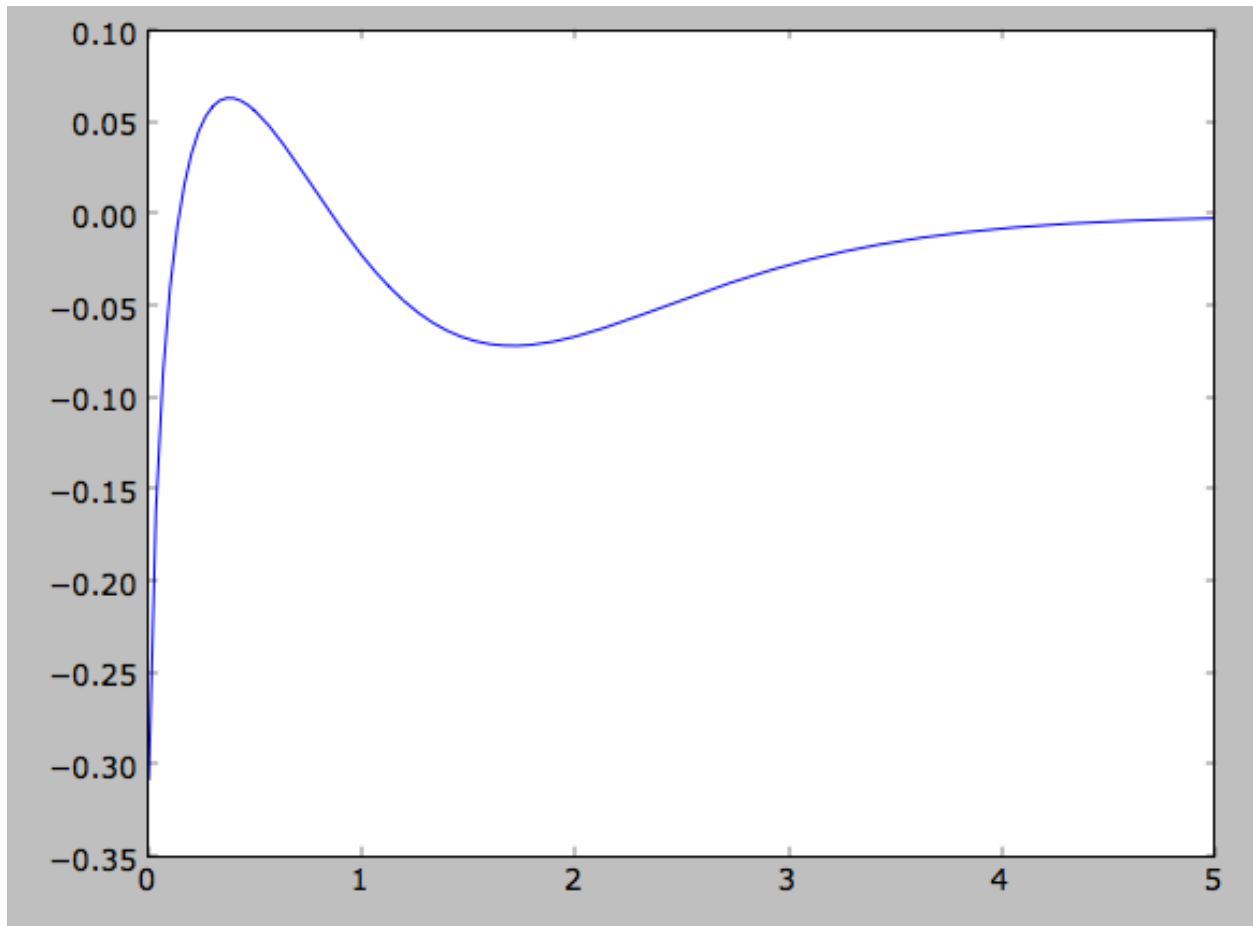
or

```
>>> plot(Rs, numpy.log(out))
```



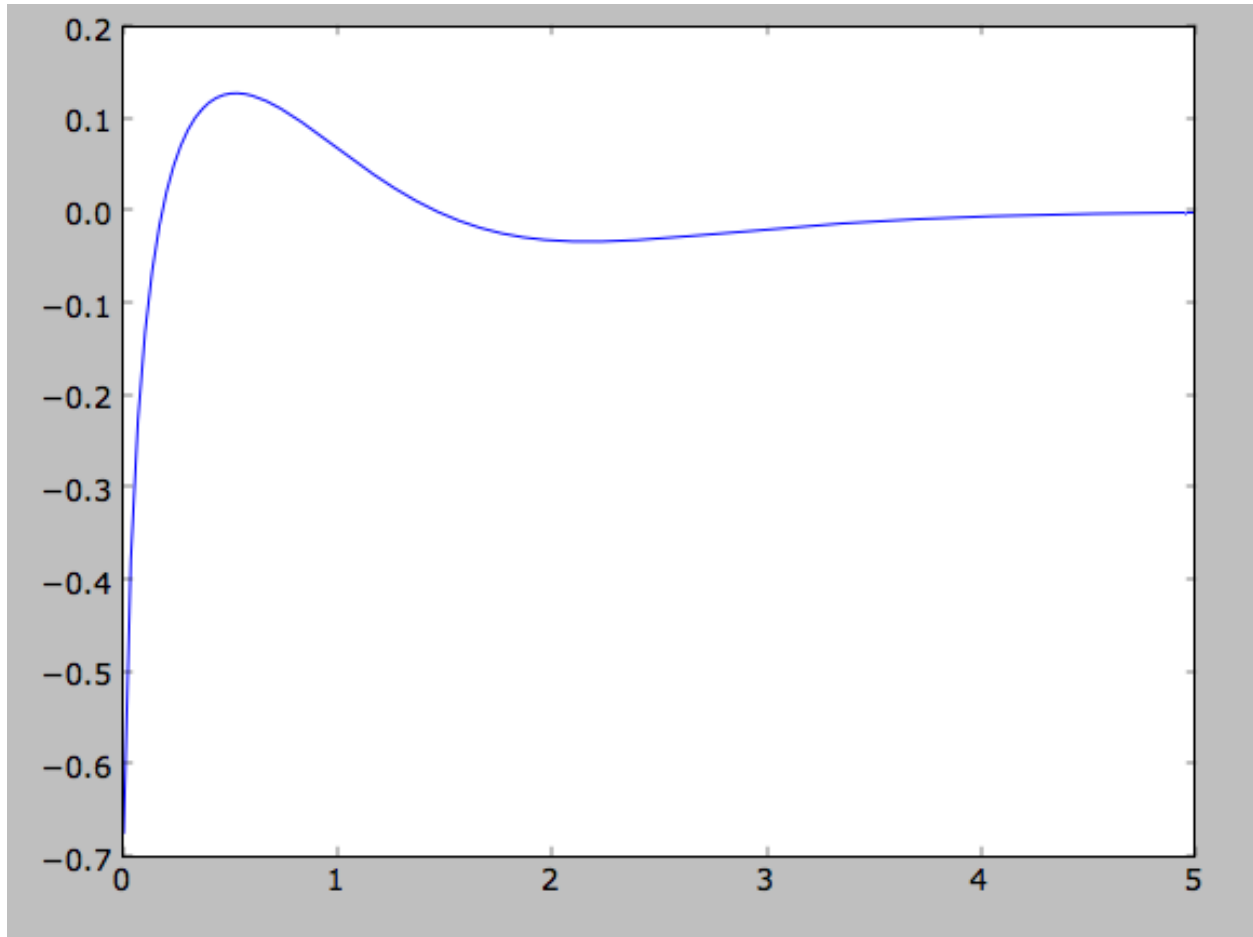
which shows the exponential behavior expected for an exponential disk. We can compare this to the input surface-mass density

```
>>> input_out= [dfc.targetSurfacemass(r) for r in Rs]
>>> plot(Rs, numpy.log(input_out)-numpy.log(out))
```



which shows that there are significant differences between the desired surface-mass density and the actual surface-mass density. We can do the same for the velocity-dispersion profile

```
>>> out= [dfc.sigmaR2(r) for r in Rs]
>>> input_out= [dfc.targetSigma2(r) for r in Rs]
>>> plot(Rs, numpy.log(input_out)-numpy.log(out))
```



That the input surface-density and velocity-dispersion profiles are not the same as the output profiles, means that estimates of DF properties based on these profiles will not be quite correct. Obviously this is the case for the surface-density and velocity-dispersion profiles themselves, which have to be explicitly calculated by integration over the DF rather than by evaluating the input profiles. This also means that estimates of the asymmetric drift based on the input profiles will be wrong. We can calculate the asymmetric drift at $R=1$ using the asymmetric drift equation derived from the Jeans equation (eq. 4.228 in Binney & Tremaine 2008), using the input surface-density and velocity dispersion profiles

```
>>> dfc.asymmetricdrift(1.)
# 0.0900000000000000024
```

which should be equal to the circular velocity minus the mean rotational velocity

```
>>> 1.-dfc.meanvT(1.)
# 0.082847230205526756
```

These are not the same in part because of the difference between the input and output surface-density and velocity-dispersion profiles (and because the `asymmetricdrift` method assumes that the ratio of the velocity dispersions squared is two using the epicycle approximation; see above).

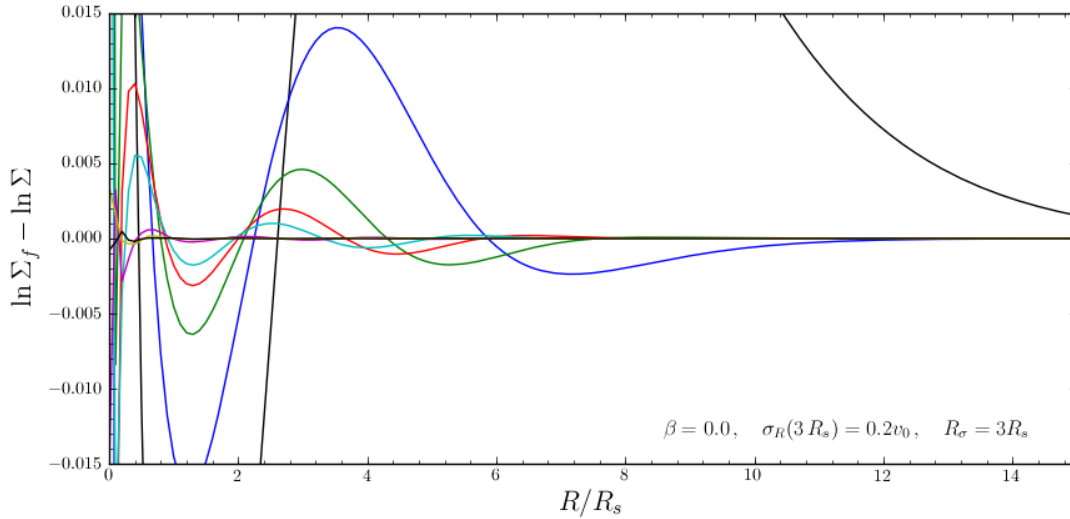
1.6.3 Using corrected disk distribution functions

As shown above, for a given surface-mass density and velocity dispersion profile, the two-dimensional disk distribution functions only do a poor job of reproducing the desired profiles. We can correct this by calculating a

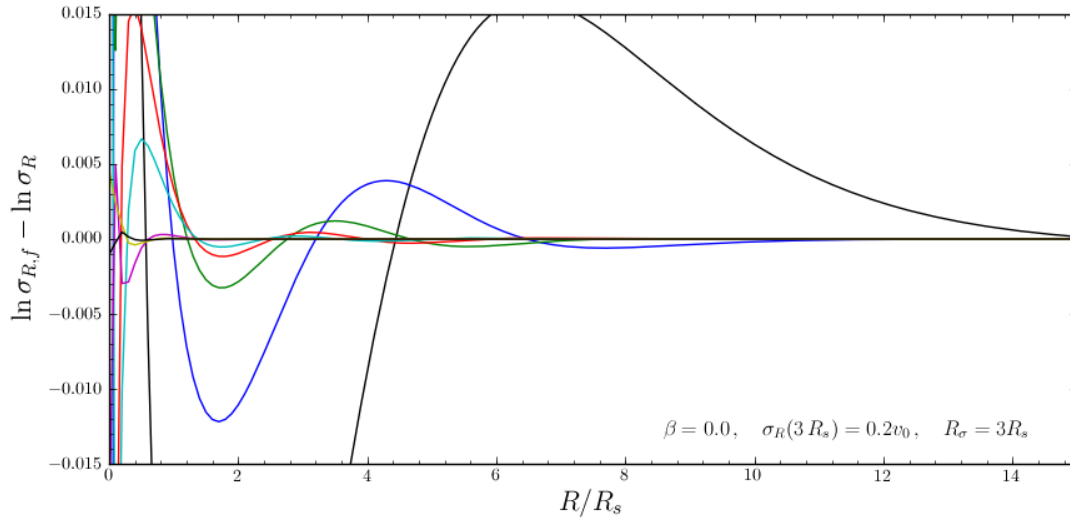
set of *corrections* to the input profiles such that the output profiles more closely resemble the desired profiles (see 1999AJ...118.1201D). galpy supports the calculation of these corrections, and comes with some pre-calculated corrections (these can be found [here](#)). For example, the following initializes a `dehne2df` with corrections up to 20th order (the default)

```
>>> dfc= dehne2df(beta=0.,correct=True)
```

The following figure shows the difference between the actual surface-mass density profile and the desired profile for 1, 2, 3, 4, 5, 10, 15, and 20 iterations



and the same for the velocity-dispersion profile



galpy will automatically save any new corrections that you calculate.

All of the methods for an uncorrected disk DF can be used for the corrected DFs as well. For example, the velocity dispersion is now

```
>>> numpy.sqrt(dfc.sigmaR2(1.))
# 0.19999985069451526
```

and the mean rotation velocity is

```
>>> dfc.meanvT(1.)
# 0.90355161181498711
```

and (correct) asymmetric drift

```
>>> 1.-dfc.meanvT(1.)
# 0.09644838818501289
```

That this still does not agree with the simple `dfc.asymmetricdrift` estimate is because of the latter's using the epicycle approximation for the ratio of the velocity dispersions.

1.6.4 Oort constants and functions

galpy also contains methods to calculate the Oort functions for two-dimensional disk distribution functions. These are known as the *Oort constants* when measured in the solar neighborhood. They are combinations of the mean velocities and derivatives thereof. galpy calculates these by direct integration over the DF and derivatives of the DF. Thus, we can calculate

```
>>> dfc= dehnendf(beta=0.)
>>> dfc.oortA(1.)
# 0.43190780889218749
>>> dfc.oortB(1.)
# -0.48524496090228575
```

The *K* and *C* Oort constants are zero for axisymmetric DFs

```
>>> dfc.oortC(1.)
# 0.0
>>> dfc.oortK(1.)
# 0.0
```

In the epicycle approximation, for a flat rotation curve $A = -B = 0.5$. The explicit calculates of *A* and *B* for warm DFs quantify how good (or bad) this approximation is

```
>>> dfc.oortA(1.)+dfc.oortB(1.)
# -0.053337152010098254
```

For the cold DF from above the approximation is much better

```
>>> dfccold= dehnendf(beta=0.,profileParams=(1./3.,1.,0.02))
>>> dfccold.oortA(1.), dfccold.oortB(1.)
# (0.49917556666144003, -0.49992824742490816)
```

1.6.5 Sampling data from the DF

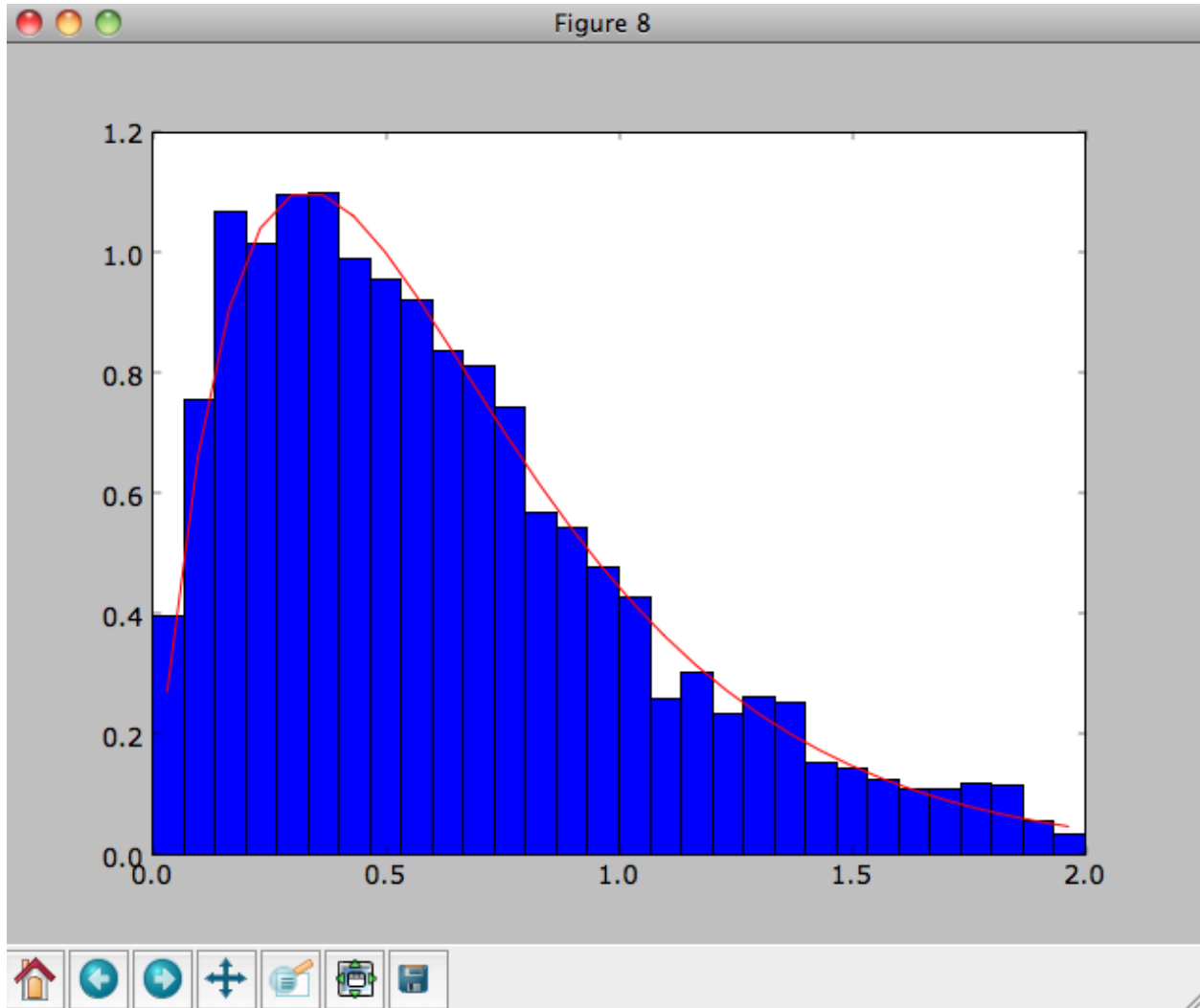
We can sample from the disk distribution functions using `sample`. `sample` can return either an energy-angular-momentum pair, or a full orbit initialization. We can sample 4000 orbits for example as (could take two minutes)

```
>>> o= dfc.sample(n=4000,returnOrbit=True,nphi=1)
```

We can then plot the histogram of the sampled radii and compare it to the input surface-mass density profile


```
>>> Rs= [e.R() for e in o]
>>> hists, bins, edges= hist(Rs,range=[0,2],normed=True,bins=30)
>>> xs= numpy.array([(bins[ii+1]+bins[ii])/2. for ii in range(len(bins)-1)])
>>> plot(xs, xs*exp(-xs*3.)*9., 'r-')
```

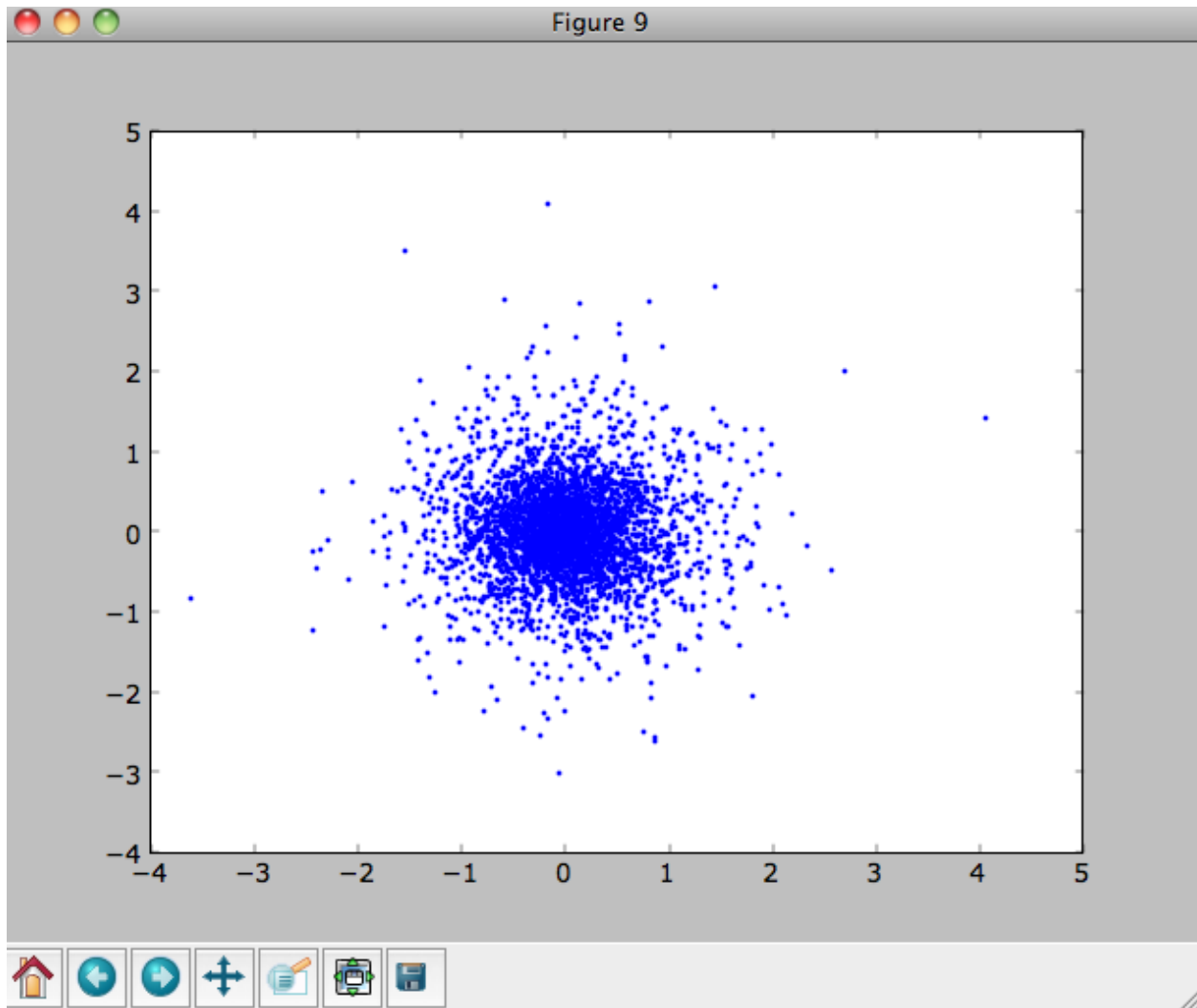
E.g.,



We can also plot the spatial distribution of the sampled disk

```
>>> xs= [e.x() for e in o]
>>> ys= [e.y() for e in o]
>>> figure()
>>> plot(xs,ys, ', ')
```

E.g.,

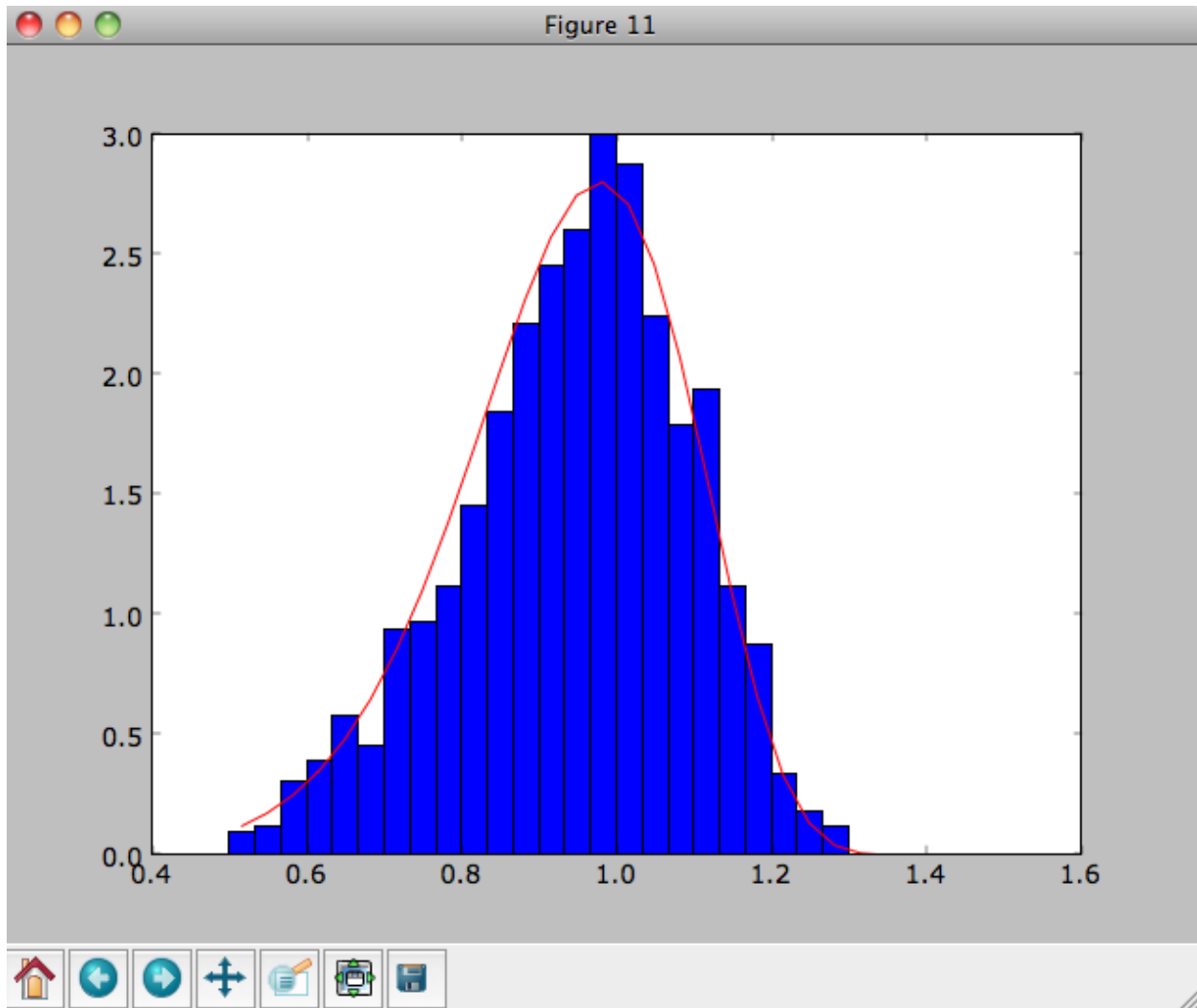


We can also sample points in a specific radial range (might take a few minutes)

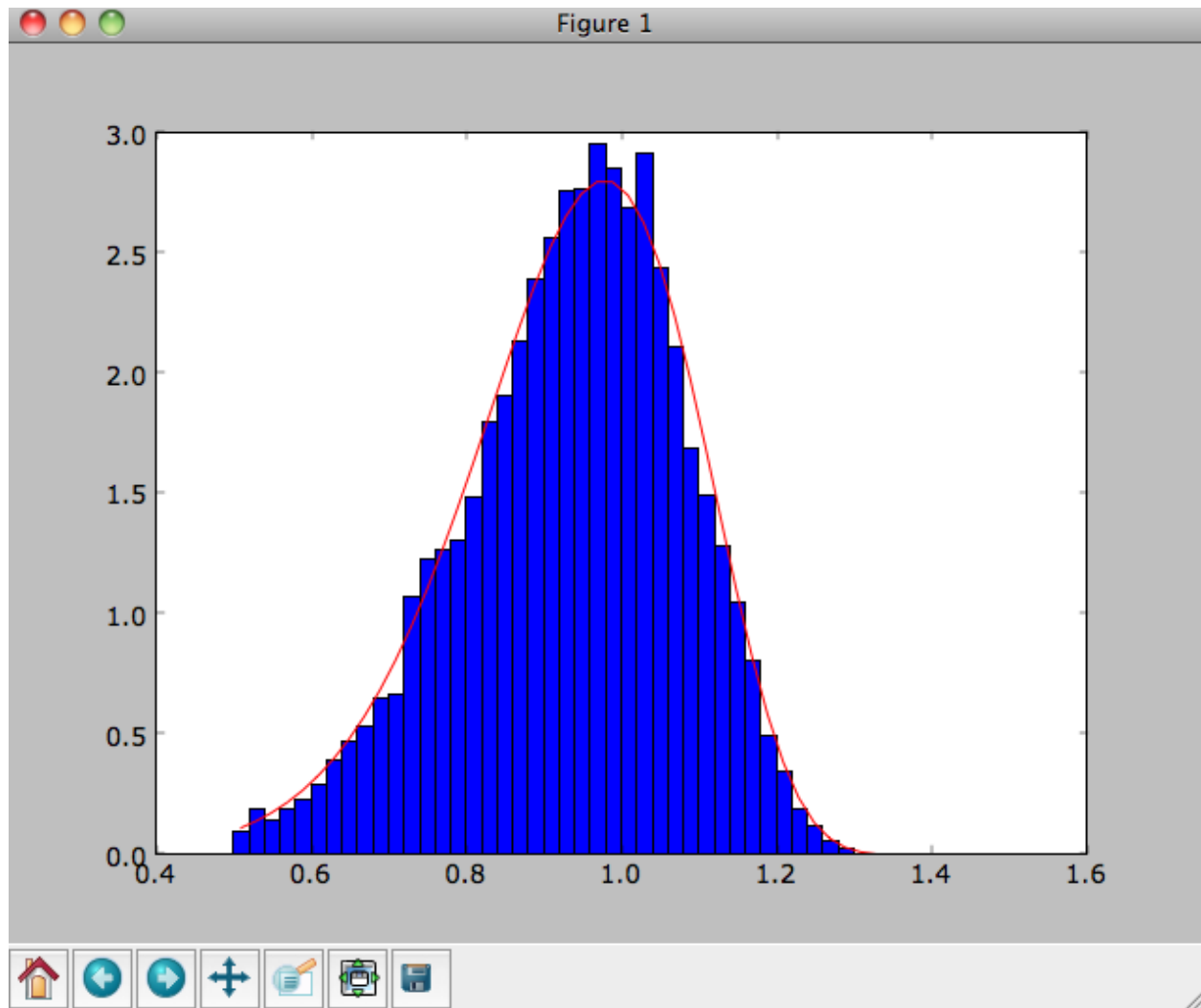
```
>>> o= dfc.sample(n=1000,returnOrbit=True,nphi=1,rrange=[0.8,1.2])
```

and we can plot the distribution of tangential velocities

```
>>> vTs= [e.vxvv[2] for e in o]
>>> hists, bins, edges= hist(vTs,range=[.5,1.5],normed=True,bins=30)
>>> xs= numpy.array([(bins[ii+1]+bins[ii])/2. for ii in range(len(bins)-1)])
>>> dfro= [dfc.Orbit([1.,0.,x])/9./numpy.exp(-3.) for x in xs]
>>> plot(xs,dfro,'r-')
```

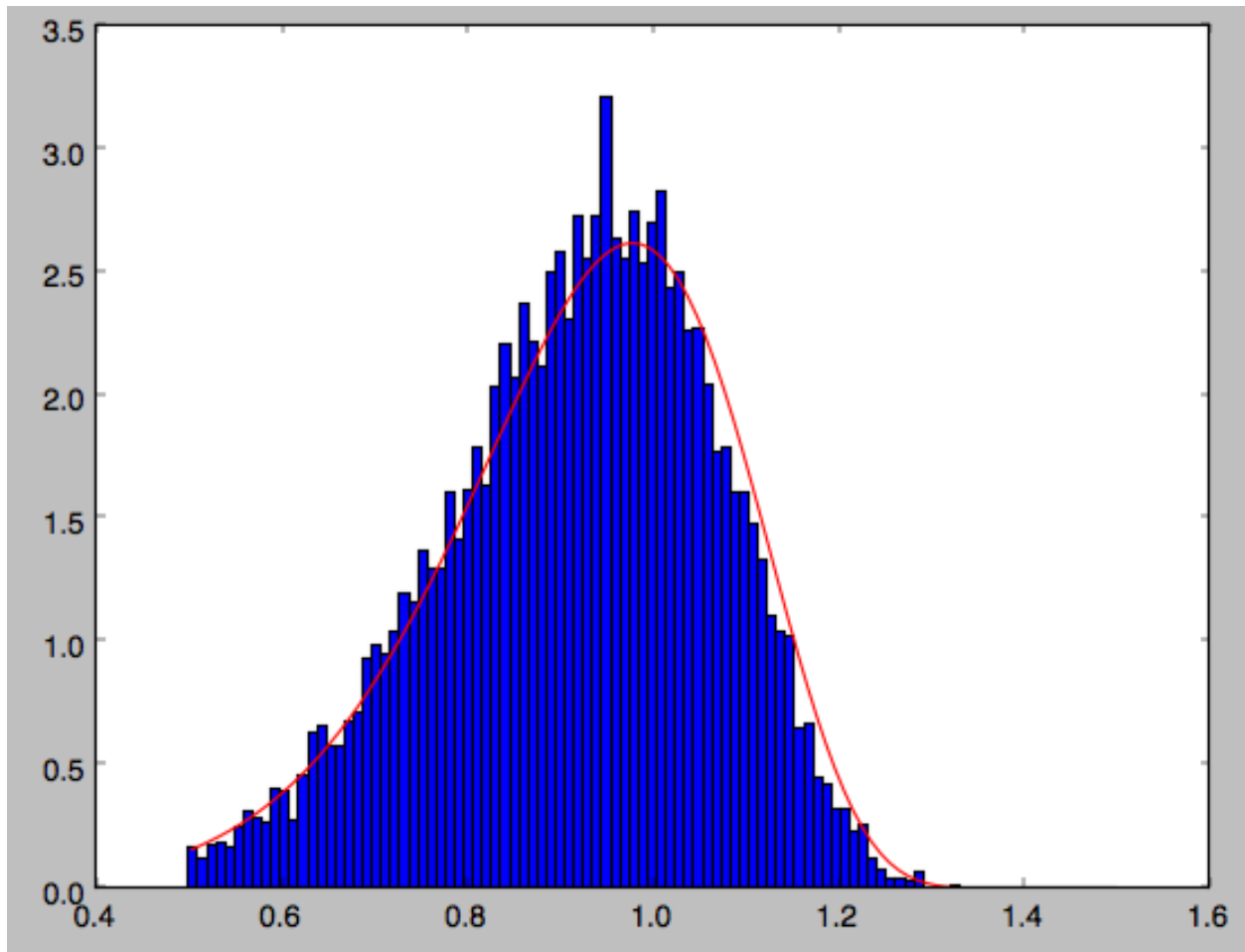


The agreement between the sampled distribution and the theoretical curve is not as good because the sampled distribution has a finite radial range. If we sample 10,000 points in `rrange=[0.95, 1.05]` the agreement is better (this takes a long time):



We can also directly sample velocities at a given radius rather than in a range of radii. Doing this for a correct DF gives

```
>>> dfc= dehndf(beta=0.,correct=True)
>>> vrvt= dfc.sampleVRVT(1.,n=10000)
>>> hists, bins, edges= hist(vrvt[:,1],range=[.5,1.5],normed=True,bins=101)
>>> xs= numpy.array([(bins[ii+1]+bins[ii])/2. for ii in range(len(bins)-1)])
>>> dfro= [dfc.Orbit([1.,0.,x]) for x in xs]
>>> plot(xs,dfro/numpy.sum(dfro)/(xs[1]-xs[0]),'r-')
```



galpy further has support for sampling along a given line of sight in the disk, which is useful for interpreting surveys consisting of a finite number of pointings. For example, we can sampled distances along a given line of sight

```
>>> ds= dfc.sampledSurfacemassLOS(30./180.*numpy.pi,n=10000)
```

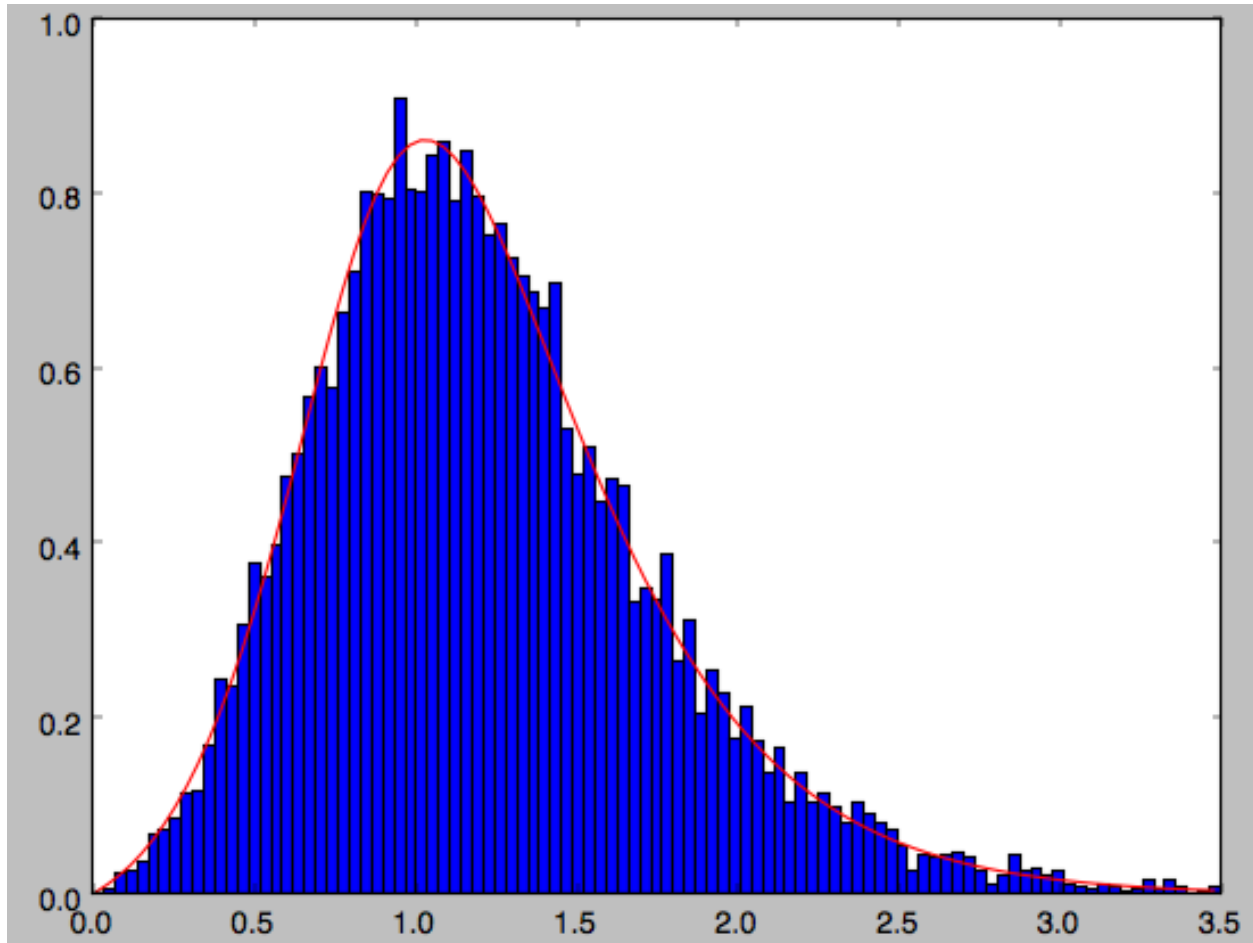
which is very fast. We can histogram these

```
>>> hists, bins, edges= hist(ds,range=[0.,3.5],normed=True,bins=101)
```

and compare it to the predicted distribution, which we can calculate as

```
>>> xs= numpy.array([(bins[ii+1]+bins[ii])/2. for ii in range(len(bins)-1)])
>>> fd= numpy.array([dfc.surfacemassLOS(d,30.) for d in xs])
>>> plot(xs,fd/numpy.sum(fd)/(xs[1]-xs[0]),'r-')
```

which shows very good agreement with the sampled distances



galpy can further sample full 4D phase-space coordinates along a given line of sight through `dfc.sampleLOS`.

1.6.6 Non-axisymmetric, time-dependent disk distribution functions

galpy also supports the evaluation of non-axisymmetric, time-dependent two-dimensional DFs. These specific DFs are constructed by assuming an initial axisymmetric steady state, described by a DF of the family discussed above, that is then acted upon by a non-axisymmetric, time-dependent perturbation. The DF at a given time and phase-space position is evaluated by integrating the orbit backwards in time in the non-axisymmetric potential until the time of the initial DF is reached. From Liouville's theorem, which states that phase-space volume is conserved along the orbit, we then know that we can evaluate the non-axisymmetric DF today as the initial DF at the initial point on the orbit. This procedure was first used by [Dehnen \(2000\)](#).

This is implemented in galpy as `galpy.df.evolveddiskdf`. Such a DF is setup by specifying the initial DF, the non-axisymmetric potential, and the time of the initial state. For example, we can look at the effect of an elliptical perturbation to the potential like that described by [Kuijken & Tremaine](#). To do this, we set up an elliptical perturbation to a logarithmic potential that is grown slowly to minimize non-adiabatic effects

```
>>> from galpy.potential import LogarithmicHaloPotential, EllipticalDiskPotential
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> ep= EllipticalDiskPotential(twophio=0.05,phib=0.,p=0.,tform=-150.,tsteady=125.)
```

This perturbation starts to be grown at `tform=-150` over a time period of `tsteady=125` time units. We will consider the effect of this perturbation on a very cold disk (velocity dispersion $\sigma_R = 0.0125 v_c$) and a warm disk ($\sigma_R = 0.15 v_c$). We set up these two initial DFs

```
>>> idfcold= dehnendf(beta=0.,profileParams=(1./3.,1.,0.0125))
>>> idfwarm= dehnendf(beta=0.,profileParams=(1./3.,1.,0.15))
```

and then set up the `evolveddiskdf`

```
>>> from galpy.df import evolveddiskdf
>>> edfcold= evolveddiskdf(idfcold,[lp,ep],to=-150.)
>>> edfwarm= evolveddiskdf(idfwarm,[lp,ep],to=-150.)
```

where we specify that the initial state is at `to=-150`.

We can now use these `evolveddiskdf` instances in much the same way as `diskdf` instances. One difference is that there is much more support for evaluating the DF on a grid (to help speed up the rather slow computations involved). Thus, we can evaluate the mean radial velocity at $R=0.9$, $\phi=22.5$ degree, and $t=0$ by using a grid

```
>>> mvrcold, gridcold= edfcold.meanvR(0.9,phi=22.5,deg=True,t=0.,grid=True,
↳returnGrid=True,gridpoints=51,nsigma=6.)
>>> mvrwarm, gridwarm= edfwarm.meanvR(0.9,phi=22.5,deg=True,t=0.,grid=True,
↳returnGrid=True,gridpoints=51)
>>> print(mvrcold, mvrwarm)
# -0.0358753028951 -0.0294763627935
```

The cold response agrees well with the analytical calculation, which predicts that this is $-0.05/\sqrt{2}$:

```
>>> print(mvrcold+0.05/sqrt(2.))
# -0.000519963835811
```

The warm response is slightly smaller in amplitude

```
>>> print(mvrwarm/mvrcold)
# 0.821633837619
```

although the numerical uncertainty in `mvrwarm` is large, because the grid is not sufficiently fine.

We can then re-use this grid in calculations of other moments of the DF, e.g.,

```
>>> print(edfcold.meanvT(0.9,phi=22.5,deg=True,t=0.,grid=gridcold))
# 0.965058551359
>>> print(edfwarm.meanvT(0.9,phi=22.5,deg=True,t=0.,grid=gridwarm))
# 0.915397094614
```

which returns the mean rotational velocity, and

```
>>> print(edfcold.vertexdev(0.9,phi=22.5,deg=True,t=0.,grid=gridcold))
# 0.0560531474616
>>> print(edfwarm.vertexdev(0.9,phi=22.5,deg=True,t=0.,grid=gridwarm))
# 0.0739164830253
```

which gives the vertex deviation in rad. The reason we have to calculate the grid out to `6nsigma` for the cold response is that the response is much bigger than the velocity dispersion of the population. This velocity dispersion is used to automatically to set the grid edges, but sometimes has to be adjusted to contain the full DF.

`evolveddiskdf` can also calculate the Oort functions, by directly calculating the spatial derivatives of the DF. These can also be calculated on a grid, such that we can do

```
>>> oortacold, gridcold, gridrcold, gridphicold= edfcold.oortA(0.9, phi=22.5, deg=True,
↳ t=0., returnGrids=True, gridpoints=51, derivGridpoints=51, grid=True, derivphiGrid=True,
↳ derivRGrid=True, nsigma=6.)
>>> oortawarm, gridwarm, gridrwarm, gridphiwarm= edfwarm.oortA(0.9, phi=22.5, deg=True,
↳ t=0., returnGrids=True, gridpoints=51, derivGridpoints=51, grid=True, derivphiGrid=True,
↳ derivRGrid=True)
>>> print(oortacold, oortawarm)
# 0.575494559999 0.526389833249
```

It is clear that these are quite different. The cold calculation is again close to the analytical prediction, which says that $A = A_{\text{axi}} + 0.05/(2\sqrt{2})$ where $A_{\text{axi}} = 1/(2 \times 0.9)$ in this case:

```
>>> print(oortacold- (0.5/0.9+0.05/2./sqrt(2.)))
# 0.0022613349141670236
```

These grids can then be re-used for the other Oort functions, for example,

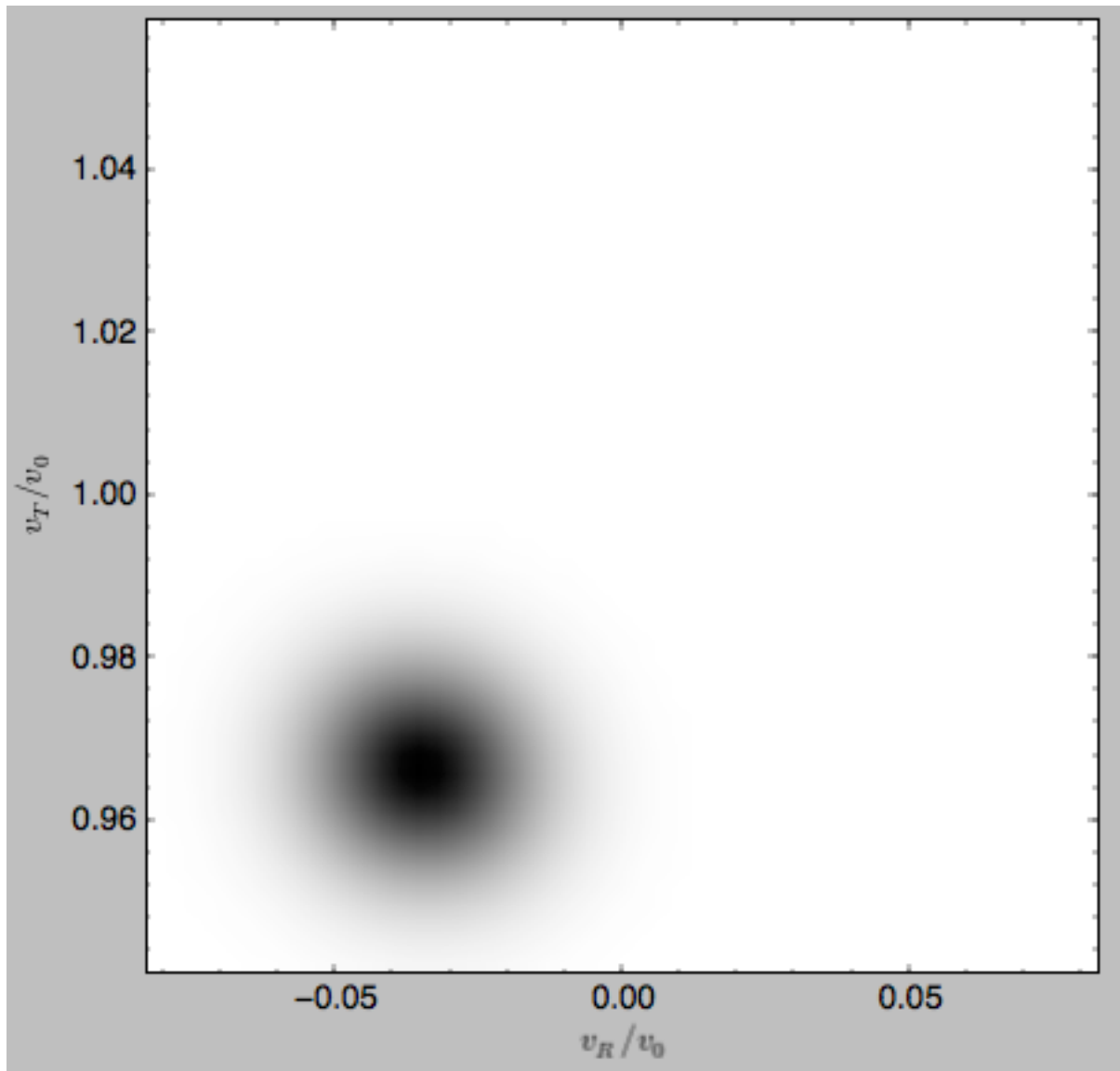
```
>>> print(edfcold.oortB(0.9, phi=22.5, deg=True, t=0., grid=gridcold,
↳ derivphiGrid=gridphicold, derivRGrid=gridrcold))
# -0.574674310521
>>> print(edfwarm.oortB(0.9, phi=22.5, deg=True, t=0., grid=gridwarm,
↳ derivphiGrid=gridphiwarm, derivRGrid=gridrwarm))
# -0.555546911144
```

and similar for `oortC` and `oortK`. These warm results should again be considered for illustration only, as the grid is not sufficiently fine to have a small numerical error.

The grids that have been calculated can also be plotted to show the full velocity DF. For example,

```
>>> gridcold.plot()
```

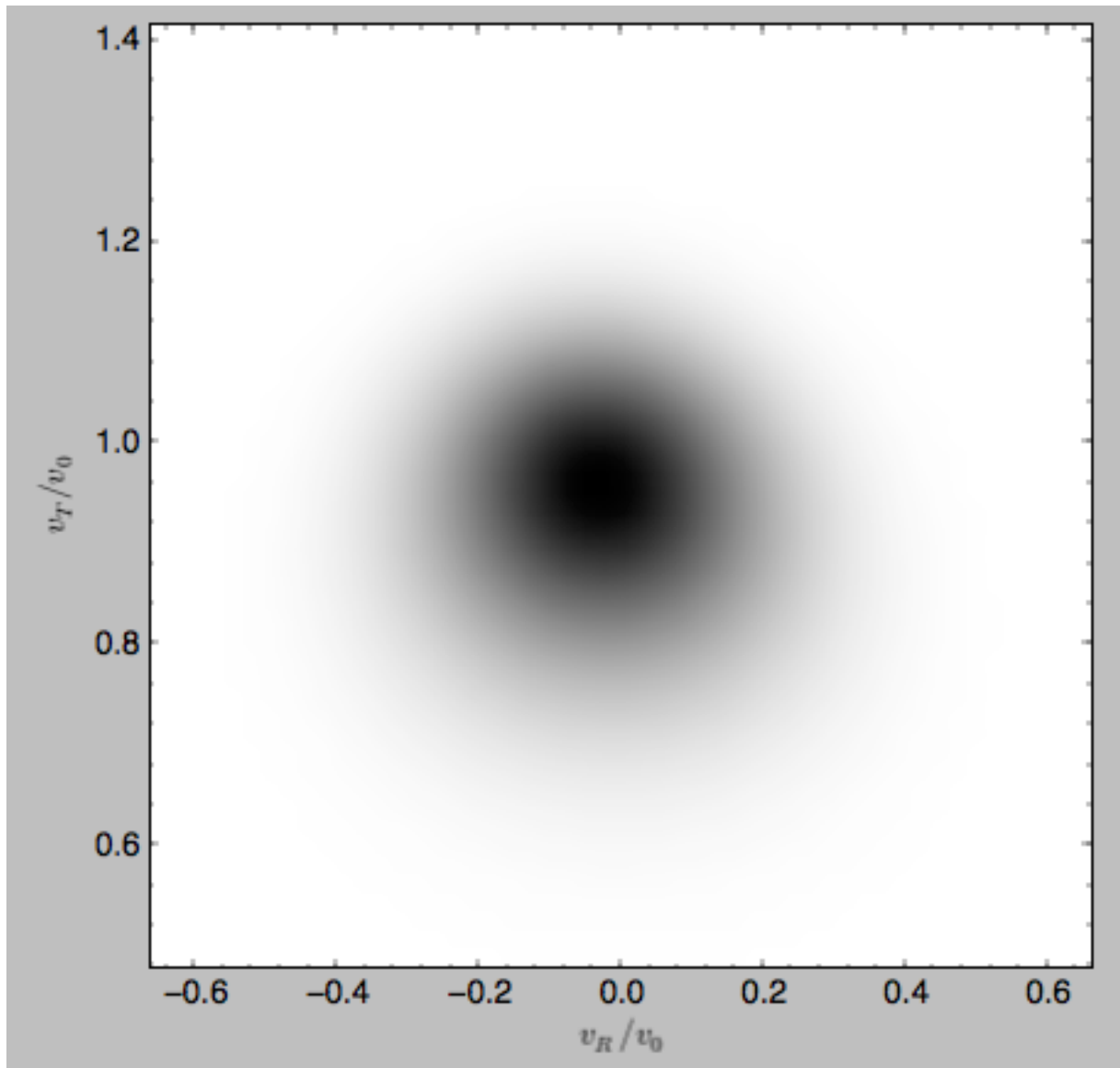
gives



which demonstrates that the DF is basically the initial DF that has been displaced (by a significant amount compared to the velocity dispersion). The warm velocity distribution is given by

```
>>> gridwarm.plot()
```

which returns



The shift of the smooth DF here is much smaller than the velocity dispersion.

1.6.7 Example: The Hercules stream in the Solar neighborhood as a result of the Galactic bar

We can combine the orbit integration capabilities of galpy with the provided distribution functions and see the effect of the Galactic bar on stellar velocities. By backward integrating orbits starting at the Solar position in a potential that includes the Galactic bar we can evaluate what the velocity distribution is that we should see today if the Galactic bar stirred up a steady-state disk. For this we initialize a flat rotation curve potential and Dehnen's bar potential

```
>>> from galpy.potential import LogarithmicHaloPotential, DehnenBarPotential
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> dp= DehnenBarPotential()
```

The Dehnen bar potential is initialized to start bar formation four bar periods before the present day and to have completely formed the bar two bar periods ago. We can integrate back to the time before bar-formation:

```
>>> ts= numpy.linspace(0,dp.tform(),1000)
```

where `dp.tform()` is the time of bar-formation (in the usual time-coordinates).

We initialize orbits on a grid in velocity space and integrate them

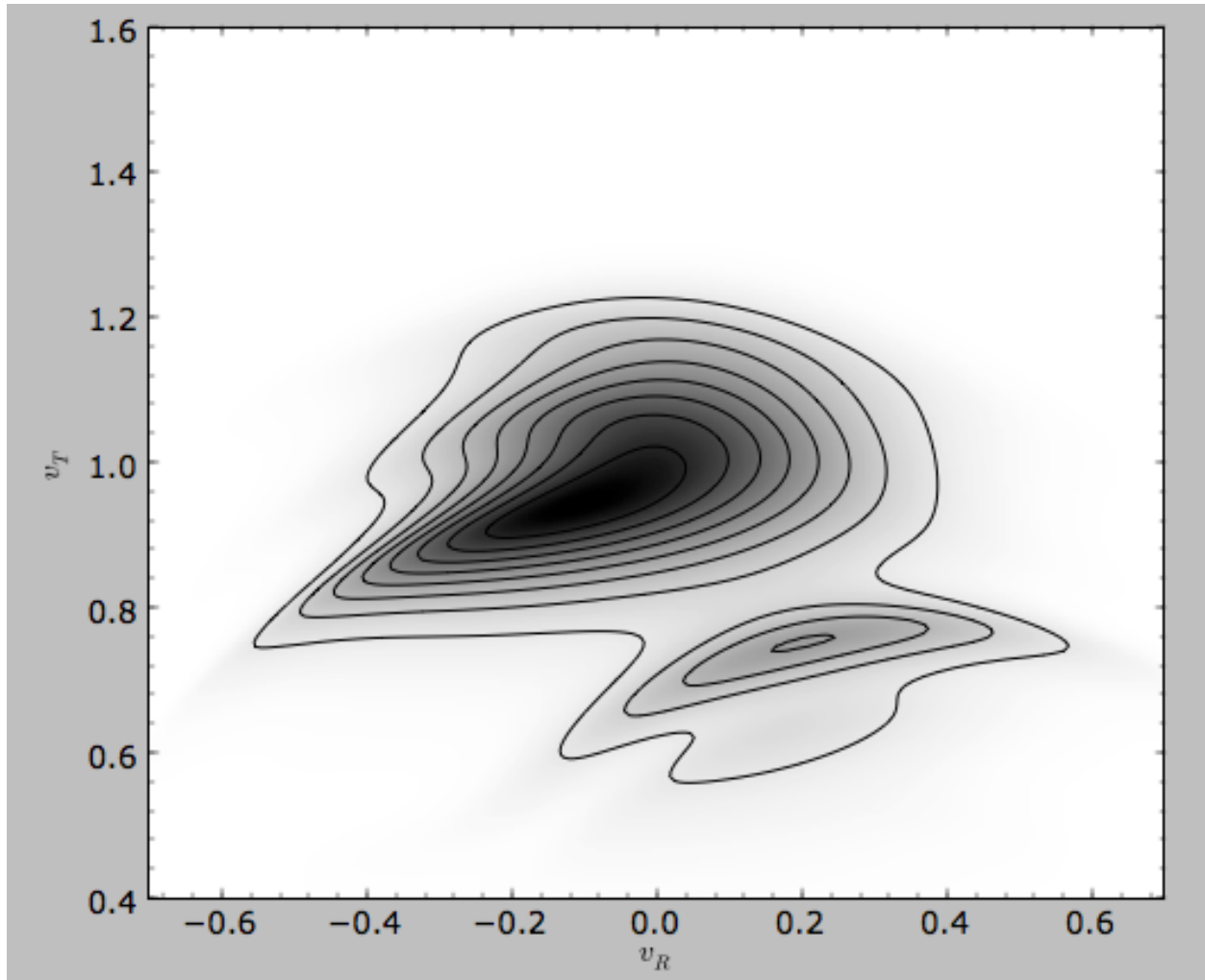
```
>>> ins=[Orbit([1.,-0.7+1.4/100*jj,1.-0.6+1.2/100*ii,0.]) for jj in range(101)]
↪ii in range(101)]
>>> int=[o.integrate(ts,[lp,dp]) for o in j] for j in ins]
```

We can then evaluate the weight of these orbits by assuming that the disk was in a steady-state before bar-formation with a Dehnen distribution function. We evaluate the Dehnen distribution function at `dp.tform()` for each of the orbits

```
>>> dfc= dehnendf(beta=0.,correct=True)
>>> out= [[dfc(o(dp.tform())) for o in j] for j in ins]
>>> out= numpy.array(out)
```

This gives

```
>>> from galpy.util.bovy_plot import bovy_dens2d
>>> bovy_dens2d(out,origin='lower',cmap='gist_yarg',contours=True,xrange=[-0.7,0.7],
↪yrange=[0.4,1.6],xlabel=r'$v_R$',ylabel=r'$v_T$')
```

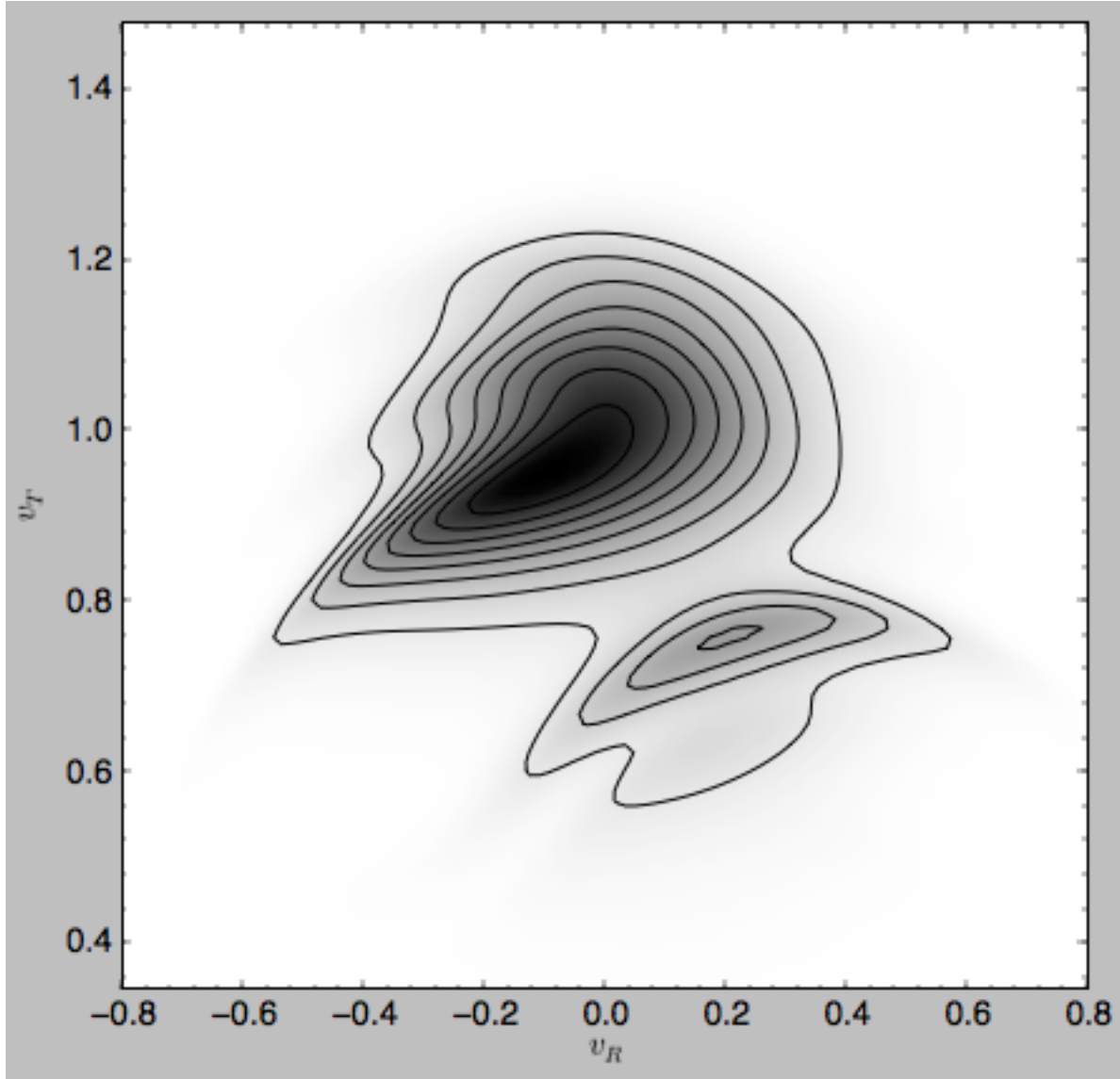


Now that galpy contains the `evolveddiskdf` described above, this whole calculation is encapsulated in this module and can be done much more easily as

```
>>> edf= evolveddiskdf(dfc,[lp,dp],to=dp.tform())
>>> mvr, grid= edf.meanvR(1.,grid=True,gridpoints=101,returnGrid=True)
```

The gridded DF can be accessed as `grid.df`, which we can plot as before

```
>>> bovy_dens2d(grid.df.T,origin='lower',cmap='gist_yarg',contours=True,xrange=[grid.
↪vRgrid[0],grid.vRgrid[-1]],yrange=[grid.vTgrid[0],grid.vTgrid[-1]],xlabel=r'$v_R$',
↪ylabel=r'$v_T$')
```



For more information see [2000AJ...119..800D](#) and [2010ApJ...725.1676B](#). Note that the x-axis in the Figure above is defined as minus the x-axis in these papers.

1.7 Action-angle coordinates

galpy can calculate actions and angles for a large variety of potentials (any time-independent potential in principle). These are implemented in a separate module `galpy.actionAngle`, and the preferred method for accessing them is through the routines in this module. There is also some support for accessing the `actionAngle` routines as methods of the `Orbit` class.

Since v1.2, galpy can also compute positions and velocities corresponding to a given set of actions and angles for axisymmetric potentials using the TorusMapper code of [Binney & McMillan \(2016\)](#). This is described in [this section](#) below. The interface for this is different than for the other action-angle classes, because the transformations are generally different.

Action-angle coordinates can be calculated for the following potentials/approximations:

- Isochrone potential
- Spherical potentials
- Adiabatic approximation
- Staeckel approximation
- A general orbit-integration-based technique

There are classes corresponding to these different potentials/approximations and actions, frequencies, and angles can typically be calculated using these three methods:

- `__call__`: returns the actions
- `actionsFreqs`: returns the actions and the frequencies
- `actionsFreqsAngles`: returns the actions, frequencies, and angles

These are not all implemented for each of the cases above yet.

The adiabatic and Staeckel approximation have also been implemented in C and using grid-based interpolation, for extremely fast action-angle calculations (see below).

1.7.1 Action-angle coordinates for the isochrone potential

The isochrone potential is the only potential for which all of the actions, frequencies, and angles can be calculated analytically. We can do this in galpy by doing

```
>>> from galpy.potential import IsochronePotential
>>> from galpy.actionAngle import actionAngleIsochrone
>>> ip= IsochronePotential(b=1.,normalize=1.)
>>> aAI= actionAngleIsochrone(ip=ip)
```

aAI is now an instance that can be used to calculate action-angle variables for the specific isochrone potential ip. Calling this instance returns (J_R, L_Z, J_z)

```
>>> aAI(1.,0.1,1.1,0.1,0.) #inputs R,vR,vT,z,vz
# (array([ 0.00713759]), array([ 1.1]), array([ 0.00553155]))
```

or for a more eccentric orbit

```
>>> aAI(1.,0.5,1.3,0.2,0.1)
# (array([ 0.13769498]), array([ 1.3]), array([ 0.02574507]))
```

Note that we can also specify phi, but this is not necessary

```
>>> aAI(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.13769498]), array([ 1.3]), array([ 0.02574507]))
```

We can likewise calculate the frequencies as well

```
>>> aAI.actionsFreqs(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.13769498]),
#  array([ 1.3]),
#  array([ 0.02574507]),
#  array([ 1.29136096]),
#  array([ 0.79093738]),
#  array([ 0.79093738]))
```

The output is $(J_R, L_Z, J_Z, \Omega_R, \Omega_\phi, \Omega_Z)$. For any spherical potential, $\Omega_\phi = \text{sgn}(L_Z)\Omega_Z$, such that the last two frequencies are the same.

We obtain the angles as well by calling

```
>>> aAI.actionsFreqsAngles(1., 0.5, 1.3, 0.2, 0.1, 0.)
# (array([ 0.13769498]),
#  array([ 1.3]),
#  array([ 0.02574507]),
#  array([ 1.29136096]),
#  array([ 0.79093738]),
#  array([ 0.79093738]),
#  array([ 0.57101518]),
#  array([ 5.96238847]),
#  array([ 1.24999949]))
```

The output here is $(J_R, L_Z, J_Z, \Omega_R, \Omega_\phi, \Omega_Z, \theta_R, \theta_\phi, \theta_Z)$.

To check that these are good action-angle variables, we can calculate them along an orbit

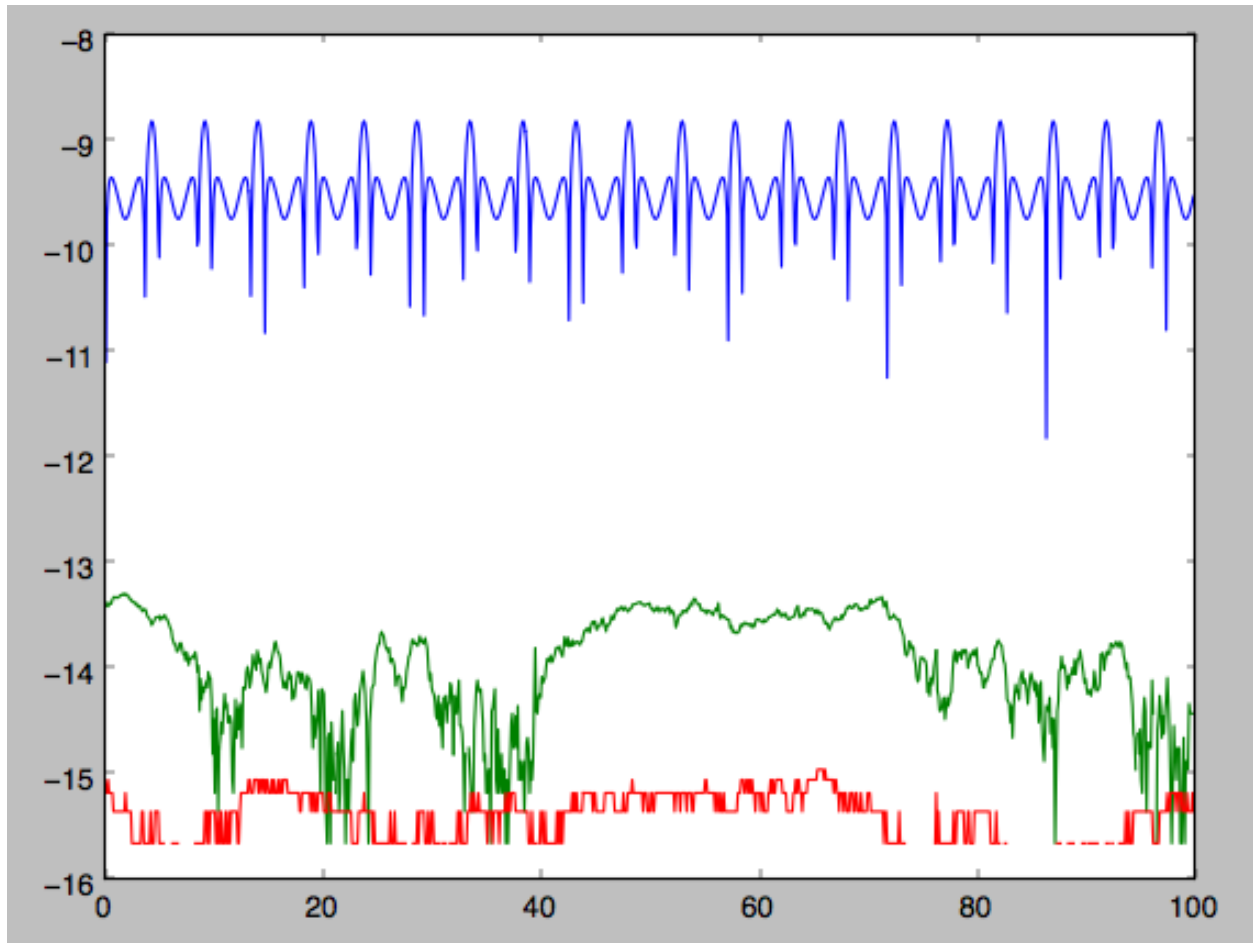
```
>>> from galpy.orbit import Orbit
>>> o= Orbit([1., 0.5, 1.3, 0.2, 0.1, 0.])
>>> ts= numpy.linspace(0., 100., 1001)
>>> o.integrate(ts, ip)
>>> jfa= aAI.actionsFreqsAngles(o.R(ts), o.vR(ts), o.vT(ts), o.z(ts), o.vz(ts), o.phi(ts))
```

which works because we can provide arrays for the R etc. inputs.

We can then check that the actions are constant over the orbit

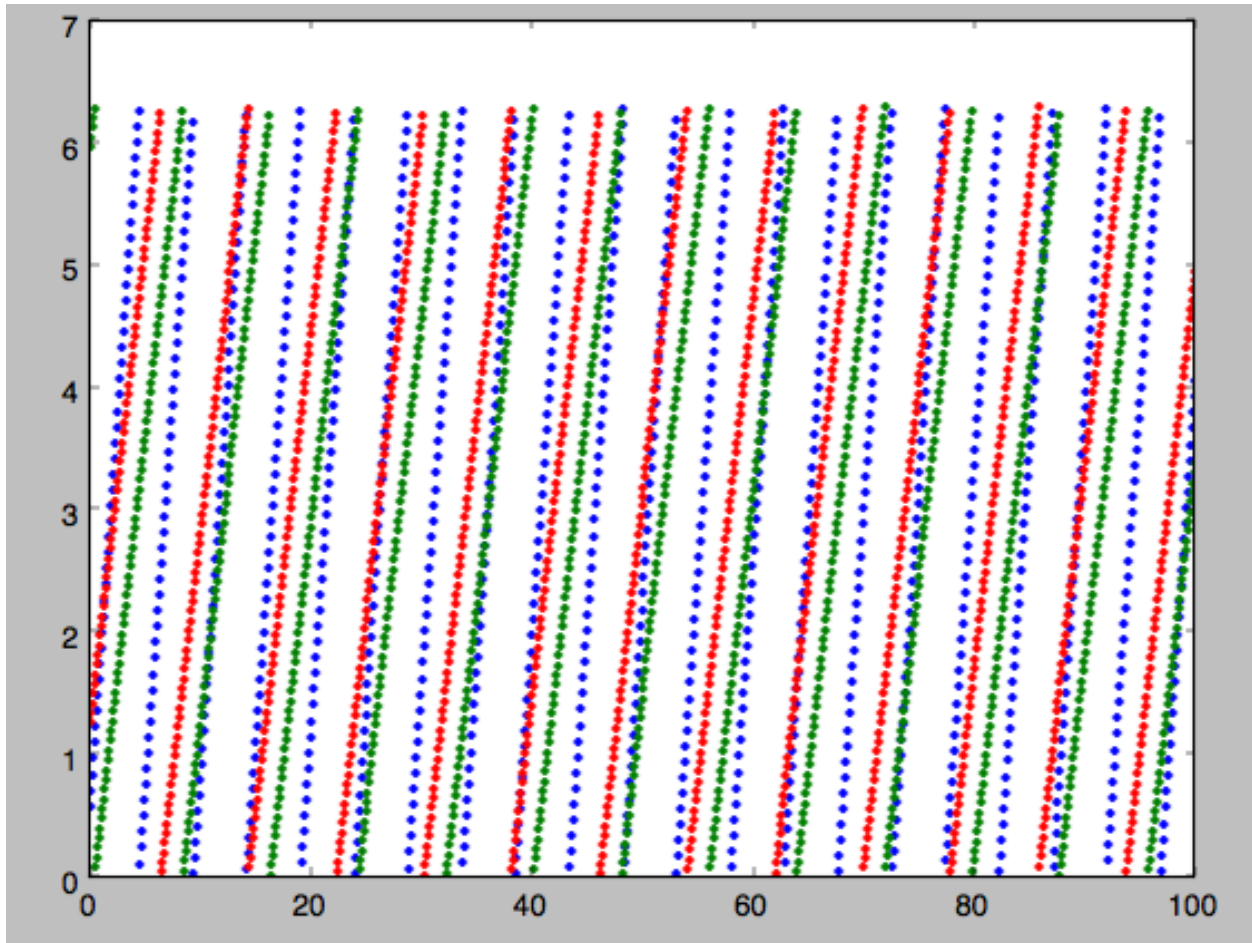
```
>>> plot(ts, numpy.log10(numpy.fabs((jfa[0]-numpy.mean(jfa[0])))))
>>> plot(ts, numpy.log10(numpy.fabs((jfa[1]-numpy.mean(jfa[1])))))
>>> plot(ts, numpy.log10(numpy.fabs((jfa[2]-numpy.mean(jfa[2])))))
```

which gives



The actions are all conserved. The angles increase linearly with time

```
>>> plot(ts, jfa[6], 'b.')
>>> plot(ts, jfa[7], 'g.')
>>> plot(ts, jfa[8], 'r.')
```

1.7.2 Action-angle coordinates for spherical potentials

Action-angle coordinates for any spherical potential can be calculated using a few orbit integrations. These are implemented in galpy in the `actionAngleSpherical` module. For example, we can do

```
>>> from galpy.potential import LogarithmicHaloPotential
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> from galpy.actionAngle import actionAngleSpherical
>>> aAS= actionAngleSpherical(pot=lp)
```

For the same eccentric orbit as above we find

```
>>> aAS(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.22022112]), array([ 1.3]), array([ 0.02574507]))
>>> aAS.actionsFreqs(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.22022112]),
#  array([ 1.3]),
#  array([ 0.02574507]),
#  array([ 0.87630459]),
#  array([ 0.60872881]),
#  array([ 0.60872881]))
>>> aAS.actionsFreqsAngles(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.22022112]),
```

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```
# array([ 1.3]),
# array([ 0.02574507]),
# array([ 0.87630459]),
# array([ 0.60872881]),
# array([ 0.60872881]),
# array([ 0.40443857]),
# array([ 5.85965048]),
# array([ 1.1472615]))
```

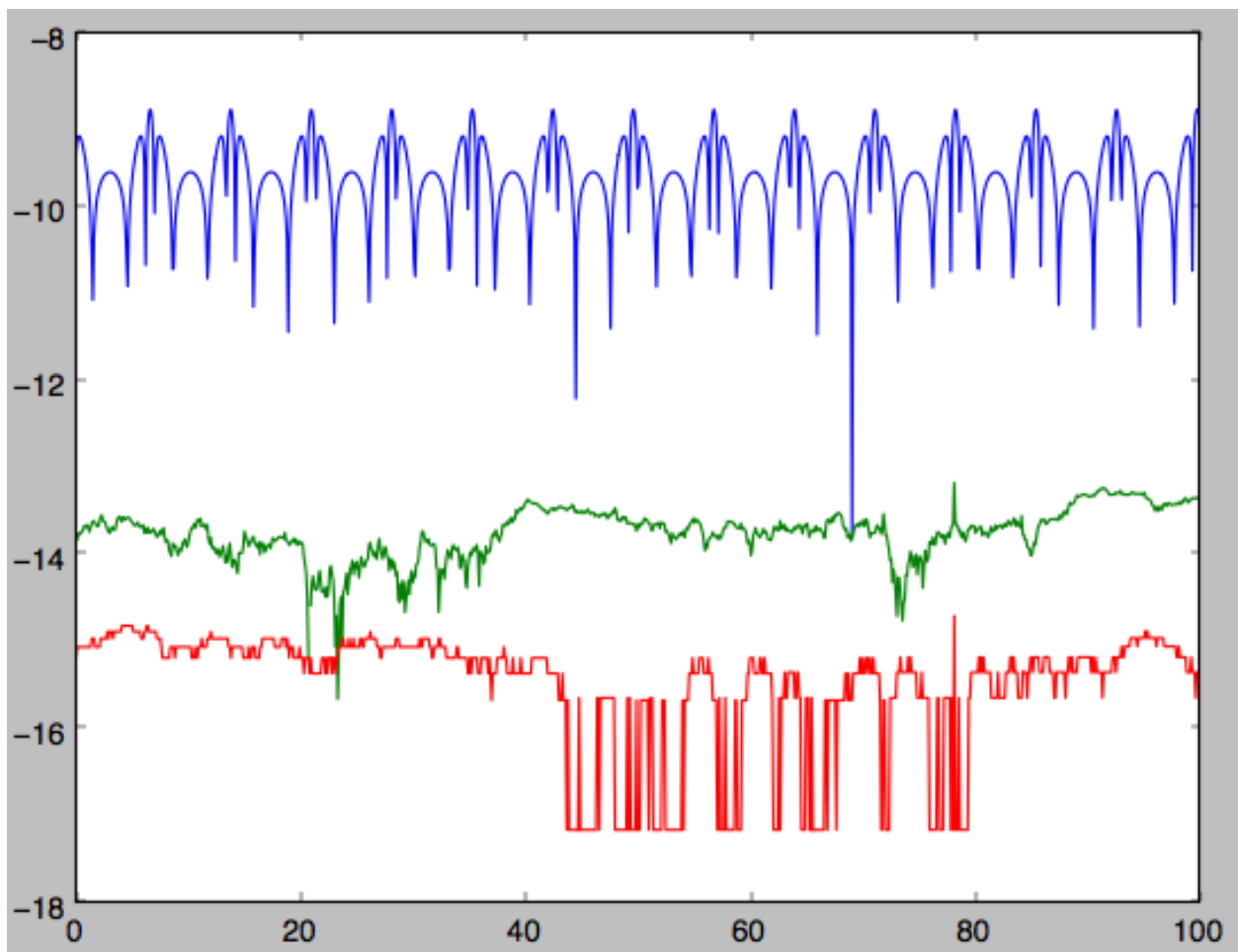
We can again check that the actions are conserved along the orbit and that the angles increase linearly with time:

```
>>> o.integrate(ts,lp)
>>> jfa= aAS.actionsFreqsAngles(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts),o.phi(ts),
↳fixed_quad=True)
```

where we use `fixed_quad=True` for a faster evaluation of the required one-dimensional integrals using Gaussian quadrature. We then plot the action fluctuations

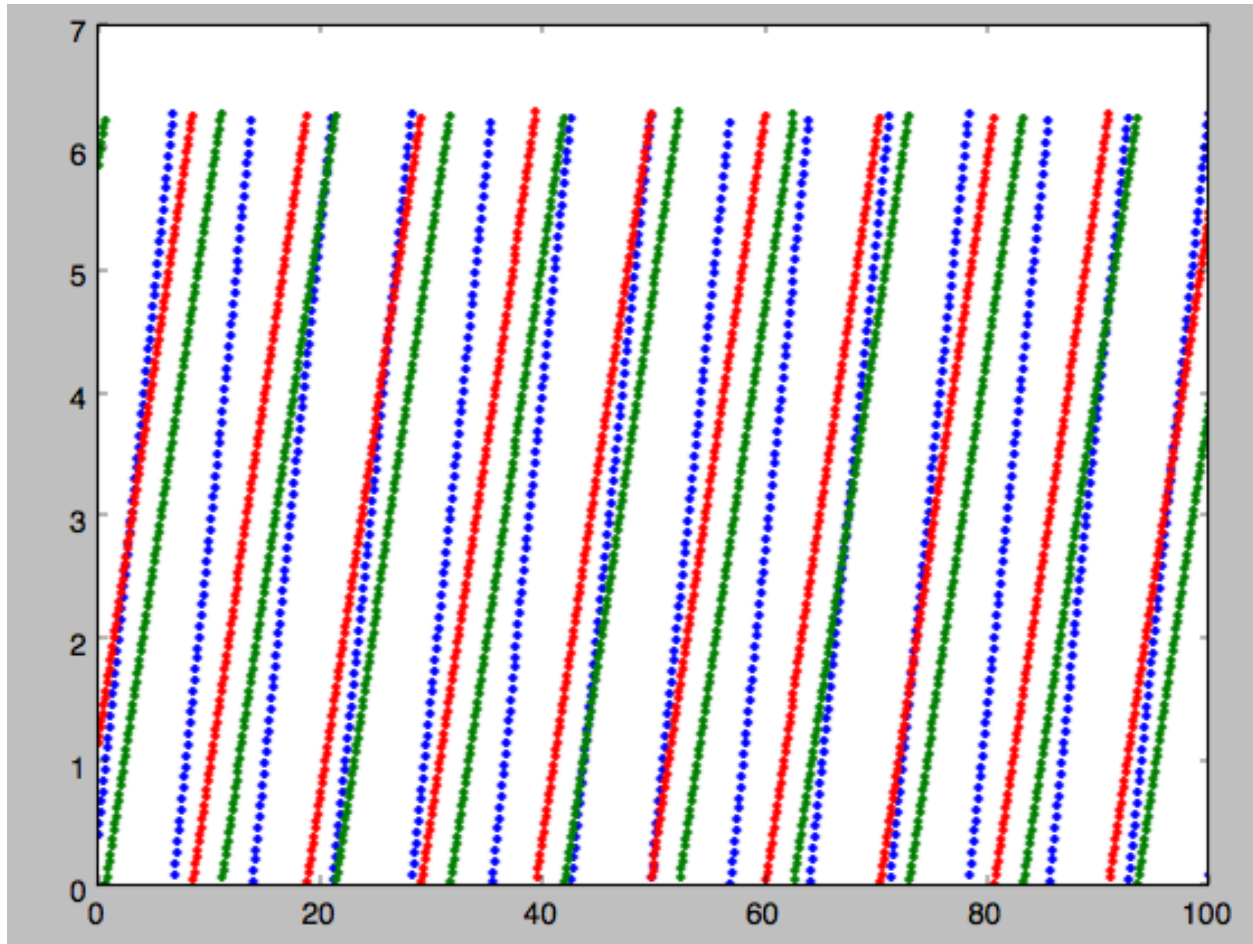
```
>>> plot(ts,numpy.log10(numpy.fabs((jfa[0]-numpy.mean(jfa[0])))))
>>> plot(ts,numpy.log10(numpy.fabs((jfa[1]-numpy.mean(jfa[1])))))
>>> plot(ts,numpy.log10(numpy.fabs((jfa[2]-numpy.mean(jfa[2])))))
```

which gives



showing that the actions are all conserved. The angles again increase linearly with time

```
>>> plot(ts, jfa[6], 'b.')
>>> plot(ts, jfa[7], 'g.')
>>> plot(ts, jfa[8], 'r.')
```



We can check the spherical action-angle calculations against the analytical calculations for the isochrone potential. Starting again from the isochrone potential used in the previous section

```
>>> ip= IsochronePotential(b=1.,normalize=1.)
>>> aAI= actionAngleIsochrone(ip=ip)
>>> aAS= actionAngleSpherical(pot=ip)
```

we can compare the actions, frequencies, and angles computed using both

```
>>> aAI.actionsFreqsAngles(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.13769498]),
# array([ 1.3]),
# array([ 0.02574507]),
# array([ 1.29136096]),
# array([ 0.79093738]),
# array([ 0.79093738]),
# array([ 0.57101518]),
# array([ 5.96238847]),
# array([ 1.24999949]))
```

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```
>>> aAS.actionsFreqsAngles(1.,0.5,1.3,0.2,0.1,0.)
# (array([ 0.13769498]),
#   array([ 1.3]),
#   array([ 0.02574507]),
#   array([ 1.29136096]),
#   array([ 0.79093738]),
#   array([ 0.79093738]),
#   array([ 0.57101518]),
#   array([ 5.96238838]),
#   array([ 1.2499994]))
```

or more explicitly comparing the two

```
>>> [r-s for r,s in zip(aAI.actionsFreqsAngles(1.,0.5,1.3,0.2,0.1,0.),aAS.
↳actionsFreqsAngles(1.,0.5,1.3,0.2,0.1,0.))]
# [array([ 6.66133815e-16]),
#   array([ 0.]),
#   array([ 0.]),
#   array([ -4.53851845e-10]),
#   array([ 4.74775219e-10]),
#   array([ 4.74775219e-10]),
#   array([ -1.65965242e-10]),
#   array([ 9.04759645e-08]),
#   array([ 9.04759649e-08))]
```

1.7.3 Action-angle coordinates using the adiabatic approximation

For non-spherical, axisymmetric potentials galpy contains multiple methods for calculating approximate action-angle coordinates. The simplest of those is the adiabatic approximation, which works well for disk orbits that do not go too far from the plane, as it assumes that the vertical motion is decoupled from that in the plane (e.g., 2010MNRAS.401.2318B).

Setup is similar as for other actionAngle objects

```
>>> from galpy.potential import MWPotential2014
>>> from galpy.actionAngle import actionAngleAdiabatic
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014)
```

and evaluation then proceeds similarly as before

```
>>> aAA(1.,0.1,1.1,0.,0.05)
# (0.01351896260559274, 1.1, 0.0004690133479435352)
```

We can again check that the actions are conserved along the orbit

```
>>> from galpy.orbit import Orbit
>>> ts=numpy.linspace(0.,100.,1001)
>>> o= Orbit([1.,0.1,1.1,0.,0.05])
>>> o.integrate(ts,MWPotential2014)
>>> js= aAA(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts))
```

This takes a while. The adiabatic approximation is also implemented in C, which leads to great speed-ups. Here is how to use it

```
>>> timeit(aAA(1.,0.1,1.1,0.,0.05))
# 10 loops, best of 3: 73.7 ms per loop
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014,c=True)
>>> timeit(aAA(1.,0.1,1.1,0.,0.05))
# 1000 loops, best of 3: 1.3 ms per loop
```

or about a *50 times* speed-up. For arrays the speed-up is even more impressive

```
>>> s= numpy.ones(100)
>>> timeit(aAA(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 10 loops, best of 3: 37.8 ms per loop
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014) #back to no C
>>> timeit(aAA(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 1 loops, best of 3: 7.71 s per loop
```

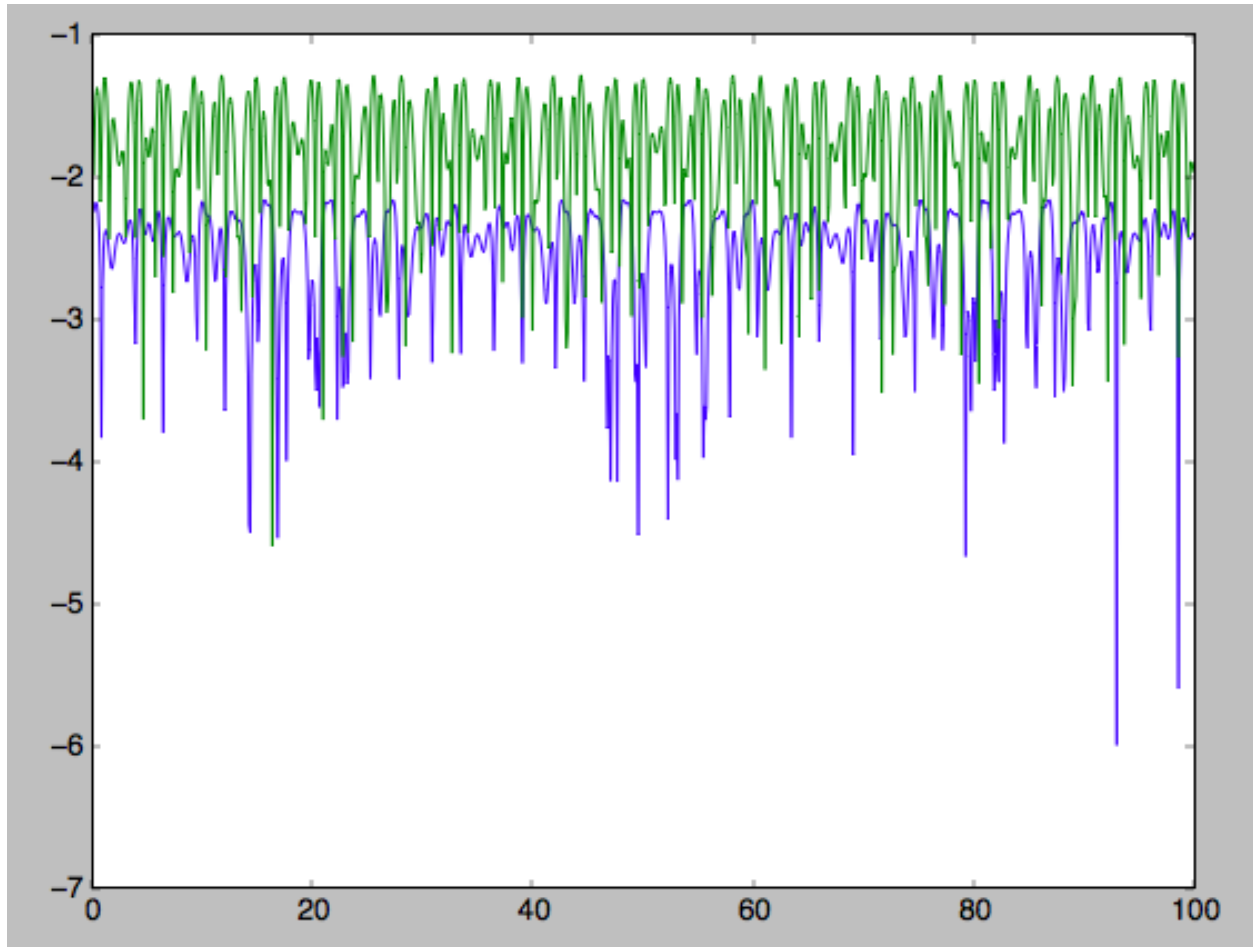
or a speed-up of 200! Back to the previous example, you can run it with `c=True` to speed up the computation

```
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014,c=True)
>>> js= aAA(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts))
```

We can plot the radial- and vertical-action fluctuation as a function of time

```
>>> plot(ts,numpy.log10(numpy.fabs((js[0]-numpy.mean(js[0]))/numpy.mean(js[0]))))
>>> plot(ts,numpy.log10(numpy.fabs((js[2]-numpy.mean(js[2]))/numpy.mean(js[2]))))
```

which gives



The radial action is conserved to about half a percent, the vertical action to two percent.

Another way to speed up the calculation of actions using the adiabatic approximation is to tabulate the actions on a grid in (approximate) integrals of the motion and evaluating new actions by interpolating on this grid. How this is done in practice is described in detail in the galpy paper. To setup this grid-based interpolation method, which is contained in `actionAngleAdiabaticGrid`, do

```
>>> from galpy.actionAngle import actionAngleAdiabaticGrid
>>> aAG= actionAngleAdiabaticGrid(pot=MWPotential2014,nR=31,nEz=31,nEr=51,nLz=51,
↪ c=True)
```

where `c=True` specifies that we use the C implementation of `actionAngleAdiabatic` for speed. We can now evaluate in the same way as before, for example

```
>>> aAA(1.,0.1,1.1,0.,0.05), aAG(1.,0.1,1.1,0.,0.05)
# ((array([ 0.01352523]), array([ 1.1]), array([ 0.00046909])),
# (0.013527010324238781, 1.1, 0.00047747359874375148))
```

which agree very well. To look at the timings, we first switch back to not using C and then list all of the relevant timings:

```
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014,c=False)
# Not using C, direct calculation
>>> timeit(aAA(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 1 loops, best of 3: 9.05 s per loop
```

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```
>>> aAA= actionAngleAdiabatic(pot=MWPotential2014,c=True)
# Using C, direct calculation
>>> timeit(aAA(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 10 loops, best of 3: 39.7 ms per loop
# Grid-based calculation
>>> timeit(aAG(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 1000 loops, best of 3: 1.09 ms per loop
```

Thus, in this example (and more generally) the grid-based calculation is significantly faster than even the direct implementation in C. The overall speed up between the direct Python version and the grid-based version is larger than 8,000; the speed up between the direct C version and the grid-based version is 36. For larger arrays of input phase-space positions, the latter speed up can increase to 150. For simpler, fully analytical potentials the speed up will be slightly less, but for `MWPotential2014` and other more complicated potentials (such as those involving a double-exponential disk), the overhead of setting up the grid is worth it when evaluating more than a few thousand actions.

The adiabatic approximation works well for orbits that stay close to the plane. The orbit we have been considering so far only reaches a height two percent of R_0 , or about 150 pc for $R_0 = 8$ kpc.

```
>>> o.zmax()*8.
# 0.17903686455491979
```

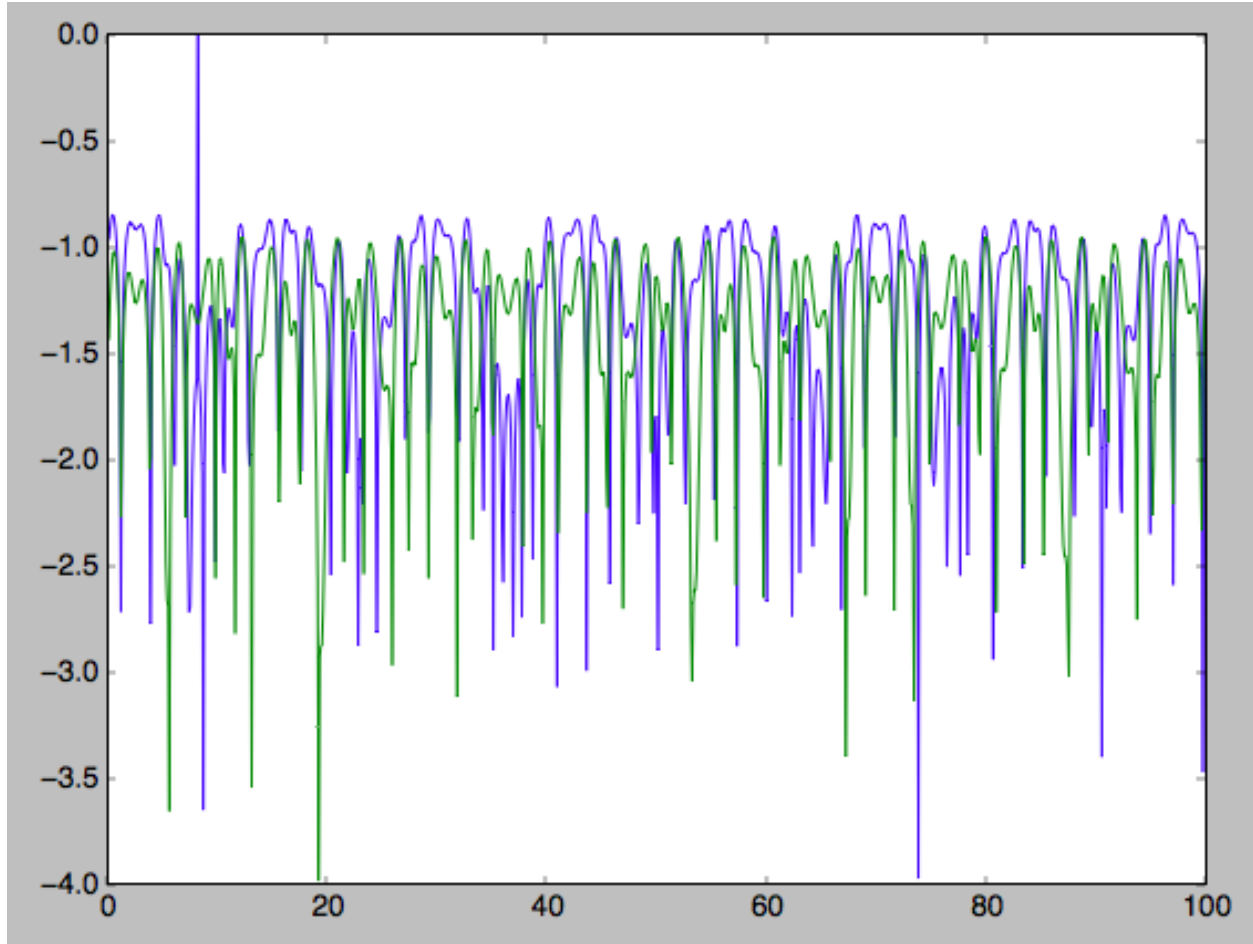
For orbits that reach distances of a kpc and more from the plane, the adiabatic approximation does not work as well. For example,

```
>>> o= Orbit([1.,0.1,1.1,0.,0.25])
>>> o.integrate(ts,MWPotential2014)
>>> o.zmax()*8.
# 1.3506059038621048
```

and we can again calculate the actions along the orbit

```
>>> js= aAA(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts))
>>> plot(ts,numpy.log10(numpy.fabs((js[0]-numpy.mean(js[0]))/numpy.mean(js[0]))))
>>> plot(ts,numpy.log10(numpy.fabs((js[2]-numpy.mean(js[2]))/numpy.mean(js[2]))))
```

which gives



The radial action is now only conserved to about ten percent and the vertical action to approximately five percent.

Warning: Frequencies and angles using the adiabatic approximation are not implemented at this time.

1.7.4 Action-angle coordinates using the Staeckel approximation

A better approximation than the adiabatic one is to locally approximate the potential as a Staeckel potential, for which actions, frequencies, and angles can be calculated through numerical integration. `galpy` contains an implementation of the algorithm of Binney (2012; [2012MNRAS.426.1324B](#)), which accomplishes the Staeckel approximation for disk-like (i.e., oblate) potentials without explicitly fitting a Staeckel potential. For all intents and purposes the adiabatic approximation is made obsolete by this new method, which is as fast and more precise. The only advantage of the adiabatic approximation over the Staeckel approximation is that the Staeckel approximation requires the user to specify a *focal length* Δ to be used in the Staeckel approximation. However, this focal length can be easily estimated from the second derivatives of the potential (see Sanders 2012; [2012MNRAS.426..128S](#)).

Starting from the second orbit example in the adiabatic section above, we first estimate a good focal length of the `MWPotential2014` to use in the Staeckel approximation. We do this by averaging (through the median) estimates at positions around the orbit (which we integrated in the example above)

```
>>> from galpy.actionAngle import estimateDeltaStaeckel
>>> estimateDeltaStaeckel(MWPotential2014,o.R(ts),o.z(ts))
# 0.40272708556203662
```


We will use $\Delta = 0.4$ in what follows. We set up the `actionAngleStaeckel` object

```
>>> from galpy.actionAngle import actionAngleStaeckel
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=False) #c=True is the_
↪ default
```

and calculate the actions

```
>>> aAS(o.R(),o.vR(),o.vT(),o.z(),o.vz())
# (0.019212848866725911, 1.1000000000000001, 0.015274597971510892)
```

The adiabatic approximation from above gives

```
>>> aAA(o.R(),o.vR(),o.vT(),o.z(),o.vz())
# (array([ 0.01686478]), array([ 1.1]), array([ 0.01590001]))
```

The `actionAngleStaeckel` calculations are sped up in two ways. First, the action integrals can be calculated using Gaussian quadrature by specifying `fixed_quad=True`

```
>>> aAS(o.R(),o.vR(),o.vT(),o.z(),o.vz(),fixed_quad=True)
# (0.01922167296633687, 1.1000000000000001, 0.015276825017286706)
```

which in itself leads to a ten times speed up

```
>>> timeit(aAS(o.R(),o.vR(),o.vT(),o.z(),o.vz(),fixed_quad=False))
# 10 loops, best of 3: 129 ms per loop
>>> timeit(aAS(o.R(),o.vR(),o.vT(),o.z(),o.vz(),fixed_quad=True))
# 100 loops, best of 3: 10.3 ms per loop
```

Second, the `actionAngleStaeckel` calculations have also been implemented in C, which leads to even greater speed-ups, especially for arrays

```
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=True)
>>> s= numpy.ones(100)
>>> timeit(aAS(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 10 loops, best of 3: 35.1 ms per loop
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=False) #back to no C
>>> timeit(aAS(1.*s,0.1*s,1.1*s,0.*s,0.05*s,fixed_quad=True))
# 1 loops, best of 3: 496 ms per loop
```

or a fifteen times speed up. The speed up is not that large because the bulge model in `MWPotential2014` requires expensive special functions to be evaluated. Computations could be sped up ten times more when using a simpler bulge model.

Similar to `actionAngleAdiabaticGrid`, we can also tabulate the actions on a grid of (approximate) integrals of the motion and interpolate over this look-up table when evaluating new actions. The details of how this look-up table is setup and used are again fully explained in the galpy paper. To use this grid-based Staeckel approximation, contained in `actionAngleStaeckelGrid`, do

```
>>> from galpy.actionAngle import actionAngleStaeckelGrid
>>> aASG= actionAngleStaeckelGrid(pot=MWPotential2014,delta=0.4,nE=51,npsi=51,nLz=61,
↪ c=True)
```

where `c=True` makes sure that we use the C implementation of the Staeckel method to calculate the grid. Because this is a fully three-dimensional grid, setting up the grid takes longer than it does for the adiabatic method (which only uses two two-dimensional grids). We can then evaluate actions as before

```
>>> aAS(o.R(),o.vR(),o.vT(),o.z(),o.vz()), aASG(o.R(),o.vR(),o.vT(),o.z(),o.vz())
# ((0.019212848866725911, 1.1000000000000001, 0.015274597971510892),
# (0.019221119033345408, 1.1000000000000001, 0.015022528662310393))
```

These actions agree very well. We can compare the timings of these methods as above

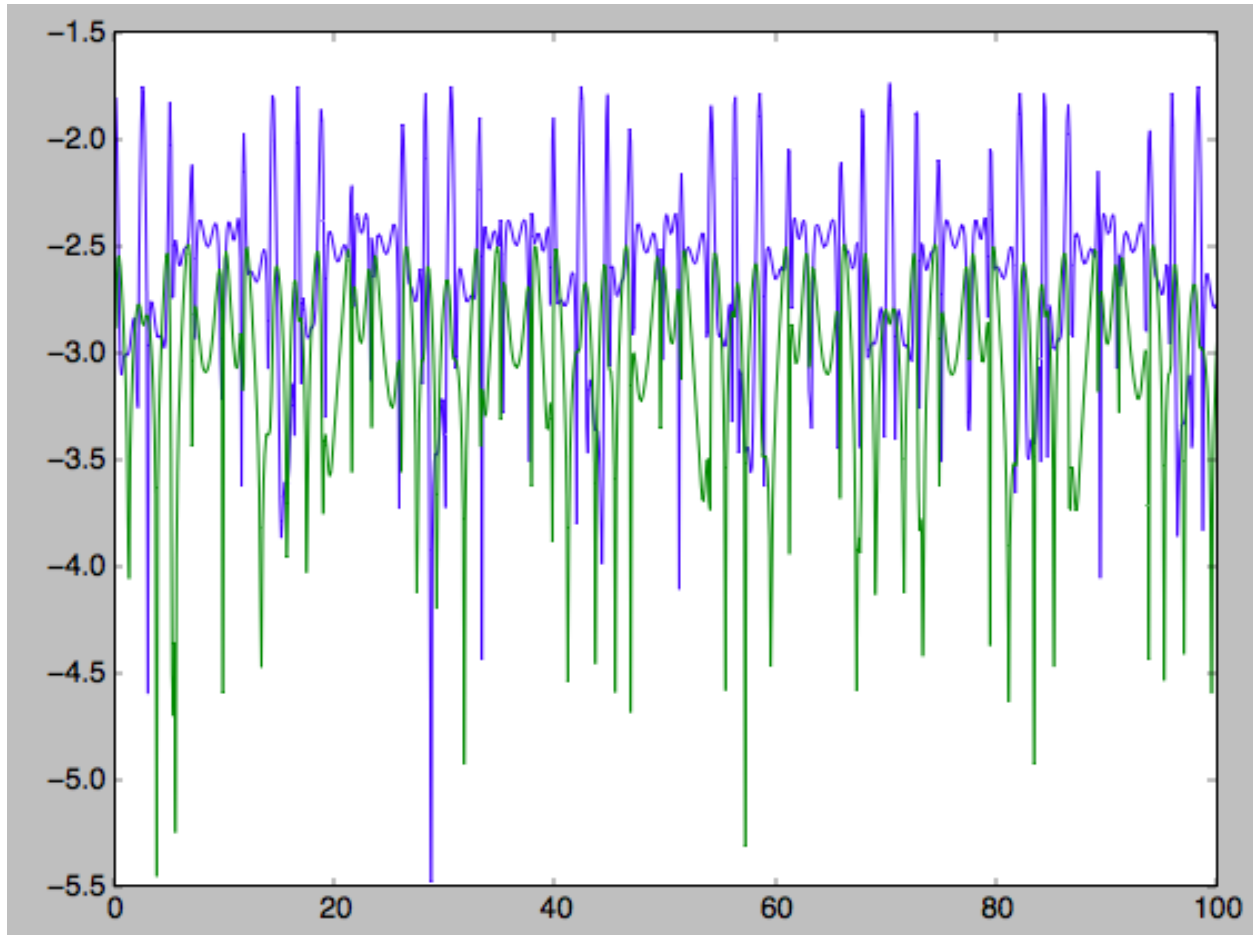
```
>>> timeit(aAS(1.*s,0.1*s,1.1*s,0.*s,0.05*s,fixed_quad=True))
# 1 loops, best of 3: 576 ms per loop # Not using C, direct calculation
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=True)
>>> timeit(aAS(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 100 loops, best of 3: 17.8 ms per loop # Using C, direct calculation
>>> timeit(aASG(1.*s,0.1*s,1.1*s,0.*s,0.05*s))
# 100 loops, best of 3: 3.45 ms per loop # Grid-based calculation
```

This demonstrates that the grid-based interpolation again leads to a significant speed up, even over the C implementation of the direct calculation. This speed up becomes more significant for larger array input, although it saturates at about 25 times (at least for MWPotential2014).

We can now go back to checking that the actions are conserved along the orbit (going back to the `c=False` version of `actionAngleStaeckel`)

```
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=False)
>>> js= aAS(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts),fixed_quad=True)
>>> plot(ts,numpy.log10(numpy.fabs((js[0]-numpy.mean(js[0]))/numpy.mean(js[0]))))
>>> plot(ts,numpy.log10(numpy.fabs((js[2]-numpy.mean(js[2]))/numpy.mean(js[2]))))
```

which gives



The radial action is now conserved to better than a percent and the vertical action to only a fraction of a percent. Clearly, this is much better than the five to ten percent errors found for the adiabatic approximation above.

For the Staeckel approximation we can also calculate frequencies and angles through the `actionsFreqs` and `actionsFreqsAngles` methods.

Warning: Frequencies and angles using the Staeckel approximation are *only* implemented in C. So use `c=True` in the setup of the `actionAngleStaeckel` object.

Warning: Angles using the Staeckel approximation in galpy are such that (a) the radial angle starts at zero at pericenter and increases then going toward apocenter; (b) the vertical angle starts at zero at $z=0$ and increases toward positive z_{max} . The latter is a different convention from that in Binney (2012), but is consistent with that in `actionAngleIsochrone` and `actionAngleSpherical`.

```
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.4,c=True)
>>> o= Orbit([1.,0.1,1.1,0.,0.25,0.]) #need to specify phi for angles
>>> aAS.actionsFreqsAngles(o.R(),o.vR(),o.vT(),o.z(),o.vz(),o.phi())
# (array([ 0.01922167]),
#  array([ 1.1]),
#  array([ 0.01527683]),
#  array([ 1.11317796]),
```

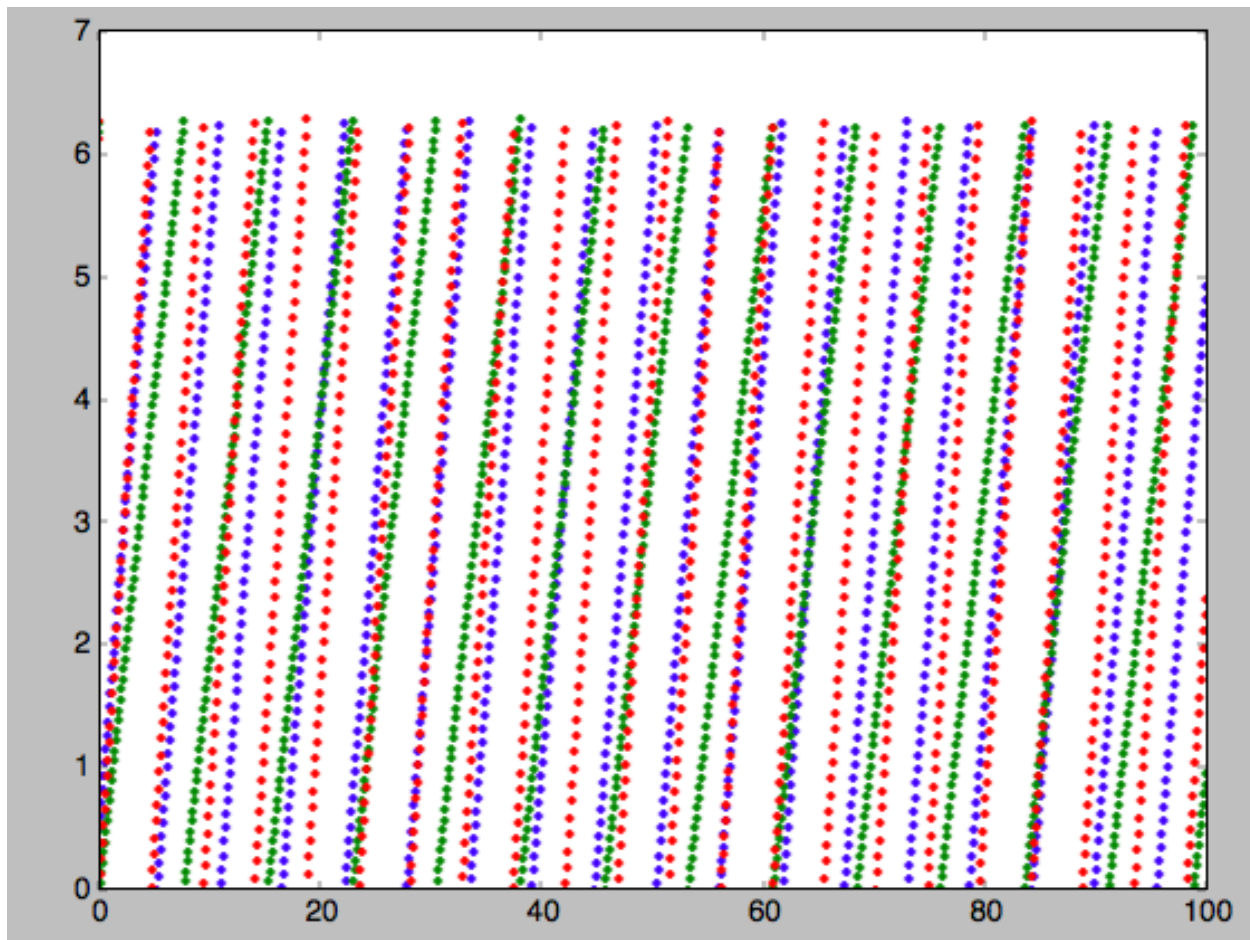
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```
# array([ 0.82538032]),
# array([ 1.34126138]),
# array([ 0.37758087]),
# array([ 6.17833493]),
# array([ 6.13368239])
```

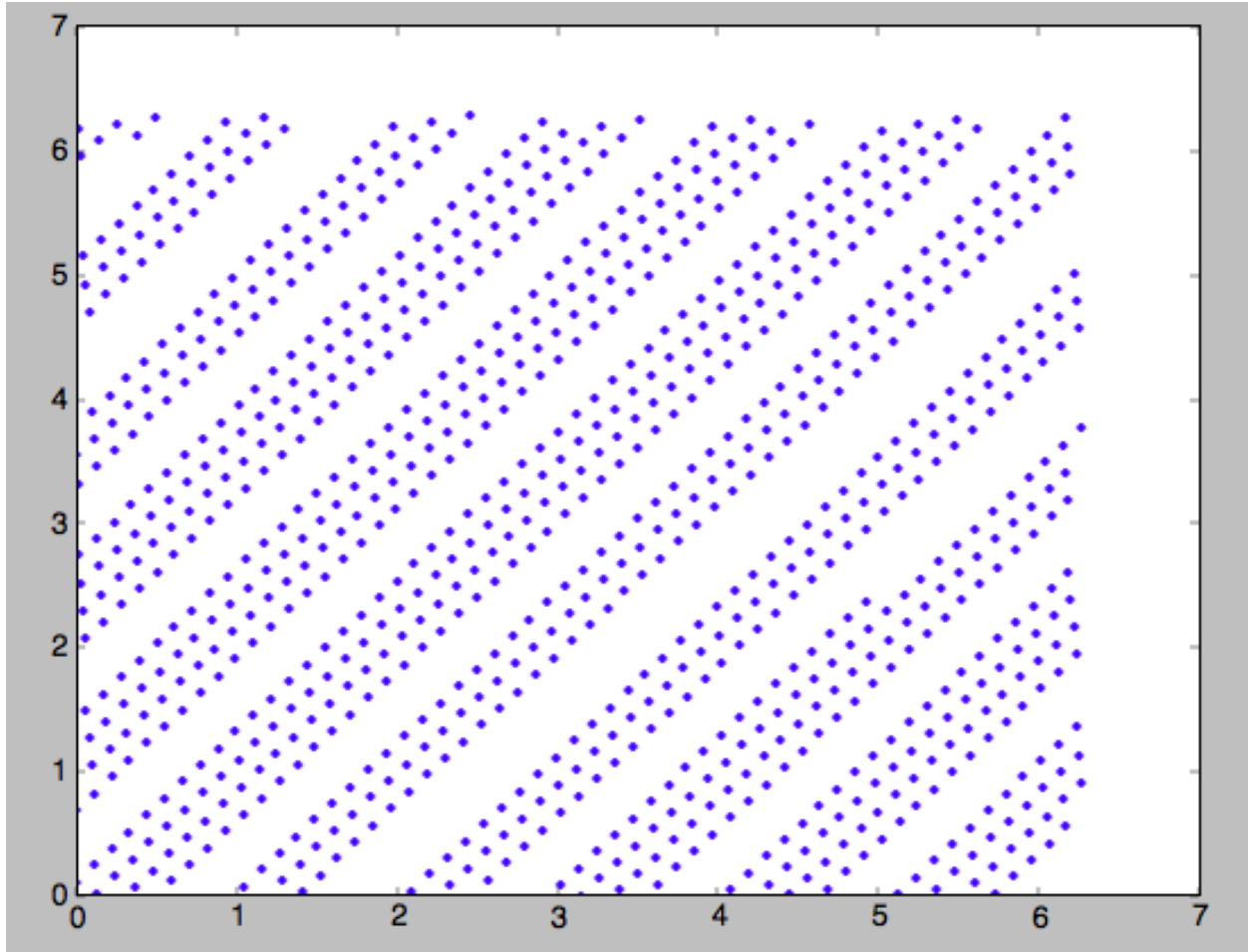
and we can check that the angles increase linearly along the orbit

```
>>> o.integrate(ts,MWPotential2014)
>>> jfa= aAS.actionsFreqsAngles(o.R(ts),o.vR(ts),o.vT(ts),o.z(ts),o.vz(ts),o.phi(ts))
>>> plot(ts,jfa[6],'b.')
>>> plot(ts,jfa[7],'g.')
>>> plot(ts,jfa[8],'r.')
>>>
```



or

```
>>> plot(jfa[6],jfa[8],'b.')
>>>
```



1.7.5 Action-angle coordinates using an orbit-integration-based approximation

The adiabatic and Staeckel approximations used above are good for stars on close-to-circular orbits, but they break down for more eccentric orbits (specifically, orbits for which the radial and/or vertical action is of a similar magnitude as the angular momentum). This is because the approximations made to the potential in these methods (that it is separable in R and z for the adiabatic approximation and that it is close to a Staeckel potential for the Staeckel approximation) break down for such orbits. Unfortunately, these methods cannot be refined to provide better approximations for eccentric orbits.

galpy contains a new method for calculating actions, frequencies, and angles that is completely general for any static potential. It can calculate the actions to any desired precision for any orbit in such potentials. The method works by employing an auxiliary isochrone potential and calculates action-angle variables by arithmetic operations on the actions and angles calculated in the auxiliary potential along an orbit (integrated in the true potential). Full details can be found in Appendix A of Bovy (2014).

We setup this method for a logarithmic potential as follows

```
>>> from galpy.actionAngle import actionAngleIsochroneApprox
>>> from galpy.potential import LogarithmicHaloPotential
>>> lp= LogarithmicHaloPotential(normalize=1.,q=0.9)
>>> aAIA= actionAngleIsochroneApprox(pot=lp,b=0.8)
```

$b=0.8$ here sets the scale parameter of the auxiliary isochrone potential; this potential can also be specified as an

IsochronePotential instance through `ip=`). We can now calculate the actions for an orbit similar to that of the GD-1 stream

```
>>> obs= numpy.array([1.56148083,0.35081535,-1.15481504,0.88719443,-0.47713334,0.
↳12019596]) #orbit similar to GD-1
>>> aAIA(*obs)
# (array([ 0.16605011]), array([-1.80322155]), array([ 0.50704439]))
```

An essential requirement of this method is that the angles calculated in the auxiliary potential go through the full range $[0, 2\pi]$. If this is not the case, galpy will raise a warning

```
>>> aAIA= actionAngleIsochroneApprox(pot=lp,b=10.8)
>>> aAIA(*obs)
# galpyWarning: Full radial angle range not covered for at least one object; actions_
↳are likely not reliable
# (array([ 0.08985167]), array([-1.80322155]), array([ 0.50849276]))
```

Therefore, some care should be taken to choosing a good auxiliary potential. galpy contains a method to estimate a decent scale parameter for the auxiliary scale parameter, which works similar to `estimateDeltaStaeckel` above except that it also gives a minimum and maximum b if multiple R and z are given

```
>>> from galpy.actionAngle import estimateBIsochrone
>>> from galpy.orbit import Orbit
>>> o= Orbit(obs)
>>> ts= numpy.linspace(0.,100.,1001)
>>> o.integrate(ts,lp)
>>> estimateBIsochrone(lp,o.R(ts),o.z(ts))
# (0.78065062339131952, 1.2265541473461612, 1.4899326335155412) #bmin,bmedian,bmax_
↳over the orbit
```

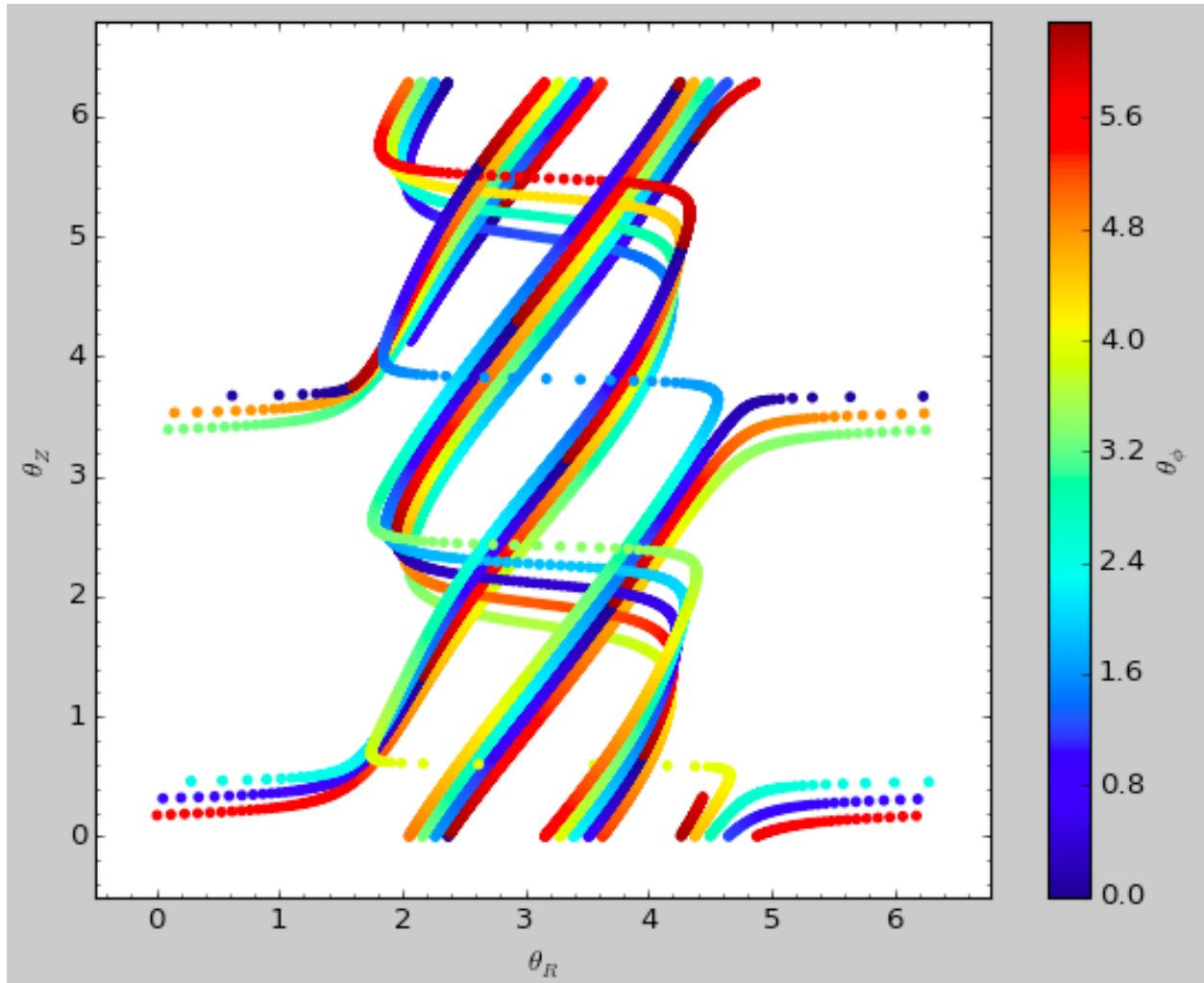
Experience shows that a scale parameter somewhere in the range returned by this function makes sure that the angles go through the full $[0, 2\pi]$ range. However, even if the angles go through the full range, the closer the angles increase to linear, the better the convergence of the algorithm is (and especially, the more accurate the calculation of the frequencies and angles is, see below). For example, for the scale parameter at the upper end of the range

```
>>> aAIA= actionAngleIsochroneApprox(pot=lp,b=1.5)
>>> aAIA(*obs)
# (array([ 0.01120145]), array([-1.80322155]), array([ 0.50788893]))
```

which does not agree with the previous calculation. We can inspect how the angles increase and how the actions converge by using the `aAIA.plot` function. For example, we can plot the radial versus the vertical angle in the auxiliary potential

```
>>> aAIA.plot(*obs,type='araz')
```

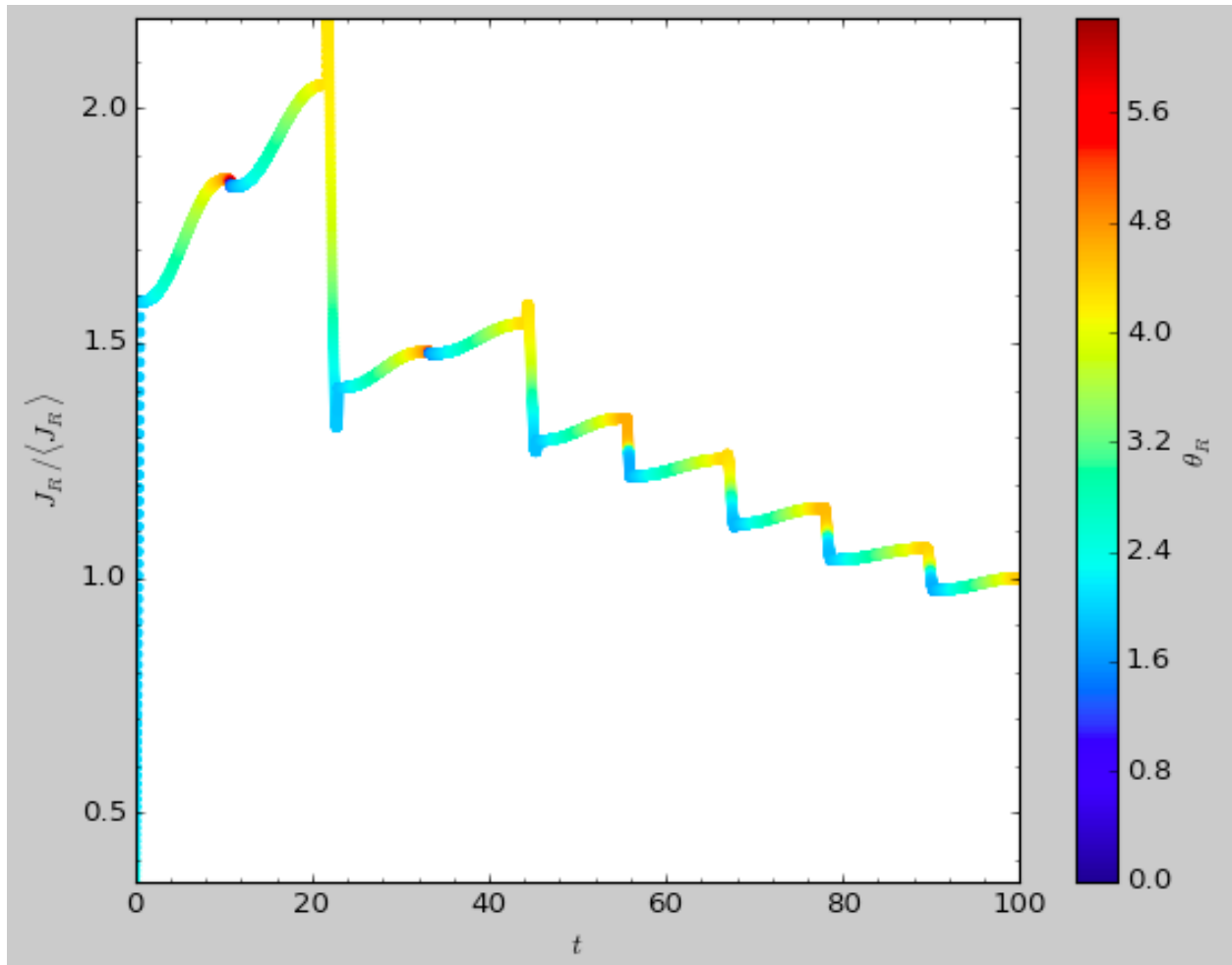
which gives



and this clearly shows that the angles increase *very* non-linearly, because the auxiliary isochrone potential used is too far from the real potential. This causes the actions to converge only very slowly. For example, for the radial action we can plot the converge as a function of integration time

```
>>> aAIA.plot(*obs, type='jr')
```

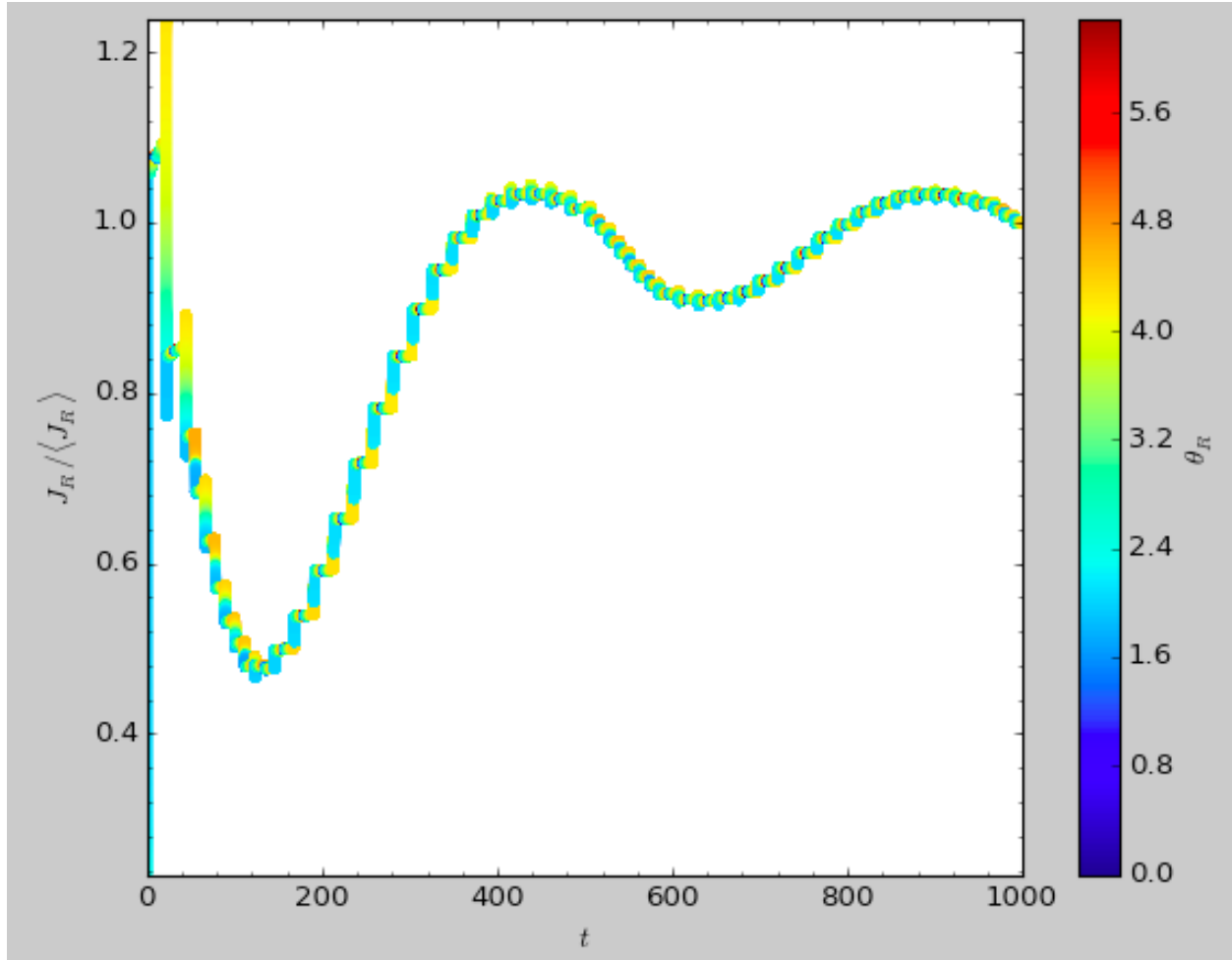
which gives



This Figure clearly shows that the radial action has not converged yet. We need to integrate *much* longer in this auxiliary potential to obtain convergence and because the angles increase so non-linearly, we also need to integrate the orbit much more finely:

```
>>> aAIA= actionAngleIsochroneApprox(pot=lp,b=1.5,tintJ=1000,ntintJ=800000)
>>> aAIA(*obs)
# (array([ 0.01711635]), array([-1.80322155]), array([ 0.51008058]))
>>> aAIA.plot(*obs,type='jr')
```

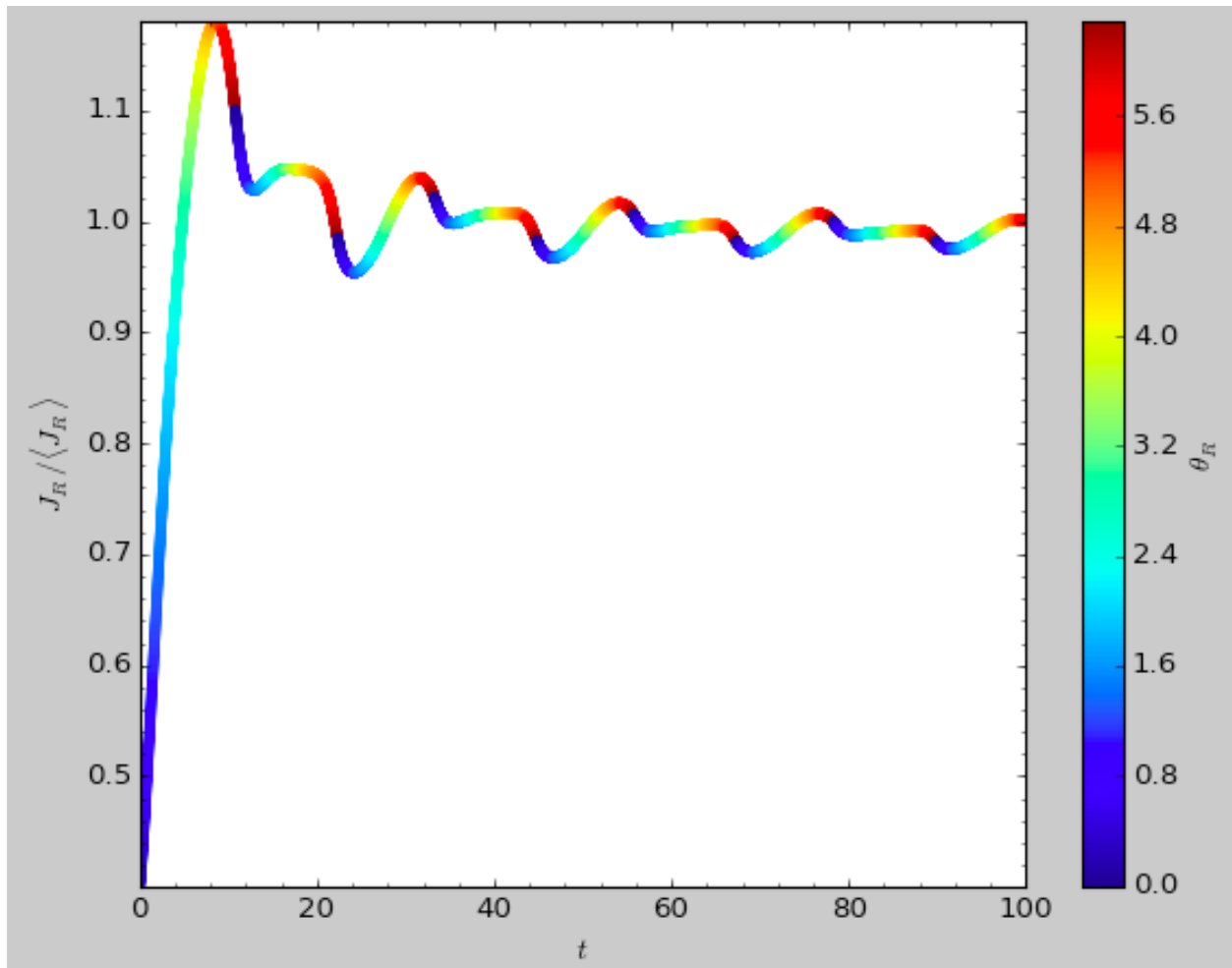
which shows slow convergence



Finding a better auxiliary potential makes convergence *much* faster and also allows the frequencies and the angles to be calculated by removing the small wiggles in the auxiliary angles vs. time (in the angle plot above, the wiggles are much larger, such that removing them is hard). The auxiliary potential used above had $b=0.8$, which shows very quick convergence and good behavior of the angles

```
>>> aAIA= actionAngleIsochroneApprox(pot=lp,b=0.8)
>>> aAIA.plot(*obs,type='jr')
```

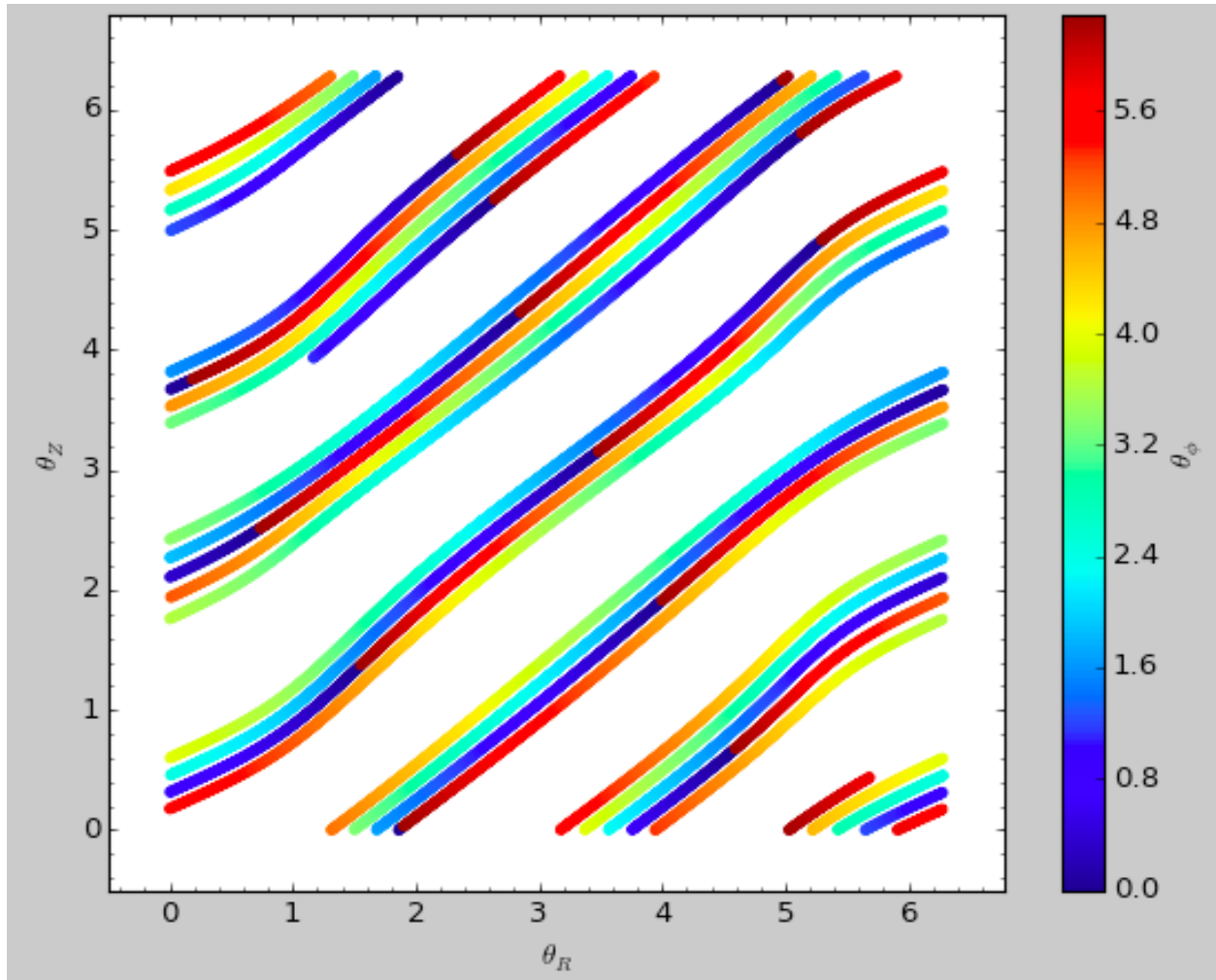
gives



and

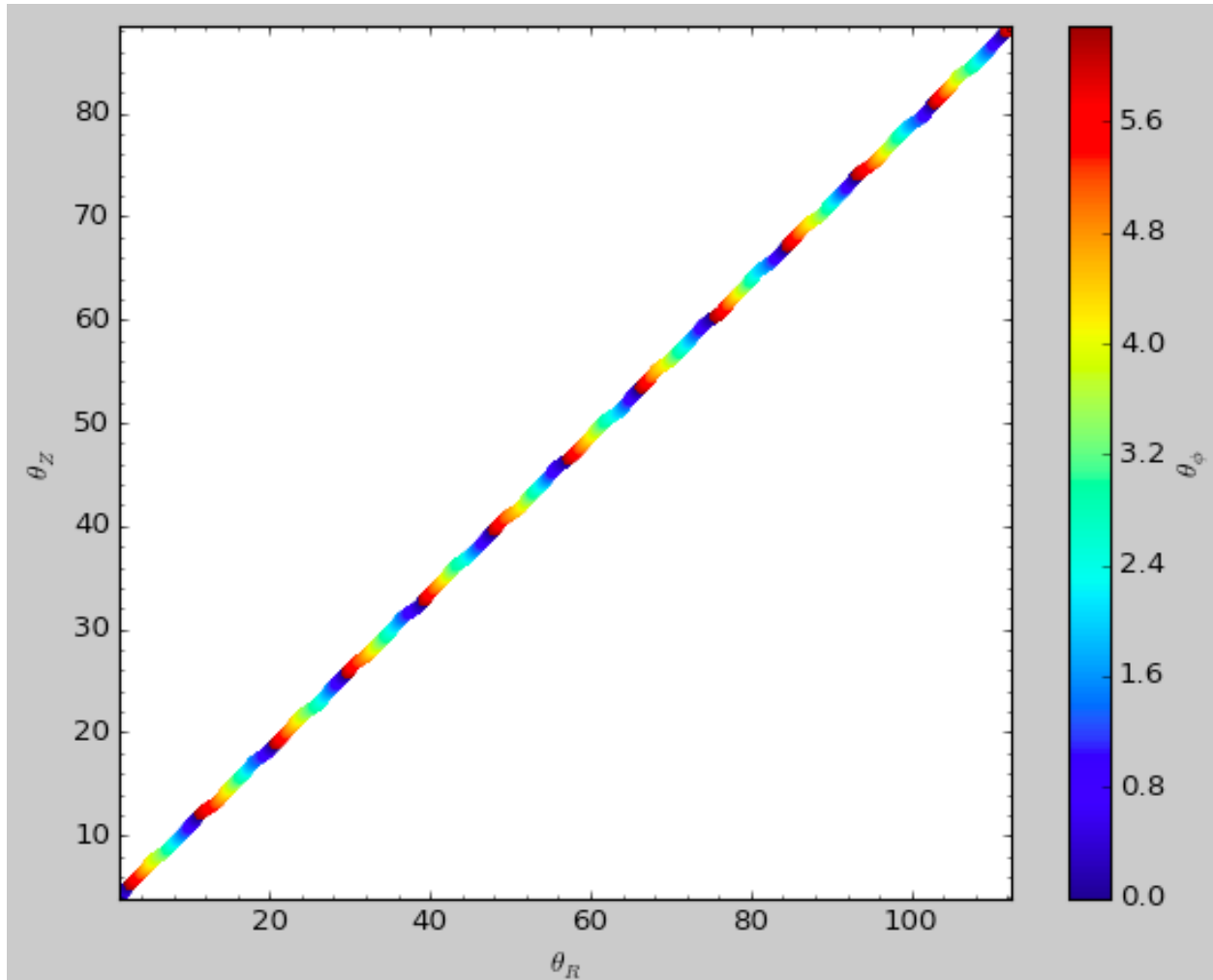
```
>>> aAIA.plot(*obs,type='araz')
```

gives



We can remove the periodic behavior from the angles, which clearly shows that they increase close-to-linear with time

```
>>> aAIA.plot(*obs, type='araz', deperiod=True)
```



We can then calculate the frequencies and the angles for this orbit as

```
>>> aAIA.actionsFreqsAngles(*obs)
# (array([ 0.16392384]),
#  array([-1.80322155]),
#  array([ 0.50999882]),
#  array([ 0.55808933]),
#  array([-0.38475753]),
#  array([ 0.42199713]),
#  array([ 0.18739688]),
#  array([ 0.3131815]),
#  array([ 2.18425661]))
```

This function takes as an argument `maxn=` the maximum n for which to remove sinusoidal wiggles. So we can raise this, for example to 4 from 3

```
>>> aAIA.actionsFreqsAngles(*obs,maxn=4)
# (array([ 0.16392384]),
#  array([-1.80322155]),
#  array([ 0.50999882]),
#  array([ 0.55808776]),
#  array([-0.38475733]),
#  array([ 0.4219968]),
```

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```
# array([ 0.18732009]),
# array([ 0.31318534]),
# array([ 2.18421296]))
```

Clearly, there is very little change, as most of the wiggles are of low n .

This technique also works for triaxial potentials, but using those requires the code to also use the azimuthal angle variable in the auxiliary potential (this is unnecessary in axisymmetric potentials as the z component of the angular momentum is conserved). We can calculate actions for triaxial potentials by specifying that `nonaxi=True`:

```
>>> aAIA(*obs,nonaxi=True)
# (array([ 0.16605011]), array([-1.80322155]), array([ 0.50704439]))
```

1.7.6 Action-angle coordinates using the TorusMapper code

All of the methods described so far allow one to compute the actions, angles, and frequencies for a given phase-space location. `galpy` also contains some support for computing the inverse transformation by using an interface to the `TorusMapper` code. Currently, this is limited to axisymmetric potentials, because the `TorusMapper` code is limited to such potentials.

The basic use of this part of `galpy` is to compute an orbit $(R, v_R, v_T, z, v_z, \phi)$ for a given torus, specified by three actions (J_R, L_Z, J_Z) and as many angles along a torus as you want. First we set up an `actionAngleTorus` object

```
>>> from galpy.actionAngle import actionAngleTorus
>>> from galpy.potential import MWPotential2014
>>> aAT= actionAngleTorus(pot=MWPotential2014)
```

To compute an orbit, we first need to compute the frequencies, which we do as follows

```
>>> jr,lz,jz= 0.1,1.1,0.2
>>> Om= aAT.Freqs(jr,lz,jz)
```

This set consists of $(\Omega_R, \Omega_\phi, \Omega_Z, \text{TMerr})$, where the last entry is the exit code of the `TorusMapper` code (will be printed as a warning when it is non-zero). Then we compute a set of angles that fall along an orbit as $\theta(t) = \theta_0 + \Omega t$ for a set of times t

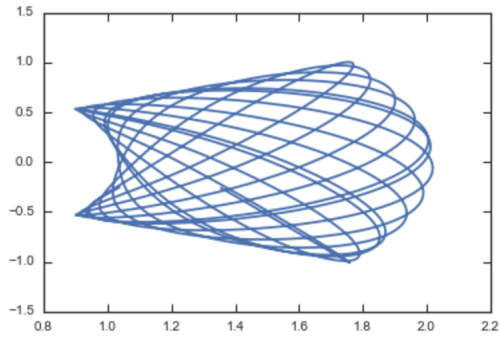
```
>>> times= numpy.linspace(0.,100.,10001)
>>> init_angle= numpy.array([1.,2.,3.])
>>> angles= numpy.tile(init_angle,(len(times),1))+Om[:3]*numpy.tile(times,(3,1)).T
```

Then we can compute the orbit by transforming the orbit in action-angle coordinates to configuration space as follows

```
>>> RvR,_,_,_,_= aAT.xvFreqs(jr,lz,jz,angles[:,0],angles[:,1],angles[:,2])
```

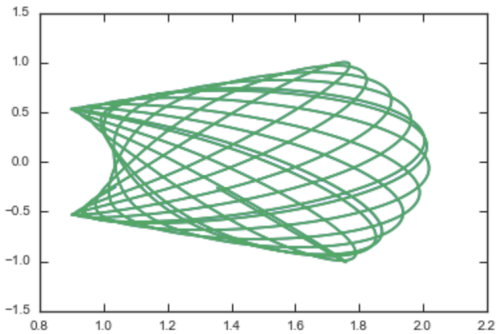
Note that the frequency is also always computed and returned by this method, because it can be obtained at zero cost. The `RvR` array has shape $(\text{ntimes}, 6)$ and the six phase-space coordinates are arranged in the usual $(R, v_R, v_T, z, v_z, \phi)$ order. The orbit in (R, Z) is then given by

```
>>> plot(RvR[:,0],RvR[:,3])
```



We can compare this to the direct numerical orbit integration. We integrate the orbit, starting at the position and velocity of the initial angle `RvR[0]`

```
>>> from galpy.orbit import Orbit
>>> orb= Orbit(RvR[0])
>>> orb.integrate(times,MWPotential2014)
>>> orb.plot(overplot=True)
```

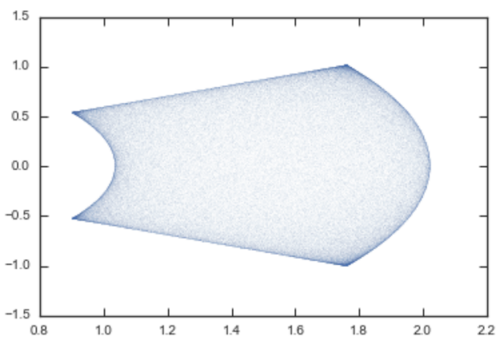


The two orbits are exactly the same.

Of course, we do not have to follow the path of an orbit to map the entire orbital torus and thus reveal the orbital building blocks of galaxies. To directly map a torus, we can do (don't worry, this doesn't take very long)

```
>>> nangles= 200001
>>> angler= numpy.random.uniform(size=nangles)*2.*numpy.pi
>>> anglep= numpy.random.uniform(size=nangles)*2.*numpy.pi
>>> anglez= numpy.random.uniform(size=nangles)*2.*numpy.pi
>>> RvR,_,_,_,_= aAT.xvFreqs(jr,lz,jz,angler,anglep,anglez)
>>> plot(RvR[:,0],RvR[:,3],',',alpha=0.02)
```

which directly shows where the orbit spends most of its time:



`actionAngleTorus` has additional methods documented on the action-angle API page for computing Hessians and Jacobians of the transformation between action-angle and configuration space coordinates.

1.7.7 Accessing action-angle coordinates for Orbit instances

While the recommended way to access the `actionAngle` routines is through the methods in the `galpy.actionAngle` modules, action-angle coordinates can also be calculated for `galpy.orbit.Orbit` instances. This is illustrated here briefly. We initialize an `Orbit` instance

```
>>> from galpy.orbit import Orbit
>>> from galpy.potential import MWPotential2014
>>> o = Orbit([1., 0.1, 1.1, 0., 0.25, 0.])
```

and we can then calculate the actions (default is to use the staeckel approximation with an automatically-estimated delta parameter, but this can be adjusted)

```
>>> o.jr(MWPotential2014), o.jp(MWPotential2014), o.jz(MWPotential2014)
# (0.018194068808944613, 1.1, 0.01540155584446606)
```

`o.jp` here gives the azimuthal action (which is the z component of the angular momentum for axisymmetric potentials). We can also use the other methods described above or adjust the parameters of the approximation (see above):

```
>>> o.jr(MWPotential2014, type='staeckel', delta=0.4), o.jp(MWPotential2014, type=
↳ 'staeckel', delta=0.4), o.jz(MWPotential2014, type='staeckel', delta=0.4)
# (0.019221672966336707, 1.1, 0.015276825017286827)
>>> o.jr(MWPotential2014, type='adiabatic'), o.jp(MWPotential2014, type='adiabatic'), o.
↳ jz(MWPotential2014, type='adiabatic')
# (0.016856430059017123, 1.1, 0.015897730620467752)
>>> o.jr(MWPotential2014, type='isochroneApprox', b=0.8), o.jp(MWPotential2014, type=
↳ 'isochroneApprox', b=0.8), o.jz(MWPotential2014, type='isochroneApprox', b=0.8)
# (0.019066091295488922, 1.1, 0.015280492319332751)
```

These two methods give very precise actions for this orbit (both are converged to about 1%) and they agree very well

```
>>> (o.jr(MWPotential2014, type='staeckel', delta=0.4) - o.jr(MWPotential2014, type=
↳ 'isochroneApprox', b=0.8)) / o.jr(MWPotential2014, type='isochroneApprox', b=0.8)
# 0.00816012408818143
>>> (o.jz(MWPotential2014, type='staeckel', delta=0.4) - o.jz(MWPotential2014, type=
↳ 'isochroneApprox', b=0.8)) / o.jz(MWPotential2014, type='isochroneApprox', b=0.8)
# 0.00023999894566772273
```

Warning: Once an action, frequency, or angle is calculated for a given type of calculation (e.g., `staeckel`), the parameters for that type are fixed in the `Orbit` instance. Call `o.resetAA()` to reset the action-angle instance used when using different parameters (i.e., different `delta=` for `staeckel` or different `b=` for `isochroneApprox`).

We can also calculate the frequencies and the angles. This requires using the `Staeckel` or `Isochrone` approximations, because frequencies and angles are currently not supported for the `adiabatic` approximation. For example, the radial frequency

```
>>> o.Or(MWPotential2014, type='staeckel', delta=0.4)
# 1.1131779637307115
>>> o.Or(MWPotential2014, type='isochroneApprox', b=0.8)
# 1.1134635974560649
```

and the radial angle

```
>>> o.wr(MWPotential2014,type='staeckel',delta=0.4)
# 0.37758086786371969
>>> o.wr(MWPotential2014,type='isochroneApprox',b=0.8)
# 0.38159809018175395
```

which again agree to 1%. We can also calculate the other frequencies, angles, as well as periods using the functions `o.Op`, `o.Oz`, `o.Wp`, `o.Wz`, `o.Tr`, `o.Tp`, `o.Tz`.

1.7.8 Example: Evidence for a Lindblad resonance in the Solar neighborhood

We can use `galpy` to calculate action-angle coordinates for a set of stars in the Solar neighborhood and look for unexplained features. For this we download the data from the Geneva-Copenhagen Survey (2009A&A...501..941H; data available at [viZier](#)). Since the velocities in this catalog are given as U,V, and W, we use the `radec` and `UVW` keywords to initialize the orbits from the raw data. For each object `ii`

```
>>> o= Orbit(vxvv[ii,:],radec=True,uvw=True,vo=220.,ro=8.)
```

We then calculate the actions and angles for each object in a flat rotation curve potential

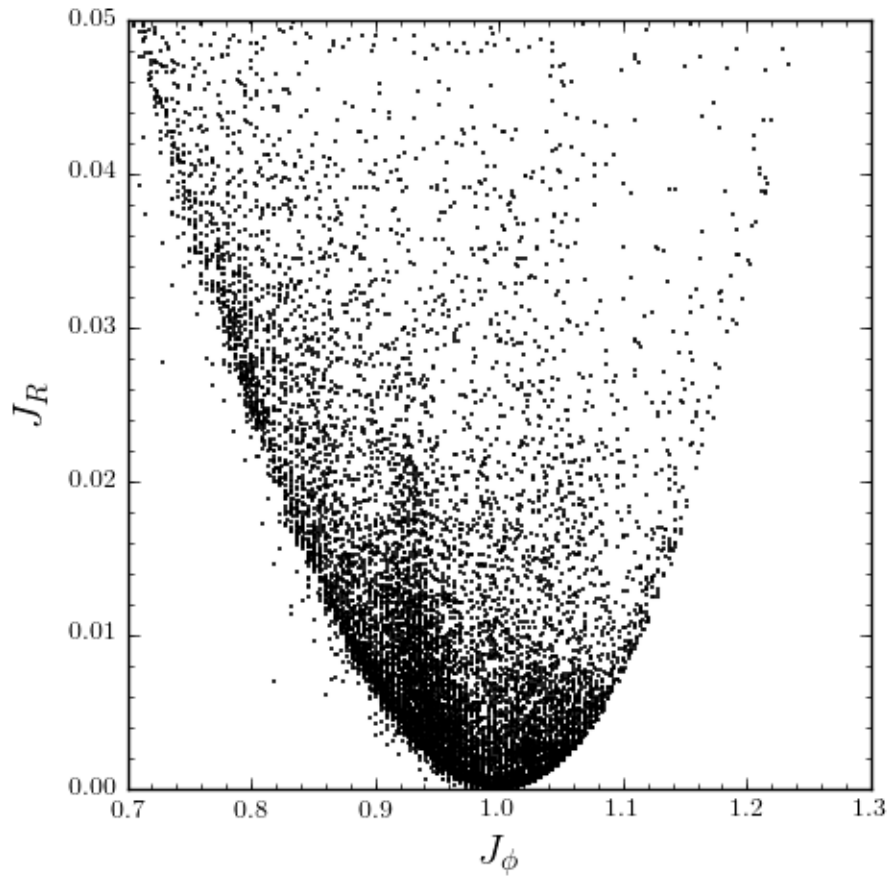
```
>>> lp= LogarithmicHaloPotential(normalize=1.)
>>> myjr[ii]= o.jr(lp)
```

etc.

Plotting the radial action versus the angular momentum

```
>>> plot.bovy_plot(myjp,myjr,'k.',ms=2.,xlabel=r'$J_{\phi}$',ylabel=r'$J_R$',
    xrange=[0.7,1.3],yrange=[0.,0.05])
```

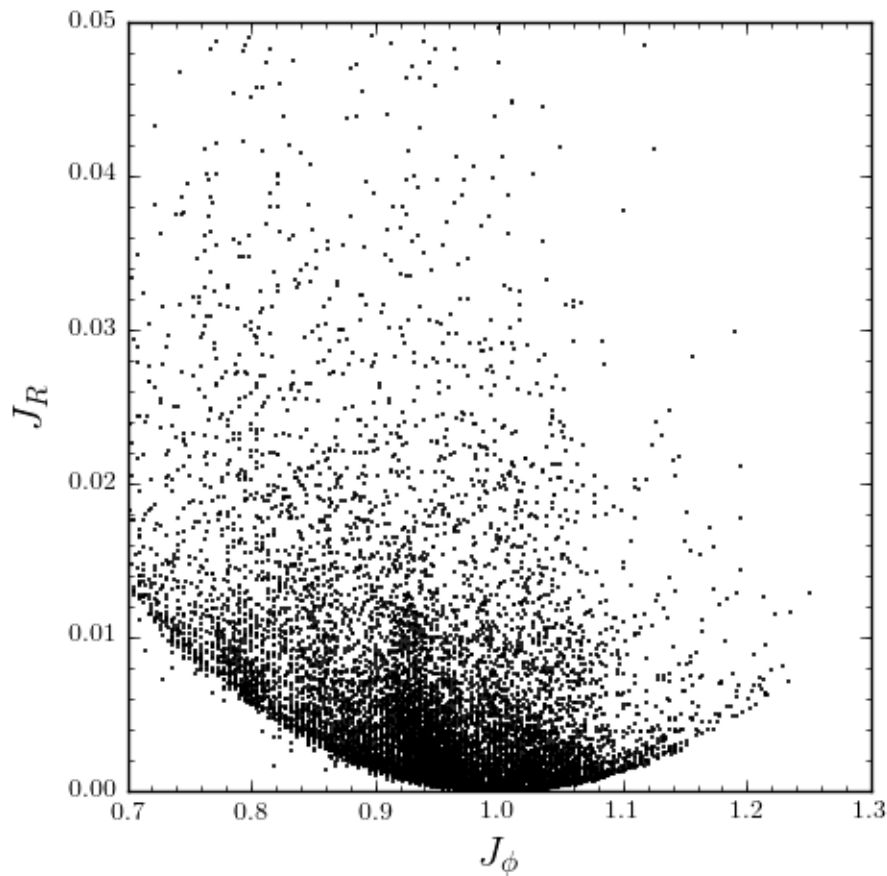
shows a feature in the distribution



If instead we use a power-law rotation curve with power-law index 1

```
>>> pp= PowerSphericalPotential(normalize=1.,alpha=-2.)
>>> myjr[ii]= o.jr(pp)
```

We find that the distribution is stretched, but the feature remains



Code for this example can be found [here](#) (note that this code uses a particular download of the GCS data set; if you use your own version, you will need to modify the part of the code that reads the data). For more information see 2010MNRAS.409..145S.

1.7.9 Example: actions in an N-body simulation

To illustrate how we can use `galpy` to calculate actions in a snapshot of an N-body simulation, we again look at the `g15784` snapshot in the `pynbody` test suite, discussed in [The potential of N-body simulations](#). Please look at that section for information on how to setup the potential of this snapshot in `galpy`. One change is that we should set `enable_c=True` in the instantiation of the `InterpSnapshotRZPotential` object

```
>>> spi= InterpSnapshotRZPotential(h1,rgrid=(numpy.log(0.01),numpy.log(20.),101),
↳logR=True,zgrid=(0.,10.,101),interpPot=True,zsym=True,enable_c=True)
>>> spi.normalize(R0=10.)
```

where we again normalize the potential to use `galpy`'s *natural units*.

We first load a pristine copy of the simulation (because the normalization above leads to some inconsistent behavior in `pynbody`)

```
>>> sc = pynbody.load('Repos/pynbody-testdata/g15784.lr.01024.gz'); hc = sc.halos();
↳hc1= hc[1]; pynbody.analysis.halo.center(hc1,mode='hyb'); pynbody.analysis.angmom.
↳faceon(hc1, cen=(0,0,0),mode='ssc'); sc.physical_units()
```

and then select particles near $R=8$ kpc by doing

```
>>> sn= pynbody.filt.BandPass('rxy','7 kpc','9 kpc')
>>> R,vR,vT,z,vz = [numpy.ascontiguousarray(hcl.s[sn][x]) for x in ('rxy','vr','vt','z'
↪ ,'vz')]
```

These have physical units, so we normalize them (the velocity normalization is the circular velocity at $R=10$ kpc, see [here](#)).

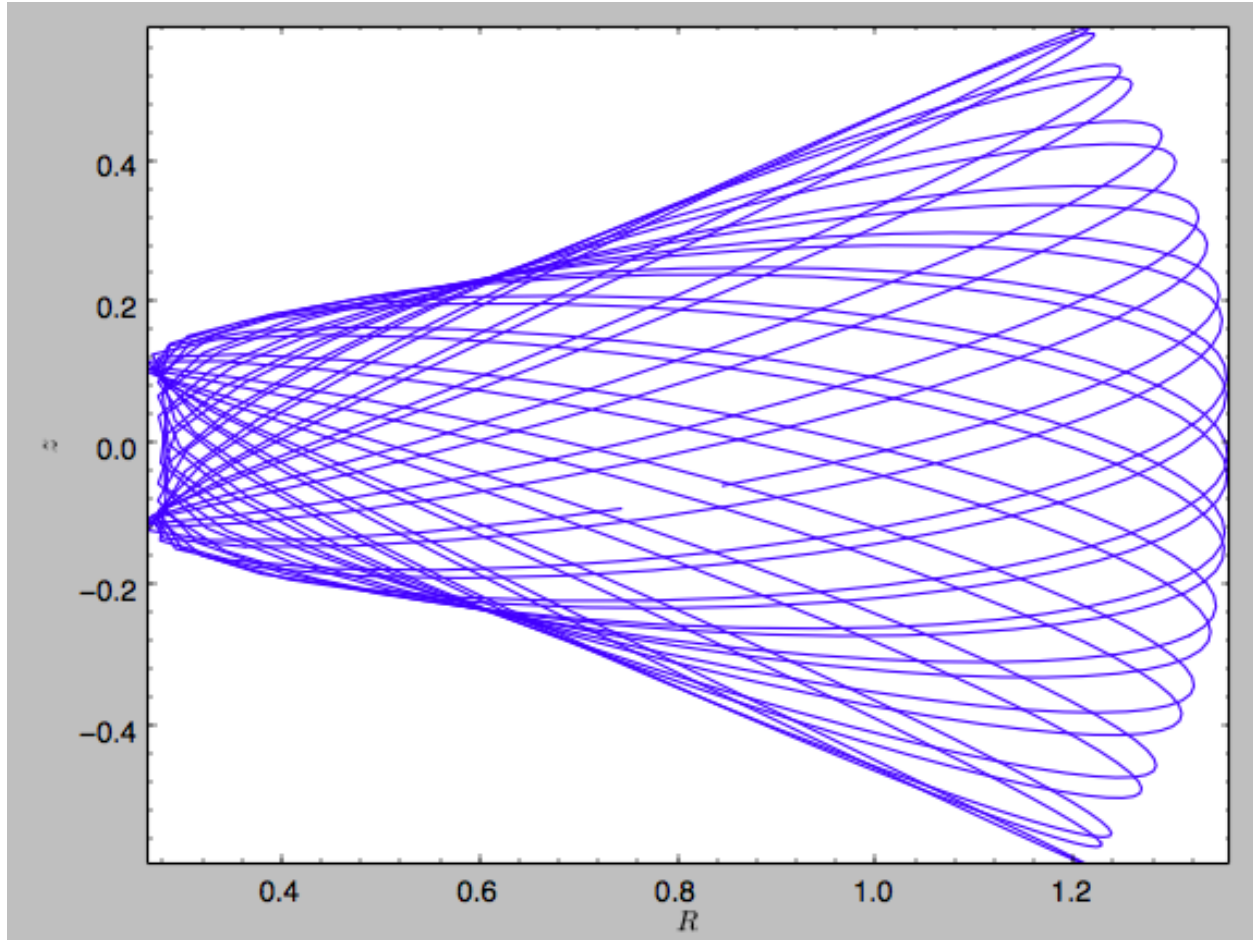
```
>>> ro, vo= 10., 294.62723076942245
>>> R/= ro
>>> z/= ro
>>> vR/= vo
>>> vT/= vo
>>> vz/= vo
```

We will calculate actions using `actionAngleStaeckel` above. We can first integrate a random orbit in this potential

```
>>> from galpy.orbit import Orbit
>>> numpy.random.seed(1)
>>> ii= numpy.random.permutation(len(R))[0]
>>> o= Orbit([R[ii],vR[ii],vT[ii],z[ii],vz[ii]])
>>> ts= numpy.linspace(0.,100.,1001)
>>> o.integrate(ts,spi)
```

This orbit looks like this

```
>>> o.plot()
```



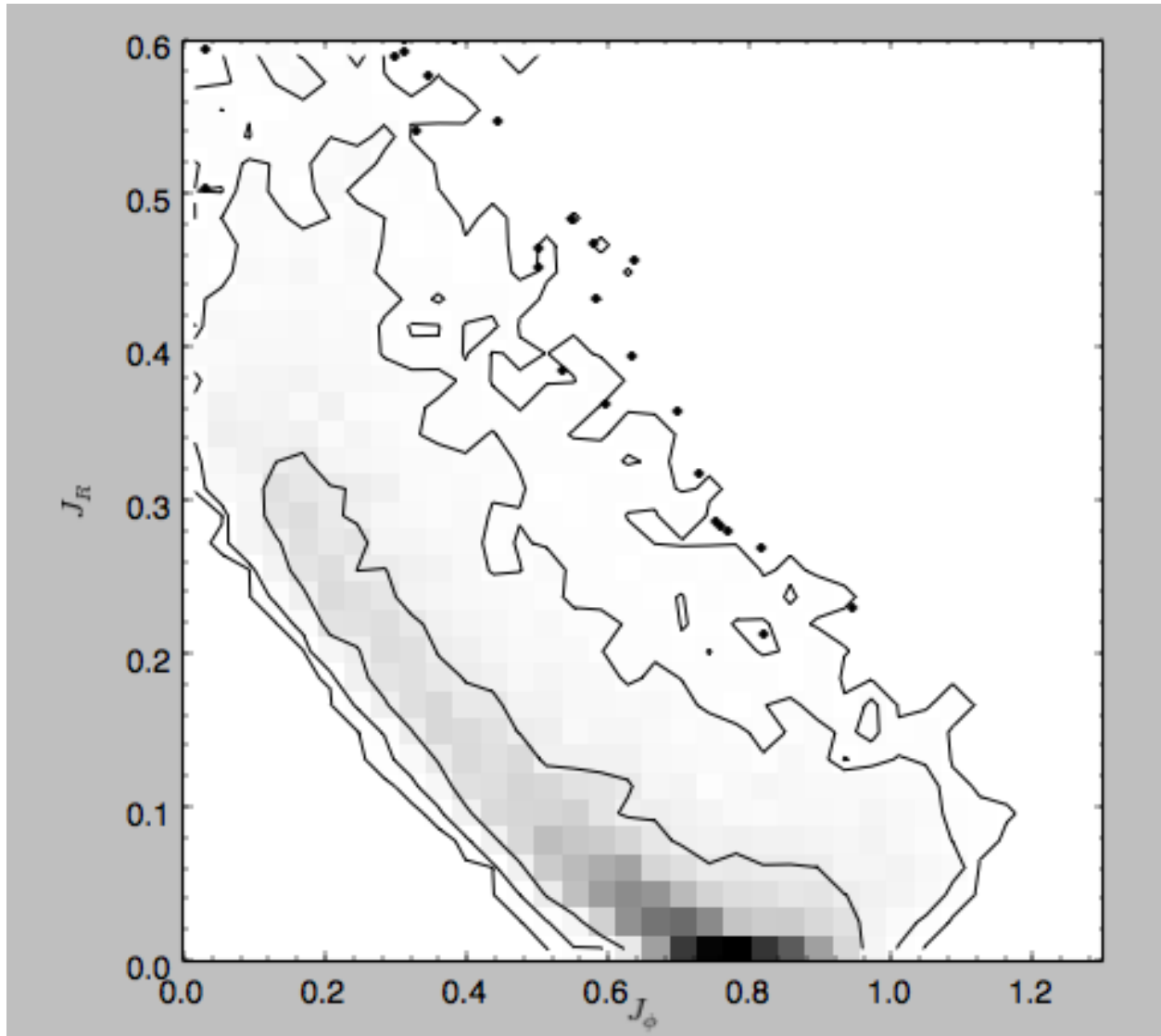
We can now calculate the actions by doing

```
>>> from galpy.actionAngle import actionAngleStaeckel
>>> aAS= actionAngleStaeckel(pot=spi,delta=0.45,c=True)
>>> jr,lz,jz= aAS(R,vR,vT,z,vz)
```

These actions are also in *natural units*; you can obtain physical units by multiplying with $ro*vo$. We can now plot these actions

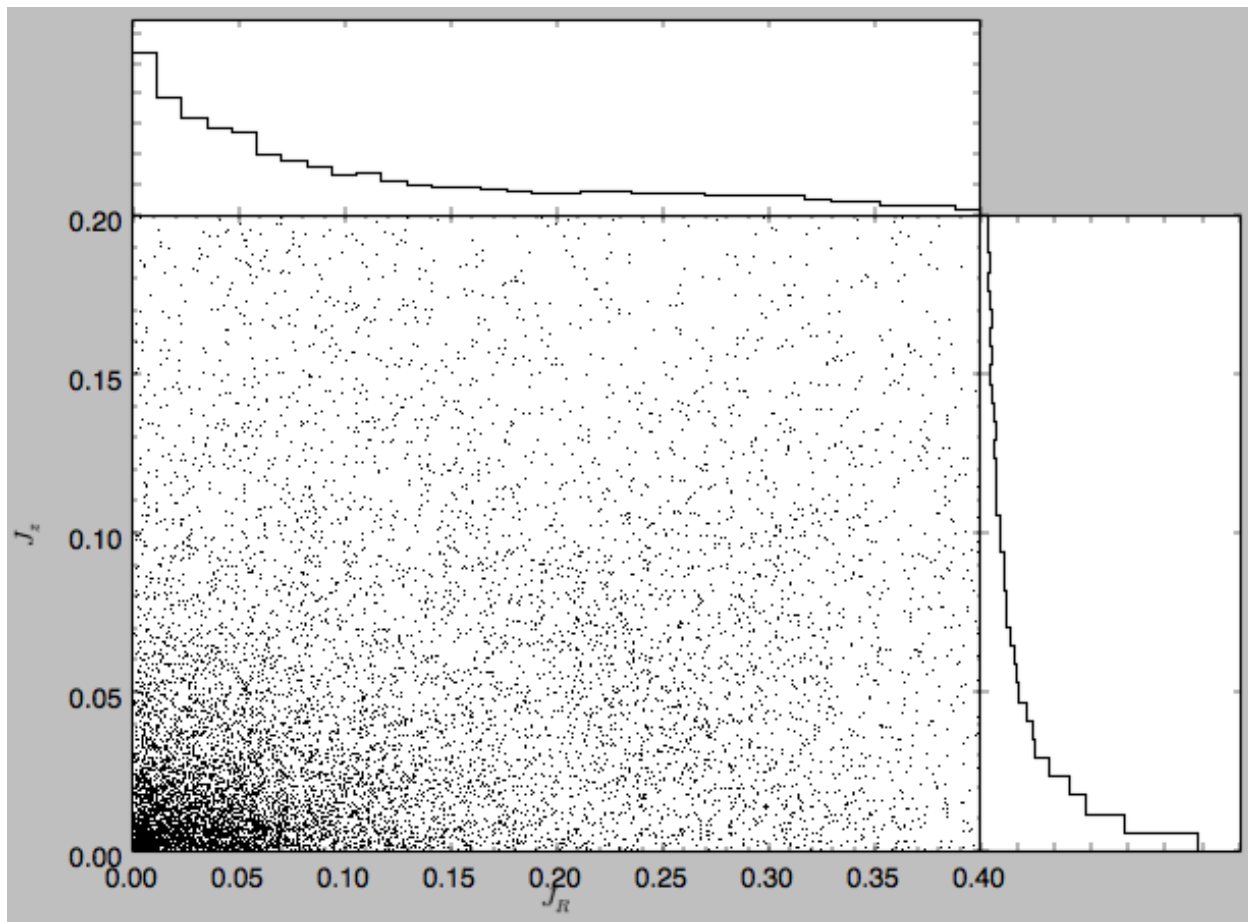
```
>>> from galpy.util import bovy_plot
>>> bovy_plot.scatterplot(lz,jr,'k.',xlabel=r'$J_\phi$',ylabel=r'$J_R$',xrange=[0.,1.
↪ 3],yrange=[0.,.6])
```

which gives



Note the similarity between this figure and the GCS figure above. The curve shape is due to the selection (low angular momentum stars can only enter the selected radial ring if they are very elliptical and therefore have large radial action) and the density gradient in angular momentum is due to the falling surface density of the disk. We can also look at the distribution of radial and vertical actions.

```
>>> bovy_plot.bovy_plot(jr, jz, 'k', xlabel=r'$J_R$', ylabel=r'$J_z$', xrange=[0., .4],
↳ yrange=[0., 0.2], onedhists=True)
```



With the other methods in the `actionAngle` module we can also calculate frequencies and angles.

1.8 Three-dimensional disk distribution functions

`galpy` contains a fully three-dimensional disk distribution: `galpy.df.quasiisothermaldf`, which is an approximately isothermal distribution function expressed in terms of action-angle variables (see [2010MNRAS.401.2318B](#) and [2011MNRAS.413.1889B](#)). Recent research shows that this distribution function provides a good model for the DF of mono-abundance sub-populations (MAPs) of the Milky Way disk (see [2013MNRAS.434..652T](#) and [2013ApJ...779..115B](#)). This distribution function family requires action-angle coordinates to evaluate the DF, so `galpy.df.quasiisothermaldf` makes heavy use of the routines in `galpy.actionAngle` (in particular those in `galpy.actionAngleAdiabatic` and `galpy.actionAngle.actionAngleStaeckel`).

1.8.1 Setting up the DF and basic properties

The quasi-isothermal DF is defined by a gravitational potential and a set of parameters describing the radial surface-density profile and the radial and vertical velocity dispersion as a function of radius. In addition, we have to provide an instance of a `galpy.actionAngle` class to calculate the actions for a given position and velocity. For example, for a `galpy.potential.MWPotential2014` potential using the adiabatic approximation for the actions, we import and define the following

```
>>> from galpy.potential import MWPotential2014
>>> from galpy.actionAngle import actionAngleAdiabatic
>>> from galpy.df import quasiisothermaldf
>>> aA= actionAngleAdiabatic(pot=MWPotential2014,c=True)
```

and then setup the `quasiisothermaldf` instance

```
>>> qdf= quasiisothermaldf(1./3.,0.2,0.1,1.,1.,pot=MWPotential2014,aA=aA,
↪cutcounter=True)
```

which sets up a DF instance with a radial scale length of $R_0/3$, a local radial and vertical velocity dispersion of $0.2 V_c(R_0)$ and $0.1 V_c(R_0)$, respectively, and a radial scale lengths of the velocity dispersions of R_0 . `cutcounter=True` specifies that counter-rotating stars are explicitly excluded (normally these are just exponentially suppressed). As for the two-dimensional disk DFs, these parameters are merely input (or target) parameters; the true density and velocity dispersion profiles calculated by evaluating the relevant moments of the DF (see below) are not exactly exponential and have scale lengths and local normalizations that deviate slightly from these input parameters. We can estimate the DF's actual radial scale length near R_0 as

```
>>> qdf.estimate_hr(1.)
# 0.32908034635647182
```

which is quite close to the input value of $1/3$. Similarly, we can estimate the scale lengths of the dispersions

```
>>> qdf.estimate_hsr(1.)
# 1.1913935820372923
>>> qdf.estimate_hsz(1.)
# 1.0506918075359255
```

The vertical profile is fully specified by the velocity dispersions and radial density / dispersion profiles under the assumption of dynamical equilibrium. We can estimate the scale height of this DF at a given radius and height as follows

```
>>> qdf.estimate_hz(1.,0.125)
# 0.021389597757156088
```

Near the mid-plane this vertical scale height becomes very large because the vertical profile flattens, e.g.,

```
>>> qdf.estimate_hz(1.,0.125/100.)
# 1.006386030587223
```

or even

```
>>> qdf.estimate_hz(1.,0.)
# 187649.98447377066
```

which is basically infinity.

1.8.2 Evaluating moments

We can evaluate various moments of the DF giving the density, mean velocities, and velocity dispersions. For example, the mean radial velocity is again everywhere zero because the potential and the DF are axisymmetric

```
>>> qdf.meanvR(1.,0.)
# 0.0
```

Likewise, the mean vertical velocity is everywhere zero

```
>>> qdf.meanvz(1.,0.)  
# 0.0
```

The mean rotational velocity has a more interesting dependence on position. Near the plane, this is the same as that calculated for a similar two-dimensional disk DF (see [Evaluating moments of the DF](#))

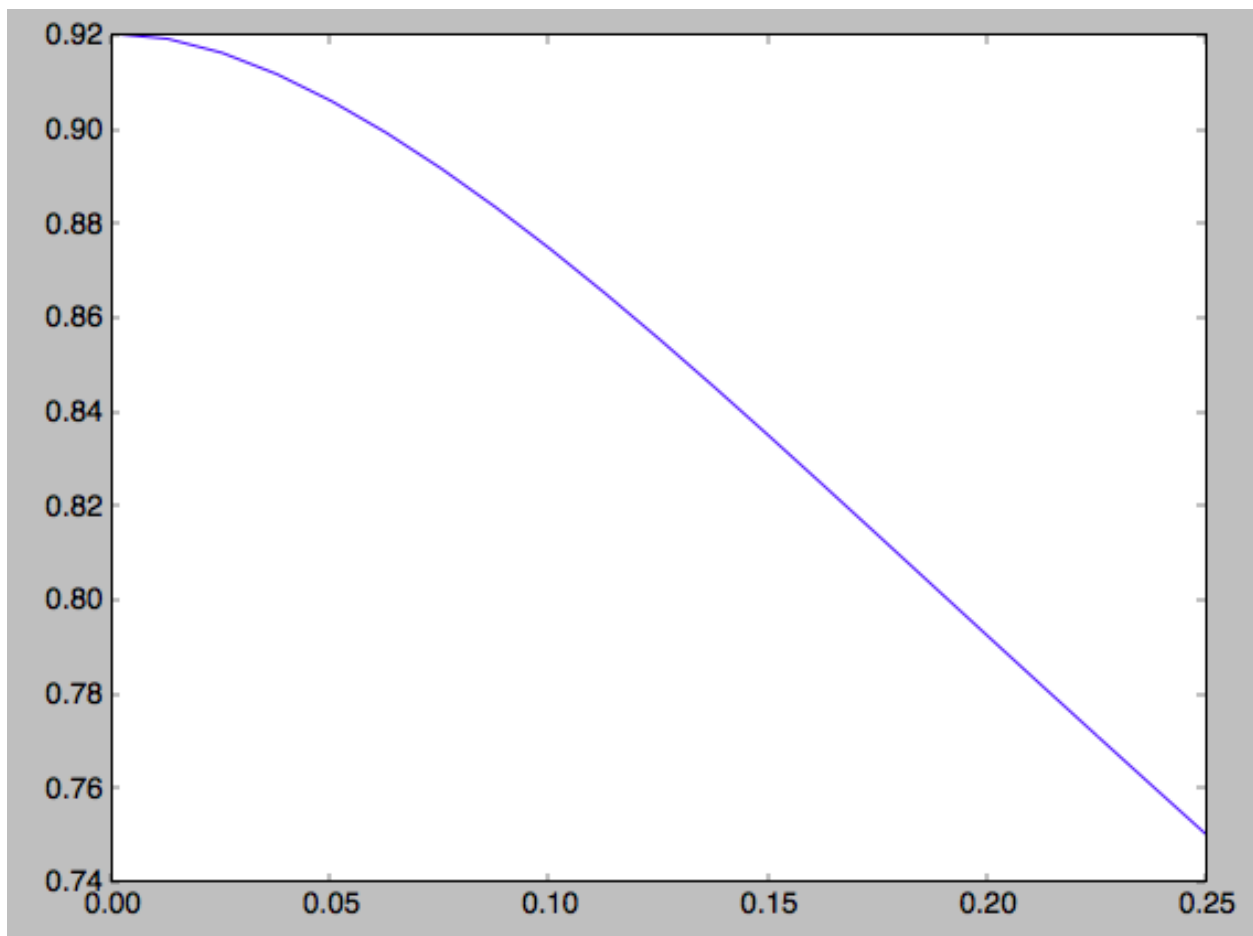
```
>>> qdf.meanvT(1.,0.)  
# 0.91988346380781227
```

However, this value decreases as one moves further from the plane. The `quasiisothermaldf` allows us to calculate the average rotational velocity as a function of height above the plane. For example,

```
>>> zs= numpy.linspace(0.,0.25,21)  
>>> mvts= numpy.array([qdf.meanvT(1.,z) for z in zs])
```

which gives

```
>>> plot(zs,mvts)
```



We can also calculate the second moments of the DF. We can check whether the radial and velocity dispersions at R_0 are close to their input values

```
>>> numpy.sqrt(qdf.sigmaR2(1.,0.))  
# 0.20807112565801389
```

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```
>>> numpy.sqrt(qdf.sigmaz2(1.,0.))
# 0.090453510526130904
```

and they are pretty close. We can also calculate the mixed R and z moment, for example,

```
>>> qdf.sigmaRz(1.,0.125)
# 0.0
```

or expressed as an angle (the *tilt of the velocity ellipsoid*)

```
>>> qdf.tilt(1.,0.125)
# 0.0
```

This tilt is zero because we are using the adiabatic approximation. As this approximation assumes that the motions in the plane are decoupled from the vertical motions of stars, the mixed moment is zero. However, this approximation is invalid for stars that go far above the plane. By using the Staeckel approximation to calculate the actions, we can model this coupling better. Setting up a `quasiisothermaldf` instance with the Staeckel approximation

```
>>> from galpy.actionAngle import actionAngleStaeckel
>>> aAS= actionAngleStaeckel(pot=MWPotential2014,delta=0.45,c=True)
>>> qdfS= quasiisothermaldf(1./3.,0.2,0.1,1.,1.,pot=MWPotential2014,aA=aAS,
    ↪ cutcounter=True)
```

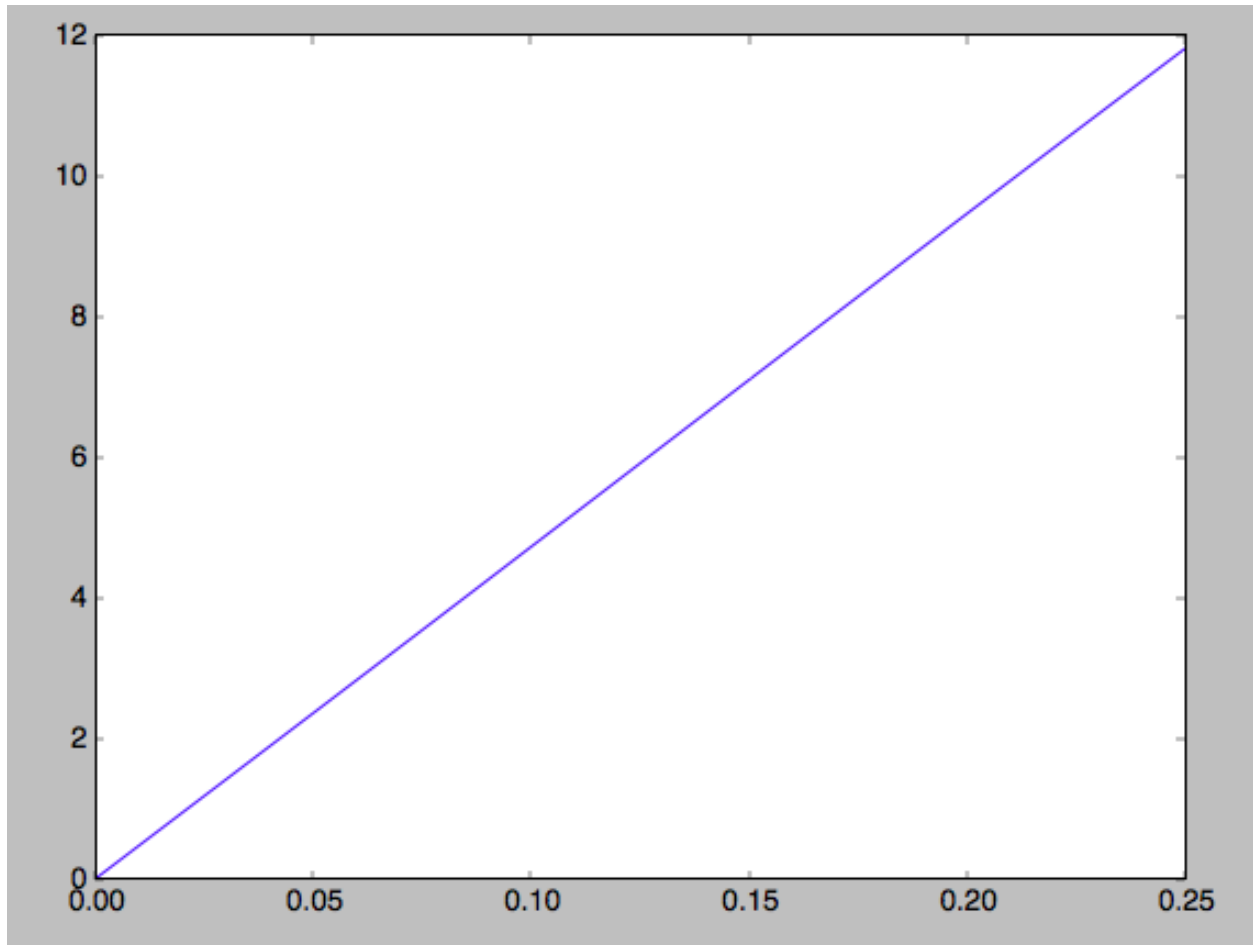
we can similarly calculate the tilt

```
>>> qdfS.tilt(1.,0.125)
# 0.10314272868452541
```

or about 5 degrees (the returned value has units of rad). As a function of height, we find

```
>>> tilts= numpy.array([qdfS.tilt(1.,z) for z in zs])
>>> plot(zs,tilts*180./numpy.pi)
```

which gives



We can also calculate the density and surface density (the zero-th velocity moments). For example, the vertical density

```
>>> densz= numpy.array([qdf.density(1.,z) for z in zs])
```

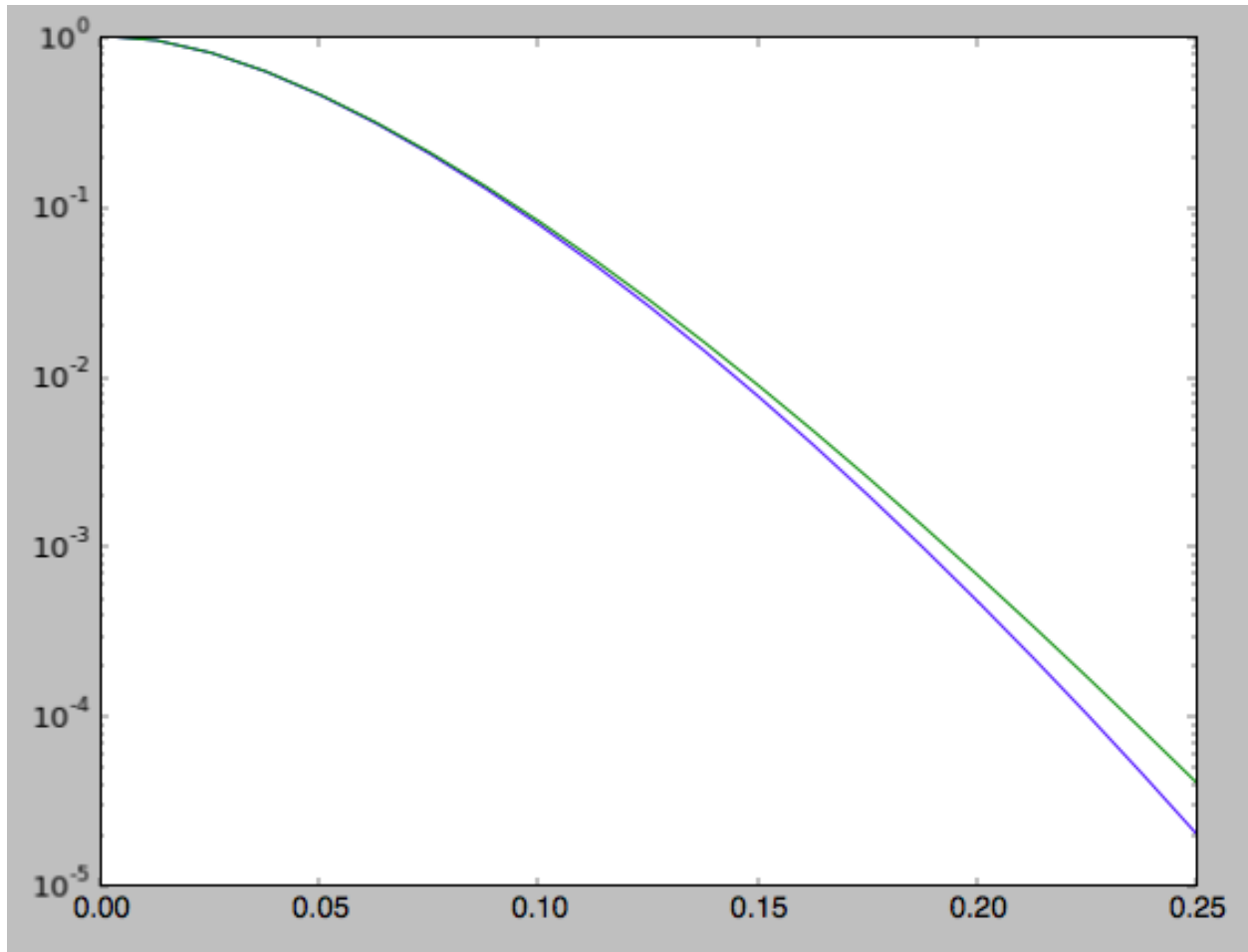
and

```
>>> denszS= numpy.array([qdfS.density(1.,z) for z in zs])
```

We can compare the vertical profiles calculated using the adiabatic and Staeckel action-angle approximations

```
>>> semilogy(zs,densz/densz[0])
>>> semilogy(zs,denszS/denszS[0])
```

which gives



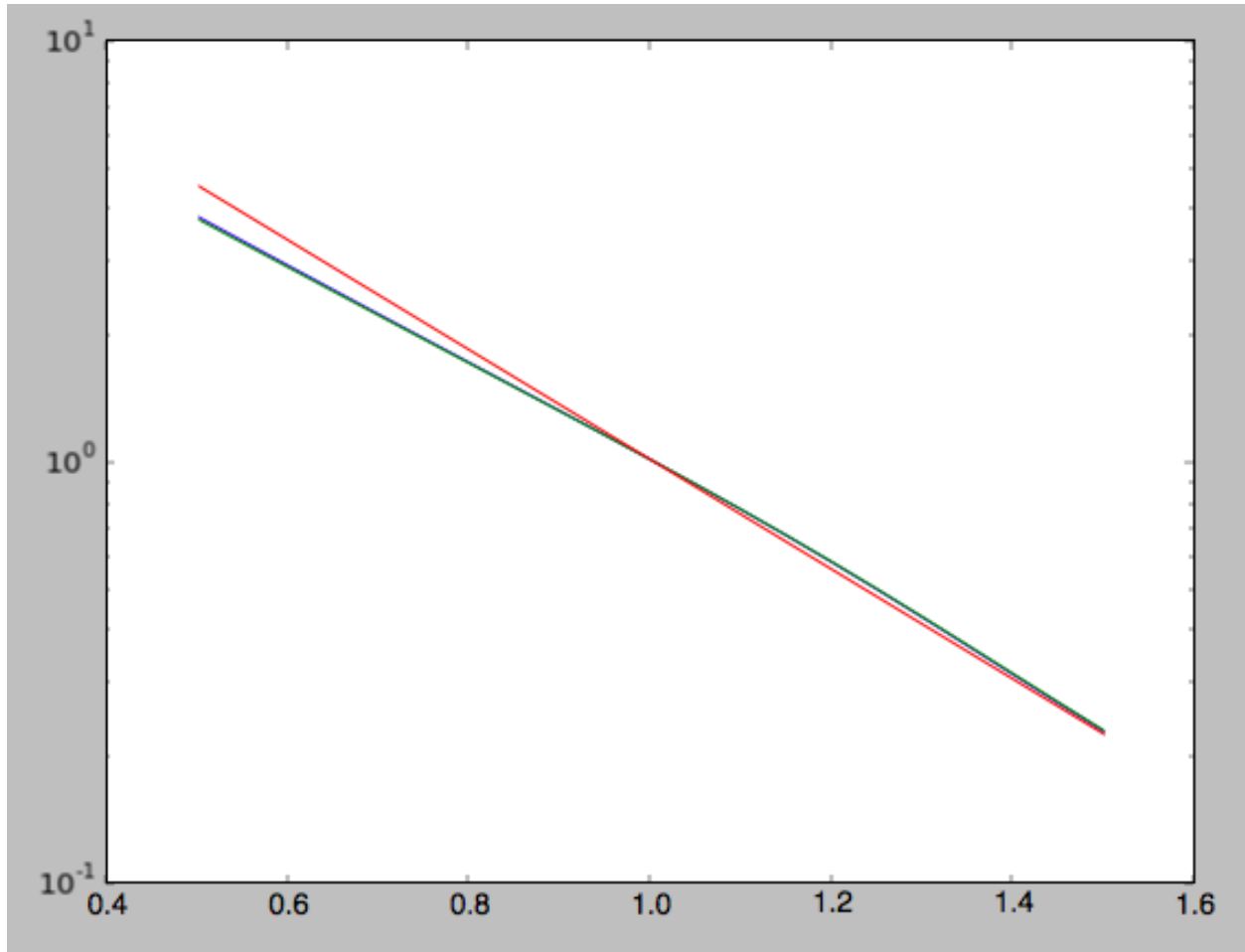
Similarly, we can calculate the radial profile of the surface density

```
>>> rs= numpy.linspace(0.5,1.5,21)
>>> surfr= numpy.array([qdf.surface_mass_z(r) for r in rs])
>>> surfrS= numpy.array([qdfS.surface_mass_z(r) for r in rs])
```

and compare them with each other and an exponential with scale length 1/3

```
>>> semilogy(rs,surfr/surfr[10])
>>> semilogy(rs,surfrS/surfrS[10])
>>> semilogy(rs,numpy.exp(-(rs-1.)/(1./3.)))
```

which gives



The two radial profiles are almost indistinguishable and are very close, if somewhat shallower, than the pure exponential profile.

General velocity moments, including all higher order moments, are implemented in `quasiisothermaldf.vmomentdensity`.

1.8.3 Evaluating and sampling the full probability distribution function

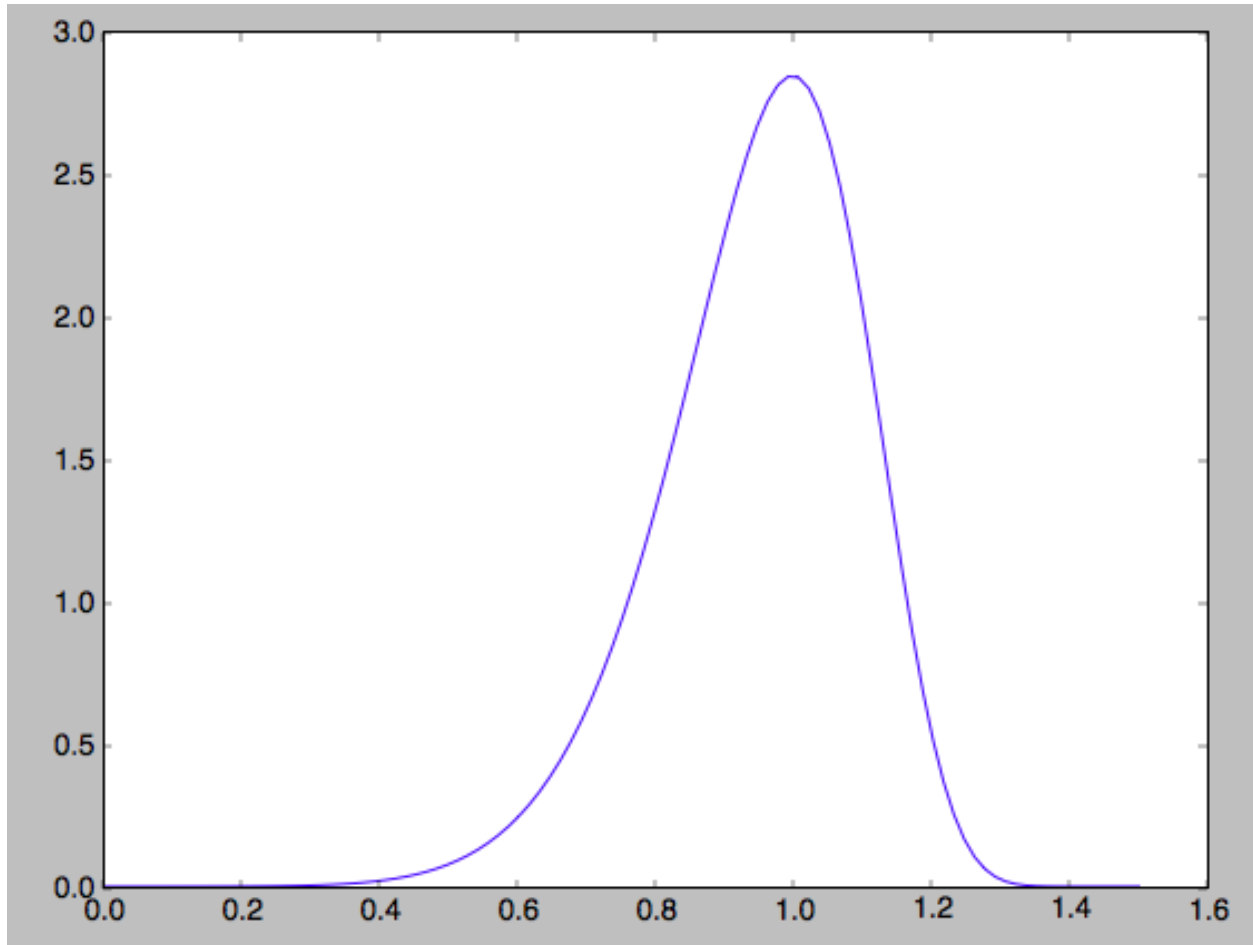
We can evaluate the distribution itself by calling the object, e.g.,

```
>>> qdf(1., 0.1, 1.1, 0.1, 0.) #input: R, vR, vT, z, vz
# array([ 16.86790643])
```

or as a function of rotational velocity, for example in the mid-plane

```
>>> vts= numpy.linspace(0., 1.5, 101)
>>> pvt= numpy.array([qdfS(1., 0., vt, 0., 0.) for vt in vts])
>>> plot(vts, pvt/numpy.sum(pvt)/(vts[1]-vts[0]))
```

which gives



This is, however, not the true distribution of rotational velocities at $R=0$ and $z=0$, because it is conditioned on zero radial and vertical velocities. We can calculate the distribution of rotational velocities marginalized over the radial and vertical velocities as

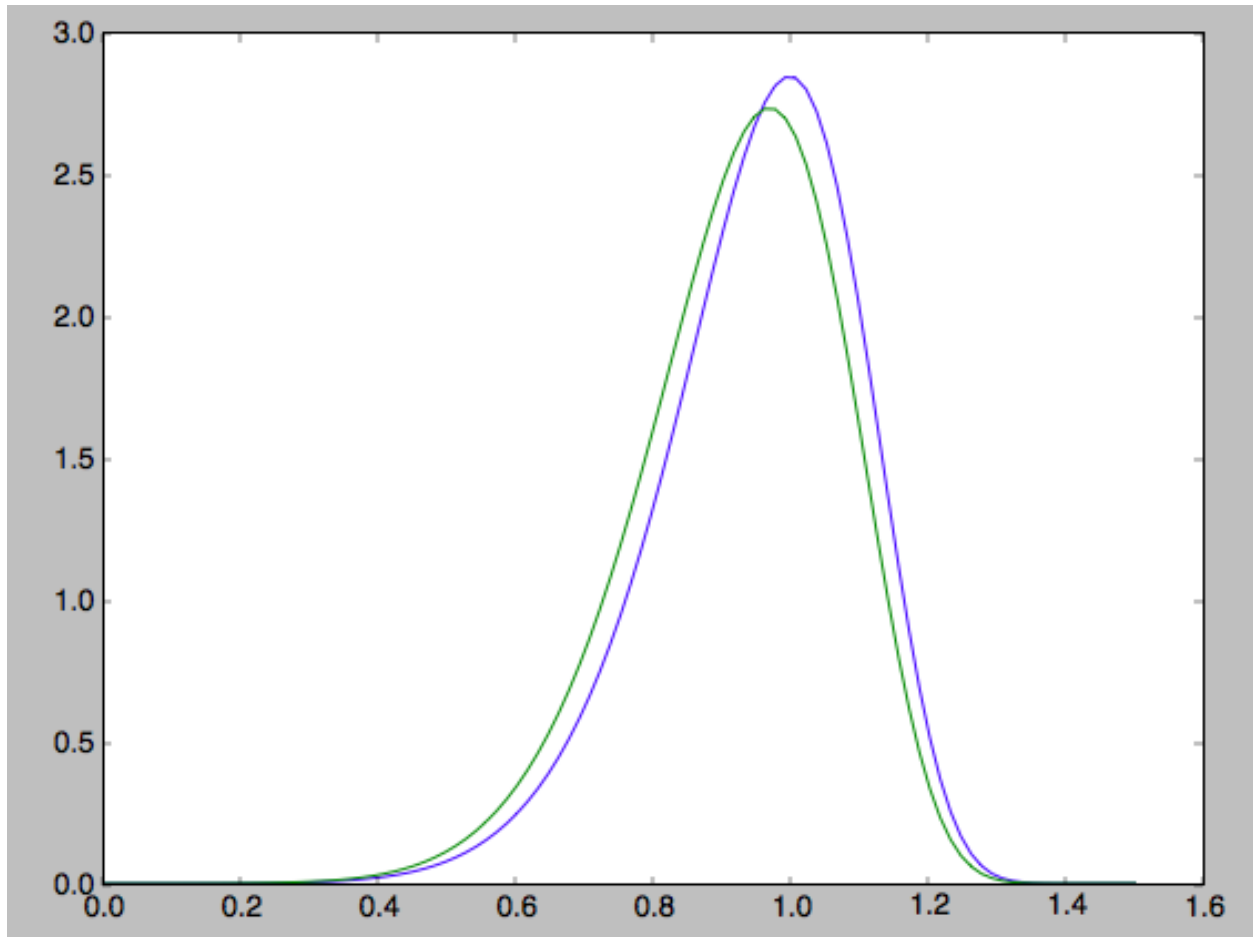
```
>>> qdfS.pvT(1.,1.,0.) #input vT,R,z
# 14.677231196899195
```

or as a function of rotational velocity

```
>>> pvt= numpy.array([qdfS.pvT(vt,1.,0.) for vt in vts])
```

overplotting this over the previous distribution gives

```
>>> plot(vts,pvt/numpy.sum(pvt)/(vts[1]-vts[0]))
```

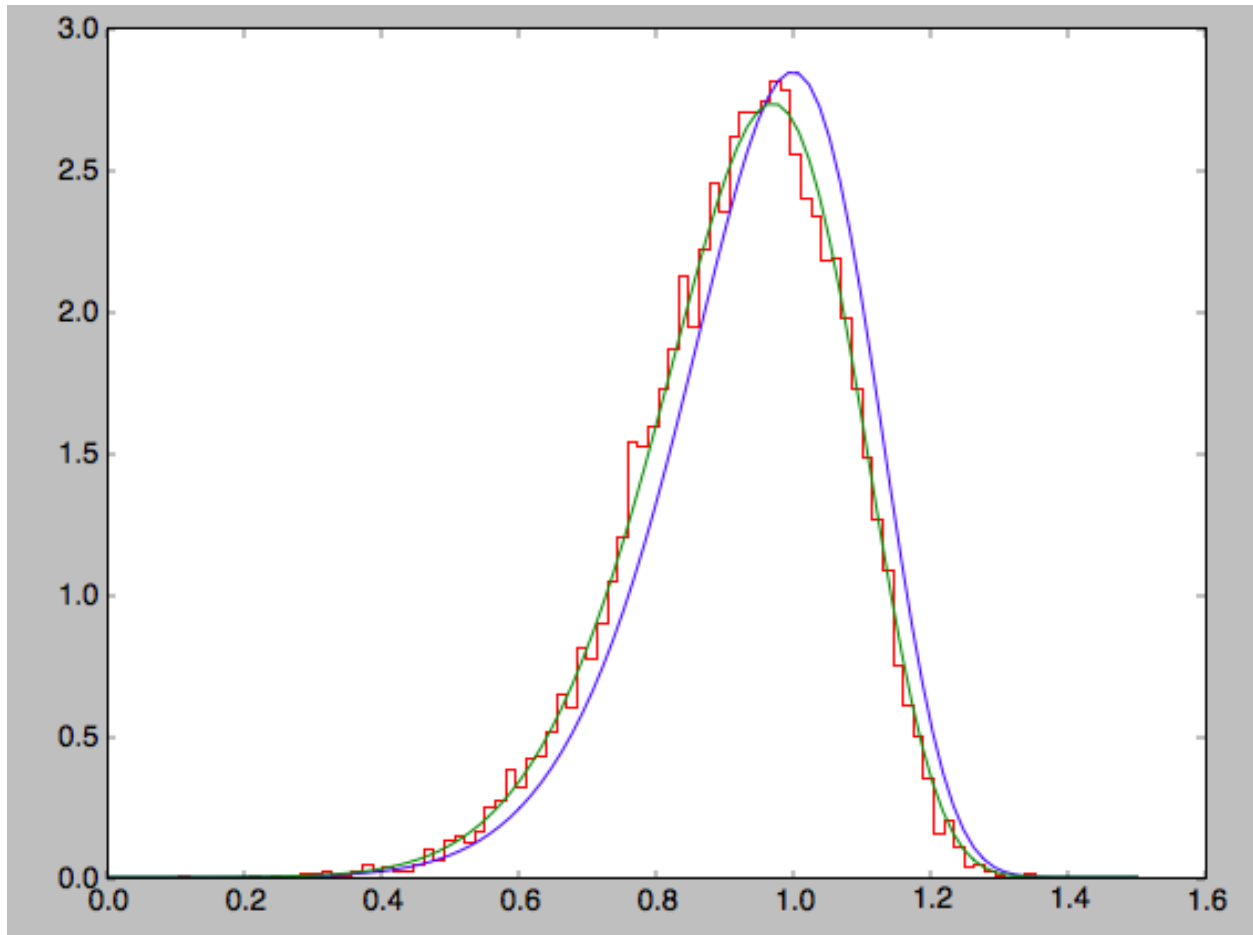


which is slightly different from the conditioned distribution. Similarly, we can calculate marginalized velocity probabilities p_{vR} , p_{vz} , p_{vRvT} , p_{vRvz} , and p_{vTvz} . These are all multiplied with the density, such that marginalizing these over the remaining velocity component results in the density.

We can sample velocities at a given location using `quasiisothermaldf.sampleV` (there is currently no support for sampling locations from the density profile, although that is rather trivial):

```
>>> vs= qdfS.sampleV(1.,0.,n=10000)
>>> hist(vs[:,1],normed=True,histtype='step',bins=101,range=[0.,1.5])
```

gives



which shows very good agreement with the green (marginalized over v_R and v_z) curve (as it should).

2.1 Dynamical modeling of tidal streams

`galpy` contains tools to model the dynamics of tidal streams, making extensive use of action-angle variables. As an example, we can model the dynamics of the following tidal stream (that of Bovy 2014; 2014ApJ...795...95B). This movie shows the disruption of a cluster on a GD-1-like orbit around the Milky Way:

The blue line is the orbit of the progenitor cluster and the black points are cluster members. The disruption is shown in an approximate orbital plane and the movie is comoving with the progenitor cluster.

Streams can be represented by simple dynamical models in action-angle coordinates. In action-angle coordinates, stream members are stripped from the progenitor cluster onto orbits specified by a set of actions (J_R, J_ϕ, J_Z), which remain constant after the stars have been stripped. This is shown in the following movie, which shows the generation of the stream in action space

The color-coding gives the angular momentum J_ϕ and the black dot shows the progenitor orbit. These actions were calculated using `galpy.actionAngle.actionAngleIsochroneApprox`. The points move slightly because of small errors in the action calculation (these are correlated, so the cloud of points moves coherently because of calculation errors). The same movie that also shows the actions of stars in the cluster can be found [here](#). This shows that the actions of stars in the cluster are not conserved (because the self-gravity of the cluster is important), but that the actions of stream members freeze once they are stripped. The angle difference between stars in a stream and the progenitor increases linearly with time, which is shown in the following movie:

where the radial and vertical angle difference with respect to the progenitor (co-moving at $(\theta_R, \theta_\phi, \theta_Z) = (\pi, \pi, \pi)$) is shown for each snapshot (the color-coding gives θ_ϕ).

One last movie provides further insight in how a stream evolves over time. The following movie shows the evolution of the stream in the two dimensional plane of frequency and angle along the stream (that is, both are projections of the three dimensional frequencies or angles onto the angle direction along the stream). The points are color-coded by the time at which they were removed from the progenitor cluster.

It is clear that disruption happens in bursts (at pericenter passages) and that the initial frequency distribution at the time of removal does not change (much) with time. However, stars removed at larger frequency difference move away from the cluster faster, such that the end of the stream is primarily made up of stars with large frequency differences

with respect to the progenitor. This leads to a gradient in the typical orbit in the stream, and the stream is on average *not* on a single orbit.

2.1.1 Modeling streams in galpy

In galpy we can model streams using the tools in `galpy.df.streamdf`. We setup a `streamdf` instance by specifying the host gravitational potential `pot=`, an `actionAngle` method (typically `galpy.actionAngle.actionAngleIsochroneApprox`), a `galpy.orbit.Orbit` instance with the position of the progenitor, a parameter related to the velocity dispersion of the progenitor, and the time since disruption began. We first import all of the necessary modules

```
>>> from galpy.df import streamdf
>>> from galpy.orbit import Orbit
>>> from galpy.potential import LogarithmicHaloPotential
>>> from galpy.actionAngle import actionAngleIsochroneApprox
>>> from galpy.util import bovy_conversion #for unit conversions
```

setup the potential and `actionAngle` instances

```
>>> lp= LogarithmicHaloPotential(normalize=1.,q=0.9)
>>> aAI= actionAngleIsochroneApprox(pot=lp,b=0.8)
```

define a progenitor `Orbit` instance

```
>>> obs= Orbit([1.56148083,0.35081535,-1.15481504,0.88719443,-0.47713334,0.12019596])
```

and instantiate the `streamdf` model

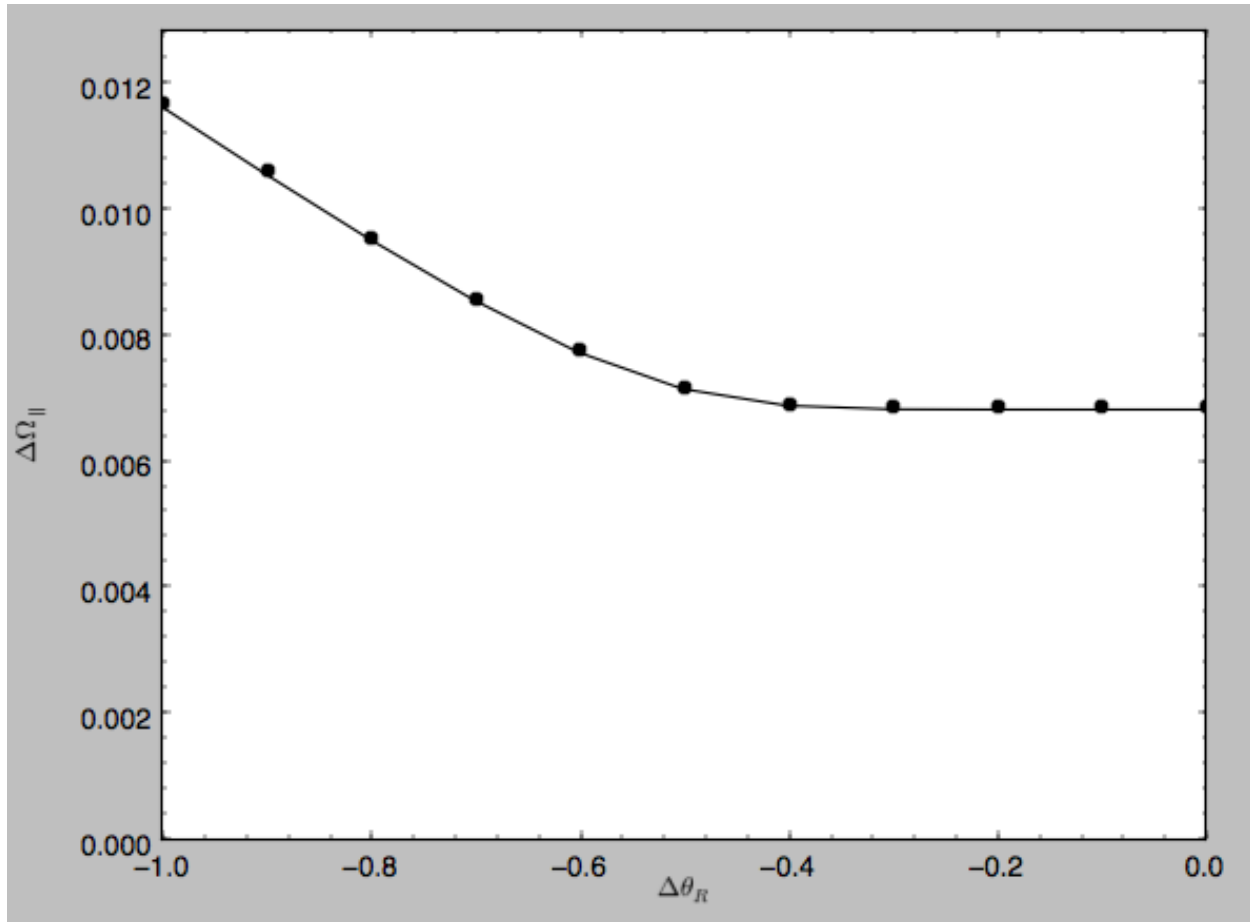
```
>>> sigv= 0.365 #km/s
>>> sdf= streamdf(sigv/220.,progenitor=obs,pot=lp,aA=aAI,leading=True,nTrackChunks=11,
↳ tdisrupt=4.5/bovy_conversion.time_in_Gyr(220.,8.))
```

for a leading stream. This runs in about half a minute on a 2011 Macbook Air.

Bovy (2014) discusses how the calculation of the track needs to be iterated for potentials where there is a large offset between the track and a single orbit. One can increase the default number of iterations by specifying `nTrackIterations=` in the `streamdf` initialization (the default is set based on the angle between the track's frequency vector and the progenitor orbit's frequency vector; you can access the number of iterations used as `sdf.nTrackIterations`). To check whether the track is calculated accurately, one can use the following

```
>>> sdf.plotCompareTrackAAModel()
```

which in this case gives



This displays the stream model's track in frequency offset (y axis) versus angle offset (x axis) as the solid line; this is the track that the model should have if it is calculated correctly. The points are the frequency and angle offset calculated from the calculated track's (\mathbf{x}, \mathbf{v}) . For a properly computed track these should line up, as they do in this figure. If they do not line up, increasing `nTrackIterations` is necessary.

We can calculate some simple properties of the stream, such as the ratio of the largest and second-to-largest eigenvalue of the Hessian $\partial\Omega/\partial\mathbf{J}$

```
>>> sdf.freqEigvalRatio(isotropic=True)
# 34.450028399901434
```

or the model's ratio of the largest and second-to-largest eigenvalue of the model frequency variance matrix

```
>>> sdf.freqEigvalRatio()
# 29.625538344985291
```

The fact that this ratio is so large means that an approximately one dimensional stream will form.

Similarly, we can calculate the angle between the frequency vector of the progenitor and of the model mean frequency vector

```
>>> sdf.misalignment()
# 0.0086441947505973005
```

which returns this angle in radians. We can also calculate the angle between the frequency vector of the progenitor and the principal eigenvector of $\partial\Omega/\partial\mathbf{J}$

```
>>> sdf.misalignment(isotropic=True)
# 0.02238411611147997
```

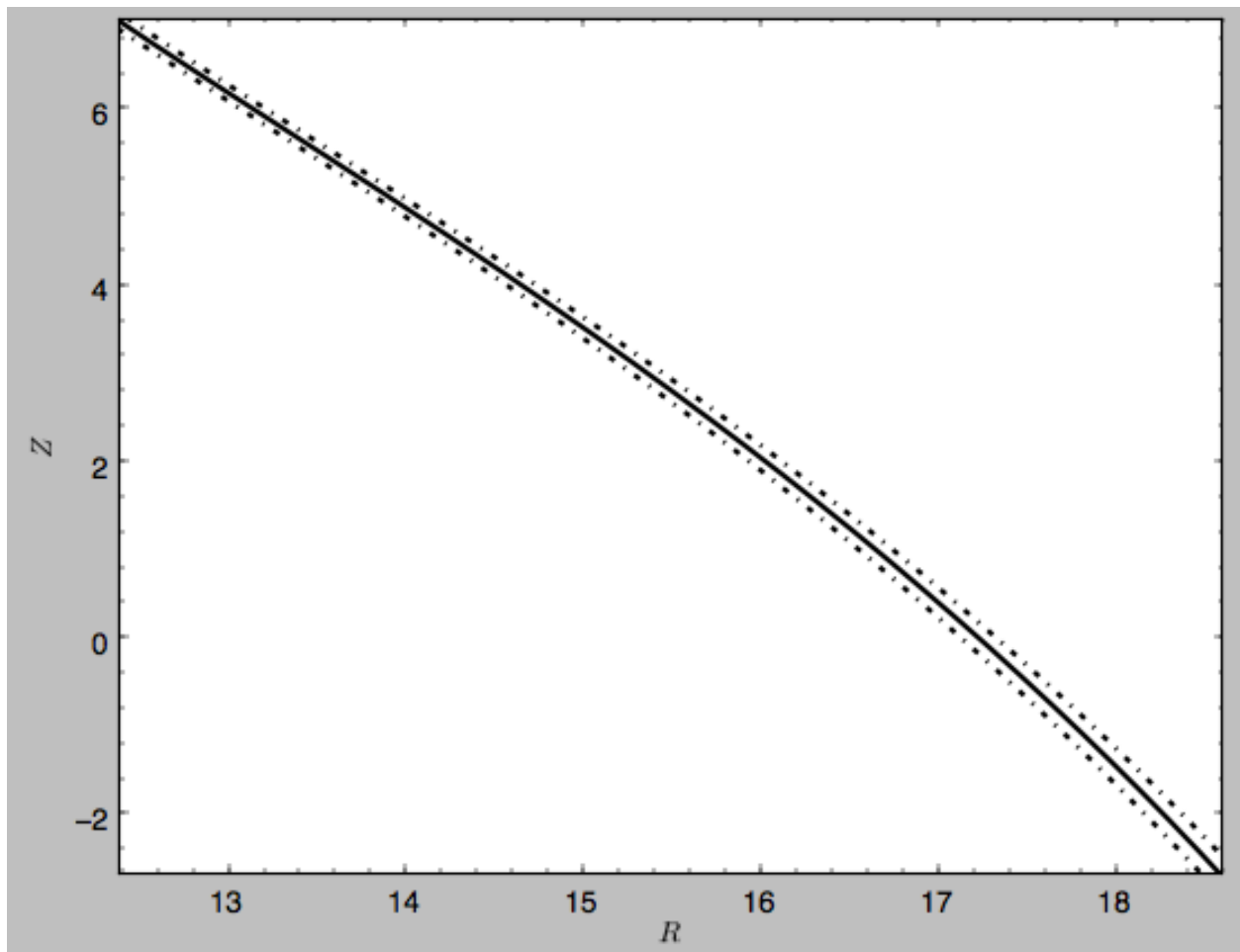
(the reason these are obtained by specifying `isotropic=True` is that these would be the ratio of the eigenvalues or the angle if we assumed that the disrupted materials action distribution were isotropic).

2.1.2 Calculating the average stream location (track)

We can display the stream track in various coordinate systems as follows

```
>>> sdf.plotTrack(d1='r', d2='z', interp=True, color='k', spread=2, overplot=False, lw=2.,
↳ scaleToPhysical=True)
```

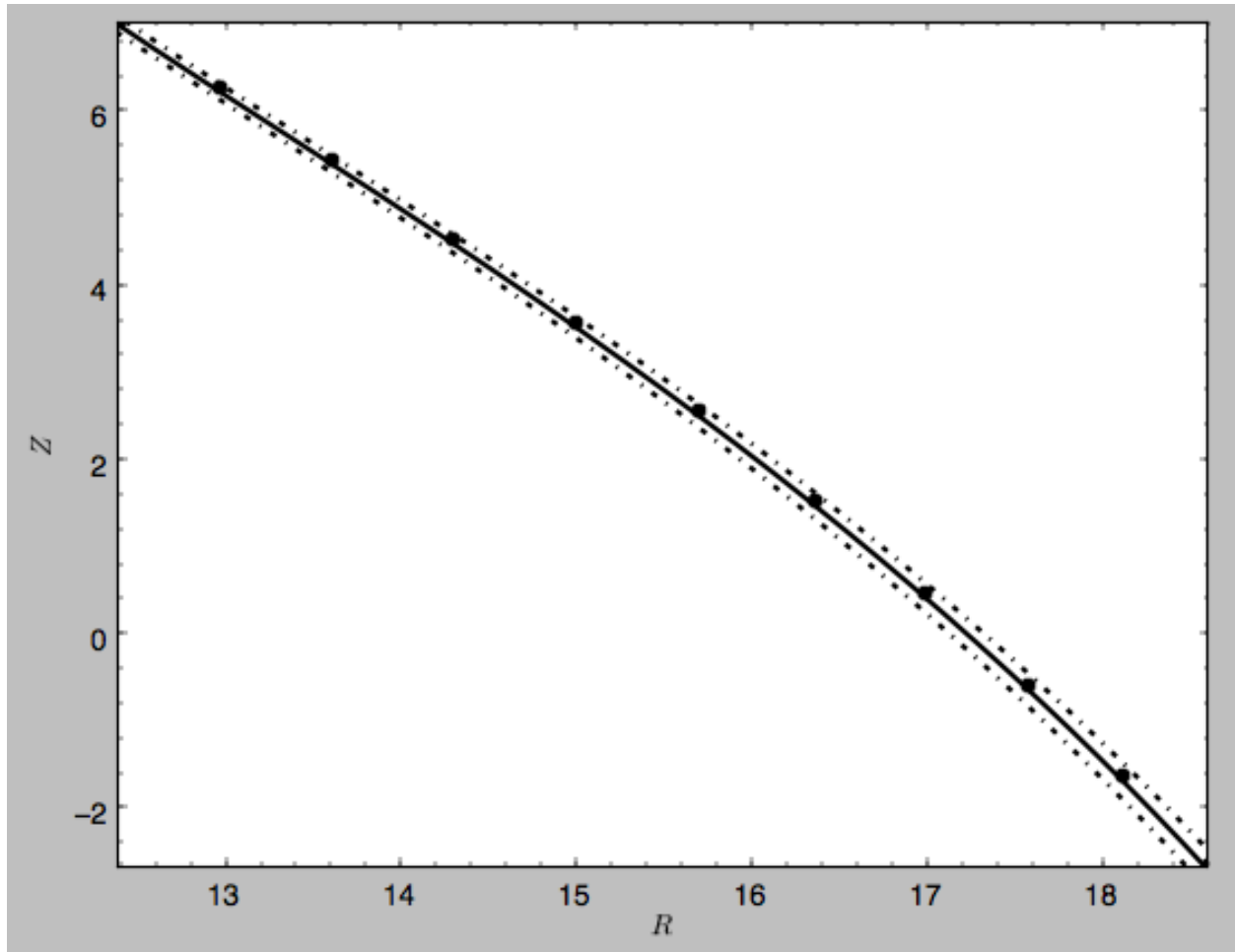
which gives



which shows the track in Galactocentric R and Z coordinates as well as an estimate of the spread around the track as the dash-dotted line. We can overplot the points along the track along which the $(x, v) \rightarrow (\Omega, \theta)$ transformation and the track position is explicitly calculated, by turning off the interpolation

```
>>> sdf.plotTrack(d1='r', d2='z', interp=False, color='k', spread=0, overplot=True, ls='none',
↳ marker='o', scaleToPhysical=True)
```

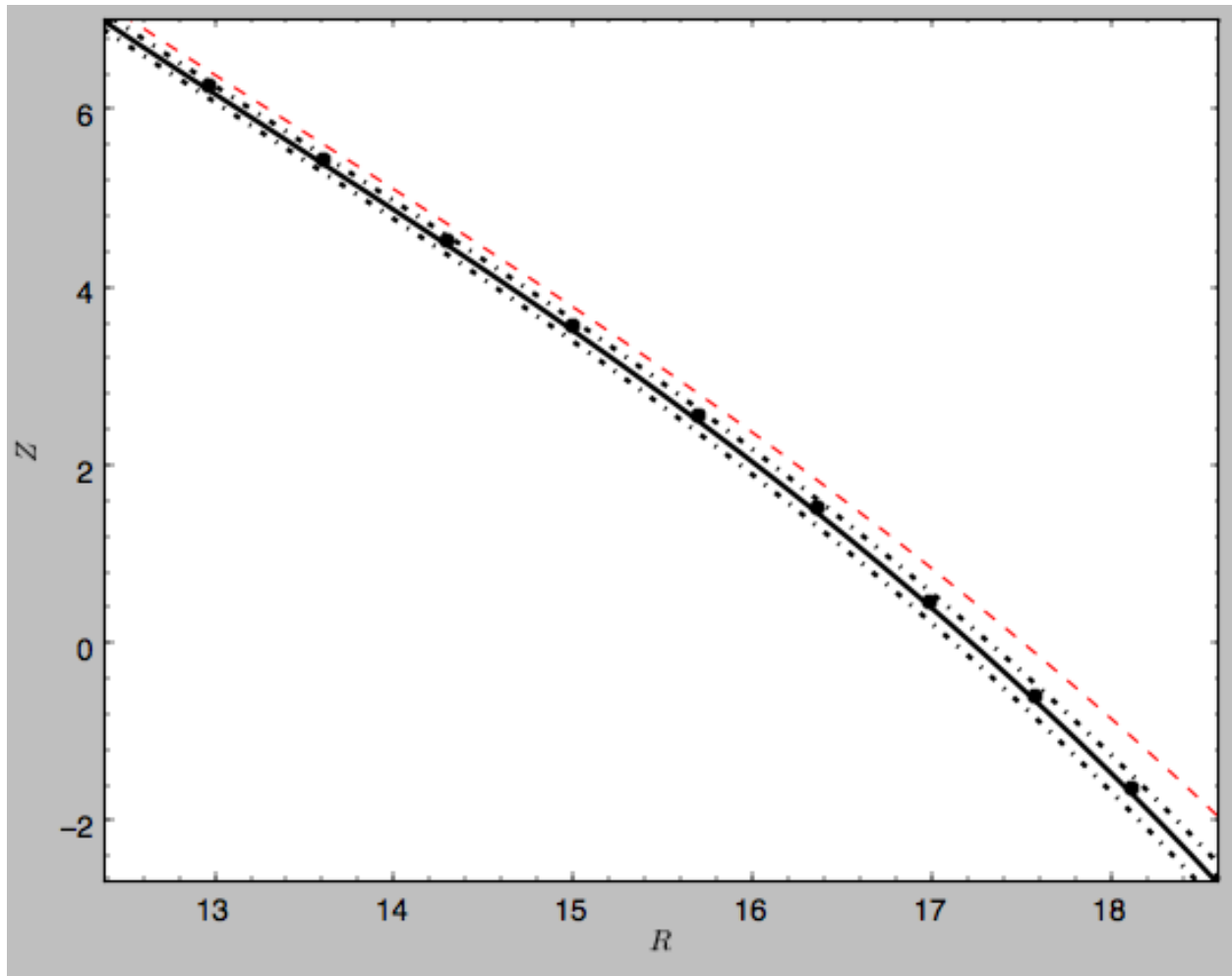
which gives



We can also overplot the orbit of the progenitor

```
>>> sdf.plotProgenitor(d1='r', d2='z', color='r', overplot=True, ls='--',
↳ scaleToPhysical=True)
```

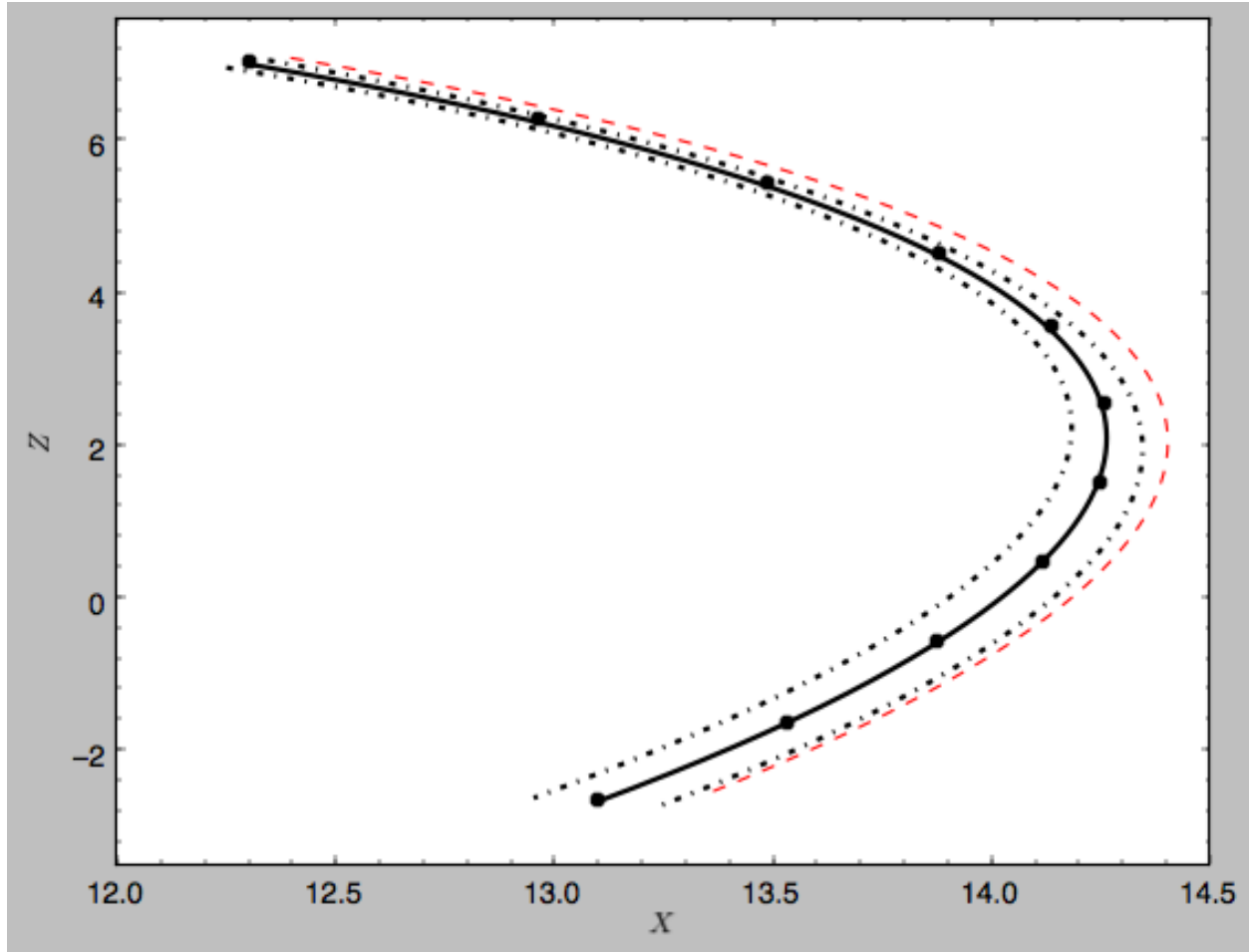
to give



We can do the same in other coordinate systems, for example X and Z (as in Figure 1 of Bovy 2014)

```
>>> sdf.plotTrack(d1='x', d2='z', interp=True, color='k', spread=2, overplot=False, lw=2.,
↳ scaleToPhysical=True)
>>> sdf.plotTrack(d1='x', d2='z', interp=False, color='k', spread=0, overplot=True, ls='none
↳ ', marker='o', scaleToPhysical=True)
>>> sdf.plotProgenitor(d1='x', d2='z', color='r', overplot=True, ls='--',
↳ scaleToPhysical=True)
>>> xlim(12., 14.5); ylim(-3.5, 7.6)
```

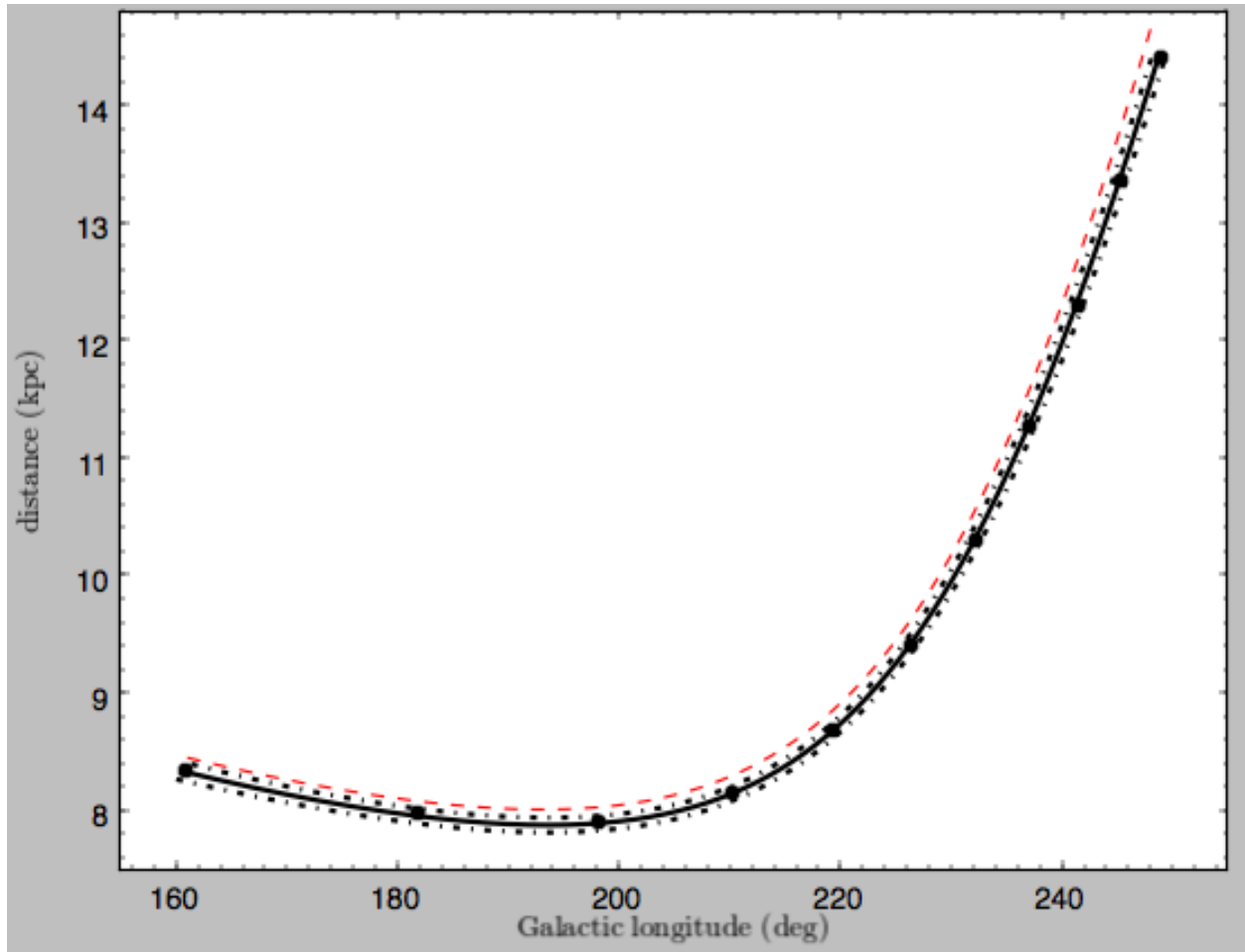
which gives



or we can calculate the track in observable coordinates, e.g.,

```
>>> sdf.plotTrack(d1='l1', d2='dist', interp=True, color='k', spread=2, overplot=False,
↳ lw=2.)
>>> sdf.plotTrack(d1='l1', d2='dist', interp=False, color='k', spread=0, overplot=True, ls=
↳ 'none', marker='o')
>>> sdf.plotProgenitor(d1='l1', d2='dist', color='r', overplot=True, ls='--')
>>> xlim(155., 255.); ylim(7.5, 14.8)
```

which displays



Coordinate transformations to physical coordinates are done using parameters set when initializing the `sdf` instance. See the help for `?streamdf` for a complete list of initialization parameters.

2.1.3 Mock stream data generation

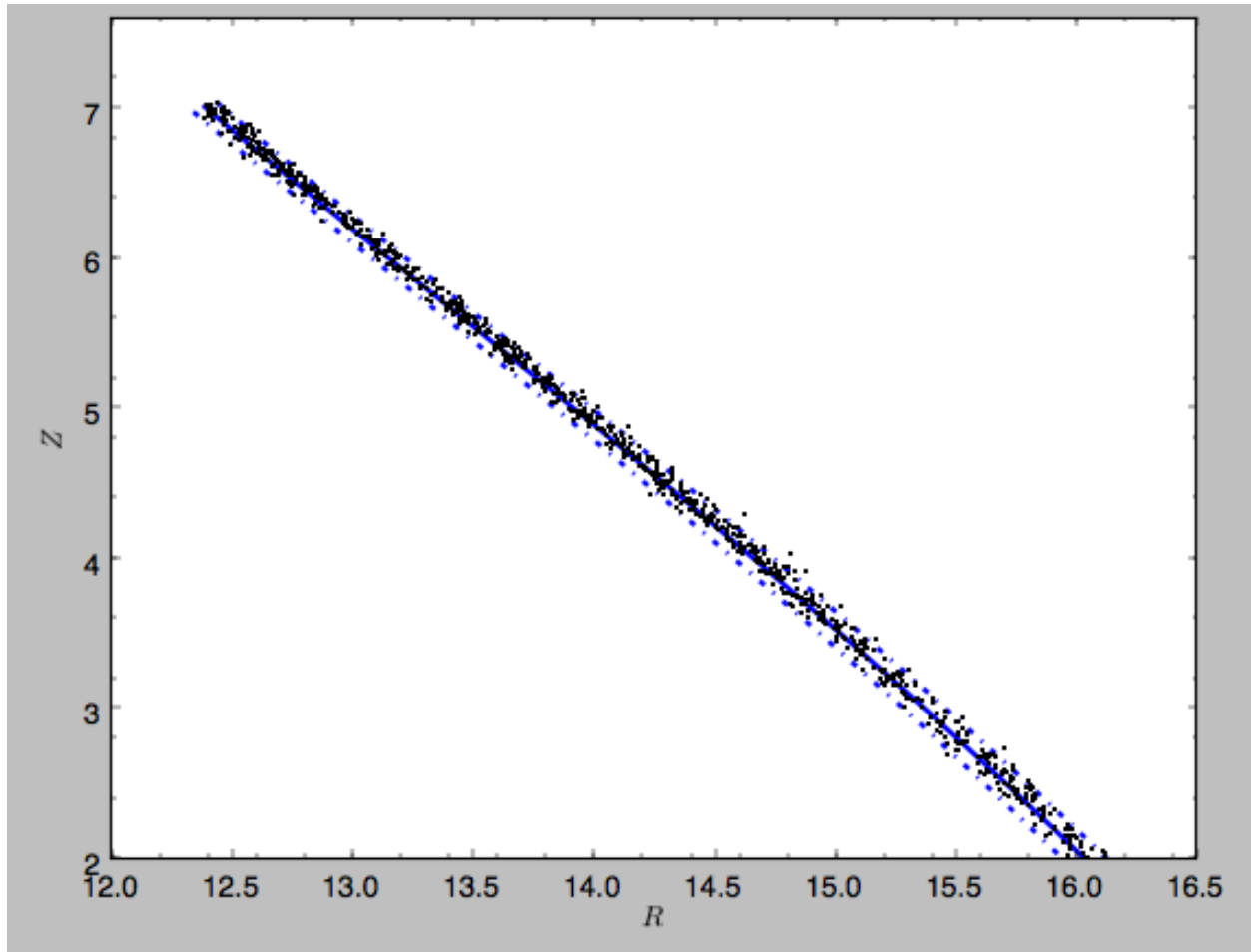
We can also easily generate mock data from the stream model. This uses `streamdf.sample`. For example,

```
>>> RvR= sdf.sample(n=1000)
```

which returns the sampled points as a set $(R, v_R, v_T, Z, v_Z, \phi)$ in natural galpy coordinates. We can plot these and compare them to the track location

```
>>> sdf.plotTrack(d1='r', d2='z', interp=True, color='b', spread=2, overplot=False, lw=2.,
↳ scaleToPhysical=True)
>>> plot(RvR[0]*8., RvR[3]*8., 'k.', ms=2.) #multiply by the physical distance scale
>>> xlim(12., 16.5); ylim(2., 7.6)
```

which gives



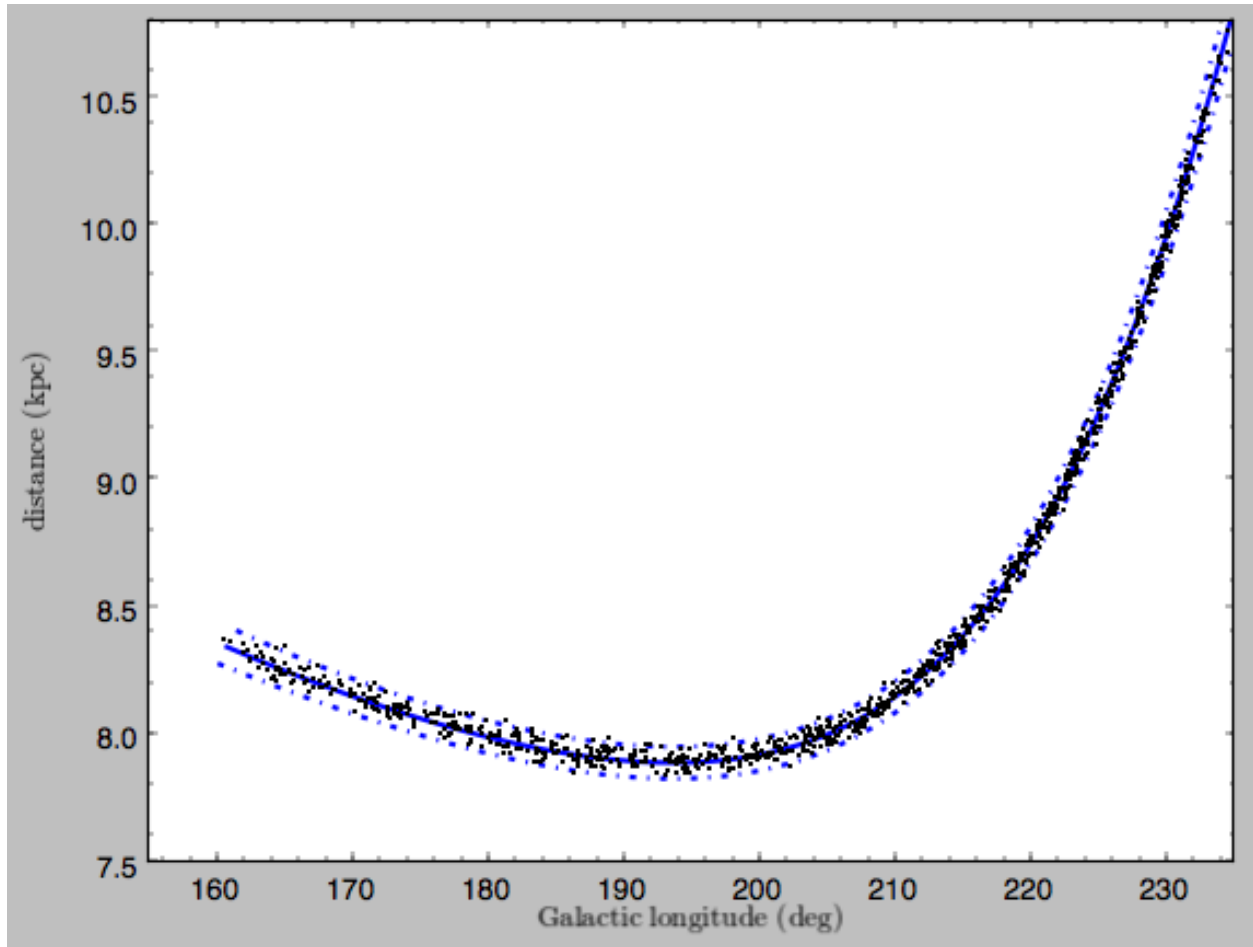
Similarly, we can generate mock data in observable coordinates

```
>>> lb= sdf.sample(n=1000,lb=True)
```

and plot it

```
>>> sdf.plotTrack(d1='l1',d2='dist',interp=True,color='b',spread=2,overplot=False,
↳lw=2.)
>>> plot(lb[0],lb[2],'k.',ms=2.)
>>> xlim(155.,235.); ylim(7.5,10.8)
```

which displays

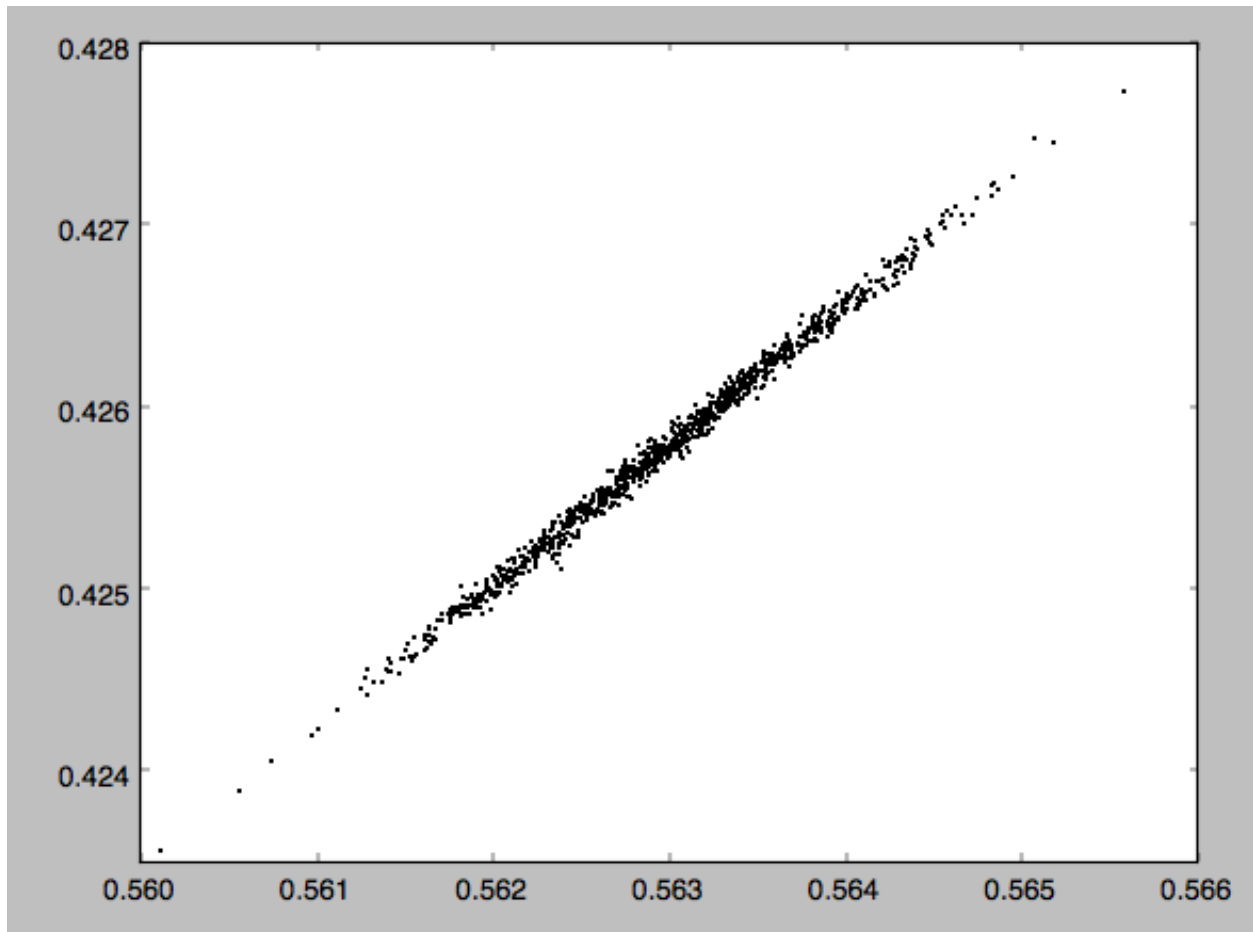


We can also just generate mock stream data in frequency-angle coordinates

```
>>> mockaA= sdf.sample(n=1000,returnAdt=True)
```

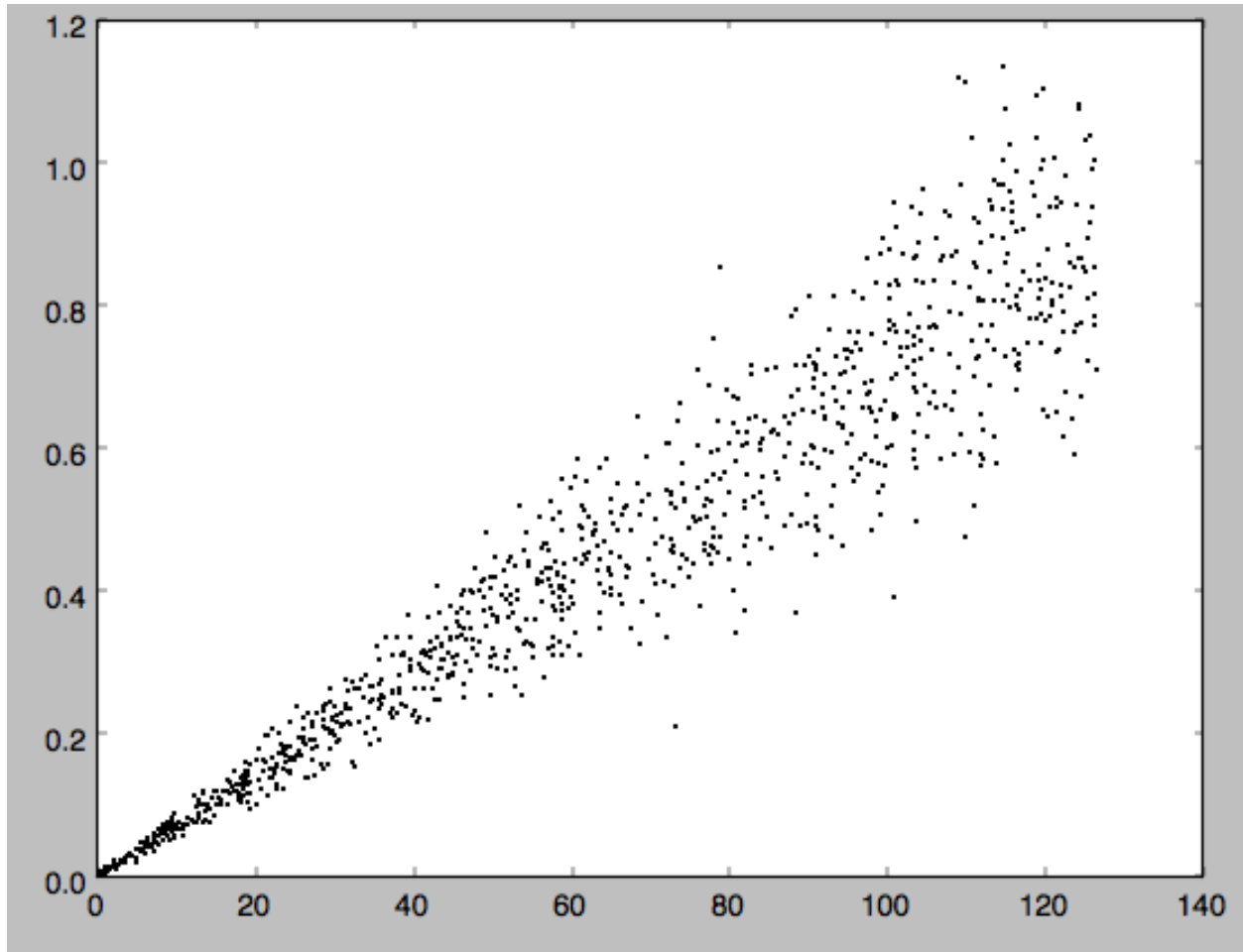
which returns a tuple with three components: an array with shape [3,N] of frequency vectors ($\Omega_R, \Omega_\phi, \Omega_Z$), an array with shape [3,N] of angle vectors ($\theta_R, \theta_\phi, \theta_Z$) and t_s , the stripping time. We can plot the vertical versus the radial frequency

```
>>> plot(mockaA[0][0],mockaA[0][2], 'k.', ms=2.)
```



or we can plot the magnitude of the angle offset as a function of stripping time

```
>>> plot(mockaA[2], numpy.sqrt(numpy.sum((mockaA[1]-numpy.tile(sdf._progenitor_angle,
↪ (1000,1)).T)**2., axis=0)), 'k.', ms=2.)
```



2.1.4 Evaluating and marginalizing the full PDF

We can also evaluate the stream PDF, the probability of a (\mathbf{x}, \mathbf{v}) phase-space position in the stream. We can evaluate the PDF, for example, at the location of the progenitor

```
>>> sdf(obs.R(), obs.vR(), obs.vT(), obs.z(), obs.vz(), obs.phi())
# array([-33.16985861])
```

which returns the natural log of the PDF. If we go to slightly higher in Z and slightly smaller in R , the PDF becomes zero

```
>>> sdf(obs.R()-0.1, obs.vR(), obs.vT(), obs.z()+0.1, obs.vz(), obs.phi())
# array([-inf])
```

because this phase-space position cannot be reached by a leading stream star. We can also marginalize the PDF over unobserved directions. For example, similar to Figure 10 in Bovy (2014), we can evaluate the PDF $p(X|Z)$ near a point on the track, say near $Z = 2$ kpc (≈ 0.25 in natural units). We first find the approximate Gaussian PDF near this point, calculated from the stream track and dispersion (see above)

```
>>> meanp, varp= sdf.gaussApprox([None, None, 2./8., None, None, None])
```

where the input is a array with entries $[X, Y, Z, v_X, v_Y, v_Z]$ and we substitute `None` for directions that we want to establish the approximate PDF for. So the above expression returns an approximation to $p(X, Y, v_X, v_Y, v_Z|Z)$. This

approximation allows us to get a sense of where the PDF peaks and what its width is

```
>>> meanp[0]*8.
# 14.267559400127833
>>> numpy.sqrt(varp[0,0])*8.
# 0.04152968631186698
```

We can now evaluate the PDF $p(X|Z)$ as a function of X near the peak

```
>>> xs= numpy.linspace(-3.*numpy.sqrt(varp[0,0]),3.*numpy.sqrt(varp[0,0]),21)+meanp[0]
>>> logps= numpy.array([sdf.callMarg([x,None,2./8.,None,None,None]) for x in xs])
>>> ps= numpy.exp(logps)
```

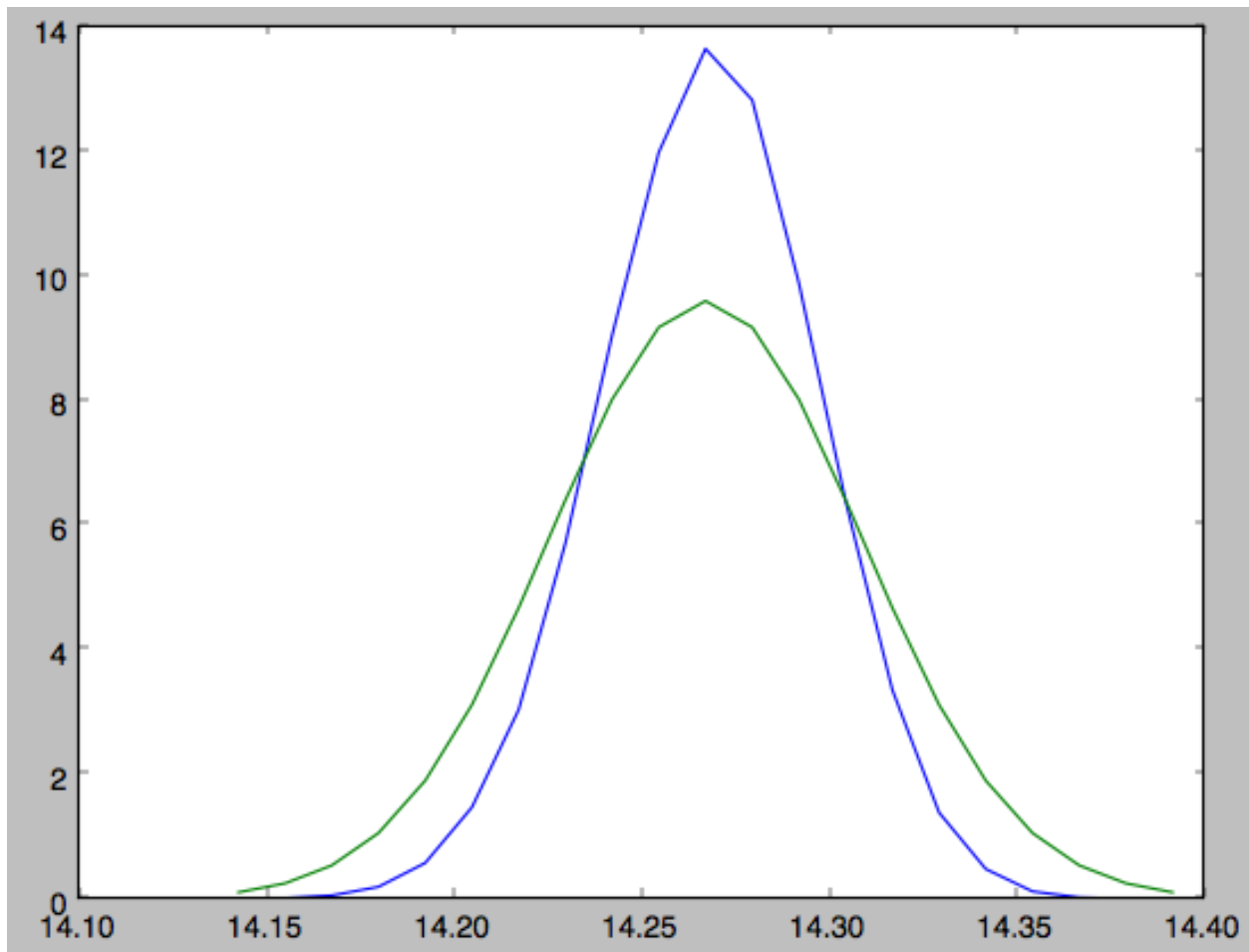
and we normalize the PDF

```
>>> ps/= numpy.sum(ps)*(xs[1]-xs[0])*8.
```

and plot it together with the Gaussian approximation

```
>>> plot(xs*8.,ps)
>>> plot(xs*8.,1./numpy.sqrt(2.*numpy.pi)/numpy.sqrt(varp[0,0])/8.*numpy.exp(-0.5*(xs-
↪ meanp[0])**2./varp[0,0]))
```

which gives



Sometimes it is hard to automatically determine the closest point on the calculated track if only one phase-space

coordinate is given. For example, this happens when evaluating $p(Z|X)$ for $X > 13$ kpc here, where there are two branches of the track in Z (see the figure of the track above). In that case, we can determine the closest track point on one of the branches by hand and then provide this closest point as the basis of PDF calculations. The following example shows how this is done for the upper Z branch at $X = 13.5$ kpc, which is near $Z = 5$ kpc (Figure 10 in Bovy 2014).

```
>>> cindx= sdf.find_closest_trackpoint(13.5/8., None, 5.32/8., None, None, None, xy=True)
```

gives the index of the closest point on the calculated track. This index can then be given as an argument for the PDF functions:

```
>>> meanp, varp= meanp, varp= sdf.gaussApprox([13.5/8., None, None, None, None, None],
↪ cindx=cindx)
```

computes the approximate $p(Y, Z, v_X, v_Y, v_Z|X)$ near the upper Z branch. In Z , this PDF has mean and dispersion

```
>>> meanp[1]*8.
# 5.4005530328542077
>>> numpy.sqrt(varp[1,1])*8.
# 0.05796023309510244
```

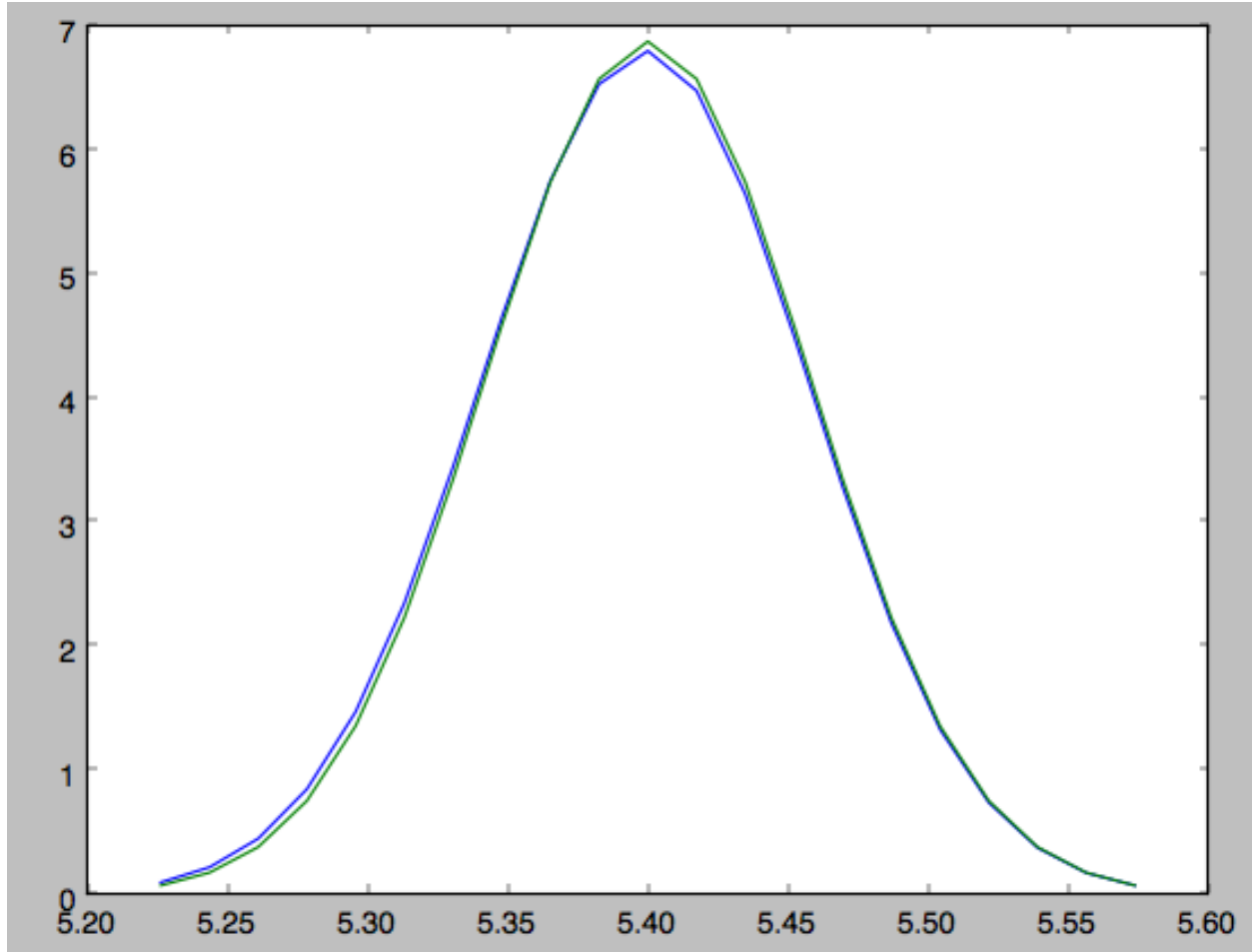
We can then evaluate $p(Z|X)$ for the upper branch as

```
>>> zs= numpy.linspace(-3.*numpy.sqrt(varp[1,1]), 3.*numpy.sqrt(varp[1,1]), 21)+meanp[1]
>>> logps= numpy.array([sdf.callMarg([13.5/8., None, z, None, None, None], cindx=cindx) for
↪ z in zs])
>>> ps= numpy.exp(logps)
>>> ps/= numpy.sum(ps)*(zs[1]-zs[0])*8.
```

and we can again plot this and the approximation

```
>>> plot(zs*8., ps)
>>> plot(zs*8., 1./numpy.sqrt(2.*numpy.pi)/numpy.sqrt(varp[1,1])/8.*numpy.exp(-0.5*(zs-
↪ meanp[1])**2./varp[1,1]))
```

which gives



The approximate PDF in this case is very close to the correct PDF. When supplying the closest track point, care needs to be taken that this really is the closest track point. Otherwise the approximate PDF will not be quite correct.

2.1.5 Modeling gaps in streams

galpy also contains tools to model the effect of impacts due to dark-matter subhalos on streams (see [Sanders, Bovy, & Erkal 2015](#)). This is implemented as a subclass `streamgapdf` of `streamdf`, because they share many of the same methods. Setting up a `streamgapdf` object requires the same arguments and keywords as setting up a `streamdf` instance (to specify the smooth underlying stream model and the Galactic potential) as well as parameters that specify the impact (impact parameter and velocity, location and time of closest approach, mass and structure of the subhalo, and helper keywords that specify how the impact should be calculated). An example used in the paper (but not that with the modifications in Sec. 6.1) is as follows. Imports:

```
>>> from galpy.df import streamdf, streamgapdf
>>> from galpy.orbit import Orbit
>>> from galpy.potential import LogarithmicHaloPotential
>>> from galpy.actionAngle import actionAngleIsochroneApprox
>>> from galpy.util import bovy_conversion
```

Parameters for the smooth stream and the potential:

```
>>> lp= LogarithmicHaloPotential(normalize=1.,q=0.9)
>>> aAI= actionAngleIsochroneApprox(pot=lp,b=0.8)
```

(continues on next page)

(continued from previous page)

```
>>> prog_unp_peri= Orbit([2.6556151742081835,
                        0.2183747276300308,
                        0.67876510797240575,
                        -2.0143395648974671,
                        -0.3273737682604374,
                        0.24218273922966019])

>>> V0, R0= 220., 8.
>>> sigv= 0.365*(10./2.)*(1./3.) # km/s
>>> tdisrupt= 10.88/bovy_conversion.time_in_Gyr(V0,R0)
```

and the parameters of the impact

```
>>> GM= 10.**-2./bovy_conversion.mass_in_1010msol(V0,R0)
>>> rs= 0.625/R0
>>> impactb= 0.
>>> subhalovel= numpy.array([6.82200571,132.7700529,149.4174464])/V0
>>> timpact= 0.88/bovy_conversion.time_in_Gyr(V0,R0)
>>> impact_angle= -2.34
```

The setup is then

```
>>> sdf_sanders15= streamgapdf(sigv/V0,progenitor=prog_unp_peri,pot=lp,aA=aAI,
                             leading=False,nTrackChunks=26,
                             nTrackIterations=1,
                             sigMeanOffset=4.5,
                             tdisrupt=tdisrupt,
                             Vnorm=V0,Rnorm=R0,
                             impactb=impactb,
                             subhalovel=subhalovel,
                             timpact=timpact,
                             impact_angle=impact_angle,
                             GM=GM,rs=rs)
```

The `streamgapdf` implementation is currently not entirely complete (for example, one cannot yet evaluate the full phase-space PDF), but the model can be sampled as in the paper above. To compare the perturbed model to the unperturbed model, we also set up an unperturbed model of the same stream

```
>>> sdf_sanders15_unp= streamdf(sigv/V0,progenitor=prog_unp_peri,pot=lp,aA=aAI,
                               leading=False,nTrackChunks=26,
                               nTrackIterations=1,
                               sigMeanOffset=4.5,
                               tdisrupt=tdisrupt,
                               Vnorm=V0,Rnorm=R0)
```

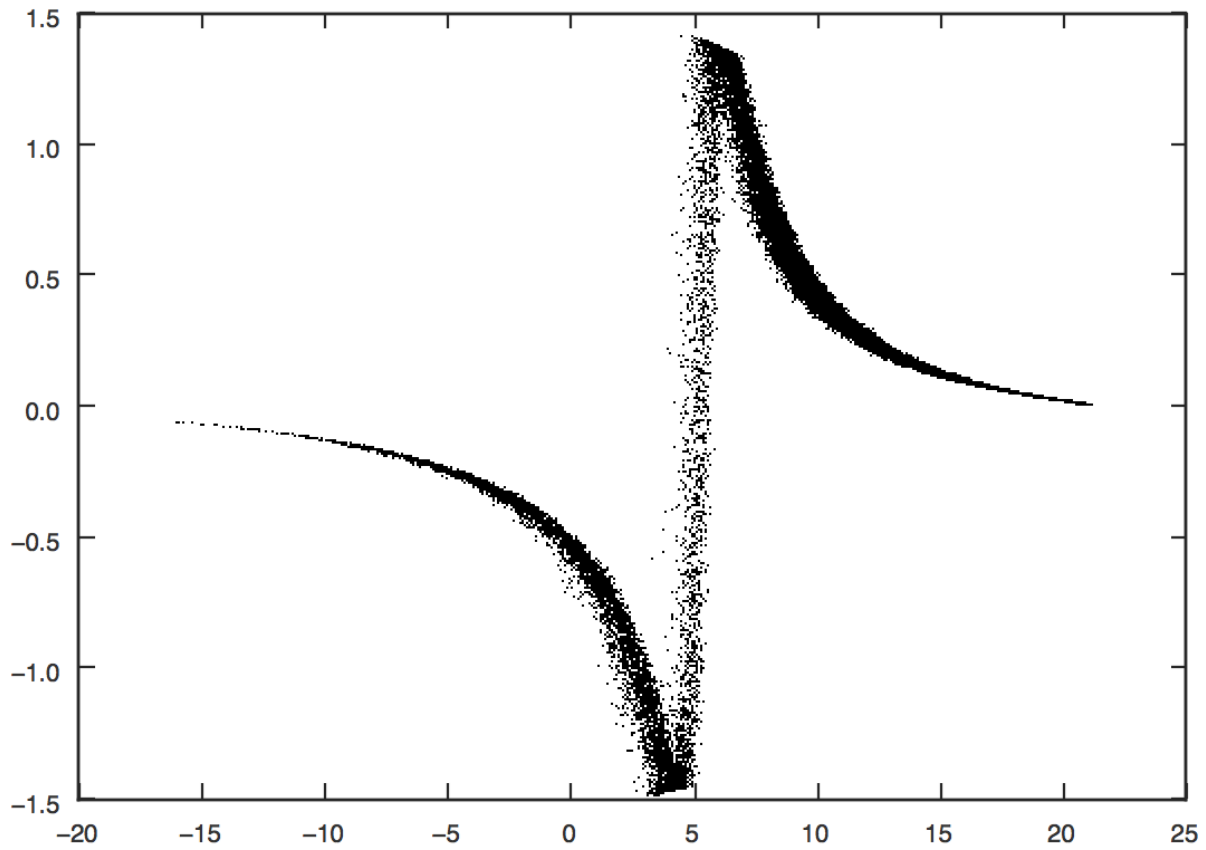
We can then sample positions and velocities for the perturbed and unperturbed production for the *same* particle by using the same random seed:

```
>>> numpy.random.seed(1)
>>> xv_mock_per= sdf_sanders15.sample(n=100000,xy=True).T
>>> numpy.random.seed(1) # should give same points
>>> xv_mock_unp= sdf_sanders15_unp.sample(n=100000,xy=True).T
```

and we can plot the offset due to the perturbation, for example,

```
>>> plot(xv_mock_unp[:,0]*R0,(xv_mock_per[:,0]-xv_mock_unp[:,0])*R0,'k,')
```

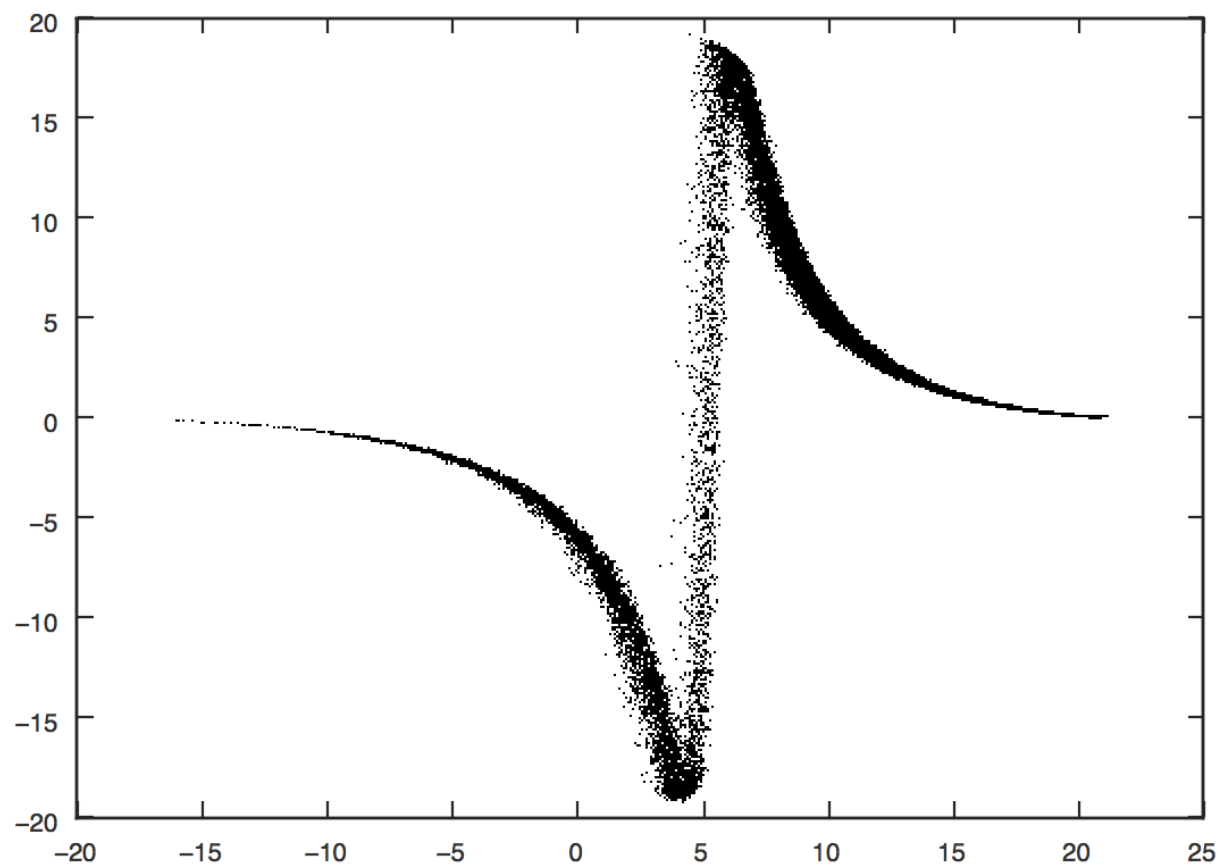
for the difference in X as a function of unperturbed X :



or

```
>>> plot(xv_mock_unp[:,0]*R0, (xv_mock_per[:,4]-xv_mock_unp[:,4])*V0, 'k,')
```

for the difference in v_Y as a function of unperturbed X :



3.1 Orbit (`galpy.orbit`)

See *Orbit initialization* for a detailed explanation on how to set up Orbit instances.

3.1.1 Class

`galpy.orbit.Orbit`

class `galpy.orbit.Orbit` (`vxvv=None`, `uvw=False`, `lb=False`, `radec=False`, `vo=None`, `ro=None`,
`zo=None`, `solarmotion=None`)

General orbit class representing an orbit

__init__ (`vxvv=None`, `uvw=False`, `lb=False`, `radec=False`, `vo=None`, `ro=None`, `zo=None`, `solarmotion=None`)

NAME:

`__init__`

PURPOSE:

Initialize an Orbit instance

INPUT:

vxvv - initial conditions

3D can be either

- 1) in Galactocentric cylindrical coordinates [`R`,`vR`,`vT`(`z`,`vz`,`phi`)]; can be Quantities
- 2) `astropy` (>`v3.0`) `SkyCoord` that includes velocities (Note that this turns *on* physical output even if `ro` and `vo` are not given)
- 3) [`ra`,`dec`,`d`,`mu_ra`, `mu_dec`,`vlos`] in [deg,deg,kpc,mas/yr,mas/yr,km/s] (all J2000.0; `mu_ra` = `mu_ra` * `cos dec`); can be Quantities; ICRS frame

- 4) [ra,dec,d,U,V,W] in [deg,deg,kpc,km/s,km/s,kms]; can be Quantities; ICRS frame
- 5) (l,b,d,mu_l, mu_b, vlos) in [deg,deg,kpc,mas/yr,mas/yr,km/s) (all J2000.0; mu_l = mu_l * cos b); can be Quantities
- 6) [l,b,d,U,V,W] in [deg,deg,kpc,km/s,km/s,kms]; can be Quantities
- 7) Unspecified: assumed to be the Sun (equivalent to vxvv= [0,0,0,0,0,0] and radec=True)

5) and 6) also work when leaving out b and mu_b/W

OPTIONAL INPUTS:

radec= if True, input is 2) (or 3) above (Note that this turns *on* physical output even if ro and vo are not given)

uvw= if True, velocities are UVW

lb= if True, input is 4) or 5) above (Note that this turns *on* physical output even if ro and vo are not given)

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

zo= offset toward the NGP of the Sun wrt the plane (kpc; can be Quantity; default = 25 pc)

solarmotion= 'hogg' or 'dehnen', or 'schoenrich', or value in [-U,V,W]; can be Quantity

If ro and/or vo are specified, outputs involving distances or velocities (whether as instance methods or in plots) will by default be displayed in the physical coordinates implied by these scales. This can be overwritten for each individual method by using use_physical=False as a keyword for the method.

OUTPUT:

instance

HISTORY:

2010-07-20 - Written - Bovy (NYU)

galpy.orbit.Orbit.from_name

classmethod `Orbit.from_name` (*name*, *vo=None*, *ro=None*, *zo=None*, *solarmotion=None*)

NAME:

from_name

PURPOSE:

given the name of an object, retrieve coordinate information for that object from SIMBAD and return a corresponding orbit

INPUT:

name - the name of the object

+standard Orbit initialization keywords:

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

zo= offset toward the NGP of the Sun wrt the plane (kpc; can be Quantity; default = 25 pc)

solarmotion= 'hogg' or 'dehnen', or 'schoenrich', or value in [-U,V,W]; can be Quantity

OUTPUT:

orbit containing the phase space coordinates of the named object

HISTORY:

2018-07-15 - Written - Mathew Bub (UofT)

3.1.2 Methods

galpy.orbit.Orbit.__add__

`Orbit.__add__` (*linOrb*)

NAME:

`__add__`

PURPOSE:

add a linear orbit and a planar orbit to make a 3D orbit

INPUT:

linear or plane orbit instance

OUTPUT:

3D orbit

HISTORY:

2010-07-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.__call__

`Orbit.__call__` (**args, **kwargs*)

NAME:

`__call__`

PURPOSE:

return the orbit at time t

INPUT:

t - desired time (can be Quantity)

rect - if true, return rectangular coordinates

OUTPUT:

an Orbit instance with initial condition set to the phase-space at time t or list of Orbit instances if multiple times are given

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.animate

`Orbit.animate(*args, **kwargs)`

NAME:

animate

PURPOSE:

animate a previously calculated orbit (with reasonable defaults)

INPUT:

d1= first dimension to plot ('x', 'y', 'R', 'vR', 'vT', 'z', 'vz', ...); can be list with up to three entries for three subplots; each entry can also be a user-defined function of time (e.g., lambda t: o.R(t) for R)

d2= second dimension to plot; can be list with up to three entries for three subplots; each entry can also be a user-defined function of time (e.g., lambda t: o.R(t) for R)

width= (600) width of output div in px

height= (400) height of output div in px

xlabel= (pre-defined labels) label for the first dimension (or list of labels if d1 is a list); should only have to be specified when using a function as d1 and can then specify as, e.g., [None,'YOUR LABEL',None] if d1 is a list of three xs and the first and last are standard entries)

ylabel= (pre-defined labels) label for the second dimension (or list of labels if d2 is a list); should only have to be specified when using a function as d2 and can then specify as, e.g., [None,'YOUR LABEL',None] if d1 is a list of three xs and the first and last are standard entries)

json_filename= (None) if set, save the data necessary for the figure in this filename (e.g., json_filename= 'orbit_data/orbit.json'); this path is also used in the output HTML, so needs to be accessible

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

IPython.display.HTML object with code to animate the orbit; can be directly shown in jupyter notebook or embedded in HTML pages; get a text version of the HTML using the `_repr_html_()` function

HISTORY:

2017-09-17-24 - Written - Bovy (UofT)

galpy.orbit.Orbit.bb

`Orbit.bb(*args, **kwargs)`

NAME:

bb

PURPOSE:

return Galactic latitude

INPUT:

t - (optional) time at which to get bb (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity) (default=[8.0,0.,0.])
OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

b(t) in deg

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.orbit.Orbit.dec

Orbit.**dec** (*args, **kwargs)

NAME:

dec

PURPOSE:

return the declination

INPUT:

t - (optional) time at which to get dec (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity) (default=[8.0,0.,0.])
OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

dec(t) in deg

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.orbit.Orbit.dist

Orbit.**dist** (*args, **kwargs)

NAME:

dist

PURPOSE:

return distance from the observer

INPUT:

t - (optional) time at which to get dist (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity) (default=[8.0,0.,0.])
OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

dist(t) in kpc

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.orbit.Orbit.E

`Orbit.E(*args, **kwargs)`

NAME:

E

PURPOSE:

calculate the energy

INPUT:

t - (optional) time at which to get the energy (can be Quantity)

pot= Potential instance or list of such instances

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

energy

HISTORY:

2010-09-15 - Written - Bovy (NYU)

galpy.orbit.Orbit.e

`Orbit.e(analytic=False, pot=None, **kwargs)`

NAME:

e

PURPOSE:

calculate the eccentricity, either numerically from the numerical orbit integration or using analytical means

INPUT:

analytic(= False) compute this analytically

pot - potential to use for analytical calculation

For 3D orbits different approximations for analytic=True are available (see the EccZmaxRperiRap method of actionAngle modules):

type= ('staeckel') type of actionAngle module to use

- 1) 'adiabatic': assuming motion splits into R and z
- 2) 'staeckel': assuming motion splits into u and v of prolate spheroidal coordinate system, exact for Staeckel potentials (incl. all spherical potentials)
- 3) 'spherical': for spherical potentials, exact

+actionAngle module setup kwargs for the corresponding actionAngle modules (actionAngleAdiabatic, actionAngleStaeckel, and actionAngleSpherical)

OUTPUT:

eccentricity

HISTORY:

2010-09-15 - Written - Bovy (NYU)

2017-12-25 - Added Staeckel approximation and made that the default - Bovy (UofT)

galpy.orbit.Orbit.ER

Orbit.ER(*args, **kwargs)

NAME:

ER

PURPOSE:

calculate the radial energy

INPUT:

t - (optional) time at which to get the radial energy (can be Quantity)

pot= Potential instance or list of such instances

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output (can be Quantity)

OUTPUT:

radial energy

HISTORY:

2013-11-30 - Written - Bovy (IAS)

galpy.orbit.Orbit.Ez

Orbit.Ez(*args, **kwargs)

NAME:

Ez

PURPOSE:

calculate the vertical energy

INPUT:

t - (optional) time at which to get the vertical energy (can be Quantity)

pot= Potential instance or list of such instances

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output (can be Quantity)

OUTPUT:

vertical energy

HISTORY:

2013-11-30 - Written - Bovy (IAS)

galpy.orbit.Orbit.fit

```
Orbit.fit(vxvv, vxvv_err=None, pot=None, radec=False, lb=False, customsky=False,
         lb_to_customsky=None, pmllpmmbb_to_customsky=None, tintJ=10, ntintJ=1000, inte-
         grate_method='dopr54_c', **kwargs)
```

NAME:

fit

PURPOSE:

fit an Orbit to data using the current orbit as the initial condition

INPUT:

vxvv - [:,6] array of positions and velocities along the orbit (if not lb=True or radec=True, these need to be in natural units [/ro,/vo], cannot be Quantities)

vxvv_err= [:,6] array of errors on positions and velocities along the orbit (if None, these are set to 0.01) (if not lb=True or radec=True, these need to be in natural units [/ro,/vo], cannot be Quantities)

pot= Potential to fit the orbit in

Keywords related to the input data:

radec= if True, input vxvv and vxvv_err are [ra,dec,d,mu_ra, mu_dec,vlos] in [deg,deg,kpc,mas/yr,mas/yr,km/s] (all J2000.0; mu_ra = mu_ra * cos dec); the attributes of the current Orbit are used to convert between these coordinates and Galactocentric coordinates

lb= if True, input vxvv and vxvv_err are [long,lat,d,mu_ll, mu_bb,vlos] in [deg,deg,kpc,mas/yr,mas/yr,km/s] (mu_ll = mu_ll * cos lat); the attributes of the current Orbit are used to convert between these coordinates and Galactocentric coordinates

customsky= if True, input vxvv and vxvv_err are [custom long,custom lat,d,mu_customll, mu_custombb,vlos] in [deg,deg,kpc,mas/yr,mas/yr,km/s] (mu_ll = mu_ll * cos lat) where custom longitude and custom latitude are a custom set of sky coordinates (e.g., ecliptic) and the proper motions are also expressed in these coordinates; you need to provide the functions lb_to_customsky and pmllpmmbb_to_customsky to convert to the custom sky coordinates (these should have the same inputs and outputs as lb_to_radec and pmllpmmbb_to_pmrpmdec); the attributes of the current Orbit are used to convert between these coordinates and Galactocentric coordinates

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer

(in kpc and km/s; entries can be Quantity) (default=Object-wide default) Cannot be an Orbit instance with the orbit of the reference point, as w/ the ra etc. functions Y is ignored and always assumed to be zero

ro= distance in kpc corresponding to R=1. (default: taken from object; can be Quantity)

vo= velocity in km/s corresponding to v=1. (default: taken from object; can be Quantity)

lb_to_customsky= function that converts l,b,degree=False to the custom sky coordinates (like lb_to_radec); needs to be given when customsky=True

pmllpmbb_to_customsky= function that converts pmll,pmbb,l,b,degree=False to proper motions in the custom sky coordinates (like pmllpmbb_to_pmrpmdec); needs to be given when customsky=True

Keywords related to the orbit integrations:

tintJ= (default: 10) time to integrate orbits for fitting the orbit (can be Quantity)

ntintJ= (default: 1000) number of time-integration points

integrate_method= (default: 'dopr54_c') integration method to use

disp= (False) display the optimizer's convergence message

OUTPUT:

max of log likelihood

HISTORY:

2014-06-17 - Written - Bovy (IAS)

galpy.orbit.Orbit.flip

Orbit.**flip** (*inplace=False*)

NAME:

flip

PURPOSE:

'flip' an orbit's initial conditions such that the velocities are minus the original velocities; useful for quick backward integration; returns a new Orbit instance

INPUT:

inplace= (False) if True, flip the orbit in-place, that is, without returning a new instance and also flip the velocities of the integrated orbit (if it exists)

OUTPUT:

Orbit instance that has the velocities of the current orbit flipped (inplace=False) or just flips all velocities of current instance (inplace=True)

HISTORY:

2014-06-17 - Written - Bovy (IAS)

2016-07-21 - Added inplace keyword - Bovy (UofT)

galpy.orbit.Orbit.integrate

Orbit.**integrate** (*t, pot, method='symplec4_c', dt=None*)

NAME:

integrate

PURPOSE:

integrate the orbit

INPUT:

t - list of times at which to output (0 has to be in this!) (can be Quantity)

pot - potential instance or list of instances

method= 'odeint' for scipy's odeint 'leapfrog' for a simple leapfrog implementation 'leapfrog_c' for a simple leapfrog implementation in C 'symplec4_c' for a 4th order symplectic integrator in C 'symplec6_c' for a 6th order symplectic integrator in C 'rk4_c' for a 4th-order Runge-Kutta integrator in C 'rk6_c' for a 6-th order Runge-Kutta integrator in C 'dopr54_c' for a Dormand-Prince integrator in C (generally the fastest)

dt= (None) if set, force the integrator to use this basic stepsize; must be an integer divisor of output stepsize (only works for the C integrators that use a fixed stepsize) (can be Quantity)

OUTPUT:

(none) (get the actual orbit using getOrbit())

HISTORY:

2010-07-10 - Written - Bovy (NYU)

2015-06-28 - Added dt keyword - Bovy (IAS)

galpy.orbit.Orbit.integrate_dxdv

Currently only supported for `planarOrbit` instances.

`Orbit.integrate_dxdv(dxdv, t, pot, method='dopr54_c', rectIn=False, rectOut=False)`

NAME:

integrate_dxdv

PURPOSE:

integrate the orbit and a small area of phase space

INPUT:

dxdv - [dR,dvR,dvT,dphi]

t - list of times at which to output (0 has to be in this!) (can be Quantity)

pot - potential instance or list of instances

method= 'odeint' for scipy's odeint

'rk4_c' for a 4th-order Runge-Kutta integrator in C

'rk6_c' for a 6-th order Runge-Kutta integrator in C

'dopr54_c' for a Dormand-Prince integrator in C (generally the fastest)

'dopr54_c' is recommended, odeint is *not* recommended

rectIn= (False) if True, input dxdv is in rectangular coordinates

rectOut= (False) if True, output dxdv (that in orbit_dxdv) is in rectangular coordinates

OUTPUT:

(none) (get the actual orbit using getOrbit_dxdv(), the orbit that is integrated alongside with dxdv is stored as usual, any previous regular orbit integration will be erased!)

HISTORY:

2010-10-17 - Written - Bovy (IAS)

2014-06-29 - Added rectIn and rectOut - Bovy (IAS)

galpy.orbit.Orbit.getOrbit

`Orbit.getOrbit()`

NAME:

`getOrbit`

PURPOSE:

return a previously calculated orbit

INPUT:

(none)

OUTPUT:

array orbit[nt,nd]

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.getOrbit_dxdv

`integrate_dxdv` is currently only supported for `planarOrbit` instances. `getOrbit_dxdv` is therefore also only supported for those types of `Orbit`.

`Orbit.getOrbit_dxdv()`

NAME:

`getOrbit_dxdv`

PURPOSE:

return a previously calculated integration of a small phase-space volume (with `integrate_dxdv`)

INPUT:

(none)

OUTPUT:

array orbit[nt,nd*2] with for each t the dxdv vector

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.helioX

`Orbit.helioX(*args, **kwargs)`

NAME:

`helioX`

PURPOSE:

return Heliocentric Galactic rectangular x-coordinate (aka “X”)

INPUT:

t - (optional) time at which to get X (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

helioX(t) in kpc

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.helioY

Orbit.**helioY**(*args, **kwargs)

NAME:

helioY

PURPOSE:

return Heliocentric Galactic rectangular y-coordinate (aka “Y”)

INPUT:

t - (optional) time at which to get Y (can be Quantity)

obs=[X,Y,Z] - (optional) position and of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

helioY(t) in kpc

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.helioZ

Orbit.**helioZ**(*args, **kwargs)

NAME:

helioZ

PURPOSE:

return Heliocentric Galactic rectangular z-coordinate (aka “Z”)

INPUT:

t - (optional) time at which to get Z (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

helioZ(t) in kpc

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.Jacobi

`Orbit.Jacobi(*args, **kwargs)`

NAME:

Jacobi

PURPOSE:

calculate the Jacobi integral $E - \Omega L$

INPUT:

t - (optional) time at which to get the Jacobi integral (can be Quantity)

OmegaP= pattern speed (can be Quantity)

pot= potential instance or list of such instances

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Jacobi integral

HISTORY:

2011-04-18 - Written - Bovy (NYU)

galpy.orbit.Orbit.jp

`Orbit.jp(pot=None, **kwargs)`

NAME:

jp

PURPOSE:

calculate the azimuthal action

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

jr

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.jr

Orbit.jr (pot=None, **kwargs)

NAME:

jr

PURPOSE:

calculate the radial action

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

jr

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.jz`Orbit.jz` (*pot=None, **kwargs*)

NAME:

jz

PURPOSE:

calculate the vertical action

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

jz

HISTORY:

2012-06-01 - Written - Bovy (IAS)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.l`Orbit.l` (**args, **kwargs*)

NAME:

l

PURPOSE:

return Galactic longitude

INPUT:

t - (optional) time at which to get l (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity) (default=[8.0,0.,0.])

OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

l(t) in deg

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.orbit.Orbit.L

`Orbit.L(*args, **kwargs)`

NAME:

L

PURPOSE:

calculate the angular momentum at time t

INPUT:

t - (optional) time at which to get the angular momentum (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

angular momentum

HISTORY:

2010-09-15 - Written - Bovy (NYU)

galpy.orbit.Orbit.Op

`Orbit.Op(pot=None, **kwargs)`

NAME:

Op

PURPOSE:

calculate the azimuthal frequency

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Op

HISTORY:

2013-11-27 - Written - Bovy (IAS)

galpy.orbit.Orbit.Or

`Orbit.Or` (*pot=None, **kwargs*)

NAME:

Or

PURPOSE:

calculate the radial frequency

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Or

HISTORY:

2013-11-27 - Written - Bovy (IAS)

galpy.orbit.Orbit.Oz

`Orbit.Oz` (*pot=None, **kwargs*)

NAME:

Oz

PURPOSE:

calculate the vertical frequency

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

- 1) 'adiabatic'
- 2) 'staeckel'
- 3) 'isochroneApprox'
- 4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Oz

HISTORY:

2013-11-27 - Written - Bovy (IAS)

galpy.orbit.Orbit.phi

Orbit.**phi** (*args, **kwargs)

NAME:

phi

PURPOSE:

return azimuth

INPUT:

t - (optional) time at which to get the azimuth (can be Quantity)

OUTPUT:

phi(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.plot

Orbit.**plot** (*args, **kwargs)

NAME:

plot

PURPOSE:

plot a previously calculated orbit (with reasonable defaults)

INPUT:

d1= first dimension to plot ('x', 'y', 'R', 'vR', 'vT', 'z', 'vz', ...); can also be a user-defined function of time (e.g., lambda t: o.R(t) for R)

d2= second dimension to plot; can also be a user-defined function of time (e.g., lambda t: o.R(t) for R)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
use_physical= use to override Object-wide default for using a physical scale for output
matplotlib.plot inputs+bovy_plot.plot inputs

OUTPUT:

sends plot to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plot3d

`Orbit.plot3d(*args, **kwargs)`

NAME:

plot3d

PURPOSE:

plot 3D aspects of an Orbit

INPUT:

d1= first dimension to plot ('x', 'y', 'R', 'vR', 'vT', 'z', 'vz', ...); can also be a user-defined function of time (e.g., lambda t: o.R(t) for R)

d2= second dimension to plot

d3= third dimension to plot

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

bovy_plot3d args and kwargs

OUTPUT:

plot

HISTORY:

2010-07-26 - Written - Bovy (NYU)

2010-09-22 - Adapted to more general framework - Bovy (NYU)

2010-01-08 - Adapted to 3D - Bovy (NYU)

galpy.orbit.Orbit.plotE

`Orbit.plotE(*args, **kwargs)`

NAME:

plotE

PURPOSE:

plot E(.) along the orbit

INPUT:

pot= Potential instance or list of instances in which the orbit was integrated
d1= plot Ez vs d1: e.g., 't', 'z', 'R', 'vR', 'vT', 'vz'
normed= if set, plot $E(t)/E(0)$ rather than $E(t)$
ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
use_physical= use to override Object-wide default for using a physical scale for output
+bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)
2014-06-16 - Changed to actually plot E rather than E/E_0 - Bovy (IAS)

galpy.orbit.Orbit.plotER

`Orbit.plotER(*args, **kwargs)`

NAME:

plotER

PURPOSE:

plot $E_R(.)$ along the orbit

INPUT:

pot= Potential instance or list of instances in which the orbit was integrated
d1= plot ER vs d1: e.g., 't', 'z', 'R'
normed= if set, plot $E(t)/E(0)$ rather than $E(t)$
ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
use_physical= use to override Object-wide default for using a physical scale for output
+bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotEz

`Orbit.plotEz(*args, **kwargs)`

NAME:

plotEz

PURPOSE:

plot $E_z(\cdot)$ along the orbit

INPUT:

pot= Potential instance or list of instances in which the orbit was integrated

d1= plot Ez vs d1: e.g., 't', 'z', 'R'

normed= if set, plot $E(t)/E(0)$ rather than $E(t)$

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

+bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotEzJz

`Orbit.plotEzJz(*args, **kwargs)`

NAME:

plotEzJz

PURPOSE:

plot $E_z(t)/\sqrt{\text{dens}(R)} / (E_z(0)/\sqrt{\text{dens}(R(0))})$ along the orbit (an approximation to the vertical action)

INPUT:

pot - Potential instance or list of instances in which the orbit was integrated

d1= plot Ez vs d1: e.g., 't', 'z', 'R'

+bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-08-08 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotphi

`Orbit.plotphi(*args, **kwargs)`

NAME:

plotphi

PURPOSE:

plot phi(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotR

`Orbit.plotR(*args, **kwargs)`

NAME:

plotR

PURPOSE:

plot R(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotvR

`Orbit.plotvR(*args, **kwargs)`

NAME:

plotvR

PURPOSE:

plot vR(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotvT

`Orbit.plotvT(*args, **kwargs)`

NAME:

plotvT

PURPOSE:

plot vT(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotvx

`Orbit.plotvx(*args, **kwargs)`

NAME:

plotvx

PURPOSE:

plot vx(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'
vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)
bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotvy

`Orbit.plotvy(*args, **kwargs)`

NAME:

plotvy

PURPOSE:

plot vy(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotvz

`Orbit.plotvz(*args, **kwargs)`

NAME:

plotvz

PURPOSE:

plot vz(.) along the orbit

INPUT: d1= plot vs d1: e.g., 't', 'z', 'R'

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotx

`Orbit.plotx(*args, **kwargs)`

NAME:

plotx

PURPOSE:

plot x(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.ploty

`Orbit.ploty(*args, **kwargs)`

NAME:

ploty

PURPOSE:

plot y(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.plotz

`Orbit.plotz(*args, **kwargs)`

NAME:

plotz

PURPOSE:

plot z(.) along the orbit

INPUT:

d1= plot vs d1: e.g., 't', 'z', 'R'

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

bovy_plot.bovy_plot inputs

OUTPUT:

figure to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.orbit.Orbit.pmbb

Orbit.**pmbb** (*args, **kwargs)

NAME:

pmbb

PURPOSE:

return proper motion in Galactic latitude (in mas/yr)

INPUT:

t - (optional) time at which to get pmbb (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) **position and velocity of observer** in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

pm_b(t) in mas/yr

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.pmdec

Orbit.**pmdec** (*args, **kwargs)

NAME:

pmdec

PURPOSE:

return proper motion in declination (in mas/yr)

INPUT:

t - (optional) time at which to get pmdec (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) **position and velocity of observer** in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantities) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

pm_dec(t) in mas/yr

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.pmll

Orbit.**pmll** (*args, **kwargs)

NAME:

pmll

PURPOSE:

return proper motion in Galactic longitude (in mas/yr)

INPUT:

t - (optional) time at which to get pmll (can be Quantity)

v obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer

in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantities) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

pm_l(t) in mas/yr

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.pmra

Orbit.**pmra** (*args, **kwargs)

NAME:

pmra

PURPOSE:

return proper motion in right ascension (in mas/yr)

INPUT:

t - (optional) time at which to get pmra (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantities) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

pm_ra(t) in mas/yr

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.r

Orbit.**r**(*args, **kwargs)

NAME:

r

PURPOSE:

return spherical radius at time t

INPUT:

t - (optional) time at which to get the radius (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

r(t)

HISTORY:

2016-04-19 - Written - Bovy (UofT)

galpy.orbit.Orbit.R

Orbit.**R**(*args, **kwargs)

NAME:

R

PURPOSE:

return cylindrical radius at time t

INPUT:

t - (optional) time at which to get the radius (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

R(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.ra

`Orbit.ra(*args, **kwargs)`

NAME:

ra

PURPOSE:

return the right ascension

INPUT:

t - (optional) time at which to get ra (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity) (default=[8.0,0.,0.]) OR Orbit object that corresponds to the orbit of the observer

(default=Object-wide default; can be Quantity)

Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

ra(t) in deg

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.orbit.Orbit.rap

`Orbit.rap(analytic=False, pot=None, **kwargs)`

NAME:

rap

PURPOSE:

calculate the apocenter radius, either numerically from the numerical orbit integration or using analytical means

INPUT:

analytic(= False) compute this analytically

pot - potential to use for analytical calculation

For 3D orbits different approximations for analytic=True are available (see the EccZmaxRperiRap method of actionAngle modules):

type= ('staeckel') type of actionAngle module to use

- 1) 'adiabatic': assuming motion splits into R and z
- 2) 'staeckel': assuming motion splits into u and v of prolate spheroidal coordinate system, exact for Staeckel potentials (incl. all spherical potentials)
- 3) 'spherical': for spherical potentials, exact

+actionAngle module setup kwargs for the corresponding actionAngle modules (actionAngleAdiabatic, actionAngleStaeckel, and actionAngleSpherical)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

R_ap

HISTORY:

2010-09-20 - Written - Bovy (NYU)

2017-12-25 - Added Staeckel approximation and made that the default - Bovy (UofT)

galpy.orbit.Orbit.resetaA

Orbit.**resetaA**(*pot=None, type=None*)

NAME:

resetaA

PURPOSE:

re-set up an actionAngle module for this Orbit

INPUT:

(none)

OUTPUT:

True if reset happened, False otherwise

HISTORY:

2014-01-06 - Written - Bovy (IAS)

galpy.orbit.Orbit.rguiding

Orbit.**rguiding**(*args, **kwargs)

NAME:

rguiding

PURPOSE:

calculate the guiding-center radius (the radius of a circular orbit with the same angular momentum)

INPUT:

pot= potential instance or list of such instances

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

R_guiding

HISTORY:

2018-08-29 - Written as thin wrapper around Potential.rl - Bovy (UofT)

galpy.orbit.Orbit.rperi

`Orbit.rperi` (*analytic=False, pot=None, **kwargs*)

NAME:

`rperi`

PURPOSE:

calculate the pericenter radius, either numerically from the numerical orbit integration or using analytical means

INPUT:

`analytic=False` compute this analytically

`pot` - potential to use for analytical calculation

For 3D orbits different approximations for `analytic=True` are available (see the `EccZmaxRperiRap` method of `actionAngle` modules):

`type=` ('staeckel') type of `actionAngle` module to use

- 1) 'adiabatic': assuming motion splits into R and z
- 2) 'staeckel': assuming motion splits into u and v of prolate spheroidal coordinate system, exact for Staeckel potentials (incl. all spherical potentials)
- 3) 'spherical': for spherical potentials, exact

+`actionAngle` module setup kwargs for the corresponding `actionAngle` modules (`actionAngleAdiabatic`, `actionAngleStaeckel`, and `actionAngleSpherical`)

`ro=` (Object-wide default) physical scale for distances to use to convert (can be `Quantity`)

`use_physical=` use to override Object-wide default for using a physical scale for output

OUTPUT:

`R_peri`

HISTORY:

2010-09-20 - Written - Bovy (NYU)

2017-12-25 - Added Staeckel approximation and made that the default - Bovy (UofT)

galpy.orbit.Orbit.setphi

`Orbit.setphi` (*phi*)

NAME:

`setphi`

PURPOSE:

set initial azimuth

INPUT:

`phi` - desired azimuth

OUTPUT:

(none)

HISTORY:

2010-08-01 - Written - Bovy (NYU)

BUGS:

Should perform check that this orbit has phi

galpy.orbit.Orbit.SkyCoord

`Orbit.SkyCoord(*args, **kwargs)`

NAME:

SkyCoord

PURPOSE:

return the position as an astropy SkyCoord

INPUT:

t - (optional) time at which to get the position (can be Quantity)

obs=[X,Y,Z] - (optional) position of observer (in kpc; entries can be Quantity)
(default=Object-wide default) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

OUTPUT:

SkyCoord(t)

HISTORY:

2015-06-02 - Written - Bovy (IAS)

galpy.orbit.Orbit.time

`Orbit.time(*args, **kwargs)`

NAME:

t

PURPOSE:

return the times at which the orbit is sampled

INPUT:

t - (default: integration times) time at which to get the time (for consistency reasons); default is to return the list of times at which the orbit is sampled (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

t(t)

HISTORY:

2014-06-11 - Written - Bovy (IAS)

galpy.orbit.Orbit.toLinear

`Orbit.toLinear()`

NAME:

toLinear

PURPOSE:

convert a 3D orbit into a 1D orbit (z)

INPUT:

(none)

OUTPUT:

linear Orbit

HISTORY:

2010-11-30 - Written - Bovy (NYU)

galpy.orbit.Orbit.toPlanar

`Orbit.toPlanar()`

NAME:

toPlanar

PURPOSE:

convert a 3D orbit into a 2D orbit

INPUT:

(none)

OUTPUT:

planar Orbit

HISTORY:

2010-11-30 - Written - Bovy (NYU)

galpy.orbit.Orbit.Tp

`Orbit.Tp(pot=None, **kwargs)`

NAME:

Tp

PURPOSE:

calculate the azimuthal period

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Tr

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.Tr

Orbit.**Tr** (*pot=None, **kwargs*)

NAME:

Tr

PURPOSE:

calculate the radial period

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Tr

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.TrTp

`Orbit.TrTp` (*pot=None, **kwargs*)

NAME:

TrTp

PURPOSE:

the 'ratio' between the radial and azimuthal period $Tr/T_{\phi} \pi$

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

OUTPUT:

$Tr/T_{\phi} \pi$

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.turn_physical_off

`Orbit.turn_physical_off()`

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2014-06-17 - Written - Bovy (IAS)

galpy.orbit.Orbit.turn_physical_on

`Orbit.turn_physical_on` (*ro=None, vo=None*)

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-01-19 - Written - Bovy (UofT)

galpy.orbit.Orbit.Tz

`Orbit.Tz` (*pot=None, **kwargs*)

NAME:

Tz

PURPOSE:

calculate the vertical period

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Tz

HISTORY:

2012-06-01 - Written - Bovy (IAS)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.U`Orbit.U(*args, **kwargs)`

NAME:

U

PURPOSE:

return Heliocentric Galactic rectangular x-velocity (aka “U”)

INPUT:

t - (optional) time at which to get U (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

U(t) in km/s

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.V`Orbit.V(*args, **kwargs)`

NAME:

V

PURPOSE:

return Heliocentric Galactic rectangular y-velocity (aka “V”)

INPUT:

t - (optional) time at which to get V (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

V(t) in km/s

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.vbb

Orbit.vbb(*args, **kwargs)

NAME:

vbb

PURPOSE:

return velocity in Galactic latitude (km/s)

INPUT:

t - (optional) time at which to get vbb (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

v_b(t) in km/s

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.vdec

Orbit.vdec(*args, **kwargs)

NAME:

vdec

PURPOSE:

return velocity in declination (km/s)

INPUT:

t - (optional) time at which to get vdec (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

v_dec(t) in km/s

HISTORY:

2011-03-27 - Written - Bovy (NYU)

galpy.orbit.Orbit.vll`Orbit.vll(*args, **kwargs)`

NAME:

vll

PURPOSE:

return the velocity in Galactic longitude (km/s)

INPUT:

t - (optional) time at which to get vll (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

v_l(t) in km/s

HISTORY:

2011-03-27 - Written - Bovy (NYU)

galpy.orbit.Orbit.vlos`Orbit.vlos(*args, **kwargs)`

NAME:

vlos

PURPOSE:

return the line-of-sight velocity (in km/s)

INPUT:

t - (optional) time at which to get vlos (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

vlos(t) in km/s

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.vphi

Orbit.**vphi** (*args, **kwargs)

NAME:

vphi

PURPOSE:

return angular velocity

INPUT:

t - (optional) time at which to get the angular velocity (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

vphi(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.vR

Orbit.**vR** (*args, **kwargs)

NAME:

vR

PURPOSE:

return radial velocity at time t

INPUT:

t - (optional) time at which to get the radial velocity

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

vR(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.vra

Orbit.**vra** (*args, **kwargs)

NAME:

vra

PURPOSE:

return velocity in right ascension (km/s)

INPUT:

`t` - (optional) time at which to get `vra` (can be Quantity)

`obs=[X,Y,Z,vx,vy,vz]` - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer `Y` is ignored and always assumed to be zero

`ro`= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

`vo`= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

`v_ra(t)` in km/s

HISTORY:

2011-03-27 - Written - Bovy (NYU)

galpy.orbit.Orbit.vT

`Orbit.vT(*args, **kwargs)`

NAME:

`vT`

PURPOSE:

return tangential velocity at time `t`

INPUT:

`t` - (optional) time at which to get the tangential velocity (can be Quantity)

`vo`= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

`use_physical`= use to override Object-wide default for using a physical scale for output

OUTPUT:

`vT(t)`

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.vx

`Orbit.vx(*args, **kwargs)`

NAME:

`vx`

PURPOSE:

return x velocity at time `t`

INPUT:

`t` - (optional) time at which to get the velocity (can be Quantity)

`vo`= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

`use_physical`= use to override Object-wide default for using a physical scale for output

OUTPUT:

$v_x(t)$

HISTORY:

2010-11-30 - Written - Bovy (NYU)

galpy.orbit.Orbit.vy

Orbit.**vy** (*args, **kwargs)

NAME:

vy

PURPOSE:

return y velocity at time t

INPUT:

t - (optional) time at which to get the velocity (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

$v_y(t)$

HISTORY:

2010-11-30 - Written - Bovy (NYU)

galpy.orbit.Orbit.vz

Orbit.**vz** (*args, **kwargs)

NAME:

vz

PURPOSE:

return vertical velocity

INPUT:

t - (optional) time at which to get the vertical velocity (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

$v_z(t)$

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.W`Orbit.W(*args, **kwargs)`

NAME:

W

PURPOSE:

return Heliocentric Galactic rectangular z-velocity (aka “W”)

INPUT:

t - (optional) time at which to get W (can be Quantity)

obs=[X,Y,Z,vx,vy,vz] - (optional) position and velocity of observer in the Galactocentric frame (in kpc and km/s) (default=[8.0,0.,0.,0.,220.,0.]; entries can be Quantity) OR Orbit object that corresponds to the orbit of the observer Y is ignored and always assumed to be zero

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

vo= (Object-wide default) physical scale for velocities to use to convert (can be Quantity)

OUTPUT:

W(t) in km/s

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.orbit.Orbit.wp`Orbit.wp(pot=None, **kwargs)`

NAME:

wp

PURPOSE:

calculate the azimuthal angle

INPUT:

pot - potential

type= (‘staeckel’) type of actionAngle module to use

1) ‘adiabatic’

2) ‘staeckel’

3) ‘isochroneApprox’

4) ‘spherical’

+actionAngle module setup kwargs

OUTPUT:

wp

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.wr

Orbit.**wr** (*pot=None, **kwargs*)

NAME:

wr

PURPOSE:

calculate the radial angle

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

OUTPUT:

wr

HISTORY:

2010-11-30 - Written - Bovy (NYU)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.wz

Orbit.**wz** (*pot=None, **kwargs*)

NAME:

wz

PURPOSE:

calculate the vertical angle

INPUT:

pot - potential

type= ('staeckel') type of actionAngle module to use

1) 'adiabatic'

2) 'staeckel'

3) 'isochroneApprox'

4) 'spherical'

+actionAngle module setup kwargs

OUTPUT:

wz

HISTORY:

2012-06-01 - Written - Bovy (IAS)

2013-11-27 - Re-written using new actionAngle modules - Bovy (IAS)

galpy.orbit.Orbit.x

`Orbit.x(*args, **kwargs)`

NAME:

x

PURPOSE:

return x

INPUT:

t - (optional) time at which to get x (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

x(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.y

`Orbit.y(*args, **kwargs)`

NAME:

y

PURPOSE:

return y

INPUT:

t - (optional) time at which to get y (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

y(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.z

`Orbit.z(*args, **kwargs)`

NAME:

z

PURPOSE:

return vertical height

INPUT:

t - (optional) time at which to get the vertical height (can be Quantity)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

z(t)

HISTORY:

2010-09-21 - Written - Bovy (NYU)

galpy.orbit.Orbit.zmax

`Orbit.zmax(analytic=False, pot=None, **kwargs)`

NAME:

zmax

PURPOSE:

calculate the maximum vertical height, either numerically from the numerical orbit integration or using analytical means

INPUT:

analytic(= False) compute this analytically

pot - potential to use for analytical calculation

For 3D orbits different approximations for analytic=True are available (see the `EccZmaxRperiRap` method of `actionAngle` modules):

type= ('staeckel') type of `actionAngle` module to use

- 1) 'adiabatic': assuming motion splits into R and z
- 2) 'staeckel': assuming motion splits into u and v of prolate spheroidal coordinate system, exact for Staeckel potentials (incl. all spherical potentials)
- 3) 'spherical': for spherical potentials, exact

+`actionAngle` module setup kwargs for the corresponding `actionAngle` modules (`actionAngleAdiabatic`, `actionAngleStaeckel`, and `actionAngleSpherical`)

ro= (Object-wide default) physical scale for distances to use to convert (can be Quantity)

use_physical= use to override Object-wide default for using a physical scale for output

OUTPUT:

Z_max

HISTORY:

2010-09-20 - Written - Bovy (NYU)

2017-12-25 - Added Staeckel approximation and made that the default - Bovy (UofT)

3.2 Potential (galpy.potential)

3.2.1 3D potentials

General instance routines

Use as `Potential-instance.method(...)`

galpy.potential.Potential.__call__

Warning: galpy potentials do *not* necessarily approach zero at infinity. To compute, for example, the escape velocity or whether or not an orbit is unbound, you need to take into account the value of the potential at infinity. E.g., $v_{\text{esc}}(r) = \sqrt{2[\Phi(\infty) - \Phi(r)]}$.

`Potential.__call__ (R, z, phi=0.0, t=0.0, dR=0, dphi=0)`

NAME:

`__call__`

PURPOSE:

evaluate the potential at (R,z,phi,t)

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

Phi(R,z,t)

HISTORY:

2010-04-16 - Written - Bovy (NYU)

galpy.potential.Potential.dens

`Potential.dens (R, z, phi=0.0, t=0.0, forcepoisson=False)`

NAME:

`dens`

PURPOSE:

evaluate the density $\rho(R,z,t)$

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

KEYWORDS:

forcepoisson= if True, calculate the density through the Poisson equation, even if an explicit expression for the density exists

OUTPUT:

$\rho(R,z,\phi,t)$

HISTORY:

2010-08-08 - Written - Bovy (NYU)

galpy.potential.Potential.dvcircdR

`Potential.dvcircdR(R, phi=None)`

NAME:

dvcircdR

PURPOSE:

calculate the derivative of the circular velocity at R wrt R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

phi= (None) azimuth to use for non-axisymmetric potentials

OUTPUT:

derivative of the circular rotation velocity wrt R

HISTORY:

2013-01-08 - Written - Bovy (IAS)

2016-06-28 - Added phi= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.Potential.epifreq

`Potential.epifreq(R)`

NAME:

epifreq

PURPOSE:

calculate the epicycle frequency at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

epicycle frequency

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.Potential.flattening

Potential.**flattening**(*R*, *z*)

NAME:

flattening

PURPOSE:

calculate the potential flattening, defined as $\sqrt{f_{\text{abs}}(z/R \ F_R/F_z)}$

INPUT:

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

OUTPUT:

flattening

HISTORY:

2012-09-13 - Written - Bovy (IAS)

galpy.potential.Potential.lindbladR

Potential.**lindbladR**(*OmegaP*, *m=2*, ***kwargs*)

NAME:

lindbladR

PURPOSE:

calculate the radius of a Lindblad resonance

INPUT:

OmegaP - pattern speed (can be Quantity)

m= order of the resonance (as in $m(\text{O-Op})=\kappa$ (negative m for outer) use m='corotation'
for corotation +scipy.optimize.brentq xtol,rtol,maxiter kwargs

OUTPUT:

radius of Linblad resonance, None if there is no resonance

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.Potential.mass

Potential.**mass** (*R, z=None, t=0.0, forceint=False*)

NAME:

mass

PURPOSE:

evaluate the mass enclosed

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z= (None) vertical height (can be Quantity)

t - time (optional; can be Quantity)

KEYWORDS:

forceint= if True, calculate the mass through integration of the density, even if an explicit expression for the mass exists

OUTPUT:

- 1) for spherical potentials: $M(<R)$ [or if z is None], when the mass is implemented explicitly, the mass enclosed within $r = \sqrt{R^2 + z^2}$ is returned when not z is None; forceint will integrate between $-z$ and z , so the two are inconsistent (If you care to have this changed, raise an issue on github)
- 2) for axisymmetric potentials: $M(<R, <f_{\text{abs}}(Z))$

HISTORY:

2014-01-29 - Written - Bovy (IAS)

galpy.potential.Potential.nemo_accname

Potential.**nemo_accname** ()

NAME:

nemo_accname

PURPOSE:

return the accname potential name for use of this potential with NEMO

INPUT:

(none)

OUTPUT:

Acceleration name

HISTORY:

2014-12-18 - Written - Bovy (IAS)

galpy.potential.Potential.nemo_accpars

Potential.**nemo_accpars** (*vo*, *ro*)

NAME:

nemo_accpars

PURPOSE:

return the accpars potential parameters for use of this potential with NEMO

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

accpars string

HISTORY:

2014-12-18 - Written - Bovy (IAS)

galpy.potential.Potential.omegac

Potential.**omegac** (*R*)

NAME:

omegac

PURPOSE:

calculate the circular angular speed at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

circular angular speed

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.Potential.phiforce

Potential.**phiforce** (*R*, *z*, *phi*=0.0, *t*=0.0)

NAME:

phiforce

PURPOSE:

evaluate the azimuthal force F_{ϕ} (*R*,*z*,*phi*,*t*)

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (rad; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

F_phi (R,z,phi,t)

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.potential.Potential.phi2deriv

Potential.**phi2deriv** (R, Z, phi=0.0, t=0.0)

NAME:

phi2deriv

PURPOSE:

evaluate the second azimuthal derivative

INPUT:

R - Galactocentric radius (can be Quantity)

Z - vertical height (can be Quantity)

phi - Galactocentric azimuth (can be Quantity)

t - time (can be Quantity)

OUTPUT:

d2Phi/dphi2

HISTORY:

2013-09-24 - Written - Bovy (IAS)

galpy.potential.Potential.plot

Potential.**plot** (t=0.0, rmin=0.0, rmax=1.5, nrs=21, zmin=-0.5, zmax=0.5, nzs=21, effective=False, Lz=None, phi=None, xy=False, xrange=None, yrange=None, justcontours=False, levels=None, cntrcolors=None, ncontours=21, savefilename=None)

NAME:

plot

PURPOSE:

plot the potential

INPUT:

t= time to plot potential at

rmin= minimum R (can be Quantity) [xmin if xy]

rmax= maximum R (can be Quantity) [ymax if xy]

nrs= grid in R
 zmin= minimum z (can be Quantity) [ymin if xy]
 zmax= maximum z (can be Quantity) [ymax if xy]
 nzs= grid in z
 phi= (None) azimuth to use for non-axisymmetric potentials
 xy= (False) if True, plot the potential in X-Y
 effective= (False) if True, plot the effective potential $\Phi + L_z^2/2/R^2$
 Lz= (None) angular momentum to use for the effective potential when effective=True
 justcontours= (False) if True, just plot contours
 savefilename - save to or restore from this savefile (pickle)
 xrange, yrange= can be specified independently from rmin,zmin, etc.
 levels= (None) contours to plot
 ncontours - number of contours when levels is None
 cntcolors= (None) colors of the contours (single color or array with length ncontours)

OUTPUT:

plot to output device

HISTORY:

2010-07-09 - Written - Bovy (NYU)
 2014-04-08 - Added effective= - Bovy (IAS)

galpy.potential.Potential.plotDensity

Potential.plotDensity (*t=0.0, rmin=0.0, rmax=1.5, nrs=21, zmin=-0.5, zmax=0.5, nzs=21, phi=None, xy=False, ncontours=21, savefilename=None, aspect=None, log=False, justcontours=False*)

NAME:

plotDensity

PURPOSE:

plot the density of this potential

INPUT:

t= time to plot potential at
 rmin= minimum R (can be Quantity) [xmin if xy]
 rmax= maximum R (can be Quantity) [ymax if xy]
 nrs= grid in R
 zmin= minimum z (can be Quantity) [ymin if xy]
 zmax= maximum z (can be Quantity) [ymax if xy]
 nzs= grid in z
 phi= (None) azimuth to use for non-axisymmetric potentials

xy= (False) if True, plot the density in X-Y
ncontours= number of contours
justcontours= (False) if True, just plot contours
savefilename= save to or restore from this savefile (pickle)
log= if True, plot the log density

OUTPUT:

plot to output device

HISTORY:

2014-01-05 - Written - Bovy (IAS)

galpy.potential.Potential.plotEscapecurve

Potential.**plotEscapecurve** (*args, **kwargs)

NAME:

plotEscapecurve

PURPOSE:

plot the escape velocity curve for this potential (in the $z=0$ plane for non-spherical potentials)

INPUT:

Range - range (can be Quantity)
grid= number of points to plot
savefilename= save to or restore from this savefile (pickle)
+bovy_plot(*args,**kwargs)

OUTPUT:

plot to output device

HISTORY:

2010-08-08 - Written - Bovy (NYU)

galpy.potential.Potential.plotRotcurve

Potential.**plotRotcurve** (*args, **kwargs)

NAME:

plotRotcurve

PURPOSE:

plot the rotation curve for this potential (in the $z=0$ plane for non-spherical potentials)

INPUT:

Range - range (can be Quantity)
grid= number of points to plot
savefilename=- save to or restore from this savefile (pickle)


```
+bovy_plot(*args,**kwargs)
```

OUTPUT:

plot to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.potential.Potential.Rphideriv

`Potential.Rphideriv` ($R, Z, \text{phi}=0.0, t=0.0$)

NAME:

Rphideriv

PURPOSE:

evaluate the mixed radial, azimuthal derivative

INPUT:

R - Galactocentric radius (can be Quantity)

Z - vertical height (can be Quantity)

phi - Galactocentric azimuth (can be Quantity)

t - time (can be Quantity)

OUTPUT:

$d^2\Phi/d\phi dR$

HISTORY:

2014-06-30 - Written - Bovy (IAS)

galpy.potential.Potential.R2deriv

`Potential.R2deriv` ($R, Z, \text{phi}=0.0, t=0.0$)

NAME:

R2deriv

PURPOSE:

evaluate the second radial derivative

INPUT:

R - Galactocentric radius (can be Quantity)

Z - vertical height (can be Quantity)

phi - Galactocentric azimuth (can be Quantity)

t - time (can be Quantity)

OUTPUT:

$d^2\Phi/dR^2$

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.Potential.r2deriv

`Potential.r2deriv` ($R, z, \text{phi}=0.0, t=0.0$)

NAME:

r2deriv

PURPOSE:

evaluate the second spherical radial derivative

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

d²phi/dr²

HISTORY:

2018-03-21 - Written - Webb (UofT)

galpy.potential.Potential.Rzderiv

`Potential.Rzderiv` ($R, Z, \text{phi}=0.0, t=0.0$)

NAME:

Rzderiv

PURPOSE:

evaluate the mixed R,z derivative

INPUT:

R - Galactocentric radius (can be Quantity)

Z - vertical height (can be Quantity)

phi - Galactocentric azimuth (can be Quantity)

t - time (can be Quantity)

OUTPUT:

d²phi/dz/dR

HISTORY:

2013-08-26 - Written - Bovy (IAS)

galpy.potential.Potential.Rforce

Potential.**Rforce** (*R, z, phi=0.0, t=0.0*)

NAME:

Rforce

PURPOSE:

evaluate cylindrical radial force $F_R(R, z)$

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

$F_R(R, z, \text{phi}, t)$

HISTORY:

2010-04-16 - Written - Bovy (NYU)

galpy.potential.Potential.rforce

Potential.**rforce** (**args, **kwargs*)

NAME:

rforce

PURPOSE:

evaluate spherical radial force $F_r(R, z)$

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

v - current velocity in cylindrical coordinates (optional, but required when including dissipative forces; can be a Quantity)

OUTPUT:

$F_r(R, z, \text{phi}, t)$

HISTORY:

2016-06-20 - Written - Bovy (UofT)

galpy.potential.Potential.rl

Potential.**rl**(lz)

NAME:

rl

PURPOSE:

calculate the radius of a circular orbit of Lz

INPUT:

lz - Angular momentum (can be Quantity)

OUTPUT:

radius

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

NOTE:

seems to take about ~0.5 ms for a Miyamoto-Nagai potential; ~0.75 ms for a MWPotential

galpy.potential.Potential.rtide

Potential.**rtide**(R, z, phi=0.0, t=0.0, M=None)

NAME:

rtide

PURPOSE:

Calculate the tidal radius for object of mass M assuming a circular orbit as

$$r_t^3 = \frac{GM_s}{\Omega^2 - d^2\Phi/dr^2}$$

where M_s is the cluster mass, Ω is the circular frequency, and Φ is the gravitational potential. For non-spherical potentials, we evaluate $\Omega^2 = (1/r)(d\Phi/dr)$ and evaluate the derivatives at the given position of the cluster.

INPUT:

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

M - (default = None) Mass of object (can be Quantity)

OUTPUT:

Tidal Radius

HISTORY:

2018-03-21 - Written - Webb (UofT)

galpy.potential.Potential.surfdens

`Potential.surfdens` ($R, z, \text{phi}=0.0, t=0.0, \text{forcepoisson}=False$)

NAME:

surfdens

PURPOSE:

evaluate the surface density $\Sigma(R, z, \phi, t) = \int_{-z}^{+z} dz' \rho(R, z', \phi, t)$

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

KEYWORDS:

forcepoisson= if True, calculate the surface density through the Poisson equation, even if an explicit expression for the surface density exists

OUTPUT:

Sigma (R,z,phi,t)

HISTORY:

2018-08-19 - Written - Bovy (UofT)

galpy.planar.Potential.toPlanar

`Potential.toPlanar` ()

NAME:

toPlanar

PURPOSE:

convert a 3D potential into a planar potential in the mid-plane

INPUT:

(none)

OUTPUT:

planarPotential

HISTORY:

unknown

galpy.potential.Potential.toVertical

`Potential.toVertical` (R)

NAME:

toVertical

PURPOSE:

convert a 3D potential into a linear (vertical) potential at R

INPUT:

R - Galactocentric radius at which to create the vertical potential (can be Quantity)

OUTPUT:

linear (vertical) potential

HISTORY

unknown

galpy.potential.Potential.ttensor

`Potential.ttensor(R, z, phi=0.0, t=0.0, eigenval=False)`

NAME:

ttensor

PURPOSE:

Calculate the tidal tensor $T_{ij} = -d(\Psi)(dx_i dx_j)$

INPUT:

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

eigenval - return eigenvalues if true (optional; boolean)

OUTPUT:

Tidal Tensor

HISTORY:

2018-03-21 - Written - Webb (UofT)

galpy.potential.Potential.turn_physical_off

`Potential.turn_physical_off()`

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.Potential.turn_physical_on

`Potential.turn_physical_on` (*ro=None, vo=None*)

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.Potential.vcirc

`Potential.vcirc` (*R, phi=None*)

NAME:

vcirc

PURPOSE:

calculate the circular velocity at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

phi= (None) azimuth to use for non-axisymmetric potentials

OUTPUT:

circular rotation velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

2016-06-15 - Added phi= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.Potential.verticalfreq

`Potential.verticalfreq` (*R*)

NAME:

verticalfreq

PURPOSE:

calculate the vertical frequency at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

vertical frequency

HISTORY:

2012-07-25 - Written - Bovy (IAS@MPIA)

galpy.potential.Potential.vesc

Potential.**vesc**(*R*)

NAME:

vesc

PURPOSE:

calculate the escape velocity at R for this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

escape velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.Potential.vterm

Potential.**vterm**(*l*, *deg=True*)

NAME:

vterm

PURPOSE:

calculate the terminal velocity at l in this potential

INPUT:

l - Galactic longitude [deg/rad; can be Quantity]

deg= if True (default), l in deg

OUTPUT:

terminal velocity

HISTORY:

2013-05-31 - Written - Bovy (IAS)

galpy.potential.Potential.z2deriv

`Potential.z2deriv` (*R*, *Z*, *phi*=0.0, *t*=0.0)

NAME:

z2deriv

PURPOSE:

evaluate the second vertical derivative

INPUT:

R - Galactocentric radius (can be Quantity)

Z - vertical height (can be Quantity)

phi - Galactocentric azimuth (can be Quantity)

t - time (can be Quantity)

OUTPUT:

$d^2\phi/dz^2$

HISTORY:

2012-07-25 - Written - Bovy (IAS@MPIA)

galpy.potential.Potential.zforce

`Potential.zforce` (*R*, *z*, *phi*=0.0, *t*=0.0)

NAME:

zforce

PURPOSE:

evaluate the vertical force F_z (*R*,*z*,*t*)

INPUT:

R - Cylindrical Galactocentric radius (can be Quantity)

z - vertical height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

F_z (*R*,*z*,*phi*,*t*)

HISTORY:

2010-04-16 - Written - Bovy (NYU)

In addition to these, the `NFWPotential` also has methods to calculate virial quantities

galpy.potential.Potential.conc

Potential.**conc** (*H=70.0, Om=0.3, overdens=200.0, wrtcrit=False, ro=None, vo=None*)

NAME:

conc

PURPOSE:

return the concentration

INPUT:

H= (default: 70) Hubble constant in km/s/Mpc

Om= (default: 0.3) Omega matter

overdens= (200) overdensity which defines the virial radius

wrtcrit= (False) if True, the overdensity is wrt the critical density rather than the mean matter density

ro= distance scale in kpc or as Quantity (default: object-wide, which if not set is 8 kpc))

vo= velocity scale in km/s or as Quantity (default: object-wide, which if not set is 220 km/s))

OUTPUT:

concentration (scale/rvir)

HISTORY:

2014-04-03 - Written - Bovy (IAS)

galpy.potential.Potential.mvir

Potential.**mvir** (*H=70.0, Om=0.3, overdens=200.0, wrtcrit=False, forceint=False, ro=None, vo=None, use_physical=False*)

NAME:

mvir

PURPOSE:

calculate the virial mass

INPUT:

H= (default: 70) Hubble constant in km/s/Mpc

Om= (default: 0.3) Omega matter

overdens= (200) overdensity which defines the virial radius

wrtcrit= (False) if True, the overdensity is wrt the critical density rather than the mean matter density

ro= distance scale in kpc or as Quantity (default: object-wide, which if not set is 8 kpc))

vo= velocity scale in km/s or as Quantity (default: object-wide, which if not set is 220 km/s))

KEYWORDS:

forceint= if True, calculate the mass through integration of the density, even if an explicit expression for the mass exists

OUTPUT:

M(<rvir)

HISTORY:

2014-09-12 - Written - Bovy (IAS)

galpy.potential.NFWPotential.rvir

```
NFWPotential.rvir(H=70.0,  Om=0.3,  overdens=200.0,  wrtcrit=False,  ro=None,  vo=None,
                  use_physical=False)
```

NAME:

 \dot{r}_{vir}

PURPOSE:

calculate the virial radius for this density distribution

INPUT:

H= (default: 70) Hubble constant in km/s/Mpc

Om= (default: 0.3) Omega matter

overdens= (200) overdensity which defines the virial radius

wrtcrit= (False) if True, the overdensity is wrt the critical density rather than the mean matter density

ro= distance scale in kpc or as Quantity (default: object-wide, which if not set is 8 kpc))

vo= velocity scale in km/s or as Quantity (default: object-wide, which if not set is 220 km/s))

OUTPUT:

virial radius

HISTORY:

2014-01-29 - Written - Bovy (IAS)

General 3D potential routines

Use as `method(...)`

galpy.potential.dvcircdR

```
galpy.potential.dvcircdR(Pot, R, phi=None)
```

NAME:

dvcircdR

PURPOSE:

calculate the derivative of the circular velocity wrt R at R in potential Pot

INPUT:

Pot - Potential instance or list of such instances

R - Galactocentric radius (can be Quantity)

phi= (None) azimuth to use for non-axisymmetric potentials

OUTPUT:

derivative of the circular rotation velocity wrt R

HISTORY:

2013-01-08 - Written - Bovy (IAS)

2016-06-28 - Added phi= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.epifreq

`galpy.potential.epifreq(Pot, R)`

NAME:

epifreq

PURPOSE:

calculate the epicycle frequency at R in the potential Pot

INPUT:

Pot - Potential instance or list thereof

R - Galactocentric radius (can be Quantity)

OUTPUT:

epicycle frequency

HISTORY:

2012-07-25 - Written - Bovy (IAS)

galpy.potential.evaluateDensities

`galpy.potential.evaluateDensities(Pot, R, z, phi=None, t=0.0, forcepoisson=False)`

NAME:

evaluateDensities

PURPOSE:

convenience function to evaluate a possible sum of densities

INPUT:

Pot - potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (can be Quantity)

t - time (can be Quantity)

forcepoisson= if True, calculate the density through the Poisson equation, even if an explicit expression for the density exists

OUTPUT:

rho(R,z)

HISTORY:

2010-08-08 - Written - Bovy (NYU)

2013-12-28 - Added forcepoisson - Bovy (IAS)

galpy.potential.evaluatephiforces

`galpy.potential.evaluatephiforces` (*Pot, R, z, phi=None, t=0.0, v=None*)

NAME:

evaluatephiforces

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT: Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

v - current velocity in cylindrical coordinates (optional, but required when including dissipative forces; can be a Quantity)

OUTPUT:

F_phi(R,z,phi,t)

HISTORY:

2010-04-16 - Written - Bovy (NYU)

2018-03-16 - Added velocity input for dissipative forces - Bovy (UofT)

galpy.potential.evaluatePotentials

Warning: galpy potentials do *not* necessarily approach zero at infinity. To compute, for example, the escape velocity or whether or not an orbit is unbound, you need to take into account the value of the potential at infinity. E.g., $v_{\text{esc}}(r) = \sqrt{2[\Phi(\infty) - \Phi(r)]}$.

`galpy.potential.evaluatePotentials` (*Pot, R, z, phi=None, t=0.0, dR=0, dphi=0*)

NAME:

evaluatePotentials

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (can be Quantity)

t - time (can be Quantity)

dR= dphi=, if set to non-zero integers, return the dR, dphi't derivative instead

OUTPUT:

Phi(R,z)

HISTORY:

2010-04-16 - Written - Bovy (NYU)

galpy.potential.evaluatephi2derivs

`galpy.potential.evaluatephi2derivs` (*Pot, R, z, phi=None, t=0.0*)

NAME:

evaluatephi2derivs

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

d2Phi/d2phi(R,z,phi,t)

HISTORY:

2018-03-28 - Written - Bovy (UofT)

galpy.potential.evaluateRphiderivs

`galpy.potential.evaluateRphiderivs` (*Pot, R, z, phi=None, t=0.0*)

NAME:

evaluateRphiderivs

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

$d^2\Phi/d^2R(R,z,\phi,t)$

HISTORY:

2012-07-25 - Written - Bovy (IAS)

galpy.potential.evaluateR2derivs

`galpy.potential.evaluateR2derivs` (*Pot, R, z, phi=None, t=0.0*)

NAME:

evaluateR2derivs

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

$d^2\Phi/d^2R(R,z,\phi,t)$

HISTORY:

2012-07-25 - Written - Bovy (IAS)

galpy.potential.evaluater2derivs

`galpy.potential.evaluater2derivs` (*Pot, R, z, phi=None, t=0.0*)

NAME:

evaluater2derivs

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

$d^2\phi/dr^2(R,z,\phi,t)$

HISTORY:

2018-03-28 - Written - Bovy (UofT)

galpy.potential.evaluateRzderivs

`galpy.potential.evaluateRzderivs` (*Pot, R, z, phi=None, t=0.0*)

NAME:

`evaluateRzderivs`

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

$d^2\Phi/dz/dR(R,z,\phi,t)$

HISTORY:

2013-08-28 - Written - Bovy (IAS)

galpy.potential.evaluateRforces

`galpy.potential.evaluateRforces` (*Pot, R, z, phi=None, t=0.0, v=None*)

NAME:

`evaluateRforce`

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity))

t - time (optional; can be Quantity)

v - current velocity in cylindrical coordinates (optional, but required when including dissipative forces; can be a Quantity)

OUTPUT:

$F_R(R, z, \phi, t)$

HISTORY:

2010-04-16 - Written - Bovy (NYU)

2018-03-16 - Added velocity input for dissipative forces - Bovy (UofT)

galpy.potential.evaluateforces

`galpy.potential.evaluateforces` (*Pot, R, z, phi=None, t=0.0, v=None*)

NAME:

evaluateforces

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

v - current velocity in cylindrical coordinates (optional, but required when including dissipative forces; can be a Quantity)

OUTPUT:

$F_r(R, z, \phi, t)$

HISTORY:

2016-06-10 - Written - Bovy (UofT)

galpy.potential.evaluateSurfaceDensities

`galpy.potential.evaluateSurfaceDensities` (*Pot, R, z, phi=None, t=0.0, forcepoisson=False*)

NAME:

evaluateSurfaceDensities

PURPOSE:

convenience function to evaluate a possible sum of surface densities

INPUT:

Pot - potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (can be Quantity)

t - time (can be Quantity)

forcepoisson= if True, calculate the surface density through the Poisson equation, even if an explicit expression for the surface density exists

OUTPUT:

Sigma(R,z)

HISTORY:

2018-08-20 - Written - Bovy (UofT)

galpy.potential.evaluatez2derivs

galpy.potential.**evaluatez2derivs** (*Pot, R, z, phi=None, t=0.0*)

NAME:

evaluatez2derivs

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials (dissipative forces in such a list are ignored)

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

OUTPUT:

d2Phi/d2z(R,z,phi,t)

HISTORY:

2012-07-25 - Written - Bovy (IAS)

galpy.potential.evaluatezforces

galpy.potential.**evaluatezforces** (*Pot, R, z, phi=None, t=0.0, v=None*)

NAME:

evaluatezforces

PURPOSE:

convenience function to evaluate a possible sum of potentials

INPUT:

Pot - a potential or list of potentials

R - cylindrical Galactocentric distance (can be Quantity)

z - distance above the plane (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

v - current velocity in cylindrical coordinates (optional, but required when including dissipative forces; can be a Quantity)

OUTPUT:

$F_z(R, z, \phi, t)$

HISTORY:

2010-04-16 - Written - Bovy (NYU)

2018-03-16 - Added velocity input for dissipative forces - Bovy (UofT)

galpy.potential.flatten

`galpy.potential.flatten` (*Pot*)

NAME:

flatten

PURPOSE:

flatten a possibly nested list of Potential instances into a flat list

INPUT:

Pot - list (possibly nested) of Potential instances

OUTPUT:

Flattened list of Potential instances

HISTORY:

2018-03-14 - Written - Bovy (UofT)

galpy.potential.flattening

`galpy.potential.flattening` (*Pot, R, z*)

NAME:

flattening

PURPOSE:

calculate the potential flattening, defined as $\sqrt{\text{fabs}(z/R F_R/F_z)}$

INPUT:

Pot - Potential instance or list thereof

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

OUTPUT:

flattening

HISTORY:

2012-09-13 - Written - Bovy (IAS)

galpy.potential.lindbladR

galpy.potential.**lindbladR** (*Pot*, *OmegaP*, *m*=2, ***kwargs*)

NAME:

lindbladR

PURPOSE:

calculate the radius of a Lindblad resonance

INPUT:

Pot - Potential instance or list of such instances

OmegaP - pattern speed (can be Quantity)

m= order of the resonance (as in m(O-Op)=kappa (negative m for outer) use m='corotation' for corotation

+scipy.optimize.brentq xtol,rtol,maxiter kwargs

OUTPUT:

radius of Linblad resonance, None if there is no resonance

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.nemo_accname

galpy.potential.**nemo_accname** (*Pot*)

NAME:

nemo_accname

PURPOSE:

return the accname potential name for use of this potential or list of potentials with NEMO

INPUT:

Pot - Potential instance or list of such instances

OUTPUT:

Acceleration name in the correct format to give to accname=

HISTORY:

2014-12-18 - Written - Bovy (IAS)

galpy.potential.nemo_accpars

galpy.potential.**nemo_accpars** (*Pot*, *vo*, *ro*)

NAME:

nemo_accpars

PURPOSE:

return the accpars potential parameters for use of this potential or list of potentials with NEMO

INPUT:

Pot - Potential instance or list of such instances

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

accpars string in the correct format to give to accpars

HISTORY:

2014-12-18 - Written - Bovy (IAS)

galpy.potential.omegac

`galpy.potential.omegac` (*Pot*, *R*)

NAME:

omegac

PURPOSE:

calculate the circular angular speed velocity at *R* in potential *Pot*

INPUT:

Pot - Potential instance or list of such instances

R - Galactocentric radius (can be Quantity)

OUTPUT:

circular angular speed

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.plotDensities

`galpy.potential.plotDensities` (*Pot*, *rmin*=0.0, *rmax*=1.5, *nrs*=21, *zmin*=-0.5, *zmax*=0.5, *nzs*=21, *phi*=None, *xy*=False, *t*=0.0, *ncontours*=21, *savefilename*=None, *aspect*=None, *log*=False, *justcontours*=False)

NAME:

plotDensities

PURPOSE:

plot the density a set of potentials

INPUT:

Pot - Potential or list of Potential instances

rmin= minimum *R* (can be Quantity) [*xmin* if *xy*]

rmax= maximum *R* (can be Quantity) [*ymin* if *xy*]

nrs= grid in *R*

zmin= minimum *z* (can be Quantity) [*ymin* if *xy*]

zmax= maximum z (can be Quantity) [ymax if xy]
nzs= grid in z
phi= (None) azimuth to use for non-axisymmetric potentials
t= (0.) time to use to evaluate potential
xy= (False) if True, plot the density in X-Y
ncontours= number of contours
justcontours= (False) if True, just plot contours
savefilename= save to or restore from this savefile (pickle)
log= if True, plot the log density

OUTPUT:

plot to output device

HISTORY:

2013-07-05 - Written - Bovy (IAS)

galpy.potential.plotEscapecurve

galpy.potential.**plotEscapecurve** (*Pot*, *args, **kwargs)

NAME:

plotEscapecurve

PURPOSE:

plot the escape velocity curve for this potential (in the z=0 plane for non-spherical potentials)

INPUT:

Pot - Potential or list of Potential instances
Rrange= Range in R to consider (can be Quantity)
grid= grid in R
savefilename= save to or restore from this savefile (pickle)
+bovy_plot.bovy_plot args and kwargs

OUTPUT:

plot to output device

HISTORY:

2010-08-08 - Written - Bovy (NYU)

galpy.potential.plotPotentials

galpy.potential.**plotPotentials** (*Pot*, *rmin*=0.0, *rmax*=1.5, *nrs*=21, *zmin*=-0.5, *zmax*=0.5, *nzs*=21, *phi*=None, *xy*=False, *t*=0.0, *effective*=False, *Lz*=None, *ncontours*=21, *savefilename*=None, *aspect*=None, *justcontours*=False, *levels*=None, *cntrcolors*=None)

NAME:

plotPotentials

PURPOSE:

plot a set of potentials

INPUT:

Pot - Potential or list of Potential instances

rmin= minimum R (can be Quantity) [xmin if xy]

rmax= maximum R (can be Quantity) [ymax if xy]

nrs= grid in R

zmin= minimum z (can be Quantity) [ymin if xy]

zmax= maximum z (can be Quantity) [ymax if xy]

nzs= grid in z

phi= (None) azimuth to use for non-axisymmetric potentials

t= (0.) time to use to evaluate potential

xy= (False) if True, plot the potential in X-Y

effective= (False) if True, plot the effective potential $\Phi + L_z^2/2/R^2$

Lz= (None) angular momentum to use for the effective potential when effective=True

justcontours= (False) if True, just plot contours

levels= (None) contours to plot

ncontours - number of contours when levels is None

cntrcolors= (None) colors of the contours (single color or array with length ncontours)

savefilename= save to or restore from this savefile (pickle)

OUTPUT:

plot to output device

HISTORY:

2010-07-09 - Written - Bovy (NYU)

galpy.potential.plotRotcurve

galpy.potential.plotRotcurve (Pot, *args, **kwargs)

NAME:

plotRotcurve

PURPOSE:

plot the rotation curve for this potential (in the $z=0$ plane for non-spherical potentials)

INPUT:

Pot - Potential or list of Potential instances

Range - Range in R to consider (needs to be in the units that you are plotting; can be Quantity)

grid= grid in R

phi= (None) azimuth to use for non-axisymmetric potentials

savefilename= save to or restore from this savefile (pickle)

+bovy_plot.bovy_plot args and kwargs

OUTPUT:

plot to output device

HISTORY:

2010-07-10 - Written - Bovy (NYU)

2016-06-15 - Added phi= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.rl

`galpy.potential.rl` (*Pot, lz*)

NAME:

rl

PURPOSE:

calculate the radius of a circular orbit of Lz

INPUT:

Pot - Potential instance or list thereof

lz - Angular momentum (can be Quantity)

OUTPUT:

radius

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

NOTE:

seems to take about ~0.5 ms for a Miyamoto-Nagai potential; ~0.75 ms for a MWPotential

galpy.potential.rtide

`galpy.potential.rtide` (*Pot, R, z, phi=0.0, t=0.0, M=None*)

NAME:

rtide

PURPOSE:

Calculate the tidal radius for object of mass M assuming a circular orbit as

$$r_t^3 = \frac{GM_s}{\Omega^2 - d^2\Phi/dr^2}$$

where M_s is the cluster mass, Ω is the circular frequency, and Φ is the gravitational potential. For non-spherical potentials, we evaluate $\Omega^2 = (1/r)(d\Phi/dr)$ and evaluate the derivatives at the given position of the cluster.

INPUT:

Pot - Potential instance or list of such instances

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

M - (default = None) Mass of object (can be Quantity)

OUTPUT:

Tidal Radius

HISTORY:

2018-03-21 - Written - Webb (UofT)

galpy.potential.ttensor

`galpy.potential.ttensor (Pot, R, z, phi=0.0, t=0.0, eigenval=False)`

NAME:

ttensor

PURPOSE:

Calculate the tidal tensor $T_{ij} = -d(\Psi)(dx_i dx_j)$

INPUT:

Pot - Potential instance or list of such instances

R - Galactocentric radius (can be Quantity)

z - height (can be Quantity)

phi - azimuth (optional; can be Quantity)

t - time (optional; can be Quantity)

eigenval - return eigenvalues if true (optional; boolean)

OUTPUT:

Tidal Tensor

HISTORY:

2018-03-21 - Written - Webb (UofT)

galpy.potential.turn_physical_off

`galpy.potential.turn_physical_off (Pot)`

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.turn_physical_on

`galpy.potential.turn_physical_on` (*Pot*, *ro=None*, *vo=None*)

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.vcirc

`galpy.potential.vcirc` (*Pot*, *R*, *phi=None*)

NAME:

vcirc

PURPOSE:

calculate the circular velocity at *R* in potential *Pot*

INPUT:

Pot - Potential instance or list of such instances

R - Galactocentric radius (can be Quantity)

phi= (None) azimuth to use for non-axisymmetric potentials

OUTPUT:

circular rotation velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

2016-06-15 - Added *phi*= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.verticalfreq

`galpy.potential.verticalfreq(Pot, R)`

NAME:

verticalfreq

PURPOSE:

calculate the vertical frequency at R in the potential Pot

INPUT:

Pot - Potential instance or list thereof

R - Galactocentric radius (can be Quantity)

OUTPUT:

vertical frequency

HISTORY:

2012-07-25 - Written - Bovy (IAS@MPIA)

galpy.potential.vesc

`galpy.potential.vesc(Pot, R)`

NAME:

vesc

PURPOSE:

calculate the escape velocity at R for potential Pot

INPUT:

Pot - Potential instances or list thereof

R - Galactocentric radius (can be Quantity)

OUTPUT:

escape velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.vterm

`galpy.potential.vterm(Pot, l, deg=True)`

NAME:

vterm

PURPOSE:

calculate the terminal velocity at l in this potential

INPUT:

Pot - Potential instance

l - Galactic longitude [deg/rad; can be Quantity]

deg= if True (default), l in deg

OUTPUT:

terminal velocity

HISTORY:

2013-05-31 - Written - Bovy (IAS)

In addition to these, the following methods are available to compute expansion coefficients for the `SCFPotential` class for a given density

galpy.potential.scf_compute_coeffs

Note: This function computes A_{cos} and A_{sin} as defined in [Hernquist & Ostriker \(1992\)](#), except that we multiply A_{cos} and A_{sin} by 2 such that the density from *Galpy's Hernquist Potential* corresponds to $A_{cos} = \delta_{0n}\delta_{0l}\delta_{0m}$ and $A_{sin} = 0$.

For a given $\rho(R, z, \phi)$ we can compute A_{cos} and A_{sin} through the following equation

$$\begin{bmatrix} A_{cos} \\ A_{sin} \end{bmatrix}_{nlm} = \frac{4a^3}{I_{nl}} \int_{\xi=0}^{\infty} \int_{\cos(\theta)=-1}^1 \int_{\phi=0}^{2\pi} (1+\xi)^2 (1-\xi)^{-4} \rho(R, z, \phi) \Phi_{nlm}(\xi, \cos(\theta), \phi) d\phi d\cos(\theta) d\xi$$

Where

$$\Phi_{nlm}(\xi, \cos(\theta), \phi) = -\frac{\sqrt{2l+1}}{a^{2l+1}} \sqrt{\frac{(l-m)!}{(l+m)!}} (1+\xi)^l (1-\xi)^{l+1} C_n^{2l+3/2}(\xi) P_{lm}(\cos(\theta)) \begin{bmatrix} \cos(m\phi) \\ \sin(m\phi) \end{bmatrix}$$

$$I_{nl} = -K_{nl} \frac{4\pi}{a^{2l+6}} \frac{\Gamma(n+4l+3)}{n!(n+2l+3/2)[\Gamma(2l+3/2)]^2} \quad K_{nl} = \frac{1}{2}n(n+4l+3) + (l+1)(2l+1)$$

P_{lm} is the Associated Legendre Polynomials whereas C_n^α is the Gegenbauer polynomial.

Also note $\xi = \frac{r-a}{r+a}$, and n, l and m are integers bounded by $0 \leq n < N$, $0 \leq l < L$, and $0 \leq m \leq l$

`galpy.potential.scf_compute_coeffs(dens, N, L, a=1.0, radial_order=None, cos_theta_order=None, phi_order=None)`

NAME:

`scf_compute_coeffs`

PURPOSE:

Numerically compute the expansion coefficients for a given triaxial density

INPUT:

`dens` - A density function that takes a parameter `R`, `z` and `phi`

`N` - size of the Nth dimension of the expansion coefficients

`L` - size of the Lth and Mth dimension of the expansion coefficients

`a` - parameter used to shift the basis functions

`radial_order` - Number of sample points of the radial integral. If `None`, `radial_order`=max(20, `N` + 3/2`L` + 1)

costheta_order - Number of sample points of the costheta integral. If None, If costheta_order=max(20, L + 1)

phi_order - Number of sample points of the phi integral. If None, If costheta_order=max(20, L + 1)

OUTPUT:

(Acos,Asin) - Expansion coefficients for density dens that can be given to SCFPotential.__init__

HISTORY:

2016-05-27 - Written - Aladdin

galpy.potential.scf_compute_coeffs_axi

Note: This function computes Acos and Asin as defined in [Hernquist & Ostriker \(1992\)](#), except that we multiply Acos by 2 such that the density from *Galpy's Hernquist Potential* corresponds to $Acos = \delta_{0n}\delta_{0l}\delta_{0m}$.

Further note that this function is a specification of *scf_compute_coeffs* where $Acos_{nlm} = 0$ at $m \neq 0$ and $Asin_{nlm} = None$

For a given $\rho(R, z)$ we can compute *Acos* and *Asin* through the following equation

$$Acos_{nlm} = \frac{8\pi a^3}{I_{nl}} \int_{\xi=0}^{\infty} \int_{\cos(\theta)=-1}^1 (1+\xi)^2 (1-\xi)^{-4} \rho(R, z) \Phi_{nlm}(\xi, \cos(\theta)) d\cos(\theta) d\xi \quad Asin_{nlm} = None$$

Where

$$\Phi_{nlm}(\xi, \cos(\theta)) = -\frac{\sqrt{2l+1}}{a^{2l+1}} (1+\xi)^l (1-\xi)^{l+1} C_n^{2l+3/2}(\xi) P_{l0}(\cos(\theta)) \delta_{m0}$$

$$I_{nl} = -K_{nl} \frac{4\pi}{a^{2l+6}} \frac{\Gamma(n+4l+3)}{n!(n+2l+3/2)[\Gamma(2l+3/2)]^2} \quad K_{nl} = \frac{1}{2}n(n+4l+3) + (l+1)(2l+1)$$

P_{lm} is the Associated Legendre Polynomials whereas C_n^α is the Gegenbauer polynomial.

Also note $\xi = \frac{r-a}{r+a}$, and n, l and m are integers bounded by $0 \leq n < N$, $0 \leq l < L$, and $m = 0$

galpy.potential.scf_compute_coeffs_axi (dens, N, L, a=1.0, radial_order=None, costheta_order=None)

NAME:

scf_compute_coeffs_axi

PURPOSE:

Numerically compute the expansion coefficients for a given axi-symmetric density

INPUT:

dens - A density function that takes a parameter R and z

N - size of the Nth dimension of the expansion coefficients

L - size of the Lth dimension of the expansion coefficients

a - parameter used to shift the basis functions

radial_order - Number of sample points of the radial integral. If None, radial_order=max(20, N + 3/2L + 1)

costheta_order - Number of sample points of the costheta integral. If None, If costheta_order=max(20, L + 1)

OUTPUT:

(Acos,Asin) - Expansion coefficients for density dens that can be given to SCFPotential.__init__

HISTORY:

2016-05-20 - Written - Aladdin

galpy.potential.scf_compute_coeffs_spherical

Note: This function computes Acos and Asin as defined in [Hernquist & Ostriker \(1992\)](#), except that we multiply Acos by 2 such that the density from *Galpy's Hernquist Potential* corresponds to $Acos = \delta_{0n}\delta_{0l}\delta_{0m}$.

Futher note that this function is a specification of *scf_compute_coeffs_axi* where $Acos_{nlm} = 0$ at $l \neq 0$

For a given $\rho(r)$ we can compute Acos and Asin through the following equation

$$Acos_{nlm} = \frac{16\pi a^3}{I_{nl}} \int_{\xi=0}^{\infty} (1+\xi)^2 (1-\xi)^{-4} \rho(r) \Phi_{nlm}(\xi) d\xi \quad Asin_{nlm} = None$$

Where

$$\Phi_{nlm}(\xi, \cos(\theta)) = -\frac{1}{2a} (1-\xi) C_n^{3/2}(\xi) \delta_{l0} \delta_{m0}$$

$$I_{n0} = -K_{n0} \frac{1}{4a} \frac{(n+2)(n+1)}{(n+3/2)} \quad K_{nl} = \frac{1}{2} n(n+3) + 1$$

C_n^α is the Gegenbauer polynomial.

Also note $\xi = \frac{r-a}{r+a}$, and n, l and m are integers bounded by $0 \leq n < N, l = m = 0$

galpy.potential.scf_compute_coeffs_spherical (dens, N, a=1.0, radial_order=None)

NAME:

scf_compute_coeffs_spherical

PURPOSE:

Numerically compute the expansion coefficients for a given spherical density

INPUT:

dens - A density function that takes a parameter R

N - size of expansion coefficients

a - parameter used to shift the basis functions

radial_order - Number of sample points of the radial integral. If None, radial_order=max(20, N + 1)

OUTPUT:

(Acos,Asin) - Expansion coefficients for density dens that can be given to SCFPotential.__init__

HISTORY:

2016-05-18 - Written - Aladdin

Specific potentials

All of the following potentials can also be modified by the specific `WrapperPotentials` listed [below](#).

Spherical potentials

Burkert potential

class galpy.potential.BurkertPotential (*amp=1.0, a=2.0, normalize=False, ro=None, vo=None*)

BurkertPotential.py: Potential with a Burkert density

$$\rho(r) = \frac{\text{amp}}{(1 + r/a) (1 + [r/a]^2)}$$

__init__ (*amp=1.0, a=2.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a Burkert-density potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass density or Gxmass density

a = scale radius (can be Quantity)

normalize - if True, normalize such that vc(1.,0.)=1., or, if given as a number, such that the force is this fraction of the force necessary to make vc(1.,0.)=1.

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2013-04-10 - Written - Bovy (IAS)

Double power-law density spherical potential

class galpy.potential.TwoPowerSphericalPotential (*amp=1.0, a=5.0, alpha=1.5, beta=3.5, normalize=False, ro=None, vo=None*)

Class that implements spherical potentials that are derived from two-power density models

$$\rho(r) = \frac{\text{amp}}{4 \pi a^3} \frac{1}{(r/a)^\alpha (1 + r/a)^{\beta-\alpha}}$$

__init__ (*amp=1.0, a=5.0, alpha=1.5, beta=3.5, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a two-power-density potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

alpha - inner power

beta - outer power

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Started - Bovy (NYU)

Jaffe potential

```
class galpy.potential.JaffePotential (amp=1.0,    a=1.0,    normalize=False,    ro=None,
                                     vo=None)
```

Class that implements the Jaffe potential

$$\rho(r) = \frac{\text{amp}}{4 \pi a^3} \frac{1}{(r/a)^2 (1 + r/a)^2}$$

```
__init__ (amp=1.0, a=1.0, normalize=False, ro=None, vo=None)
```

NAME:

__init__

PURPOSE:

Initialize a Jaffe potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Written - Bovy (NYU)

Hernquist potential

class galpy.potential.HernquistPotential (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

Class that implements the Hernquist potential

$$\rho(r) = \frac{\text{amp}}{4\pi a^3} \frac{1}{(r/a)(1+r/a)^3}$$

__init__ (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

Initialize a Hernquist potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Written - Bovy (NYU)

Isochrone potential

class galpy.potential.IsochronePotential (*amp=1.0, b=1.0, normalize=False, ro=None, vo=None*)

Class that implements the Isochrone potential

$$\Phi(r) = -\frac{\text{amp}}{b + \sqrt{b^2 + r^2}}$$

with $\text{amp} = GM$ the total mass.

__init__ (*amp=1.0, b=1.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize an isochrone potential

INPUT:

amp= amplitude to be applied to the potential, the total mass (default: 1); can be a Quantity with units of mass density or Gxmass density

b= scale radius of the isochrone potential (can be Quantity)

normalize= if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2013-09-08 - Written - Bovy (IAS)

Kepler potential

class galpy.potential.**KeplerPotential** (amp=1.0, normalize=False, ro=None, vo=None)

Class that implements the Kepler (point mass) potential

$$\Phi(r) = -\frac{\text{amp}}{r}$$

with $\text{amp} = GM$ the mass.

__init__ (amp=1.0, normalize=False, ro=None, vo=None)

NAME:

__init__

PURPOSE:

initialize a Kepler, point-mass potential

INPUT:

amp - amplitude to be applied to the potential, the mass of the point mass (default: 1); can be a Quantity with units of mass density or Gxmass density

alpha - inner power

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-10 - Written - Bovy (NYU)

NFW potential

```
class galpy.potential.NFWPotential(amp=1.0, a=1.0, normalize=False, conc=None,  
mvir=None, vo=None, ro=None, H=70.0, Om=0.3,  
overdens=200.0, wrtcrit=False)
```

Class that implements the NFW potential

$$\rho(r) = \frac{\text{amp}}{4\pi a^3} \frac{1}{(r/a)(1+r/a)^2}$$

```
__init__(amp=1.0, a=1.0, normalize=False, conc=None, mvir=None, vo=None, ro=None, H=70.0,  
Om=0.3, overdens=200.0, wrtcrit=False)
```

NAME:

`__init__`

PURPOSE:

Initialize a NFW potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

`a` - scale radius (can be Quantity)

`normalize` - if True, normalize such that $vc(1,0)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1,0)=1.$

Alternatively, NFW potentials can be initialized using

`conc`= concentration

`mvir`= virial mass in 10^{12} Msolar

in which case you also need to supply the following keywords

`H`= (default: 70) Hubble constant in km/s/Mpc

`Om`= (default: 0.3) Omega matter

`overdens`= (200) overdensity which defines the virial radius

`wrtcrit`= (False) if True, the overdensity is wrt the critical density rather than the mean matter density

`ro` , `vo`= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Written - Bovy (NYU)

2014-04-03 - Initialization w/ concentration and mass - Bovy (IAS)

Plummer potential

class galpy.potential.PlummerPotential (*amp=1.0, b=0.8, normalize=False, ro=None, vo=None*)

Class that implements the Plummer potential

$$\Phi(R, z) = -\frac{\text{amp}}{\sqrt{R^2 + z^2 + b^2}}$$

with $\text{amp} = GM$ the total mass.

__init__ (*amp=1.0, b=0.8, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a Plummer potential

INPUT:

amp - amplitude to be applied to the potential, the total mass (default: 1); can be a Quantity with units of mass or Gxmass

b - scale parameter (can be Quantity)

normalize - if True, normalize such that $\text{vc}(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $\text{vc}(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2015-06-15 - Written - Bovy (IAS)

Power-law density spherical potential

class galpy.potential.PowerSphericalPotential (*amp=1.0, alpha=1.0, normalize=False, r1=1.0, ro=None, vo=None*)

Class that implements spherical potentials that are derived from power-law density models

$$\rho(r) = \frac{\text{amp}}{r_1^3} \left(\frac{r_1}{r}\right)^\alpha$$

__init__ (*amp=1.0, alpha=1.0, normalize=False, r1=1.0, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a power-law-density potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

alpha - inner power

r1= (1.) reference radius for amplitude (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-10 - Written - Bovy (NYU)

Power-law density spherical potential with an exponential cut-off

```
class galpy.potential.PowerSphericalPotentialwCutoff (amp=1.0, alpha=1.0, rc=1.0,  
                                                    normalize=False, r1=1.0,  
                                                    ro=None, vo=None)
```

Class that implements spherical potentials that are derived from power-law density models

$$\rho(r) = \text{amp} \left(\frac{r_1}{r} \right)^\alpha \exp \left(-(r/rc)^2 \right)$$

```
__init__ (amp=1.0, alpha=1.0, rc=1.0, normalize=False, r1=1.0, ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a power-law-density potential

INPUT:

amp= amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass density or Gxmass density

alpha= inner power

rc= cut-off radius (can be Quantity)

r1= (1.) reference radius for amplitude (can be Quantity)

normalize= if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2013-06-28 - Written - Bovy (IAS)

Pseudo-isothermal potential

class galpy.potential.PseudoIsothermalPotential (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

Class that implements the pseudo-isothermal potential

$$\rho(r) = \frac{\text{amp}}{4\pi a^3} \frac{1}{1 + (r/a)^2}$$

__init__ (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a pseudo-isothermal potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - core radius (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2015-12-04 - Started - Bovy (UofT)

Spherical Shell Potential

class galpy.potential.SphericalShellPotential (*amp=1.0, a=0.75, normalize=False, ro=None, vo=None*)

Class that implements the potential of an infinitesimally-thin, spherical shell

$$\rho(r) = \frac{\text{amp}}{4\pi a^2} \delta(r - a)$$

with $\text{amp} = GM$ the mass of the shell.

__init__ (*amp=1.0, a=0.75, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a spherical shell potential

INPUT:

amp - mass of the shell (default: 1); can be a Quantity with units of mass or Gxmass

a= (0.75) radius of the shell (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$; note that because the force is always zero at $r < a$, this does not work if $a > 1$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2018-08-04 - Written - Bovy (UofT)

Axisymmetric potentials

Double exponential disk potential

```
class galpy.potential.DoubleExponentialDiskPotential (amp=1.0,
                                                    hr=0.3333333333333333,
                                                    hz=0.0625,      maxiter=20,
                                                    tol=0.001,   normalize=False,
                                                    ro=None, vo=None, new=True,
                                                    kmaxFac=2.0, glorder=10)
```

Class that implements the double exponential disk potential

$$\rho(R, z) = \text{amp} \exp(-R/h_R - |z|/h_z)$$

```
__init__ (amp=1.0, hr=0.3333333333333333, hz=0.0625, maxiter=20, tol=0.001, normalize=False,
         ro=None, vo=None, new=True, kmaxFac=2.0, glorder=10)
```

NAME:

__init__

PURPOSE:

initialize a double-exponential disk potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass density or Gxmass density

hr - disk scale-length (can be Quantity)

hz - scale-height (can be Quantity)

tol - relative accuracy of potential-evaluations

maxiter - scipy.integrate keyword

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

DoubleExponentialDiskPotential object

HISTORY:

2010-04-16 - Written - Bovy (NYU)

2013-01-01 - Re-implemented using faster integration techniques - Bovy (IAS)

Flattened Power-law potential

Flattening is in the potential as in [Evans \(1994\)](#) rather than in the density

```
class galpy.potential.FlattenedPowerPotential(amp=1.0, alpha=0.5, q=0.9, core=1e-08, normalize=False, r1=1.0, ro=None, vo=None)
```

Class that implements a power-law potential that is flattened in the potential (NOT the density)

$$\Phi(R, z) = -\frac{\text{amp } r_1^\alpha}{\alpha (R^2 + (z/q)^2 + \text{core}^2)^{\alpha/2}}$$

and the same as LogarithmicHaloPotential for $\alpha = 0$

See Figure 1 in [Evans \(1994\)](#) for combinations of alpha and q that correspond to positive densities

```
__init__(amp=1.0, alpha=0.5, q=0.9, core=1e-08, normalize=False, r1=1.0, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a flattened power-law potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of velocity-squared

alpha - power

q - flattening

core - core radius (can be Quantity)

r1= (1.) reference radius for amplitude (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2013-01-09 - Written - Bovy (IAS)

Interpolated axisymmetric potential

The `interpRZPotential` class provides a general interface to generate interpolated instances of general three-dimensional, axisymmetric potentials or lists of such potentials. This interpolated potential can be used in any function where other three-dimensional galpy potentials can be used. This includes functions that use C to speed up calculations, if the `interpRZPotential` instance was set up with `enable_c=True`. Initialize as

```
>>> from galpy import potential
>>> ip= potential.interpRZPotential(potential.MWPotential,interpPot=True)
```

which sets up an interpolation of the potential itself only. The potential and all different forces and functions (`dens`, `vcirc`, `epifreq`, `verticalfreq`, `dvcircdR`) are interpolated separately and one needs to specify that these need to be interpolated separately (so, for example, one needs to set `interpRforce=True` to interpolate the radial force, or `interpvcirc=True` to interpolate the circular velocity).

When points outside the grid are requested within the python code, the instance will fall back on the original (non-interpolated) potential. However, when the potential is used purely in C, like during orbit integration in C or during action-angle evaluations in C, there is no way for the potential to fall back onto the original potential and nonsense or NaNs will be returned. Therefore, when using `interpRZPotential` in C, one must make sure that the whole relevant part of the (R, z) plane is covered. One more time:

Warning: When an interpolated potential is used purely in C, like during orbit integration in C or during action-angle evaluations in C, there is no way for the potential to fall back onto the original potential and nonsense or NaNs will be returned. Therefore, when using `interpRZPotential` in C, one must make sure that the whole relevant part of the (R, z) plane is covered.

```
class galpy.potential.interpRZPotential (RZPot=None,          rgrid=(-4.605170185988091,
2.995732273553991, 101), zgrid=(0.0, 1.0,
101), logR=True, interpPot=False, interpRforce=False, interpzforce=False, interpDens=False,
interpvcirc=False, interpdvcircdR=False, interpepifreq=False, interpverticalfreq=False,
ro=None, vo=None, use_c=False, enable_c=False, zsym=True, numcores=None)
```

Class that interpolates a given potential on a grid for fast orbit integration

```
__init__ (RZPot=None, rgrid=(-4.605170185988091, 2.995732273553991, 101), zgrid=(0.0, 1.0,
101), logR=True, interpPot=False, interpRforce=False, interpzforce=False, interpDens=False,
interpvcirc=False, interpdvcircdR=False, interpepifreq=False, interpverticalfreq=False,
ro=None, vo=None, use_c=False, enable_c=False, zsym=True, numcores=None)
```

NAME:

```
__init__
```

PURPOSE:

Initialize an `interpRZPotential` instance

INPUT:

RZPot - RZPotential to be interpolated

rgrid - R grid to be given to `linspace` as in `rs= linspace(*rgrid)`

zgrid - z grid to be given to `linspace` as in `zs= linspace(*zgrid)`

logR - if True, rgrid is in the log of R so `logrs= linspace(*rgrid)`

interpPot, interpRforce, interpzforce, interpDens, interpvcirc, interpepfreq, interpverticalfreq, interpdvircdr= if True, interpolate these functions

use_c= use C to speed up the calculation of the grid

enable_c= enable use of C for interpolations

zsym= if True (default), the potential is assumed to be symmetric around $z=0$ (so you can use, e.g., `zgrid=(0.,1.,101)`).

numcores= if set to an integer, use this many cores (only used for vcirc, dvircdR, epifreq, and verticalfreq; NOT NECESSARILY FASTER, TIME TO MAKE SURE)

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

instance

HISTORY:

2010-07-21 - Written - Bovy (NYU)

2013-01-24 - Started with new implementation - Bovy (IAS)

Interpolated SnapshotRZ potential

This class is built on the `interpRZPotential` class; see the documentation of that class [here](#) for additional information on how to setup objects of the `InterpSnapshotRZPotential` class.

```
class galpy.potential.InterpSnapshotRZPotential(s, ro=None, vo=None,
                                                rgrid=(-4.605170185988091,
                                                2.995732273553991, 101), zgrid=(0.0,
                                                1.0, 101), interpepfreq=False,
                                                interpverticalfreq=False, interp-
                                                Pot=True, enable_c=True, logR=True,
                                                zsym=True, numcores=None, naz-
                                                imuths=4, use_pkdgrav=False)
```

Interpolated axisymmetrized potential extracted from a simulation output (see `interpRZPotential` and `SnapshotRZPotential`)

```
__init__(s, ro=None, vo=None, rgrid=(-4.605170185988091, 2.995732273553991, 101), zgrid=(0.0,
1.0, 101), interpepfreq=False, interpverticalfreq=False, interpPot=True, enable_c=True,
logR=True, zsym=True, numcores=None, nazimuths=4, use_pkdgrav=False)
```

NAME:

`__init__`

PURPOSE:

Initialize an `InterpSnapshotRZPotential` instance

INPUT:

s - a simulation snapshot loaded with `pynbody`

rgrid - R grid to be given to `linspace` as in `rs= linspace(*rgrid)`

zgrid - z grid to be given to `linspace` as in `zs= linspace(*zgrid)`

logR - if True, rgrid is in the log of R so `logrs= linspace(*rgrid)`

interpPot, interpepfreq, interpverticalfreq= if True, interpolate these functions (interpPot=True also interpolates the R and zforce)

enable_c= enable use of C for interpolations

zsym= if True (default), the potential is assumed to be symmetric around $z=0$ (so you can use, e.g., `zgrid=(0.,1.,101)`).

numcores= if set to an integer, use this many cores

nazimuths= (4) number of azimuths to average over

use_pkdgrav= (False) use PKDGRAV to calculate the snapshot's potential and forces (CURRENTLY NOT IMPLEMENTED)

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

instance

HISTORY:

2013 - Written - Rok Roskar (ETH)

2014-11-24 - Edited for merging into main galpy - Bovy (IAS)

Kuzmin disk potential

class `galpy.potential.KuzminDiskPotential` (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

Class that implements the Kuzmin Disk potential

$$\Phi(R, z) = -\frac{\text{amp}}{\sqrt{R^2 + (a + |z|)^2}}$$

with $\text{amp} = GM$ the total mass.

__init__ (*amp=1.0, a=1.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a Kuzmin disk Potential

INPUT:

amp - amplitude to be applied to the potential, the total mass (default: 1); can be a Quantity with units of mass density or Gxmass density

a - scale length (can be Quantity)

normalize - if True, normalize such that $\text{vc}(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $\text{vc}(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

KuzminDiskPotential object

HISTORY:

2016-05-09 - Written - Aladdin

Kuzmin-Kutuzov Staechel potential

class galpy.potential.**KuzminKutuzovStaeckelPotential** (*amp=1.0, ac=5.0, Delta=1.0, normalize=False, ro=None, vo=None*)

Class that implements the Kuzmin-Kutuzov Staechel potential

$$\Phi(R, z) = -\frac{\text{amp}}{\sqrt{\lambda} + \sqrt{\nu}}$$

(see, e.g., [Batsleer & Dejonghe 1994](#))

__init__ (*amp=1.0, ac=5.0, Delta=1.0, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a Kuzmin-Kutuzov Staechel potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass density or Gxmass density

ac - axis ratio of the coordinate surfaces; (a/c) = sqrt(-alpha) / sqrt(-gamma) (default: 5.)

Delta - focal distance that defines the spheroidal coordinate system (default: 1.); Delta=sqrt(gamma-alpha) (can be Quantity)

normalize - if True, normalize such that vc(1.,0.)=1., or, if given as a number, such that the force is this fraction of the force necessary to make vc(1.,0.)=1.

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2015-02-15 - Written - Trick (MPIA)

Logarithmic halo potential

class galpy.potential.**LogarithmicHaloPotential** (*amp=1.0, core=1e-08, q=1.0, b=None, normalize=False, ro=None, vo=None*)

Class that implements the logarithmic potential

$$\Phi(R, z) = \frac{\text{amp}}{2} \ln \left[R^2 + \left(\frac{z}{q} \right)^2 + \text{core}^2 \right]$$

Alternatively, the potential can be made triaxial by adding a parameter *b*

$$\Phi(x, y, z) = \frac{\text{amp}}{2} \ln \left[x^2 + \left(\frac{y}{b} \right)^2 + \left(\frac{z}{q} \right)^2 + \text{core}^2 \right]$$

With these definitions, $\sqrt{\text{amp}}$ is the circular velocity at $r \gg \text{core}$ at $(y, z) = (0, 0)$.

`__init__` (*amp=1.0, core=1e-08, q=1.0, b=None, normalize=False, ro=None, vo=None*)

NAME:

`__init__`

PURPOSE:

initialize a logarithmic potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of velocity-squared

core - core radius at which the logarithm is cut (can be Quantity)

q - potential flattening $(z/q)^{**2}$.

b= (None) if set, shape parameter in y-direction ($y \rightarrow y/b$; see definition)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1$.

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-04-02 - Started - Bovy (NYU)

Miyamoto-Nagai potential

`class galpy.potential.MiyamotoNagaiPotential` (*amp=1.0, a=1.0, b=0.1, normalize=False, ro=None, vo=None*)

Class that implements the Miyamoto-Nagai potential

$$\Phi(R, z) = -\frac{\text{amp}}{\sqrt{R^2 + (a + \sqrt{z^2 + b^2})^2}}$$

with $\text{amp} = GM$ the total mass.

`__init__` (*amp=1.0, a=1.0, b=0.1, normalize=False, ro=None, vo=None*)

NAME:

`__init__`

PURPOSE:

initialize a Miyamoto-Nagai potential

INPUT:

amp - amplitude to be applied to the potential, the total mass (default: 1); can be a Quantity with units of mass or Gxmass

a - scale length (can be Quantity)

b - scale height (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Started - Bovy (NYU)

Three Miyamoto-Nagai disk approximation to an exponential disk

```
class galpy.potential.MN3ExponentialDiskPotential(amp=1.0,  
                                                  hr=0.3333333333333333,  
                                                  hz=0.0625, sech=False, pos-  
                                                  dens=False, normalize=False,  
                                                  ro=None, vo=None)  
class that implements the three Miyamoto-Nagai approximation to a radially-exponential disk potential of Smith et al. 2015
```

$$\rho(R, z) = \text{amp} \exp(-R/h_R - |z|/h_z)$$

or

$$\rho(R, z) = \text{amp} \exp(-R/h_R) \text{sech}^2(-|z|/h_z)$$

depending on whether `sech=True` or not. This density is approximated using three Miyamoto-Nagai disks

```
__init__(amp=1.0, hr=0.3333333333333333, hz=0.0625, sech=False, posdens=False, normal-  
         ize=False, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a 3MN approximation to an exponential disk potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass density or Gxmass density

hr - disk scale-length (can be Quantity)

hz - scale-height (can be Quantity)

sech= (False) if True, hz is the scale height of a sech vertical profile (default is exponential vertical profile)

posdens= (False) if True, allow only positive density solutions (Table 2 in Smith et al. rather than Table 1)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

MN3ExponentialDiskPotential object

HISTORY:

2015-02-07 - Written - Bovy (IAS)

Razor-thin exponential disk potential

```
class galpy.potential.RazorThinExponentialDiskPotential (amp=1.0,  
                                                    hr=0.3333333333333333,  
                                                    maxiter=20, tol=0.001,  
                                                    normalize=False,  
                                                    ro=None, vo=None,  
                                                    new=True, glorder=100)
```

Class that implements the razor-thin exponential disk potential

$$\rho(R, z) = \text{amp} \exp(-R/h_R) \delta(z)$$

```
__init__ (amp=1.0, hr=0.3333333333333333, maxiter=20, tol=0.001, normalize=False, ro=None,  
          vo=None, new=True, glorder=100)
```

NAME:

`__init__`

PURPOSE:

initialize a razor-thin-exponential disk potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1); can be a Quantity with units of surface-mass or Gxsurface-mass

`hr` - disk scale-length (can be Quantity)

`tol` - relative accuracy of potential-evaluations

`maxiter` - scipy.integrate keyword

`normalize` - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

`ro=`, `vo=` distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

RazorThinExponentialDiskPotential object

HISTORY:

2012-12-27 - Written - Bovy (IAS)

Ring potential

class galpy.potential.RingPotential (*amp=1.0, a=0.75, normalize=False, ro=None, vo=None*)

Class that implements the potential of an infinitesimally-thin, circular ring

$$\rho(R, z) = \frac{\text{amp}}{2\pi R_0} \delta(R - R_0) \delta(z)$$

with $\text{amp} = GM$ the mass of the ring.

__init__ (*amp=1.0, a=0.75, normalize=False, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a circular ring potential

INPUT:

amp - mass of the ring (default: 1); can be a Quantity with units of mass or Gxmass

a (0.75) radius of the ring (can be Quantity)

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$; note that because the force is always positive at $r < a$, this does not work if $a > 1$

ro-, *vo*= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2018-08-04 - Written - Bovy (UofT)

Axisymmetrized N-body snapshot potential

class galpy.potential.SnapshotRZPotential (*s, num_threads=None, nazimuths=4, ro=None, vo=None*)

Class that implements an axisymmetrized version of the potential of an N-body snapshot (requires [pynbody](#))

_evaluate, *_Rforce*, and *_zforce* calculate a hash for the array of points that is passed in by the user. The hash and corresponding potential/force arrays are stored – if a subsequent request matches a previously computed hash, the previous results are returned and not recalculated.

__init__ (*s, num_threads=None, nazimuths=4, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

Initialize a SnapshotRZ potential object

INPUT:

s - a simulation snapshot loaded with pynbody
 num_threads= (4) number of threads to use for calculation
 nazimuths= (4) number of azimuths to average over
 ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

instance

HISTORY:

2013 - Written - Rok Roskar (ETH)

2014-11-24 - Edited for merging into main galpy - Bovy (IAS)

Ellipsoidal triaxial potentials

galpy has very general support for implementing triaxial (or the oblate and prolate special cases) of ellipsoidal potentials through the general `EllipsoidalPotential` class. These potentials have densities that are uniform on ellipsoids, thus only functions of $m^2 = x^2 + \frac{y^2}{b^2} + \frac{z^2}{c^2}$. New potentials of this type can be implemented by inheriting from this class and implementing the `_mdens(self,m)`, `_psi(self,m)`, and `_mdens_deriv` functions for the density, its integral with respect to m^2 , and its derivative with respect to m, respectively. For adding a C implementation, follow similar steps (use `PerfectEllipsoidPotential` as an example to follow).

Perfect Ellipsoid potential

```
class galpy.potential.PerfectEllipsoidPotential (amp=1.0,  a=5.0,  b=1.0,  c=1.0,
                                                zvec=None,  pa=None,  glorder=50,  normalize=False,  ro=None,
                                                vo=None)
```

Potential of the perfect ellipsoid (de Zeeuw 1985):

$$\rho(x, y, z) = \frac{\text{amp } a}{\pi^2 b c} \frac{1}{(m^2 + a^2)^2}$$

where $\text{amp} = GM$ is the total mass and $m^2 = x^2 + y^2/b^2 + z^2/c^2$.

```
__init__ (amp=1.0,  a=5.0,  b=1.0,  c=1.0,  zvec=None,  pa=None,  glorder=50,  normalize=False,
         ro=None,  vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a perfect ellipsoid potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units that depend on the specific spheroidal potential

a - scale radius (can be Quantity)

b - y-to-x axis ratio of the density

c - z-to-x axis ratio of the density

`zvec=` (None) If set, a unit vector that corresponds to the z axis

`pa=` (None) If set, the position angle of the x axis (rad or Quantity)

`glorder=` (50) if set, compute the relevant force and potential integrals with Gaussian quadrature of this order

`ro=`, `vo=` distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2018-08-06 - Started - Bovy (UofT)

Double power-law density triaxial potential

```
class galpy.potential.TwoPowerTriaxialPotential (amp=1.0, a=5.0, alpha=1.5, beta=3.5,  
                                                b=1.0, c=1.0, zvec=None, pa=None,  
                                                glorder=50, normalize=False,  
                                                ro=None, vo=None)
```

Class that implements triaxial potentials that are derived from two-power density models

$$\rho(x, y, z) = \frac{\text{amp}}{4 \pi a^3} \frac{1}{(m/a)^\alpha (1 + m/a)^{\beta-\alpha}}$$

with

$$m^2 = x'^2 + \frac{y'^2}{b^2} + \frac{z'^2}{c^2}$$

and (x', y', z') is a rotated frame wrt (x, y, z) specified by parameters `zvec` and `pa` which specify (a) `zvec`: the location of the z' axis in the (x, y, z) frame and (b) `pa`: the position angle of the x' axis wrt the \hat{x} axis, that is, the x axis after rotating to `zvec`.

Note that this general class of potentials does *not* automatically revert to the special `TriaxialNFWPotential`, `TriaxialHernquistPotential`, or `TriaxialJaffePotential` when using their (α, β) values (like `TwoPowerSphericalPotential`).

```
__init__ (amp=1.0, a=5.0, alpha=1.5, beta=3.5, b=1.0, c=1.0, zvec=None, pa=None, glorder=50,  
         normalize=False, ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a triaxial two-power-density potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or `Gxmass`

`a` - scale radius (can be Quantity)

`alpha` - inner power ($0 \leq \alpha < 3$)

`beta` - outer power ($\beta > 2$)

`b` - y -to- x axis ratio of the density

c - z-to-x axis ratio of the density

zvec= (None) If set, a unit vector that corresponds to the z axis

pa= (None) If set, the position angle of the x axis (rad or Quantity)

glorder= (50) if set, compute the relevant force and potential integrals with Gaussian quadrature of this order

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2016-05-30 - Started - Bovy (UofT)

2018-08-07 - Re-written using the general EllipsoidalPotential class - Bovy (UofT)

Triaxial Jaffe potential

```
class galpy.potential.TriaxialJaffePotential (amp=1.0, a=2.0, b=1.0, c=1.0, zvec=None,
                                             pa=None, normalize=False, glorder=50,
                                             ro=None, vo=None)
```

Class that implements the Jaffe potential

$$\rho(x, y, z) = \frac{\text{amp}}{4 \pi a^3} \frac{1}{(m/a)^2 (1 + m/a)^2}$$

with

$$m^2 = x'^2 + \frac{y'^2}{b^2} + \frac{z'^2}{c^2}$$

and (x', y', z') is a rotated frame wrt (x, y, z) specified by parameters `zvec` and `pa` which specify (a) `zvec`: the location of the z' axis in the (x, y, z) frame and (b) `pa`: the position angle of the x' axis wrt the \hat{x} axis, that is, the x axis after rotating to `zvec`.

```
__init__ (amp=1.0, a=2.0, b=1.0, c=1.0, zvec=None, pa=None, normalize=False, glorder=50,
         ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

Initialize a Jaffe potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

b - y-to-x axis ratio of the density

c - z-to-x axis ratio of the density

`zvec=` (None) If set, a unit vector that corresponds to the z axis

`pa=` (None) If set, the position angle of the x axis

`glorder=` (50) if set, compute the relevant force and potential integrals with Gaussian quadrature of this order

`normalize` - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

`ro=`, `vo=` distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Written - Bovy (UofT)

2018-08-07 - Re-written using the general EllipsoidalPotential class - Bovy (UofT)

Triaxial Hernquist potential

```
class galpy.potential.TriaxialHernquistPotential(amp=1.0, a=2.0, normalize=False,  
                                                b=1.0, c=1.0, zvec=None, pa=None,  
                                                glorder=50, ro=None, vo=None)
```

Class that implements the triaxial Hernquist potential

$$\rho(x, y, z) = \frac{\text{amp}}{4\pi a^3} \frac{1}{(m/a)(1+m/a)^3}$$

with

$$m^2 = x'^2 + \frac{y'^2}{b^2} + \frac{z'^2}{c^2}$$

and (x', y', z') is a rotated frame wrt (x, y, z) specified by parameters `zvec` and `pa` which specify (a) `zvec`: the location of the z' axis in the (x, y, z) frame and (b) `pa`: the position angle of the x' axis wrt the \hat{x} axis, that is, the x axis after rotating to `zvec`.

```
__init__(amp=1.0, a=2.0, normalize=False, b=1.0, c=1.0, zvec=None, pa=None, glorder=50,  
         ro=None, vo=None)
```

NAME:

__init__

PURPOSE:

Initialize a triaxial Hernquist potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

`a` - scale radius (can be Quantity)

`b` - y-to-x axis ratio of the density

`c` - z-to-x axis ratio of the density

`zvec=` (None) If set, a unit vector that corresponds to the z axis

pa= (None) If set, the position angle of the x axis

glorder= (50) if set, compute the relevant force and potential integrals with Gaussian quadrature of this order

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2010-07-09 - Written - Bovy (UofT)

2018-08-07 - Re-written using the general EllipsoidalPotential class - Bovy (UofT)

Triaxial NFW potential

```
class galpy.potential.TriaxialNFWPotential (amp=1.0, a=2.0, b=1.0, c=1.0, zvec=None,
                                           pa=None, normalize=False, conc=None,
                                           mvir=None, glorder=50, vo=None,
                                           ro=None, H=70.0, Om=0.3, overdens=200.0,
                                           wrtcrit=False)
```

Class that implements the triaxial NFW potential

$$\rho(x, y, z) = \frac{\text{amp}}{4 \pi a^3} \frac{1}{(m/a) (1 + m/a)^2}$$

with

$$m^2 = x'^2 + \frac{y'^2}{b^2} + \frac{z'^2}{c^2}$$

and (x', y', z') is a rotated frame wrt (x, y, z) specified by parameters `zvec` and `pa` which specify (a) `zvec`: the location of the z' axis in the (x, y, z) frame and (b) `pa`: the position angle of the x' axis wrt the \hat{x} axis, that is, the x axis after rotating to `zvec`.

```
__init__ (amp=1.0, a=2.0, b=1.0, c=1.0, zvec=None, pa=None, normalize=False, conc=None,
          mvir=None, glorder=50, vo=None, ro=None, H=70.0, Om=0.3, overdens=200.0,
          wrtcrit=False)
```

NAME:

__init__

PURPOSE:

Initialize a triaxial NFW potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

b - y-to-x axis ratio of the density

c - z-to-x axis ratio of the density

`zvec=` (None) If set, a unit vector that corresponds to the z axis

`pa=` (None) If set, the position angle of the x axis

`glorder=` (50) if set, compute the relevant force and potential integrals with Gaussian quadrature of this order

`normalize` - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

Alternatively, NFW potentials can be initialized using

`conc=` concentration

`mvir=` virial mass in 10^{12} Msolar

in which case you also need to supply the following keywords

`H=` (default: 70) Hubble constant in km/s/Mpc

`Om=` (default: 0.3) Omega matter

`overdens=` (200) overdensity which defines the virial radius

`wrtcrit=` (False) if True, the overdensity is wrt the critical density rather than the mean matter density

`ro=`, `vo=` distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2016-05-30 - Written - Bovy (UofT)

2018-08-06 - Re-written using the general `EllipsoidalPotential` class - Bovy (UofT)

Note that the Ferrers potential listed below is a potential of this type, but it is currently not implemented using the `EllipsoidalPotential` class.

Spiral, bar, other triaxial, and miscellaneous potentials

Dehnen bar potential

```
class galpy.potential.DehnenBarPotential(amp=1.0, omegab=None, rb=None, chi=0.8,
                                         rolr=0.9, barphi=0.4363323129985824, tform=-
                                         4.0, tsteady=None, beta=0.0, alpha=0.01,
                                         Af=None, ro=None, vo=None)
```

Class that implements the Dehnen bar potential (Dehnen 2000), generalized to 3D following Monari et al. (2016)

$$\Phi(R, z, \phi) = A_b(t) \cos(2(\phi - \Omega_b t)) \left(\frac{R}{r}\right)^2 \times \begin{cases} -(R_b/r)^3, & \text{for } r \geq R_b \\ (r/R_b)^3 - 2, & \text{for } r \leq R_b. \end{cases}$$

where $r^2 = R^2 + z^2$ is the spherical radius and

$$A_b(t) = A_f \left(\frac{3}{16} \xi^5 - \frac{5}{8} \xi^3 + \frac{15}{16} \xi + \frac{1}{2} \right), \xi = 2 \frac{t/T_b - t_{\text{form}}}{T_{\text{steady}}} - 1, \text{ if } t_{\text{form}} \leq \frac{t}{T_b} \leq t_{\text{form}} + T_{\text{steady}}$$

and

$$A_b(t) = \begin{cases} 0, & \frac{t}{T_b} < t_{\text{form}} \\ A_f, & \frac{t}{T_b} > t_{\text{form}} + T_{\text{steady}} \end{cases}$$

where

$$T_b = \frac{2\pi}{\Omega_b}$$

is the bar period and the strength can also be specified using α

$$\alpha = 3 \frac{A_f}{v_0^2} \left(\frac{R_b}{r_0} \right)^3.$$

`__init__` (*amp=1.0, omegab=None, rb=None, chi=0.8, rolr=0.9, barphi=0.4363323129985824, tform=-4.0, tsteady=None, beta=0.0, alpha=0.01, Af=None, ro=None, vo=None*)

NAME:

`__init__`

PURPOSE:

initialize a Dehnen bar potential

INPUT:

amp - amplitude to be applied to the potential (default: 1., see alpha or Ab below)

barphi - angle between sun-GC line and the bar's major axis (in rad; default=25 degree; or can be Quantity)

tform - start of bar growth / bar period (default: -4)

tsteady - time from tform at which the bar is fully grown / bar period (default: -tform/2, st the perturbation is fully grown at tform/2)

Either provide:

a) rolr - radius of the Outer Lindblad Resonance for a circular orbit (can be Quantity)

chi - fraction $R_{\text{bar}} / R_{\text{CR}}$ (corotation radius of bar)

alpha - relative bar strength (default: 0.01)

beta - power law index of rotation curve (to calculate OLR, etc.)

b) omegab - rotation speed of the bar (can be Quantity)

rb - bar radius (can be Quantity)

Af - bar strength (can be Quantity)

OUTPUT:

(none)

HISTORY:

2010-11-24 - Started - Bovy (NYU)

2017-06-23 - Converted to 3D following Monari et al. (2016) - Bovy (UofT/CCA)

Ferrers potential

```
class galpy.potential.FerrersPotential (amp=1.0, a=1.0, n=2, b=0.35, c=0.2375,  
omegab=0.0, pa=0.0, normalize=False, ro=None,  
vo=None)
```

Class that implements triaxial Ferrers potential for the ellipsoidal density profile with the short axis along the z-direction

$$\rho(x, y, z) = \frac{\text{amp}}{\pi^{1.5} a^3 b c} \frac{\Gamma(n + \frac{5}{2})}{\Gamma(n + 1)} (1 - (m/a)^2)^n$$

with

$$m^2 = x'^2 + \frac{y'^2}{b^2} + \frac{z'^2}{c^2}$$

and (x', y', z') is a rotated frame wrt (x, y, z) so that the major axis is aligned with x' .

Note that this potential has not yet been optimized for speed and has no C implementation, so orbit integration is currently slow.

```
__init__ (amp=1.0, a=1.0, n=2, b=0.35, c=0.2375, omegab=0.0, pa=0.0, normalize=False, ro=None,  
vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a Ferrers potential

INPUT:

amp - total mass of the ellipsoid determines the amplitude of the potential; can be a Quantity with units of mass or Gxmass

a - scale radius (can be Quantity)

n - power of Ferrers density ($n > 0$)

b - y-to-x axis ratio of the density

c - z-to-x axis ratio of the density

omgab - rotation speed of the ellipsoid (can be Quantity)

pa= (None) If set, the position angle of the x axis (rad or Quantity)

normalize - if True, normalize such that $vc(1., 0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1., 0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

Moving object potential

```
class galpy.potential.MovingObjectPotential(orbit, amp=1.0, GM=0.06, ro=None,
                                             vo=None, softening=None, soft-
                                             ening_model='plummer', soft-
                                             ening_length=0.01)
```

Class that implements the potential coming from a moving object

$$\Phi(R, z, \phi, t) = -\text{amp } GM S(d)$$

where d is the distance between (R, z, ϕ) and the moving object at time t and $S(\cdot)$ is a softening kernel. In the case of Plummer softening, this kernel is

$$S(d) = \frac{1}{\sqrt{d^2 + \text{softening_length}^2}}$$

Plummer is currently the only implemented softening.

```
__init__(orbit, amp=1.0, GM=0.06, ro=None, vo=None, softening=None, soft-
         ening_model='plummer', softening_length=0.01)
```

NAME:

```
__init__
```

PURPOSE:

initialize a MovingObjectPotential

INPUT:

orbit - the Orbit of the object (Orbit object)

amp= - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

GM - 'mass' of the object (degenerate with *amp*, don't use both); can be a Quantity with units of mass or Gxmass

Softening: either provide

a) *softening*= with a ForceSoftening-type object

b) *softening_model*= type of softening to use ('plummer')

softening_length= (optional; can be Quantity)

ro-, *vo*= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2011-04-10 - Started - Bovy (NYU)

Softened-needle bar potential

class galpy.potential.**SoftenedNeedleBarPotential** (*amp=1.0, a=4.0, b=0.0, c=1.0, normalize=False, pa=0.4, omegab=1.8, ro=None, vo=None*)

Class that implements the softened needle bar potential from [Long & Murali \(1992\)](#)

$$\Phi(x, y, z) = \frac{\text{amp}}{2a} \ln \left(\frac{x - a + T_-}{x + a + T_+} \right)$$

where

$$T_{\pm} = \sqrt{(a \pm x)^2 + y^2 + (b + \sqrt{z^2 + c^2})^2}$$

For a prolate bar, set *b* to zero.

__init__ (*amp=1.0, a=4.0, b=0.0, c=1.0, normalize=False, pa=0.4, omegab=1.8, ro=None, vo=None*)

NAME:

__init__

PURPOSE:

initialize a softened-needle bar potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass

a= (4.) Bar half-length (can be Quantity)

b= (1.) Triaxial softening length (can be Quantity)

c= (1.) Prolate softening length (can be Quantity)

pa= (0.4) The position angle of the x axis (rad or Quantity)

omegab= (1.8) Pattern speed (can be Quantity)

normalize - if True, normalize such that *vc*(1.,0.)=1., or, if given as a number, such that the force is this fraction of the force necessary to make *vc*(1.,0.)=1.

ro=, *vo*= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

(none)

HISTORY:

2016-11-02 - Started - Bovy (UofT)

Spiral arms potential

class galpy.potential.**SpiralArmsPotential** (*amp=1, ro=None, vo=None, amp_units='density', N=2, alpha=0.2, r_ref=1, phi_ref=0, Rs=0.3, H=0.125, omega=0, Cs=[1]*)

Class that implements the spiral arms potential from ([Cox and Gomez 2002](#)). Should be used to modulate an

existing potential (density is positive in the arms, negative outside).

$$\Phi(R, \phi, z) = -4\pi G H \rho_0 \exp\left(-\frac{R - r_{ref}}{R_s}\right) \sum \frac{C_n}{K_n D_n} \cos(n\gamma) \operatorname{sech}^{B_n}\left(\frac{K_n z}{B_n}\right)$$

where

$$K_n = \frac{nN}{R \sin(\alpha)}$$

$$B_n = K_n H (1 + 0.4 K_n H)$$

$$D_n = \frac{1 + K_n H + 0.3 (K_n H)^2}{1 + 0.3 K_n H}$$

and

$$\gamma = N \left[\phi - \phi_{ref} - \frac{\ln(R/r_{ref})}{\tan(\alpha)} \right]$$

The default of $C_n = [1]$ gives a sinusoidal profile for the potential. An alternative from Cox and Gomez (2002) creates a density that behaves approximately as a cosine squared in the arms but is separated by a flat interarm region by setting

$$C_n = \left[\frac{8}{3\pi}, \frac{1}{2}, \frac{8}{15\pi} \right]$$

__init__ (*amp=1, ro=None, vo=None, amp_units='density', N=2, alpha=0.2, r_ref=1, phi_ref=0, Rs=0.3, H=0.125, omega=0, Cs=[1]*)

NAME: __init__

PURPOSE: initialize a spiral arms potential

INPUT:

amp amplitude to be applied to the potential (default: 1); can be a Quantity with units of density. ($amp = 4\pi G \rho_0$)

ro distance scales for translation into internal units (default from configuration file)

vo velocity scales for translation into internal units (default from configuration file)

N number of spiral arms

alpha pitch angle of the logarithmic spiral arms in radians (can be Quantity)

r_ref fiducial radius where $\rho = \rho_0$ (r_0 in the paper by Cox and Gomez) (can be Quantity)

phi_ref reference angle ($\phi_p(r_0)$ in the paper by Cox and Gomez) (can be Quantity)

Rs radial scale length of the drop-off in density amplitude of the arms (can be Quantity)

H scale height of the stellar arm perturbation (can be Quantity)

Cs list of constants multiplying the $\cos(n\gamma)$ terms

omega rotational pattern speed of the spiral arms (can be Quantity)

OUTPUT: (none)

HISTORY: Started - 2017-05-12 Jack Hong (UBC)

Completed - 2017-07-04 Jack Hong (UBC)

All galpy potentials can also be made to rotate using the `SolidBodyRotationWrapperPotential` listed in the section on wrapper potentials [below](#).

General Poisson solvers for disks and halos

Disk potential using SCF basis-function-expansion

```
class galpy.potential.DiskSCFPotential (amp=1.0,    normalize=False,    dens=<function
DiskSCFPotential.<lambda>>,    Sigma={'amp':
1.0, 'h': 0.3333333333333333, 'type': 'exp'},
hz={'h': 0.037037037037037035, 'type':
'exp'},    Sigma_amp=None,    dSigma_dR=None,
d2Sigma_dR2=None,    Hz=None,    dHzdz=None,
N=10, L=10, a=1.0, radial_order=None, cos-
theta_order=None, phi_order=None, ro=None,
vo=None)
```

Class that implements a basis-function-expansion technique for solving the Poisson equation for disk (+halo) systems. We solve the Poisson equation for a given density $\rho(R, \phi, z)$ by introducing K helper function pairs $[\Sigma_i(R), h_i(z)]$, with $h_i(z) = d^2 H(z)/dz^2$ and search for solutions of the form

$$\Phi(R, \phi, z) = \Phi_{\text{ME}}(R, \phi, z) + 4\pi G \sum_i \Sigma_i(r) H_i(z),$$

where r is the spherical radius $r^2 = R^2 + z^2$. We can solve for $\Phi_{\text{ME}}(R, \phi, z)$ by solving

$$\frac{\Delta \Phi_{\text{ME}}(R, \phi, z)}{4\pi G} = \rho(R, \phi, z) - \sum_i \left\{ \Sigma_i(r) h_i(z) + \frac{d^2 \Sigma_i(r)}{dr^2} H_i(z) + \frac{2}{r} \frac{d\Sigma_i(r)}{dr} \left[H_i(z) + z \frac{dH_i(z)}{dz} \right] \right\}.$$

We solve this equation by using the *SCFPotential* class and methods (*scf_compute_coeffs_axi* or *scf_compute_coeffs* depending on whether $\rho(R, \phi, z)$ is axisymmetric or not). This technique works very well if the disk portion of the potential can be exactly written as $\rho_{\text{disk}} = \sum_i \Sigma_i(R) h_i(z)$, because the effective density on the right-hand side of this new Poisson equation is then not ‘disky’ and can be well represented using spherical harmonics. But the technique is general and can be used to compute the potential of any disk+halo potential; the closer the disk is to $\rho_{\text{disk}} \approx \sum_i \Sigma_i(R) h_i(z)$, the better the technique works.

This technique was introduced by Kuijken & Dubinski (1995) and was popularized by Dehnen & Binney (1998). The current implementation is a slight generalization of the technique in those papers and uses the SCF approach of Hernquist & Ostriker (1992) to solve the Poisson equation for $\Phi_{\text{ME}}(R, \phi, z)$ rather than solving it on a grid using spherical harmonics and interpolating the solution (as done in Dehnen & Binney 1998).

```
__init__ (amp=1.0,    normalize=False,    dens=<function    DiskSCFPotential.<lambda>>,
Sigma={'amp': 1.0, 'h': 0.3333333333333333, 'type': 'exp'}, hz={'h':
0.037037037037037035, 'type': 'exp'}, Sigma_amp=None, dSigma_dR=None,
d2Sigma_dR2=None, Hz=None, dHzdz=None, N=10, L=10, a=1.0, radial_order=None,
costheta_order=None, phi_order=None, ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a DiskSCF Potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); cannot have units currently

normalize - if True, normalize such that $vc(1.,0.)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1.,0.)=1.$

ro=, vo= distance and velocity scales for translation into internal units (default from configuration file)

dens= function of $R, z, [\text{phi optional}]$ that gives the density [in natural units, cannot return a Quantity currently]

N=, L=, a=, radial_order=, costheta_order=, phi_order= keywords setting parameters for SCF solution for Phi_ME (see [scf_compute_coeffs_axi](#) or [scf_compute_coeffs](#) depending on whether $\rho(R, \phi, z)$ is axisymmetric or not)

Either:

(a) Sigma= Dictionary of surface density (example: { 'type': 'exp', 'h': 1./3., 'amp': 1., 'Rhole': 0. } for $\text{amp} \times \exp(-R_{\text{hole}}/R - R/h)$)

hz= Dictionary of vertical profile, either 'exp' or 'sech2' (example { 'type': 'exp', 'h': 1./27. } for $\exp(-|z|/h)/[2h]$, sech2 is $\text{sech}^2(z/[2h])/[4h]$)

(b) Sigma= function of R that gives the surface density

dSigmadR= function that gives $d \text{Sigma} / d R$

d2SigmadR2= function that gives $d^2 \text{Sigma} / d R^2$

Sigma_amp= amplitude to apply to all Sigma functions

hz= function of z that gives the vertical profile

Hz= function of z such that $d^2 \text{Hz}(z) / d z^2 = \text{hz}$

dHzdz= function of z that gives $d \text{Hz}(z) / d z$

In both of these cases lists of arguments can be given for multiple disk components; can't mix (a) and (b) in these lists; if hz is a single item the same vertical profile is assumed for all Sigma

OUTPUT:

DiskSCFPotential object

HISTORY:

2016-12-26 - Written - Bovy (UofT)

Hernquist & Ostriker Self-Consistent-Field-type potential

class galpy.potential.SCFPotential (amp=1.0, Acos=array([[[[1]]]]), Asin=None, a=1.0, normalize=False, ro=None, vo=None)

Class that implements the [Hernquist & Ostriker \(1992\)](#) Self-Consistent-Field-type potential. Note that we divide the amplitude by 2 such that $A_{cos} = \delta_{0n} \delta_{0l} \delta_{0m}$ and $A_{sin} = 0$ corresponds to [Galpy's Hernquist Potential](#).

$$\rho(r, \theta, \phi) = \frac{\text{amp}}{2} \sum_{n=0}^{\infty} \sum_{l=0}^{\infty} \sum_{m=0}^l N_{lm} P_{lm}(\cos(\theta)) \tilde{\rho}_{nl}(r) (A_{cos,nlm} \cos(m\phi) + A_{sin,nlm} \sin(m\phi))$$

where

$$\tilde{\rho}_{nl}(r) = \frac{K_{nl}}{\sqrt{\pi}} \frac{(ar)^l}{(r/a)(a+r)^{2l+3}} C_n^{2l+3/2}(\xi)$$

$$\Phi(r, \theta, \phi) = \sum_{n=0}^{\infty} \sum_{l=0}^{\infty} \sum_{m=0}^l N_{lm} P_{lm}(\cos(\theta)) \tilde{\Phi}_{nl}(r) (A_{cos,nlm} \cos(m\phi) + A_{sin,nlm} \sin(m\phi))$$

where

$$\tilde{\Phi}_{nl}(r) = -\sqrt{4\pi} K_{nl} \frac{(ar)^l}{(a+r)^{2l+1}} C_n^{2l+3/2}(\xi)$$

where

$$\xi = \frac{r-a}{r+a} \quad N_{lm} = \sqrt{\frac{2l+1}{4\pi} \frac{(l-m)!}{(l+m)!} (2-\delta_{m0})} \quad K_{nl} = \frac{1}{2} n(n+4l+3) + (l+1)(2l+1)$$

and P_{lm} is the Associated Legendre Polynomials whereas C_n^α is the Gegenbauer polynomial.

`__init__` (*amp=1.0, Acos=array([[[1]]])*, *Asin=None*, *a=1.0, normalize=False, ro=None, vo=None*)
NAME:

`__init__`

PURPOSE:

initialize a SCF Potential

INPUT:

amp - amplitude to be applied to the potential (default: 1); can be a Quantity with units of mass or Gxmass

Acos - The real part of the expansion coefficient (NxLxL matrix, or optionally NxLx1 if *Asin=None*)

Asin - The imaginary part of the expansion coefficient (NxLxL matrix or None)

a - scale length (can be Quantity)

normalize - if True, normalize such that $vc(1,0)=1.$, or, if given as a number, such that the force is this fraction of the force necessary to make $vc(1,0)=1.$

ro, *vo*= distance and velocity scales for translation into internal units (default from configuration file)

OUTPUT:

SCFPotential object

HISTORY:

2016-05-13 - Written - Aladdin

Dissipative forces

Chandrasekhar dynamical friction

```
class galpy.potential.ChandrasekharDynamicalFrictionForce (amp=1.0,    GMs=0.1,
                                                         gamma=1.0,
                                                         rhm=0.0,  dens=None,
                                                         sigmar=None,
                                                         const_lnLambda=False,
                                                         minr=0.0001,
                                                         maxr=25.0,   nr=501,
                                                         ro=None, vo=None)
```

Class that implements the Chandrasekhar dynamical friction force

$$\mathbf{F}(\mathbf{x}, \mathbf{v}) = -2\pi [G M] [G \rho(\mathbf{x})] \ln[1 + \Lambda^2] \left[\operatorname{erf}(X) - \frac{2X}{\sqrt{\pi}} \exp(-X^2) \right] \frac{\mathbf{v}}{|\mathbf{v}|^3}$$

on a mass (e.g., a satellite galaxy or a black hole) M at position \mathbf{x} moving at velocity \mathbf{v} through a background density ρ . The quantity X is the usual $X = |\mathbf{v}|/[\sqrt{2}\sigma_r(r)]$. The factor Λ that goes into the Coulomb logarithm is taken to be

$$\Lambda = \frac{r/\gamma}{\max(r_{\text{hm}}, GM/|\mathbf{v}|^2)},$$

where γ is a constant. This γ should be the absolute value of the logarithmic slope of the density $\gamma = |d \ln \rho / d \ln r|$, although for $\gamma < 1$ it is advisable to set $\gamma = 1$. Implementation here roughly follows [2016MNRAS.463..858P](#) and earlier work.

```
__init__ (amp=1.0,    GMs=0.1,    gamma=1.0,    rhm=0.0,    dens=None,    sigmar=None,
          const_lnLambda=False, minr=0.0001, maxr=25.0, nr=501, ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

initialize a Chandrasekhar Dynamical Friction force

INPUT:

amp - amplitude to be applied to the potential (default: 1)

GMs - satellite mass; can be a Quantity with units of mass or Gxmass; can be adjusted after initialization by setting obj.GMs= where obj is your ChandrasekharDynamicalFrictionForce instance

rhm - half-mass radius of the satellite (set to zero for a black hole; can be a Quantity); can be adjusted after initialization by setting obj.rhm= where obj is your ChandrasekharDynamicalFrictionForce instance

gamma - Free-parameter in Λ

dens - Potential instance or list thereof that represents the density [default: LogarithmicHaloPotential(normalize=1.,q=1.)]

sigmar= (None) function that gives the velocity dispersion as a function of r (has to be in natural units!); if None, computed from the dens potential using the spherical Jeans equation (in galpy.df.jeans) assuming zero anisotropy

cont_lnLambda= (False) if set to a number, use a constant $\ln(\Lambda)$ instead with this value

minr= (0.0001) minimum r at which to apply dynamical friction: at $r < \text{minr}$, friction is set to zero (can be a Quantity)

Interpolation:

maxr= (25) maximum r for which sigmar gets interpolated; for best performance set this to the maximum r you will consider (can be a Quantity)

nr= (501) number of radii to use in the interpolation of sigmar

You can check that sigmar is interpolated correctly by comparing the methods sigmar [the interpolated version] and sigmar_orig [the original or directly computed version]

OUTPUT:

(none)

HISTORY:

2011-12-26 - Started - Bovy (NYU)

2018-03-18 - Re-started: updated to r dependent Lambda form and integrated into galpy framework - Bovy (UofT)

2018-07-23 - Calculate sigmar from the Jeans equation and interpolate it; allow GMs and rhm to be set on the fly - Bovy (UofT)

In addition to these classes, a simple Milky-Way-like potential fit to data on the Milky Way is included as `galpy.potential.MWPotential2014` (see the `galpy` paper for details). Note that this potential assumes a circular velocity of 220 km/s at the solar radius at 8 kpc; see [arXiv/1412.3451](https://arxiv.org/abs/1412.3451) for full information on how this potential was fit. This potential is defined as

```
>>> bp= PowerSphericalPotentialwCutoff(alpha=1.8,rc=1.9/8.,normalize=0.05)
>>> mp= MiyamotoNagaiPotential(a=3./8.,b=0.28/8.,normalize=.6)
>>> np= NFWPotential(a=16/8.,normalize=.35)
>>> MWPotential2014= [bp,mp,np]
```

and can thus be used like any list of Potentials. If one wants to add the supermassive black hole at the Galactic center, this can be done by

```
>>> from galpy.potential import KeplerPotential
>>> from galpy.util import bovy_conversion
>>> MWPotential2014wBH= [MWPotential2014,KeplerPotential(amp=4*10**6./bovy_conversion.
↳mass_in_msol(220.,8.))]
```

for a black hole with a mass of $4 \times 10^6 M_\odot$ (this works because a list of Potential instances can contain a nested list of Potential instances in versions ≥ 1.4). If you want to take into account dynamical friction for, say, an object of mass $5 \times 10^{10} M_\odot$ and a half-mass radius of 5 kpc, do

```
>>> from galpy.potential import ChandrasekharDynamicalFrictionForce
>>> from astropy import units
>>> cdf= ChandrasekharDynamicalFrictionForce(GMs=5.*10.**10.*units.Msun,
                                             rhm=5.*units.kpc,
                                             dens=MWPotential2014)
>>> MWPotential2014wDF= [MWPotential2014,cdf]
```

where we have specified the parameters of the dynamical friction with units; alternatively, convert them directly to galpy natural units as

```
>>> cdf= ChandrasekharDynamicalFrictionForce(GMs=5.*10.**10./bovy_conversion.mass_in_
↳msol(220.,8.),
                                             rhm=5./8.,
                                             dens=MWPotential2014)
>>> MWPotential2014wDF= [MWPotential2014,cdf]
```


As explained in [this section](#), *without* this black hole or dynamical friction, `MWPotential2014` can be used with Dehnen’s `gyrfalcON` code using `accname=PowSphwCut+MiyamotoNagai+NFW` and `accpars=0,1001.79126907,1.8,1.9#0,306770.418682,3.0,0.28#0,16.0,162.958241887`.

An older version `galpy.potential.MWPotential` of a similar potential that was *not* fit to data on the Milky Way is defined as

```
>>> mp= MiyamotoNagaiPotential(a=0.5,b=0.0375,normalize=.6)
>>> np= NFWPotential(a=4.5,normalize=.35)
>>> hp= HernquistPotential(a=0.6/8,normalize=0.05)
>>> MWPotential= [mp,np,hp]
```

`galpy.potential.MWPotential2014` **supersedes** `galpy.potential.MWPotential`.

3.2.2 2D potentials

General instance routines

Use as `Potential-instance.method(...)`

`galpy.potential.planarPotential.__call__`

`planarPotential.__call__` (*R*, *phi*=0.0, *t*=0.0, *dR*=0, *dphi*=0)

NAME:

`__call__`

PURPOSE:

evaluate the potential

INPUT:

R - Cylindrica radius (can be Quantity)

phi= azimuth (optional; can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

$\Phi(R, \phi, t)$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

`galpy.potential.planarPotential.phiforce`

`planarPotential.phiforce` (*R*, *phi*=0.0, *t*=0.0)

NAME:

`phiforce`

PURPOSE:

evaluate the phi force

INPUT:

R - Cylindrical radius (can be Quantity)
phi= azimuth (optional; can be Quantity)
t= time (optional; can be Quantity)

OUTPUT:

$F_{\phi}(R,(\phi,t))$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.planarPotential.Rforce

`planarPotential.Rforce (R, phi=0.0, t=0.0)`

NAME:

Rforce

PURPOSE:

evaluate the radial force

INPUT:

R - Cylindrical radius (can be Quantity)
phi= azimuth (optional; can be Quantity)
t= time (optional; can be Quantity)

OUTPUT:

$F_R(R,(\phi,t))$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.planarPotential.turn_physical_off

`planarPotential.turn_physical_off()`

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.planarPotential.turn_physical_on

`planarPotential.turn_physical_on(ro=None, vo=None)`

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

General axisymmetric potential instance routines

Use as `Potential-instance.method(...)`

galpy.potential.planarAxiPotential.epifreq

`Potential.epifreq(R)`

NAME:

epifreq

PURPOSE:

calculate the epicycle frequency at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

epicycle frequency

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.planarAxiPotential.lindbladR

`Potential.lindbladR(OmegaP, m=2, **kwargs)`

NAME:

lindbladR

PURPOSE:

calculate the radius of a Lindblad resonance

INPUT:

OmegaP - pattern speed (can be Quantity)

m= order of the resonance (as in m(O-Op)=kappa (negative m for outer) use m='corotation'
for corotation +scipy.optimize.brentq xtol,rtol,maxiter kwargs

OUTPUT:

radius of Linblad resonance, None if there is no resonance

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.planarAxiPotential.omegac

Potential.omegac(R)

NAME:

omegac

PURPOSE:

calculate the circular angular speed at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

circular angular speed

HISTORY:

2011-10-09 - Written - Bovy (IAS)

galpy.potential.planarAxiPotential.plot

planarAxiPotential.plot(*args,**kwargs)

NAME: plot

PURPOSE: plot the potential

INPUT: Rrange - range (can be Quantity) grid - number of points to plot savefilename - save to or restore from
this savefile (pickle) +bovy_plot(*args,**kwargs)

OUTPUT: plot to output device

HISTORY: 2010-07-13 - Written - Bovy (NYU)

galpy.potential.planarAxiPotential.plotEscapecurve

planarAxiPotential.plotEscapecurve(*args,**kwargs)

NAME:

plotEscapecurve

PURPOSE:

plot the escape velocity curve for this potential

INPUT:

Rrange - range (can be Quantity)

grid - number of points to plot

savefilename - save to or restore from this savefile (pickle)

+bovy_plot(*args,**kwargs)

OUTPUT:

plot to output device

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.planarAxiPotential.plotRotcurve

planarAxiPotential.plotRotcurve(*args,**kwargs)

NAME:

plotRotcurve

PURPOSE:

plot the rotation curve for this potential

INPUT:

Rrange - range (can be Quantity)

grid - number of points to plot

savefilename - save to or restore from this savefile (pickle)

+bovy_plot(*args,**kwargs)

OUTPUT:

plot to output device

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.planarAxiPotential.vcirc

Potential.vcirc(R, phi=None)

NAME:

vcirc

PURPOSE:

calculate the circular velocity at R in this potential

INPUT:

R - Galactocentric radius (can be Quantity)

phi= (None) azimuth to use for non-axisymmetric potentials

OUTPUT:

circular rotation velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

2016-06-15 - Added phi= keyword for non-axisymmetric potential - Bovy (UofT)

galpy.potential.planarAxiPotential.vesc

Potential.**vesc** (*R*)

NAME:

vesc

PURPOSE:

calculate the escape velocity at *R* for this potential

INPUT:

R - Galactocentric radius (can be Quantity)

OUTPUT:

escape velocity

HISTORY:

2011-10-09 - Written - Bovy (IAS)

General 2D potential routines

Use as `method(...)`

galpy.potential.evaluateplanarphiforces

galpy.potential.**evaluateplanarphiforces** (*Pot*, *R*, *phi=None*, *t=0.0*)

NAME:

evaluateplanarphiforces

PURPOSE:

evaluate the phiforce of a (list of) planarPotential instance(s)

INPUT:

Pot - (list of) planarPotential instance(s)

R - Cylindrical radius (can be Quantity)

phi= azimuth (optional; can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

$F_{\phi}(R, \phi, t)$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.evaluateplanarPotentials

`galpy.potential.evaluateplanarPotentials` (*Pot, R, phi=None, t=0.0, dR=0, dphi=0*)

NAME:

`evaluateplanarPotentials`

PURPOSE:

evaluate a (list of) planarPotential instance(s)

INPUT:

Pot - (list of) planarPotential instance(s)

R - Cylindrical radius (can be Quantity)

phi= azimuth (optional; can be Quantity)

t= time (optional; can be Quantity)

dR=, dphi= if set to non-zero integers, return the dR,dphi't derivative instead

OUTPUT:

$\Phi(R, \phi, t)$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.evaluateplanarRforces

`galpy.potential.evaluateplanarRforces` (*Pot, R, phi=None, t=0.0*)

NAME:

`evaluateplanarRforces`

PURPOSE:

evaluate the Rforce of a (list of) planarPotential instance(s)

INPUT:

Pot - (list of) planarPotential instance(s)

R - Cylindrical radius (can be Quantity)

phi= azimuth (optional can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

$F_R(R, \phi, t)$

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.evaluateplanarR2derivs

galpy.potential.**evaluateplanarR2derivs** (*Pot, R, phi=None, t=0.0*)

NAME:

evaluateplanarR2derivs

PURPOSE:

evaluate the second radial derivative of a (list of) planarPotential instance(s)

INPUT:

Pot - (list of) planarPotential instance(s)

R - Cylindrical radius (can be Quantity)

phi= azimuth (optional; can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

F_R(R(,phi,t))

HISTORY:

2010-10-09 - Written - Bovy (IAS)

galpy.potential.LinShuReductionFactor

galpy.potential.**LinShuReductionFactor** (*axiPot, R, sigmar, nonaxiPot=None, k=None, m=None, OmegaP=None*)

NAME:

LinShuReductionFactor

PURPOSE:

Calculate the Lin & Shu (1966) reduction factor: the reduced linear response of a kinematically-warm stellar disk to a perturbation

INPUT:

axiPot - The background, axisymmetric potential

R - Cylindrical radius (can be Quantity)

sigmar - radial velocity dispersion of the population (can be Quantity)

Then either provide:

1) m= m in the perturbation's m x phi (number of arms for a spiral)

k= wavenumber (see Binney & Tremaine 2008)

OmegaP= pattern speed (can be Quantity)

2) nonaxiPot= a non-axisymmetric Potential instance (such as SteadyLogSpiralPotential) that has functions that return OmegaP, m, and wavenumber

OUTPUT:

reduction factor

HISTORY:

2014-08-23 - Written - Bovy (IAS)

galpy.potential.plotplanarPotentials

`galpy.potential.plotplanarPotentials` (*Pot*, *args, **kwargs)

NAME:

plotplanarPotentials

PURPOSE:

plot a planar potential

INPUT:

Range - range (can be Quantity)

xrange, yrange - if relevant (can be Quantity)

grid, gridx, gridy - number of points to plot

savefilename - save to or restore from this savefile (pickle)

ncontours - number of contours to plot (if applicable)

+bovy_plot(*args,**kwargs) or bovy_dens2d(**kwargs)

OUTPUT:

plot to output device

HISTORY:

2010-07-13 - Written - Bovy (NYU)

Specific potentials

All of the 3D potentials above can be used as two-dimensional potentials in the mid-plane.

galpy.potential.toPlanarPotential

`galpy.potential.toPlanarPotential` (*Pot*)

NAME:

toPlanarPotential

PURPOSE:

convert an Potential to a planarPotential in the mid-plane ($z=0$)

INPUT:

Pot - Potential instance or list of such instances (existing planarPotential instances are just copied to the output)

OUTPUT:

planarPotential instance(s)

HISTORY:

2016-06-11 - Written - Bovy (UofT)

galpy.potential.RZToplanarPotential

galpy.potential.RZToplanarPotential (RZPot)

NAME:

RZToplanarPotential

PURPOSE:

convert an RZPotential to a planarPotential in the mid-plane (z=0)

INPUT:

RZPot - RZPotential instance or list of such instances (existing planarPotential instances are just copied to the output)

OUTPUT:

planarPotential instance(s)

HISTORY:

2010-07-13 - Written - Bovy (NYU)

In addition, a two-dimensional bar potential, two spiral potentials, the [Henon & Heiles \(1964\)](#) potential, and some static non-axisymmetric perturbations are included

Cos(m phi) disk potential

Generalization of the *lopsided* and *elliptical* disk potentials to any m and to allow for a break radius within which the radial dependence of the potential changes from R^p to R^{-p} .

class galpy.potential.CosmphiDiskPotential (amp=1.0, phib=0.4363323129985824, p=1.0, phio=0.01, m=4, r1=1.0, rb=None, cp=None, sp=None, ro=None, vo=None)

Class that implements the disk potential

$$\Phi(R, \phi) = \text{amp } \phi_0 \cos [m (\phi - \phi_b)] \times \begin{cases} \left(\frac{R}{R_1}\right)^p, & \text{for } R \geq R_b \\ \left[2 - \left(\frac{R_b}{R}\right)^p\right] \times \left(\frac{R_b}{R_1}\right)^p, & \text{for } R \leq R_b. \end{cases}$$

This potential can be grown between t_{form} and $t_{\text{form}} + T_{\text{steady}}$ in a similar way as DehnenBarPotential by wrapping it with a DehnenSmoothWrapperPotential

__init__ (amp=1.0, phib=0.4363323129985824, p=1.0, phio=0.01, m=4, r1=1.0, rb=None, cp=None, sp=None, ro=None, vo=None)

NAME:

__init__

PURPOSE:

initialize an cosmphi disk potential

INPUT:

amp= amplitude to be applied to the potential (default: 1.), degenerate with phio below, but kept for overall consistency with potentials

m= cos(m * (phi - phib)), integer

p= power-law index of the $\phi(R) = (R/R_0)^p$ part

$r1 = (1.)$ normalization radius for the amplitude (can be Quantity); $\text{amp} \times \text{phio}$ is only the potential at $(R, \phi) = (r1, \text{pib})$ when $r1 > r_b$; otherwise more complicated

$r_b = (\text{None})$ if set, break radius for power-law: potential R^p at $R > R_b$, R^{-p} at $R < R_b$, potential and force continuous at R_b

Either:

a) $\text{phib} =$ angle (in rad; default=25 degree; or can be Quantity)

$\text{phio} =$ potential perturbation (in terms of phio/v_o^2 if $v_o=1$ at $R_o=1$; or can be Quantity with units of velocity-squared)

b) $\text{cp}, \text{sp} = m * \text{phio} * \cos(m * \text{phib}), m * \text{phio} * \sin(m * \text{phib})$; can be Quantity with units of velocity-squared)

OUTPUT:

(none)

HISTORY:

2011-10-27 - Started - Bovy (IAS)

2017-09-16 - Added break radius r_b - Bovy (UofT)

Elliptical disk potential

Like in [Kuijken & Tremaine](#). See [galpy.potential.CosmphiDiskPotential](#) for a more general version that allows for a break radius within which the radial dependence of the potential changes from R^p to R^{-p} (elliptical disk corresponds to $m=2$).

```
class galpy.potential.EllipticalDiskPotential (amp=1.0,    phib=0.4363323129985824,
                                             p=1.0,      twophio=0.01,    r1=1.0,
                                             tform=None,  tsteady=None,  cp=None,
                                             sp=None,    ro=None,    vo=None)
```

Class that implements the Elliptical disk potential of Kuijken & Tremaine (1994)

$$\Phi(R, \phi) = \text{amp} \phi_0 \left(\frac{R}{R_1} \right)^p \cos(2(\phi - \phi_b))$$

This potential can be grown between t_{form} and $t_{\text{form}} + T_{\text{steady}}$ in a similar way as [DehnenBarPotential](#), but times are given directly in galpy time units

```
__init__ (amp=1.0,  phib=0.4363323129985824,  p=1.0,  twophio=0.01,  r1=1.0,  tform=None,
          tsteady=None,  cp=None,  sp=None,  ro=None,  vo=None)
```

NAME:

__init__

PURPOSE:

initialize an Elliptical disk potential

$\text{phi}(R, \phi) = \text{phio} (R/R_o)^p \cos[2(\phi - \text{phib})]$

INPUT:

$\text{amp} =$ amplitude to be applied to the potential (default: 1.), see twophio below

$t_{\text{form}} =$ start of growth (to smoothly grow this potential (can be Quantity)

$t_{\text{steady}} =$ time delay at which the perturbation is fully grown (default: 2.; can be Quantity)

p= power-law index of the $\phi(R) = (R/R_0)^p$ part

r1= (1.) normalization radius for the amplitude (can be Quantity)

Either:

a) phib= angle (in rad; default=25 degree; or can be Quantity)

twophio= potential perturbation (in terms of $2\phi_0/v_0^2$ if $v_0=1$ at $R_0=1$; can be Quantity with units of velocity-squared)

b) cp, sp= twophio * cos(2phib), twophio * sin(2phib) (can be Quantity with units of velocity-squared)

OUTPUT:

(none)

HISTORY:

2011-10-19 - Started - Bovy (IAS)

Henon-Heiles potential

class galpy.potential.HenonHeilesPotential (amp=1.0, ro=None, vo=None)

Class that implements a the [Henon & Heiles \(1964\)](#) potential

$$\Phi(R, \phi) = \frac{\text{amp}}{2} \left[R^2 + \frac{2 R^3}{3} \sin(3 \phi) \right]$$

__init__ (amp=1.0, ro=None, vo=None)

NAME:

__init__

PURPOSE:

initialize a Henon-Heiles potential

INPUT:

amp - amplitude to be applied to the potential (default: 1.)

OUTPUT:

(none)

HISTORY:

2017-10-16 - Written - Bovy (UofT)

Lopsided disk potential

Like in [Kuijken & Tremaine](#), but for $m=1$. See [galpy.potential.CosmphiDiskPotential](#) for a more general version that allows for a break radius within which the radial dependence of the potential changes from R^p to R^{-p} (lopsided disk corresponds to $m=1$).

class galpy.potential.LopsidedDiskPotential (amp=1.0, phib=0.4363323129985824,
p=1.0, phio=0.01, r1=1.0, cp=None,
sp=None, ro=None, vo=None)

Class that implements the disk potential

$$\Phi(R, \phi) = \text{amp} \phi_0 \left(\frac{R}{R_1} \right)^p \cos(\phi - \phi_b)$$

Special case of CosmphiDiskPotential with m=1; see documentation for CosmphiDiskPotential

__init__ (*amp=1.0, phib=0.4363323129985824, p=1.0, phio=0.01, r1=1.0, cp=None, sp=None, ro=None, vo=None*)

NAME:

`__init__`

PURPOSE:

initialize an cosmphi disk potential

INPUT:

amp= amplitude to be applied to the potential (default: 1.), degenerate with phio below, but kept for overall consistency with potentials

m= cos(m * (phi - phib)), integer

p= power-law index of the phi(R) = (R/Ro)^p part

r1= (1.) normalization radius for the amplitude (can be Quantity); amp x phio is only the potential at (R,phi) = (r1,pib) when r1 > rb; otherwise more complicated

rb= (None) if set, break radius for power-law: potential R^p at R > Rb, R^-p at R < Rb, potential and force continuous at Rb

Either:

a) phib= angle (in rad; default=25 degree; or can be Quantity)

phio= potential perturbation (in terms of phio/vo^2 if vo=1 at Ro=1; or can be Quantity with units of velocity-squared)

b) cp, sp= m * phio * cos(m * phib), m * phio * sin(m * phib); can be Quantity with units of velocity-squared)

OUTPUT:

(none)

HISTORY:

2011-10-27 - Started - Bovy (IAS)

2017-09-16 - Added break radius rb - Bovy (UofT)

Steady-state logarithmic spiral potential

```
class galpy.potential.SteadyLogSpiralPotential (amp=1.0,      omegas=0.65,      A=-  
                                              0.035,      alpha=-7.0,      m=2,  
                                              gamma=0.7853981633974483,  
                                              p=None, tform=None, tsteady=None,  
                                              ro=None, vo=None)
```

Class that implements a steady-state spiral potential

$$\Phi(R, \phi) = \frac{\text{amp} \times A}{\alpha} \cos(\alpha \ln R - m(\phi - \Omega_s t - \gamma))$$

Can be grown in a similar way as the DehnenBarPotential, but using $T_s = 2\pi/\Omega_s$ to normalize t_{form} and T_{steady} .

```
__init__(amp=1.0, omegas=0.65, A=-0.035, alpha=-7.0, m=2, gamma=0.7853981633974483,  
         p=None, tform=None, tsteady=None, ro=None, vo=None)
```

NAME:

__init__

PURPOSE:

initialize a logarithmic spiral potential

INPUT:

amp - amplitude to be applied to the potential (default: 1., A below)

gamma - angle between sun-GC line and the line connecting the peak of the spiral pattern at the Solar radius (in rad; default=45 degree; or can be Quantity)

A - amplitude (alpha*potential-amplitude; default=0.035; can be Quantity)

omegas= - pattern speed (default=0.65; can be Quantity)

m= number of arms

Either provide:

a) alpha=

b) p= pitch angle (rad; can be Quantity)

tform - start of spiral growth / spiral period (default: -Infinity)

tsteady - time from tform at which the spiral is fully grown / spiral period (default: 2 periods)

OUTPUT:

(none)

HISTORY:

2011-03-27 - Started - Bovy (NYU)

Transient logarithmic spiral potential

```
class galpy.potential.TransientLogSpiralPotential(amp=1.0, omegas=0.65, A=-  
                                                    0.035, alpha=-7.0, m=2,  
                                                    gamma=0.7853981633974483,  
                                                    p=None, sigma=1.0, to=0.0,  
                                                    ro=None, vo=None)
```

Class that implements a steady-state spiral potential

$$\Phi(R, \phi) = \frac{\text{amp}(t)}{\alpha} \cos(\alpha \ln R - m(\phi - \Omega_s t - \gamma))$$

where

$$\text{amp}(t) = \text{amp} \times A \exp\left(-\frac{[t - t_0]^2}{2\sigma^2}\right)$$

```
__init__(amp=1.0, omegas=0.65, A=-0.035, alpha=-7.0, m=2, gamma=0.7853981633974483,  
         p=None, sigma=1.0, to=0.0, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a transient logarithmic spiral potential localized around to

INPUT:

amp - amplitude to be applied to the potential (default: 1., A below)

gamma - angle between sun-GC line and the line connecting the peak of the spiral pattern at the Solar radius (in rad; default=45 degree; can be Quantity)

A - amplitude (alpha*potential-amplitude; default=0.035; can be Quantity)

omegas= - pattern speed (default=0.65; can be Quantity)

m= number of arms

to= time at which the spiral peaks (can be Quantity)

sigma= “spiral duration” (sigma in Gaussian amplitude; can be Quantity)

Either provide:

a) alpha=

b) p= pitch angle (rad; can be Quantity)

OUTPUT:

(none)

HISTORY:

2011-03-27 - Started - Bovy (NYU)

3.2.3 1D potentials

General instance routines

Use as `Potential-instance.method(...)`

galpy.potential.linearPotential.__call__

`linearPotential.__call__(x, t=0.0)`

NAME: `__call__`

PURPOSE:

evaluate the potential

INPUT:

x - position (can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

Phi(x,t)

HISTORY:

2010-07-12 - Written - Bovy (NYU)

galpy.potential.linearPotential.force

`linearPotential.force(x, t=0.0)`

NAME:

force

PURPOSE:

evaluate the force

INPUT:

x - position (can be Quantity)

t= time (optional; can be Quantity)

OUTPUT:

$F(x,t)$

HISTORY:

2010-07-12 - Written - Bovy (NYU)

galpy.potential.linearPotential.plot

`linearPotential.plot(t=0.0, min=-15.0, max=15, ns=21, savefilename=None)`

NAME:

plot

PURPOSE:

plot the potential

INPUT:

t - time to evaluate the potential at

min - minimum x

max - maximum x

ns - grid in x

savefilename - save to or restore from this savefile (pickle)

OUTPUT:

plot to output device

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.linearPotential.turn_physical_off

```
linearPotential.turn_physical_off()
```

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

galpy.potential.linearPotential.turn_physical_on

```
linearPotential.turn_physical_on(ro=None, vo=None)
```

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-01-30 - Written - Bovy (UofT)

General 1D potential routines

Use as `method(...)`

galpy.potential.evaluateLinearForces

```
galpy.potential.evaluateLinearForces(Pot, x, t=0.0)
```

NAME:

evaluateLinearForces

PURPOSE:

evaluate the forces due to a list of potentials

INPUT:

Pot - (list of) linearPotential instance(s)
x - evaluate forces at this position (can be Quantity)
t - time to evaluate at (can be Quantity)

OUTPUT:

force(x,t)

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.evaluatelinearPotentials

`galpy.potential.evaluatelinearPotentials` (*Pot, x, t=0.0*)

NAME:

evaluatelinearPotentials

PURPOSE:

evaluate the sum of a list of potentials

INPUT:

Pot - (list of) linearPotential instance(s)
x - evaluate potentials at this position (can be Quantity)
t - time to evaluate at (can be Quantity)

OUTPUT:

pot(x,t)

HISTORY:

2010-07-13 - Written - Bovy (NYU)

galpy.potential.plotlinearPotentials

`galpy.potential.plotlinearPotentials` (*Pot, t=0.0, min=-15.0, max=15, ns=21, savefilename=None*)

NAME:

plotlinearPotentials

PURPOSE:

plot a combination of potentials

INPUT:

t - time to evaluate potential at
min - minimum x
max - maximum x
ns - grid in x
savefilename - save to or restore from this savefile (pickle)

OUTPUT:

plot to output device

HISTORY:

2010-07-13 - Written - Bovy (NYU)

Specific potentials

Vertical Kuijken & Gilmore potential

class `galpy.potential.KGPotential` ($K=1.15$, $F=0.03$, $D=1.8$, $amp=1.0$, $ro=None$, $vo=None$)

Class representing the Kuijken & Gilmore (1989) potential

$$\Phi(x) = \text{amp} \left(K \left(\sqrt{x^2 + D^2} - D \right) + F x^2 \right)$$

__init__ ($K=1.15$, $F=0.03$, $D=1.8$, $amp=1.0$, $ro=None$, $vo=None$)

NAME:

`__init__`

PURPOSE:

Initialize a KGPotential

INPUT:

K = K parameter ($= 2\pi\Sigma_{\text{disk}}$; specify either as surface density or directly as force [i.e., including $2\pi G$]; can be Quantity)

F = F parameter ($= 4\pi\rho_{\text{halo}}$; specify either as density or directly as second potential derivative [i.e., including $4\pi G$]; can be Quantity)

D = D parameter (natural units or Quantity length units)

amp - an overall amplitude

OUTPUT:

instance

HISTORY:

2010-07-12 - Written - Bovy (NYU)

One-dimensional potentials can also be derived from 3D axisymmetric potentials as the vertical potential at a certain Galactocentric radius

`galpy.potential.RZToverticalPotential`

`galpy.potential.RZToverticalPotential` ($RZPot$, R)

NAME:

`RZToverticalPotential`

PURPOSE:

convert a RZPotential to a vertical potential at a given R

INPUT:

RZPot - RZPotential instance or list of such instances

R - Galactocentric radius at which to evaluate the vertical potential (can be Quantity)

OUTPUT:

(list of) linearPotential instance(s)

HISTORY:

2010-07-21 - Written - Bovy (NYU)

3.2.4 Potential wrappers

Gravitational potentials in `galpy` can also be modified using wrappers, for example, to change their amplitude as a function of time. These wrappers can be applied to *any* `galpy` potential (although whether they can be used in `C` depends on whether the wrapper *and* all of the potentials that it wraps are implemented in `C`). Multiple wrappers can be applied to the same potential.

Specific wrappers

Corotating rotation wrapper potential

```
class galpy.potential.CorotatingRotationWrapperPotential (amp=1.0, pot=None,  
                                                         vpo=1.0, beta=0.0,  
                                                         to=0.0, pa=0.0,  
                                                         ro=None, vo=None)
```

Potential wrapper class that implements rotation with fixed $R \times$ pattern-speed around the z -axis. Can be used to make spiral structure that is everywhere co-rotating. The potential is rotated by replacing

$$\phi \rightarrow \phi + \frac{V_p(R)}{R} \times (t - t_0) + pa$$

with $V_p(R)$ the circular velocity curve, t_0 a reference time—time at which the potential is unchanged by the wrapper—and pa the position angle at $t = 0$. The circular velocity is parameterized as

$$V_p(R) = V_{p,0} \left(\frac{R}{R_0} \right)^\beta.$$

```
__init__ (amp=1.0, pot=None, vpo=1.0, beta=0.0, to=0.0, pa=0.0, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a CorotatingRotationWrapper Potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1.)

`pot` - Potential instance or list thereof; this potential is made to rotate around the z axis by the wrapper

`vpo`= (1.) amplitude of the circular-velocity curve (can be a Quantity)

`beta`= (0.) power-law amplitude of the circular-velocity curve

to= (0.) reference time at which the potential == pot

pa= (0.) the position angle (can be a Quantity)

OUTPUT:

(none)

HISTORY:

2018-02-21 - Started - Bovy (UofT)

Dehnen-like smoothing wrapper potential

```
class galpy.potential.DehnenSmoothWrapperPotential(amp=1.0, pot=None, tform=-4.0, tsteady=None, ro=None, vo=None)
```

Potential wrapper class that implements the growth of a gravitational potential following [Dehnen \(2000\)](#). The amplitude A applied to a potential wrapped by an instance of this class is changed as

$$A(t) = amp \left(\frac{3}{16} \xi^5 - \frac{5}{8} \xi^3 + \frac{15}{16} \xi + \frac{1}{2} \right)$$

where

$$\xi = \begin{cases} 0 & t < t_{\text{form}} \\ 2 \left(\frac{t - t_{\text{form}}}{t_{\text{steady}}} \right) - 1, & t_{\text{form}} \leq t \leq t_{\text{form}} + t_{\text{steady}} \\ 1 & t > t_{\text{form}} + t_{\text{steady}} \end{cases}$$

```
__init__(amp=1.0, pot=None, tform=-4.0, tsteady=None, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a DehnenSmoothWrapper Potential

INPUT:

amp - amplitude to be applied to the potential (default: 1.)

pot - Potential instance or list thereof; the amplitude of this will be grown by this wrapper

tform - start of growth (can be a Quantity)

tsteady - time from tform at which the potential is fully grown (default: -tform/2, st the perturbation is fully grown at tform/2; can be a Quantity)

OUTPUT:

(none)

HISTORY:

2017-06-26 - Started - Bovy (UofT)

Gaussian-modulated amplitude wrapper potential

```
class galpy.potential.GaussianAmplitudeWrapperPotential (amp=1.0,      pot=None,  
                                                         to=0.0,      sigma=1.0,  
                                                         ro=None, vo=None)
```

Potential wrapper class that allows the amplitude of a Potential object to be modulated as a Gaussian. The amplitude A applied to a potential wrapped by an instance of this class is changed as

$$A(t) = amp \exp\left(-\frac{[t - t_0]^2}{2\sigma^2}\right)$$

```
__init__ (amp=1.0, pot=None, to=0.0, sigma=1.0, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a GaussianAmplitudeWrapper Potential

INPUT:

`amp` - amplitude to be applied to the potential (default: 1.)

`pot` - Potential instance or list thereof; this potential is made to rotate around the z axis by the wrapper

`to` = (0.) time at which the Gaussian peaks

`sigma` = (1.) standard deviation of the Gaussian (can be a Quantity)

OUTPUT:

(none)

HISTORY:

2018-02-21 - Started - Bovy (UofT)

Solid-body rotation wrapper potential

```
class galpy.potential.SolidBodyRotationWrapperPotential (amp=1.0,      pot=None,  
                                                         omega=1.0,      pa=0.0,  
                                                         ro=None, vo=None)
```

Potential wrapper class that implements solid-body rotation around the z-axis. Can be used to make a bar or other perturbation rotate. The potential is rotated by replacing

$$\phi \rightarrow \phi + \Omega \times t + pa$$

with Ω the fixed pattern speed and `pa` the position angle at $t = 0$.

```
__init__ (amp=1.0, pot=None, omega=1.0, pa=0.0, ro=None, vo=None)
```

NAME:

`__init__`

PURPOSE:

initialize a SolidBodyRotationWrapper Potential

INPUT:

amp - amplitude to be applied to the potential (default: 1.)

pot - Potential instance or list thereof; this potential is made to rotate around the z axis by the wrapper

omega= (1.) the pattern speed (can be a Quantity)

pa= (0.) the position angle (can be a Quantity)

OUTPUT:

(none)

HISTORY:

2017-08-22 - Started - Bovy (UofT)

3.3 actionAngle (galpy.actionAngle)

3.3.1 (x, v) → (J, O, a)

General instance routines

Not necessarily supported for all different types of actionAngle calculations. These have extra arguments for different actionAngle modules, so check the documentation of the module-specific functions for more info (e.g., `?actionAngleIsochrone.__call__`)

galpy.actionAngle.actionAngle.__call__

actionAngle.__call__ (*args, **kwargs)

NAME:

__call__

PURPOSE:

evaluate the actions (jr,lz,jz)

INPUT:

Either:

a) R,vR,vT,z,vz[,phi]:

1) floats: phase-space value for single object (phi is optional) (each can be a Quantity)

2) numpy.ndarray: [N] phase-space values for N objects (each can be a Quantity)

b) Orbit instance: initial condition used if that's it, orbit(t) if there is a time given as well as the second argument

OUTPUT:

(jr,lz,jz)

HISTORY:

2014-01-03 - Written for top level - Bovy (IAS)

galpy.actionAngle.actionAngle.actionsFreqs

`actionAngle.actionsFreqs(*args, **kwargs)`

NAME:

actionsFreqs

PURPOSE:

evaluate the actions and frequencies (jr,lz,jz,Omegar,Omegaphi,Omegaz)

INPUT:

Either:

a) R,vR,vT,z,vz[,phi]:

1) floats: phase-space value for single object (phi is optional) (each can be a Quantity)

2) numpy.ndarray: [N] phase-space values for N objects (each can be a Quantity)

b) Orbit instance: initial condition used if that's it, orbit(t) if there is a time given as well as the second argument

OUTPUT:

(jr,lz,jz,Omegar,Omegaphi,Omegaz)

HISTORY:

2014-01-03 - Written for top level - Bovy (IAS)

galpy.actionAngle.actionAngle.actionsFreqsAngles

`actionAngle.actionsFreqsAngles(*args, **kwargs)`

NAME:

actionsFreqsAngles

PURPOSE:

evaluate the actions, frequencies, and angles (jr,lz,jz,Omegar,Omegaphi,Omegaz,angler,anglephi,anglez)

INPUT:

Either:

a) R,vR,vT,z,vz,phi:

1) floats: phase-space value for single object (phi is optional) (each can be a Quantity)

2) numpy.ndarray: [N] phase-space values for N objects (each can be a Quantity)

b) Orbit instance: initial condition used if that's it, orbit(t) if there is a time given as well as the second argument

OUTPUT:

(jr,lz,jz,Omegar,Omegaphi,Omegaz,angler,anglephi,anglez)

HISTORY:

2014-01-03 - Written for top level - Bovy (IAS)

galpy.actionAngle.actionAngle.EccZmaxRperiRap`actionAngle.EccZmaxRperiRap(*args, **kwargs)`

NAME:

EccZmaxRperiRap

PURPOSE:

evaluate the eccentricity, maximum height above the plane, peri- and apocenter

INPUT:

Either:

a) R,vR,vT,z,vz[,phi]:

1) floats: phase-space value for single object (phi is optional) (each can be a Quantity)

2) numpy.ndarray: [N] phase-space values for N objects (each can be a Quantity)

b) Orbit instance: initial condition used if that's it, orbit(t) if there is a time given as well as the second argument

OUTPUT:

(e,zmax,rperi,rap)

HISTORY:

2017-12-12 - Written - Bovy (UofT)

galpy.actionAngle.actionAngle.turn_physical_off`actionAngle.turn_physical_off()`

NAME:

turn_physical_off

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2017-06-05 - Written - Bovy (UofT)

galpy.actionAngle.actionAngle.turn_physical_on`actionAngle.turn_physical_on(ro=None, vo=None)`

NAME:

turn_physical_on

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

ro= reference distance (kpc; can be Quantity)

vo= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-06-05 - Written - Bovy (UofT)

Specific actionAngle modules

actionAngleIsochrone

class galpy.actionAngle.actionAngleIsochrone(*args, **kwargs)

Action-angle formalism for the isochrone potential, on the Jphi, Jtheta system of Binney & Tremaine (2008)

__init__(*args, **kwargs)

NAME: __init__

PURPOSE: initialize an actionAngleIsochrone object

INPUT: Either:

b= scale parameter of the isochrone parameter (can be Quantity)

ip= instance of a IsochronePotential

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

OUTPUT:

instance

HISTORY: 2013-09-08 - Written - Bovy (IAS)

actionAngleSpherical

class galpy.actionAngle.actionAngleSpherical(*args, **kwargs)

Action-angle formalism for spherical potentials

__init__(*args, **kwargs)

NAME:

__init__

PURPOSE:

initialize an actionAngleSpherical object

INPUT:

pot= a Spherical potential

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

OUTPUT:

instance

HISTORY:

2013-12-28 - Written - Bovy (IAS)

actionAngleAdiabatic

```
class galpy.actionAngle.actionAngleAdiabatic(*args, **kwargs)
```

Action-angle formalism for axisymmetric potentials using the adiabatic approximation

```
__init__(*args, **kwargs)
```

NAME:

__init__

PURPOSE:

initialize an actionAngleAdiabatic object

INPUT:

pot= potential or list of potentials (planarPotentials)

gamma= (default=1.) replace L_z by $L_z + \gamma J_z$ in effective potential

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

OUTPUT:

instance

HISTORY:

2012-07-26 - Written - Bovy (IAS@MPIA)

actionAngleAdiabaticGrid

```
class galpy.actionAngle.actionAngleAdiabaticGrid(pot=None, zmax=1.0, gamma=1.0,
                                                  Rmax=5.0, nR=16, nEz=16, nEr=31,
                                                  nLz=31, numcores=1, **kwargs)
```

Action-angle formalism for axisymmetric potentials using the adiabatic approximation, grid-based interpolation

```
__init__(pot=None, zmax=1.0, gamma=1.0, Rmax=5.0, nR=16, nEz=16, nEr=31, nLz=31, num-
        cores=1, **kwargs)
```

NAME: __init__

PURPOSE: initialize an actionAngleAdiabaticGrid object

INPUT:

pot= potential or list of potentials
zmax= zmax for building Ez grid
Rmax = Rmax for building grids
gamma= (default=1.) replace Lz by Lz+gamma Jz in effective potential
nEz=, nEr=, nLz, nR= grid size
numcores= number of cpus to use to parallelize
c= if True, use C to calculate actions
ro= distance from vantage point to GC (kpc; can be Quantity)
vo= circular velocity at ro (km/s; can be Quantity)
+scipy.integrate.quad keywords

OUTPUT:

instance

HISTORY:

2012-07-27 - Written - Bovy (IAS@MPIA)

actionAngleStaeckel

class galpy.actionAngle.actionAngleStaeckel (*args, **kwargs)

Action-angle formalism for axisymmetric potentials using Binney (2012)'s Staeckel approximation

__init__ (*args, **kwargs)

NAME: __init__

PURPOSE: initialize an actionAngleStaeckel object

INPUT: pot= potential or list of potentials (3D)

delta= focus (can be Quantity)

useu0 - use u0 to calculate dV (NOT recommended)

c= if True, always use C for calculations

order= (10) number of points to use in the Gauss-Legendre numerical integration of the relevant action, frequency, and angle integrals

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

OUTPUT:

instance

HISTORY:

2012-11-27 - Written - Bovy (IAS)

actionAngleStaeckelGrid

```
class galpy.actionAngle.actionAngleStaeckelGrid(pot=None, delta=None, Rmax=5.0,  
                                                nE=25, npsi=25, nLz=30, num-  
                                                cores=1, interpecc=False, **kwargs)
```

Action-angle formalism for axisymmetric potentials using Binney (2012)'s Staeckel approximation, grid-based interpolation

```
__init__(pot=None, delta=None, Rmax=5.0, nE=25, npsi=25, nLz=30, numcores=1, inter-  
         pecc=False, **kwargs)
```

NAME: `__init__`

PURPOSE: initialize an actionAngleStaeckelGrid object

INPUT: `pot`= potential or list of potentials

`delta`= focus of prolate confocal coordinate system (can be Quantity)

`Rmax` = `Rmax` for building grids (natural units)

`nE`, `npsi`, `nLz`= grid size

`interpecc`= (False) if True, also interpolate the approximate eccentricity, `zmax`, `rperi`, and `rapo`

`numcores`= number of cpus to use to parallllize

`ro`= distance from vantage point to GC (kpc; can be Quantity)

`vo`= circular velocity at `ro` (km/s; can be Quantity)

OUTPUT:

instance

HISTORY:

2012-11-29 - Written - Bovy (IAS)

2017-12-15 - Written - Bovy (UofT)

actionAngleIsochroneApprox

```
class galpy.actionAngle.actionAngleIsochroneApprox(*args, **kwargs)
```

Action-angle formalism using an isochrone potential as an approximate potential and using a Fox & Binney (2014?) like algorithm to calculate the actions using orbit integrations and a torus-machinery-like angle-fit to get the angles and frequencies (Bovy 2014)

```
__init__(*args, **kwargs)
```

NAME: `__init__`

PURPOSE: initialize an actionAngleIsochroneApprox object

INPUT:

Either:

`b`= scale parameter of the isochrone parameter (can be Quantity)

`ip`= instance of a IsochronePotential

`aAI`= instance of an actionAngleIsochrone

pot= potential to calculate action-angle variables for
tintJ= (default: 100) time to integrate orbits for to estimate actions (can be Quantity)
ntintJ= (default: 10000) number of time-integration points
integrate_method= (default: 'dopr54_c') integration method to use
dt= (None) orbit.integrate dt keyword (for fixed stepsize integration)
maxn= (default: 3) Default value for all methods when using a grid in vec(n) up to this n (zero-based)
ro= distance from vantage point to GC (kpc; can be Quantity)
vo= circular velocity at ro (km/s; can be Quantity)

OUTPUT:

instance

HISTORY: 2013-09-10 - Written - Bovy (IAS)

3.3.2 (**J**, **a**) → (**x**, **v**, **O**)

General instance routines

Warning: While the `actionAngleTorus` code below can compute the Jacobian and Hessian of the (**J**, **a**) → (**x**, **v**, **O**) transformation, the accuracy of these does not appear to be very good using the current interface to the `TorusMapper` code, so care should be taken when using these.

Currently, only the interface to the `TorusMapper` code supports going from (**J**, **a**) → (**x**, **v**, **O**). Instance methods are

`galpy.actionAngle.actionAngleTorus.__call__`

`actionAngleTorus.__call__(jr, jphi, jz, angler, anglephi, anglez, **kwargs)`

NAME:

`__call__`

PURPOSE:

evaluate the phase-space coordinates (**x**,**v**) for a number of angles on a single torus

INPUT:

jr - radial action (scalar)

jphi - azimuthal action (scalar)

jz - vertical action (scalar)

angler - radial angle (array [N])

anglephi - azimuthal angle (array [N])

anglez - vertical angle (array [N])

tol= (object-wide value) goal for $|\mathbf{dJ}/\mathbf{J}|$ along the torus

OUTPUT:

[R,vR,vT,z,vz,phi]

HISTORY:

2015-08-07 - Written - Bovy (UofT)

galpy.actionAngle.actionAngleTorus.Freqs

`actionAngleTorus.Freqs(jr, jphi, jz, **kwargs)`

NAME:

Freqs

PURPOSE:

return the frequencies corresponding to a torus

INPUT:

jr - radial action (scalar)

jphi - azimuthal action (scalar)

jz - vertical action (scalar)

tol= (object-wide value) goal for $|dJ/J|$ along the torus

OUTPUT:

(OmegaR, Omegaphi, Omegaz)

HISTORY:

2015-08-07 - Written - Bovy (UofT)

galpy.actionAngle.actionAngleTorus.hessianFreqs

`actionAngleTorus.hessianFreqs(jr, jphi, jz, **kwargs)`

NAME:

hessianFreqs

PURPOSE:

return the Hessian $d\Omega/dJ$ and frequencies Ω corresponding to a torus

INPUT:

jr - radial action (scalar)

jphi - azimuthal action (scalar)

jz - vertical action (scalar)

tol= (object-wide value) goal for $|dJ/J|$ along the torus

dJ= (object-wide value) action difference when computing derivatives (Hessian or Jacobian)

nosym= (False) if True, don't explicitly symmetrize the Hessian (good to check errors)

OUTPUT:

(dO/dJ, Omegar, Omegaphi, Omegaz, Autofit error message)

HISTORY:

2016-07-15 - Written - Bovy (UofT)

galpy.actionAngle.actionAngleTorus.xvFreqs

`actionAngleTorus.xvFreqs(jr, jphi, jz, angler, anglephi, anglez, **kwargs)`

NAME:

xvFreqs

PURPOSE:

evaluate the phase-space coordinates (x,v) for a number of angles on a single torus as well as the frequencies

INPUT:

jr - radial action (scalar)

jphi - azimuthal action (scalar)

jz - vertical action (scalar)

angler - radial angle (array [N])

anglephi - azimuthal angle (array [N])

anglez - vertical angle (array [N])

tol= (object-wide value) goal for $|dJ/J|$ along the torus

OUTPUT:

([R,vR,vT,z,vz,phi],OmegaR,Omegaphi,Omegaz,AutoFit error message)

HISTORY:

2015-08-07 - Written - Bovy (UofT)

galpy.actionAngle.actionAngleTorus.xvJacobianFreqs

`actionAngleTorus.xvJacobianFreqs(jr, jphi, jz, angler, anglephi, anglez, **kwargs)`

NAME:

xvJacobianFreqs

PURPOSE:

return [R,vR,vT,z,vz,phi], the Jacobian $d[R,vR,vT,z,vz,phi] / d(J,angle)$, the Hessian d^2O/dJ^2 , and frequencies Omega corresponding to a torus at multiple sets of angles

INPUT:

jr - radial action (scalar)

jphi - azimuthal action (scalar)

jz - vertical action (scalar)

angler - radial angle (array [N])

anglephi - azimuthal angle (array [N])

anglez - vertical angle (array [N])

tol= (object-wide value) goal for $|dJ/J|$ along the torus

dJ= (object-wide value) action difference when computing derivatives (Hessian or Jacobian)

nosym= (False) if True, don't explicitly symmetrize the Hessian (good to check errors)

OUTPUT:

```
([R,vR,vT,z,vz,phi], [N,6] array
  d[R,vR,vT,z,vz,phi]/d[J,angle], -> (N,6,6) array
  dO/dJ, -> (3,3) array
  Omegar,Omegaphi,Omegaz, [N] arrays
  Autofit error message)
```

HISTORY:

2016-07-19 - Written - Bovy (UofT)

Specific actionAngle modules

actionAngleTorus

class galpy.actionAngle.actionAngleTorus(*args, **kwargs)

Action-angle formalism using the Torus machinery

__init__(*args, **kwargs)

NAME:

__init__

PURPOSE:

initialize an actionAngleTorus object

INPUT:

pot= potential or list of potentials (3D)

tol= default tolerance to use when fitting tori (**ldJ/J**)

dJ= default action difference when computing derivatives (Hessian or Jacobian)

OUTPUT:

instance

HISTORY:

2015-08-07 - Written - Bovy (UofT)

3.4 DF (galpy.df)

galpy.df contains tools for dealing with distribution functions of stars in galaxies. It mainly contains a number of classes that define different types of distribution function, but galpy.df.jeans also has some tools for solving the Jeans equations for equilibrium systems.

3.4.1 Jeans modeling tools (`galpy.df.jeans`)

`galpy.df.jeans.sigmar`

`galpy.df.jeans.sigmar` (*Pot, r, dens=None, beta=0.0*)

NAME:

`sigmar`

PURPOSE:

Compute the radial velocity dispersion using the spherical Jeans equation

INPUT:

`Pot` - potential or list of potentials (evaluated at $R=r/\sqrt{2}$, $z=r/\sqrt{2}$, sphericity not checked)

`r` - Galactocentric radius (can be Quantity)

`dens=` (None) tracer density profile (function of `r`); if None, the density is assumed to be that corresponding to the potential

`beta=` (0.) anisotropy; can be a constant or a function of `r`

OUTPUT:

`sigma_r(r)`

HISTORY:

2018-07-05 - Written - Bovy (UofT)

`galpy.df.jeans.sigmalos`

`galpy.df.jeans.sigmalos` (*Pot, R, dens=None, surfdens=None, beta=0.0, sigma_r=None*)

NAME:

`sigmalos`

PURPOSE:

Compute the line-of-sight velocity dispersion using the spherical Jeans equation

INPUT:

`Pot` - potential or list of potentials (evaluated at $R=r/\sqrt{2}$, $z=r/\sqrt{2}$, sphericity not checked)

`R` - Galactocentric projected radius (can be Quantity)

`dens=` (None) tracer density profile (function of `r`); if None, the density is assumed to be that corresponding to the potential

`surfdens=` (None) tracer surface density profile (value at `R` or function of `R`); if None, the surface density is assumed to be that corresponding to the density

`beta=` (0.) anisotropy; can be a constant or a function of `r`

`sigma_r=` (None) if given, the solution of the spherical Jeans equation `sigma_r(r)` (used instead of solving the Jeans equation as part of this routine)

OUTPUT:

`sigma_los(R)`

HISTORY:

2018-08-27 - Written - Bovy (UofT)

3.4.2 General instance routines for all df classes

galpy.actionAngle.actionAngle.turn_physical_off

`actionAngle.turn_physical_off()`

NAME:

`turn_physical_off`

PURPOSE:

turn off automatic returning of outputs in physical units

INPUT:

(none)

OUTPUT:

(none)

HISTORY:

2017-06-05 - Written - Bovy (UofT)

galpy.actionAngle.actionAngle.turn_physical_on

`actionAngle.turn_physical_on(ro=None, vo=None)`

NAME:

`turn_physical_on`

PURPOSE:

turn on automatic returning of outputs in physical units

INPUT:

`ro`= reference distance (kpc; can be Quantity)

`vo`= reference velocity (km/s; can be Quantity)

OUTPUT:

(none)

HISTORY:

2016-06-05 - Written - Bovy (UofT)

3.4.3 Two-dimensional, axisymmetric disk distribution functions

Distribution function for orbits in the plane of a galactic disk.

General instance routines

galpy.df.diskdf.__call__

diskdf.__call__ (*args, **kwargs)

NAME:

__call__

PURPOSE:

evaluate the distribution function

INPUT:

either an orbit instance, a list of such instances, or E,Lz

- 1) Orbit instance or list: a) Orbit instance alone: use vxvv member b) Orbit instance + t: call the Orbit instance (for list, each instance is called at t)
- 2) E - energy ($/v_o^2$; or can be Quantity) L - angular momentum ($/r_o/v_o$; or can be Quantity)
- 3) array vxvv [3/4,nt] [must be in natural units $/v_o/r_o$; use Orbit interface for physical-unit input]

KWARGS:

marginalizeVperp - marginalize over perpendicular velocity (only supported with 1a) for single orbits above)

marginalizeVlos - marginalize over line-of-sight velocity (only supported with 1a) for single orbits above)

nsigma= number of sigma to integrate over when marginalizing

+scipy.integrate.quad keywords

OUTPUT:

DF(orbit/E,L)

HISTORY:

2010-07-10 - Written - Bovy (NYU)

galpy.df.diskdf.asymmetricdrift

diskdf.asymmetricdrift (R)

NAME:

asymmetricdrift

PURPOSE:

estimate the asymmetric drift (vc-mean-vphi) from an approximation to the Jeans equation

INPUT:

R - radius at which to calculate the asymmetric drift (can be Quantity)

OUTPUT:

asymmetric drift at R

HISTORY:

2011-04-02 - Written - Bovy (NYU)

galpy.df.diskdf.kurtosisvR

diskdf.**kurtosisvR**(*R, romberg=False, nsigma=None, phi=0.0*)

NAME:

kurtosisvR

PURPOSE:

calculate excess kurtosis in vR at R by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vR \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

kurtosisvR

HISTORY:

2011-12-07 - Written - Bovy (NYU)

galpy.df.diskdf.kurtosisvT

diskdf.**kurtosisvT**(*R, romberg=False, nsigma=None, phi=0.0*)

NAME:

kurtosisvT

PURPOSE:

calculate excess kurtosis in vT at R by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vR \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

kurtosisvT

HISTORY:

2011-12-07 - Written - Bovy (NYU)

galpy.df.diskdf.meanvR

diskdf.**meanvR**(*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

meanvR

PURPOSE:

calculate $\langle vR \rangle$ at *R* by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vR \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

$\langle vR \rangle$ at *R*

HISTORY:

2011-03-30 - Written - Bovy (NYU)

galpy.df.diskdf.meanvT

diskdf.**meanvT**(*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

meanvT

PURPOSE:

calculate $\langle vT \rangle$ at *R* by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vT \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

$\langle vT \rangle$ at *R*

HISTORY:

2011-03-30 - Written - Bovy (NYU)

galpy.df.diskdf.oortA

`diskdf.oortA` (*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

oortA

PURPOSE:

calculate the Oort function A

INPUT:

R - radius at which to calculate A (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

Oort A at R

HISTORY:

2011-04-19 - Written - Bovy (NYU)

BUGS:

could be made more efficient, e.g., surfacemass is calculated multiple times

galpy.df.diskdf.oortB

`diskdf.oortB` (*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

oortB

PURPOSE:

calculate the Oort function B

INPUT:

R - radius at which to calculate B (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

Oort B at R

HISTORY:

2011-04-19 - Written - Bovy (NYU)

BUGS:

could be made more efficient, e.g., surfacemass is calculated multiple times

galpy.df.diskdf.oortC

`diskdf.oortC(R, romberg=False, nsigma=None, phi=0.0)`

NAME:

oortC

PURPOSE:

calculate the Oort function C

INPUT:

R - radius at which to calculate C (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

Oort C at R

HISTORY:

2011-04-19 - Written - Bovy (NYU)

BUGS:

could be made more efficient, e.g., surfacemass is calculated multiple times we know this is zero, but it is calculated anyway (bug or feature?)

galpy.df.diskdf.oortK

`diskdf.oortK(R, romberg=False, nsigma=None, phi=0.0)`

NAME:

oortK

PURPOSE:

calculate the Oort function K

INPUT:

R - radius at which to calculate K (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

Oort K at R

HISTORY:

2011-04-19 - Written - Bovy (NYU)

BUGS:

could be made more efficient, e.g., surfacemass is calculated multiple times we know this is zero, but it is calculated anyway (bug or feature?)

galpy.df.diskdf.sigma2surfacemass

`diskdf.sigma2surfacemass` (*R*, *romberg=False*, *nsigma=None*, *relative=False*)

NAME:

sigma2surfacemass

PURPOSE:

calculate the product σ_R^2 x surface-mass at R by marginalizing over velocity

INPUT:

R - radius at which to calculate the σ_R^2 x surfacemass density (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

σ_R^2 x surface-mass at R

HISTORY:

2010-03-XX - Written - Bovy (NYU)

galpy.df.diskdf.sigma2

`diskdf.sigma2` (*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

sigma2

PURPOSE:

calculate σ_R^2 at R by marginalizing over velocity

INPUT:

R - radius at which to calculate σ_R^2 density (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

sigma_R^2 at R

HISTORY:

2010-03-XX - Written - Bovy (NYU)

galpy.df.diskdf.sigmaR2

diskdf.**sigmaR2** (*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

sigmaR2 (duplicate of sigma2 for consistency)

PURPOSE:

calculate sigma_R^2 at R by marginalizing over velocity

INPUT:

R - radius at which to calculate sigma_R^2 (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

sigma_R^2 at R

HISTORY:

2011-03-30 - Written - Bovy (NYU)

galpy.df.diskdf.sigmaT2

diskdf.**sigmaT2** (*R*, *romberg=False*, *nsigma=None*, *phi=0.0*)

NAME:

sigmaT2

PURPOSE:

calculate sigma_T^2 at R by marginalizing over velocity

INPUT:

R - radius at which to calculate sigma_T^2 (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

sigma_T^2 at R

HISTORY:

2011-03-30 - Written - Bovy (NYU)

galpy.df.diskdf.skewvR

`diskdf.skewvR(R, romberg=False, nsigma=None, phi=0.0)`

NAME:

skewvR

PURPOSE:

calculate skew in vR at R by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vR \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

skewvR

HISTORY:

2011-12-07 - Written - Bovy (NYU)

galpy.df.diskdf.skewvT

`diskdf.skewvT(R, romberg=False, nsigma=None, phi=0.0)`

NAME:

skewvT

PURPOSE:

calculate skew in vT at R by marginalizing over velocity

INPUT:

R - radius at which to calculate $\langle vR \rangle$ (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

skewvT

HISTORY:

2011-12-07 - Written - Bovy (NYU)

galpy.df.diskdf.surfacemass

diskdf.**surfacemass** (*R*, *romberg=False*, *nsigma=None*, *relative=False*)

NAME:

surfacemass

PURPOSE:

calculate the surface-mass at R by marginalizing over velocity

INPUT:

R - radius at which to calculate the surfacemass density (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

romberg - if True, use a romberg integrator (default: False)

OUTPUT:

surface mass at R

HISTORY:

2010-03-XX - Written - Bovy (NYU)

galpy.df.diskdf.surfacemassLOS

diskdf.**surfacemassLOS** (*d*, *l*, *deg=True*, *target=True*, *romberg=False*, *nsigma=None*, *relative=None*)

NAME:

surfacemassLOS

PURPOSE:

evaluate the surface mass along the LOS given l and d

INPUT:

d - distance along the line of sight (can be Quantity)

l - Galactic longitude (in deg, unless deg=False; can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

KEYWORDS:

target= if True, use target surfacemass (default)

romberg - if True, use a romberg integrator (default: False)

deg= if False, l is in radians

OUTPUT:

Sigma(d,l) x d

HISTORY:

2011-03-24 - Written - Bovy (NYU)

galpy.df.diskdf.targetSigma2

`diskdf.targetSigma2 (R, log=False)`

NAME:

targetSigma2

PURPOSE:

evaluate the target $\Sigma_R(R)$

INPUT:

R - radius at which to evaluate (can be Quantity)

OUTPUT:

target $\Sigma_R(R)$

log - if True, return the log (default: False)

HISTORY:

2010-03-28 - Written - Bovy (NYU)

galpy.df.diskdf.targetSurfacemass

`diskdf.targetSurfacemass (R, log=False)`

NAME:

targetSurfacemass

PURPOSE:

evaluate the target surface mass at R

INPUT:

R - radius at which to evaluate (can be Quantity)

log - if True, return the log (default: False)

OUTPUT:

$\Sigma(R)$

HISTORY:

2010-03-28 - Written - Bovy (NYU)

galpy.df.diskdf.targetSurfacemassLOS

`diskdf.targetSurfacemassLOS (d, l, log=False, deg=True)`

NAME:

targetSurfacemassLOS

PURPOSE:

evaluate the target surface mass along the LOS given l and d

INPUT:

d - distance along the line of sight (can be Quantity)

l - Galactic longitude (in deg, unless deg=False; can be Quantity)

deg= if False, l is in radians

log - if True, return the log (default: False)

OUTPUT:

Sigma(d,l) x d

HISTORY:

2011-03-23 - Written - Bovy (NYU)

galpy.df.diskdf._vmomentsurfacemass

`diskdf._vmomentsurfacemass` (*R*, *n*, *m*, *romberg=False*, *nsigma=None*, *relative=False*, *phi=0.0*, *deg=None*, *riv=None*)

Non-physical version of vmomentsurfacemass, otherwise the same

Sampling routines

galpy.df.diskdf.sample

`diskdf.sample` (*n=1*, *rrange=None*, *returnROrbit=True*, *returnOrbit=False*, *nphi=1.0*, *los=None*, *losdeg=True*, *nsigma=None*, *maxd=None*, *target=True*)

NAME:

sample

PURPOSE:

sample n*nphi points from this DF

INPUT:

n - number of desired sample (specifying this rather than calling this routine n times is more efficient)

rrange - if you only want samples in this range, set this keyword (only works when asking for an (RZ)Orbit) (can be Quantity)

returnROrbit - if True, return a planarROrbit instance: [R,vR,vT] (default)

returnOrbit - if True, return a planarOrbit instance (including phi)

nphi - number of azimuths to sample for each E,L

los= line of sight sampling along this line of sight (can be Quantity)

losdeg= los in degrees? (default=True)

target= if True, use target surface mass and sigma2 profiles (default=True)

nsigma= number of sigma to rejection-sample on

maxd= maximum distance to consider (for the rejection sampling)

OUTPUT:

n*nphi list of [[E,Lz],...] or list of planar(R)Orbits

CAUTION: lists of EL need to be post-processed to account for the κ/ω_R discrepancy

HISTORY:

2010-07-10 - Started - Bovy (NYU)

galpy.df.diskdf.sampledSurfacemassLOS

`diskdf.sampledSurfacemassLOS(l, n=1, maxd=None, target=True)`

NAME:

sampledSurfacemassLOS

PURPOSE:

sample a distance along the line of sight

INPUT:

l - Galactic longitude (in rad; can be Quantity)

n = number of distances to sample

$maxd$ = maximum distance to consider (for the rejection sampling) (can be Quantity)

$target$ = if True, sample from the ‘target’ surface mass density, rather than the actual surface mass density (default=True)

OUTPUT:

list of samples

HISTORY:

2011-03-24 - Written - Bovy (NYU)

hhgalpy.df.diskdf.sampleLOS

`diskdf.sampleLOS(los, n=1, deg=True, maxd=None, nsigma=None, targetSurfmass=True, targetSigma2=True)`

NAME:

sampleLOS

PURPOSE:

sample along a given LOS

INPUT:

los - line of sight (in deg, unless $deg=False$; can be Quantity)

n = number of desired samples

deg = los in degrees? (default=True)

targetSurfmass, targetSigma2 = if True, use target surface mass and sigma2 profiles, respectively (there is not much p (default=True)

OUTPUT:

returns list of Orbits

BUGS: target=False uses target distribution for derivatives (this is a detail)

HISTORY:

2011-03-24 - Started - Bovy (NYU)

galpy.df.diskdf.sampleVRVT

diskdf.**sampleVRVT** (*R, n=1, nsigma=None, target=True*)

NAME:

sampleVRVT

PURPOSE:

sample a radial and azimuthal velocity at R

INPUT:

R - Galactocentric distance (can be Quantity)

n= number of distances to sample

nsigma= number of sigma to rejection-sample on

target= if True, sample using the 'target' sigma_R rather than the actual sigma_R (default=True)

OUTPUT:

list of samples

BUGS:

should use the fact that vR and vT separate

HISTORY:

2011-03-24 - Written - Bovy (NYU)

Specific distribution functions

Dehnen DF

```
class galpy.df.dehnendf (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>,
                        profileParams=(0.3333333333333333, 1.0, 0.2), correct=False, beta=0.0,
                        **kwargs)
```

Dehnen's 'new' df

```
__init__ (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>, profileParams=(0.3333333333333333, 1.0, 0.2), correct=False, beta=0.0, **kwargs)
```

NAME: __init__

PURPOSE: Initialize a Dehnen 'new' DF

INPUT:

surfaceSigma - instance or class name of the target surface density and sigma_R profile (default: both exponential)

profileParams - parameters of the surface and sigma_R profile: (xD,xS,Sro) where

xD - disk surface mass scalelength (can be Quantity)

xS - disk velocity dispersion scalelength (can be Quantity)

Sro - disk velocity dispersion at Ro (can be Quantity)

Directly given to the 'surfaceSigmaProfile class, so could be anything that class takes

beta - power-law index of the rotation curve

correct - if True, correct the DF

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

+DFcorrection kwargs (except for those already specified)

OUTPUT:

instance

HISTORY:

2010-03-10 - Written - Bovy (NYU)

Schwarzschild DF

```
class galpy.df.schwarzschilddf (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>,
                                profileParams=(0.3333333333333333, 1.0, 0.2), correct=False,
                                beta=0.0, **kwargs)
```

Schwarzschild's df

```
__init__ (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>, profileParams=(0.3333333333333333, 1.0, 0.2), correct=False, beta=0.0, **kwargs)
```

NAME: `__init__`

PURPOSE: Initialize a Schwarzschild DF

INPUT:

surfaceSigma - instance or class name of the target surface density and sigma_R profile (default: both exponential)

profileParams - parameters of the surface and sigma_R profile: (xD,xS,Sro) where

xD - disk surface mass scalelength (can be Quantity)

xS - disk velocity dispersion scalelength (can be Quantity)

Sro - disk velocity dispersion at Ro (can be Quantity)

Directly given to the 'surfaceSigmaProfile class, so could be anything that class takes

beta - power-law index of the rotation curve

correct - if True, correct the DF

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

+DFcorrection kwargs (except for those already specified)

OUTPUT:

instance

HISTORY:

2017-09-17 - Written - Bovy (UofT)

Shu DF

```
class galpy.df.shudf (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>,
                      profileParams=(0.3333333333333333, 1.0, 0.2), correct=False, beta=0.0,
                      **kwargs)
```

Shu's df (1969)

```
__init__ (surfaceSigma=<class 'galpy.df.surfaceSigmaProfile.expSurfaceSigmaProfile'>, profileParams=(0.3333333333333333, 1.0, 0.2), correct=False, beta=0.0, **kwargs)
```

NAME: `__init__`

PURPOSE: Initialize a Shu DF

INPUT:

surfaceSigma - instance or class name of the target surface density and sigma_R profile (default: both exponential)

profileParams - parameters of the surface and sigma_R profile: (xD,xS,Sro) where

xD - disk surface mass scalelength (can be Quantity)

xS - disk velocity dispersion scalelength (can be Quantity)

Sro - disk velocity dispersion at Ro (can be Quantity)

Directly given to the 'surfaceSigmaProfile class, so could be anything that class takes

beta - power-law index of the rotation curve

correct - if True, correct the DF

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

+DFcorrection kwargs (except for those already specified)

OUTPUT:

instance

HISTORY:

2010-05-09 - Written - Bovy (NYU)

3.4.4 Two-dimensional, non-axisymmetric disk distribution functions

Distribution function for orbits in the plane of a galactic disk in non-axisymmetric potentials. These are calculated using the technique of [Dehnen 2000](#), where the DF at the current time is obtained as the evolution of an initially-axisymmetric DF at time t_0 in the non-axisymmetric potential until the current time.

General instance routines

galpy.df.evolveddiskdf.__call__

evolveddiskdf.__call__(*args, **kwargs)

NAME:

__call__

PURPOSE:

evaluate the distribution function

INPUT:

Orbit instance:

- a) Orbit instance alone: use initial state and t=0
- b) Orbit instance + t: Orbit instance *NOT* called (i.e., Orbit's initial condition is used, call Orbit yourself), t can be Quantity

If t is a list of t, DF is returned for each t, times must be in descending order and equally spaced (does not work with marginalize...)

marginalizeVperp - marginalize over perpendicular velocity (only supported with 1a) above) + nsigma, +scipy.integrate.quad keywords

marginalizeVlos - marginalize over line-of-sight velocity (only supported with 1a) above) + nsigma, +scipy.integrate.quad keywords

log= if True, return the log (not for deriv, bc that can be negative)

integrate_method= method argument of orbit.integrate

deriv= None, 'R', or 'phi': calculates derivative of the moment wrt R or phi **not with the marginalize options**

OUTPUT:

DF(orbit,t)

HISTORY:

2011-03-30 - Written - Bovy (NYU)

2011-04-15 - Added list of times option - Bovy (NYU)

The DF of a two-dimensional, non-axisymmetric disk

class galpy.df.evolveddiskdf(initdf, pot, to=0.0)

Class that represents a diskdf as initial DF + subsequent secular evolution

__init__(initdf, pot, to=0.0)

NAME:

__init__

PURPOSE:

initialize

INPUT:

initdf - the df at the start of the evolution (at to) (units are transferred)

pot - potential to integrate orbits in

to= initial time (time at which initdf is evaluated; orbits are integrated from current t back to to)
(can be Quantity)

OUTPUT:

instance

HISTORY:

2011-03-30 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.meanvR

`evolveddiskdf.meanvR(R, t=0.0, nsigma=None, deg=False, phi=0.0, epsrel=0.01, epsabs=1e-05, grid=None, gridpoints=101, returnGrid=False, surfacemass=None, hierarchygrid=False, nlevels=2, integrate_method='dopr54_c')`

NAME:

meanvR

PURPOSE:

calculate the mean vR of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment(/ro) (can be Quantity)

phi= azimuth (rad unless deg=True; can be Quantity)

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)

surfacemass= if set use this pre-calculated surfacemass

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid object (default=False)

hierarchygrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

mean vR

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.meanvT

`evolveddiskdf.meanvT` (*R*, *t*=0.0, *nsigma*=None, *deg*=False, *phi*=0.0, *epsrel*=0.01, *epsabs*=1e-05, *grid*=None, *gridpoints*=101, *returnGrid*=False, *surfacemass*=None, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

meanvT

PURPOSE:

calculate the mean vT of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment (can be Quantity)

phi= azimuth (rad unless deg=True; can be Quantity)

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)

surfacemass= if set use this pre-calculated surfacemass

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid object (default=False)

hierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

mean vT

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.oortA

`evolveddiskdf.oortA` (*R*, *t*=0.0, *nsigma*=None, *deg*=False, *phi*=0.0, *epsrel*=0.01, *epsabs*=1e-05, *grid*=None, *gridpoints*=101, *returnGrids*=False, *derivRGrid*=None, *derivphiGrid*=None, *derivGridpoints*=101, *derivHierarchgrid*=False, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

oortA

PURPOSE:

calculate the Oort function A at (R,phi,t)

INPUT:

R - radius at which to calculate A (can be Quantity)
phi= azimuth (rad unless deg=True; can be Quantity)
t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)
nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)
deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity
epsrel, epsabs - scipy.integrate keywords
grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid
derivRGrid, derivphiGrid= if set to True, build a grid and use that to evaluate integrals of the derivatives of the DF; if set to a grid-objects (such as returned by this procedure), use this grid
gridpoints= number of points to use for the grid in 1D (default=101)
derivGridpoints= number of points to use for the grid in 1D (default=101)
returnGrid= if True, return the grid objects (default=False)
hierarchgrid= if True, use a hierarchical grid (default=False)
derivHierarchgrid= if True, use a hierarchical grid (default=False)
nlevels= number of hierarchical levels for the hierarchical grid
integrate_method= orbit.integrate method argument

OUTPUT:

Oort A at R,phi,t

HISTORY:

2011-10-16 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.oortB

`evolveddiskdf.oortB`(*R*, *t*=0.0, *nsigma*=None, *deg*=False, *phi*=0.0, *epsrel*=0.01, *epsabs*=1e-05, *grid*=None, *gridpoints*=101, *returnGrids*=False, *derivRGrid*=None, *derivphiGrid*=None, *derivGridpoints*=101, *derivHierarchgrid*=False, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

oortB

PURPOSE:

calculate the Oort function B at (R,phi,t)

INPUT:

R - radius at which to calculate B (can be Quantity)
phi= azimuth (rad unless deg=True; can be Quantity)
t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)
nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

derivRGrid, derivphiGrid= if set to True, build a grid and use that to evaluate integrals of the derivatives of the DF; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

derivGridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid objects (default=False)

hierarchgrid= if True, use a hierarchical grid (default=False)

derivHierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

Oort B at R,phi,t

HISTORY:

2011-10-16 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.oortC

`evolveddiskdf.oortC(R, t=0.0, nsigma=None, deg=False, phi=0.0, epsrel=0.01, epsabs=1e-05, grid=None, gridpoints=101, returnGrids=False, derivRGrid=None, derivphiGrid=None, derivGridpoints=101, derivHierarchgrid=False, hierarchgrid=False, nlevels=2, integrate_method='dopr54_c')`

NAME:

`oortC`

PURPOSE:

calculate the Oort function C at (R,phi,t)

INPUT:

R - radius at which to calculate C (can be Quantity)

phi= azimuth (rad unless deg=True; can be Quantity)

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

derivRGrid, derivphiGrid= if set to True, build a grid and use that to evaluate integrals of the derivatives of the DF; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)
derivGridpoints= number of points to use for the grid in 1D (default=101)
returnGrid= if True, return the grid objects (default=False)
hierarchgrid= if True, use a hierarchical grid (default=False)
derivHierarchgrid= if True, use a hierarchical grid (default=False)
nlevels= number of hierarchical levels for the hierarchical grid
integrate_method= orbit.integrate method argument

OUTPUT:

Oort C at R,phi,t

HISTORY:

2011-10-16 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.oortK

`evolveddiskdf.oortK(R, t=0.0, nsigma=None, deg=False, phi=0.0, epsrel=0.01, epsabs=1e-05, grid=None, gridpoints=101, returnGrids=False, derivRGrid=None, derivphiGrid=None, derivGridpoints=101, derivHierarchgrid=False, hierarchgrid=False, nlevels=2, integrate_method='dopr54_c')`

NAME:

oortK

PURPOSE:

calculate the Oort function K at (R,phi,t)

INPUT:

R - radius at which to calculate K (can be Quantity)
phi= azimuth (rad unless deg=True; can be Quantity)
t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)
nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)
deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity
epsrel, epsabs - scipy.integrate keywords
grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid
derivRGrid, derivphiGrid= if set to True, build a grid and use that to evaluate integrals of the derivatives of the DF; if set to a grid-objects (such as returned by this procedure), use this grid
gridpoints= number of points to use for the grid in 1D (default=101)
derivGridpoints= number of points to use for the grid in 1D (default=101)
returnGrid= if True, return the grid objects (default=False)
hierarchgrid= if True, use a hierarchical grid (default=False)
derivHierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

Oort K at R,phi,t

HISTORY:

2011-10-16 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.sigmaR2

evolveddiskdf.**sigmaR2**(*R*, *t*=0.0, *nsigma*=None, *deg*=False, *phi*=0.0, *epsrel*=0.01, *epsabs*=1e-05, *grid*=None, *gridpoints*=101, *returnGrid*=False, *surface-mass*=None, *meanvR*=None, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

sigmaR2

PURPOSE:

calculate the radial variance of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment (can be Quantity)

phi= azimuth (rad unless deg=True; can be Quantity)

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)

surfacemass, meanvR= if set use this pre-calculated surfacemass and mean vR

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid object (default=False)

hierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

variance of vR

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.sigmaRT

`evolveddiskdf.sigmaRT` (*R*, *t*=0.0, *nsigma*=None, *deg*=False, *epsrel*=0.01, *epsabs*=1e-05, *phi*=0.0, *grid*=None, *gridpoints*=101, *returnGrid*=False, *surfacemass*=None, *meanvR*=None, *meanvT*=None, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

sigmaRT

PURPOSE:

calculate the radial-tangential co-variance of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment (can be Quantity)

phi= azimuth (rad unless deg=True; can be Quantity)

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)

surfacemass, meanvR, meanvT= if set use this pre-calculated surfacemass and mean vR and vT

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid object (default=False)

hierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

covariance of vR and vT

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.sigmaT2

`evolveddiskdf.sigmaT2` (*R*, *t*=0.0, *nsigma*=None, *deg*=False, *phi*=0.0, *epsrel*=0.01, *epsabs*=1e-05, *grid*=None, *gridpoints*=101, *returnGrid*=False, *surface-mass*=None, *meanvT*=None, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

sigmaT2

PURPOSE:

calculate the tangential variance of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment (can be Quantity)
 phi= azimuth (rad unless deg=True; can be Quantity)
 t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)
 surfacemass, meanvT= if set use this pre-calculated surfacemass and mean tangential velocity
 nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)
 deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity
 epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)
 grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid
 gridpoints= number of points to use for the grid in 1D (default=101)
 returnGrid= if True, return the grid object (default=False)
 hierarchgrid= if True, use a hierarchical grid (default=False)
 nlevels= number of hierarchical levels for the hierarchical grid
 integrate_method= orbit.integrate method argument

OUTPUT:

variance of vT

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.vertexdev

`evolveddiskdf.vertexdev` (*R*, *t*=0.0, *nsigma*=None, *deg*=False, *epsrel*=0.01, *epsabs*=1e-05, *phi*=0.0, *grid*=None, *gridpoints*=101, *returnGrid*=False, *sigmaR2*=None, *sigmaT2*=None, *sigmaRT*=None, *surfacemass*=None, *hierarchgrid*=False, *nlevels*=2, *integrate_method*='dopr54_c')

NAME:

vertexdev

PURPOSE:

calculate the vertex deviation of the velocity distribution at (R,phi)

INPUT:

R - radius at which to calculate the moment (can be Quantity)
 phi= azimuth (rad unless deg=True; can be Quantity)
 t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced) (can be Quantity)
 sigmaR2, sigmaT2, sigmaRT= if set the vertex deviation is simply calculated using these
 nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous)

deg= azimuth is in degree (default=False); do not set this when giving phi as a Quantity

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid

gridpoints= number of points to use for the grid in 1D (default=101)

returnGrid= if True, return the grid object (default=False)

hierarchgrid= if True, use a hierarchical grid (default=False)

nlevels= number of hierarchical levels for the hierarchical grid

integrate_method= orbit.integrate method argument

OUTPUT:

vertex deviation in rad

HISTORY:

2011-03-31 - Written - Bovy (NYU)

galpy.df.evolveddiskdf.vmomentsurfacemass

```
evolveddiskdf.vmoment_surfacemass(R, n, m, t=0.0, nsigma=None, deg=False, epsrel=0.01,
                                   epsabs=1e-05, phi=0.0, grid=None, gridpoints=101,
                                   returnGrid=False, hierarchgrid=False, nlevels=2,
                                   print_progress=False, integrate_method='dopr54_c',
                                   deriv=None)
```

NAME:

vmomentsurfacemass

PURPOSE:

calculate the an arbitrary moment of the velocity distribution at (R,phi) times the surfacemass

INPUT:

R - radius at which to calculate the moment (in natural units)

phi= azimuth (rad unless deg=True)

n - vR^n

m - vT^m

t= time at which to evaluate the DF (can be a list or ndarray; if this is the case, list needs to be in descending order and equally spaced)

nsigma - number of sigma to integrate the velocities over (based on an estimate, so be generous, but not too generous)

deg= azimuth is in degree (default=False)

epsrel, epsabs - scipy.integrate keywords (the integration calculates the ratio of this vmoment to that of the initial DF)

grid= if set to True, build a grid and use that to evaluate integrals; if set to a grid-objects (such as returned by this procedure), use this grid; if this was created for a list of times, moments are calculated for each time

gridpoints= number of points to use for the grid in 1D (default=101)
 returnGrid= if True, return the grid object (default=False)
 hierarchygrid= if True, use a hierarchical grid (default=False)
 nlevels= number of hierarchical levels for the hierarchical grid
 print_progress= if True, print progress updates
 integrate_method= orbit.integrate method argument
 deriv= None, 'R', or 'phi': calculates derivative of the moment wrt R or phi **onny with grid options**

OUTPUT:

$\langle vR^n vT^m \times \text{surface-mass} \rangle$ at R,phi (no support for units)

COMMENT:

grid-based calculation is the only one that is heavily tested (although the test suite also tests the direct calculation)

HISTORY:

2011-03-30 - Written - Bovy (NYU)

3.4.5 Three-dimensional disk distribution functions

Distribution functions for orbits in galactic disks, including the vertical motion for stars reaching large heights above the plane. Currently only the *quasi-isothermal DF*.

General instance routines

galpy.df.quasiisothermaldf.__call__

quasiisothermaldf.__call__ (*args, **kwargs)

NAME:

__call__

PURPOSE:

return the DF

INPUT:

Either:

a)(jr,lz,jz) tuple; each can be a Quantity

where: jr - radial action lz - z-component of angular momentum jz - vertical action

b) R,vR,vT,z,vz

c) Orbit instance: initial condition used if that's it, orbit(t) if there is a time given as well

log= if True, return the natural log

+scipy.integrate.quadrature kwargs

func= function of (jr,lz,jz) to multiply f with (useful for moments)

OUTPUT:

value of DF

HISTORY:

2012-07-25 - Written - Bovy (IAS@MPIA)

NOTE:

For Miyamoto-Nagai/adiabatic approximation this seems to take about 30 ms / evaluation in the extended Solar neighborhood For a MWPotential/adiabatic approximation this takes about 50 ms / evaluation in the extended Solar neighborhood

For adiabatic-approximation grid this seems to take about 0.67 to 0.75 ms / evaluation in the extended Solar neighborhood (includes some out of the grid)

up to 200x faster when called with vector R,vR,vT,z,vz

galpy.df.quasiisothermaldf.density

`quasiisothermaldf.density` (*R*, *z*, *nsigma=None*, *mc=False*, *nmc=10000*, *gl=True*, *ngl=10*,
***kwargs*)

NAME:

density

PURPOSE:

calculate the density at R,z by marginalizing over velocity

INPUT:

R - radius at which to calculate the density (can be Quantity)

z - height at which to calculate the density (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

density at (R,z)

HISTORY:

2012-07-26 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.estimate_hr

`quasiisothermaldf.estimate_hr` (*R*, *z=0.0*, *dR=1e-08*, ***kwargs*)

NAME:

estimate_hr

PURPOSE:

estimate the exponential scale length at R

INPUT:

R - Galactocentric radius (can be Quantity)

z= height (default: 0 pc) (can be Quantity)

dR- range in R to use (can be Quantity)

density kwargs

OUTPUT:

estimated hR

HISTORY:

2012-09-11 - Written - Bovy (IAS)

2013-01-28 - Re-written - Bovy

galpy.df.quasiisothermaldf.estimate_hsr

`quasiisothermaldf.estimate_hsr(R, z=0.0, dR=1e-08, **kwargs)`

NAME:

estimate_hsr

PURPOSE:

estimate the exponential scale length of the radial dispersion at R

INPUT:

R - Galactocentric radius (can be Quantity)

z= height (default: 0 pc) (can be Quantity)

dR- range in R to use (can be Quantity)

density kwargs

OUTPUT:

estimated hsr

HISTORY:

2013-03-08 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.estimate_hsz

`quasiisothermaldf.estimate_hsz(R, z=0.0, dR=1e-08, **kwargs)`

NAME:

estimate_hsz

PURPOSE:

estimate the exponential scale length of the vertical dispersion at R

INPUT:

R - Galactocentric radius (can be Quantity)

z= height (default: 0 pc) (can be Quantity)

dR- range in R to use (can be Quantity)

density kwargs

OUTPUT:

estimated hsz

HISTORY:

2013-03-08 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.estimate_hz

quasiisothermaldf.**estimate_hz**(R, z, dz=1e-08, **kwargs)

NAME:

estimate_hz

PURPOSE:

estimate the exponential scale height at R

INPUT:

R - Galactocentric radius (can be Quantity)

dz - z range to use (can be Quantity)

density kwargs

OUTPUT:

estimated hz

HISTORY:

2012-08-30 - Written - Bovy (IAS)

2013-01-28 - Re-written - Bovy

galpy.df.quasiisothermaldf._jmomentdensity

quasiisothermaldf.**_jmomentdensity**(R, z, n, m, o, nsigma=None, mc=True, nmc=10000, _returnnmc=False, _vrs=None, _vts=None, _vzs=None, **kwargs)

Non-physical version of jmomentdensity, otherwise the same

galpy.df.quasiisothermaldf.meanjr

quasiisothermaldf.**meanjr**(R, z, nsigma=None, mc=True, nmc=10000, **kwargs)

NAME:

meanjr

PURPOSE:

calculate the mean radial action by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

OUTPUT:

meanjr

HISTORY:

2012-08-09 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.meanjz

quasiisothermaldf.**meanjz** (*R*, *z*, *nsigma=None*, *mc=True*, *nmc=10000*, ***kwargs*)

NAME:

meanjz

PURPOSE:

calculate the mean vertical action by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

OUTPUT:

meanjz

HISTORY:

2012-08-09 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.meanlz

quasiisothermaldf.**meanlz** (*R*, *z*, *nsigma=None*, *mc=True*, *nmc=10000*, ***kwargs*)

NAME:

meanlz

PURPOSE:

calculate the mean angular momentum by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

OUTPUT:

meanlz

HISTORY:

2012-08-09 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.meanvR

quasiisothermaldf.**meanvR**(R, z, nsigma=None, mc=False, nmc=10000, gl=True, ngl=10, **kwargs)

NAME:

meanvR

PURPOSE:

calculate the mean radial velocity by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

meanvR

HISTORY:

2012-12-23 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.meanvT

`quasiisothermaldf.meanvT` (*R*, *z*, *nsigma=None*, *mc=False*, *nmc=10000*, *gl=True*, *ngl=10*, ***kwargs*)

NAME:

meanvT

PURPOSE:

calculate the mean rotational velocity by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

`scipy.integrate.tplquad` *kwargs* *epsabs* and *epsrel*

mc= if True, calculate using Monte Carlo integration

nmc= if *mc*, use *nmc* samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if *gl*, use *ngl*-th order Gauss-Legendre integration for each dimension

OUTPUT:

meanvT

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.meanvz

`quasiisothermaldf.meanvz` (*R*, *z*, *nsigma=None*, *mc=False*, *nmc=10000*, *gl=True*, *ngl=10*, ***kwargs*)

NAME:

meanvz

PURPOSE:

calculate the mean vertical velocity by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

`scipy.integrate.tplquad` *kwargs* *epsabs* and *epsrel*

mc= if True, calculate using Monte Carlo integration

nmc= if *mc*, use *nmc* samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

meanvz

HISTORY:

2012-12-23 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.pvR

quasiisothermaldf.**pvr**(vR, R, z, gl=True, ngl=20, nsigma=4.0, vTmax=1.5)

NAME:

pvr

PURPOSE:

calculate the marginalized vR probability at this location (NOT normalized by the density)

INPUT:

vR - radial velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

nsigma - sets integration limits to $[-1,+1]*nsigma*\sigma_z(R)$ for integration over vz (default: 4)

vTmax - sets integration limits to $[0,vTmax]$ for integration over vT (default: 1.5)

OUTPUT:

p(vR,R,z)

HISTORY:

2012-12-22 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.pvRvT

quasiisothermaldf.(vR, vT, R, z, gl=True, ngl=20, nsigma=4.0)

NAME:

pvRvT

PURPOSE:

calculate the marginalized (vR,vT) probability at this location (NOT normalized by the density)

INPUT:

vR - radial velocity (can be Quantity)

vT - tangential velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

nsigma - sets integration limits to $[-1, +1] * \text{nsigma} * \text{sigma}_z(R)$ for integration over v_z (default: 4)

OUTPUT:

$p(v_R, v_T, R, z)$

HISTORY:

2013-01-02 - Written - Bovy (IAS) 2018-01-12 - Added Gauss-Legendre integration prefactor

nsigma/2 - Trick (MPA)

galpy.df.quasiisothermaldf.pvRvz

quasiisothermaldf.**pvRvz** ($v_R, v_z, R, z, gl=True, ngl=20, vTmax=1.5$)

NAME:

pvR

PURPOSE:

calculate the marginalized (v_R, v_z) probability at this location (NOT normalized by the density)

INPUT:

v_R - radial velocity (can be Quantity)

v_z - vertical velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

$vTmax$ - sets integration limits to $[0, vTmax]$ for integration over v_T (default: 1.5)

OUTPUT:

$p(v_R, v_z, R, z)$

HISTORY:

2013-01-02 - Written - Bovy (IAS) 2018-01-12 - Added Gauss-Legendre integration prefactor $vTmax/2$ - Trick (MPA)

galpy.df.quasiisothermaldf.pvT

quasiisothermaldf.**pvT** ($v_T, R, z, gl=True, ngl=20, nsigma=4.0$)

NAME:

pvT

PURPOSE:

calculate the marginalized v_T probability at this location (NOT normalized by the density)

INPUT:

vT - tangential velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

nsigma - sets integration limits to $[-1, +1] * \text{nsigma} * \text{sigma}(R)$ for integration over v_z and v_R (default: 4)

OUTPUT:

p(vT,R,z)

HISTORY:

2012-12-22 - Written - Bovy (IAS) 2018-01-12 - Added Gauss-Legendre integration prefactor
nsigma^{2/4} - Trick (MPA)

galpy.df.quasiisothermaldf.pvTvz

quasiisothermaldf.pvTvz (vT, vz, R, z, gl=True, ngl=20, nsigma=4.0)

NAME:

pvTvz

PURPOSE:

calculate the marginalized (vT,vz) probability at this location (NOT normalized by the density)

INPUT:

vT - tangential velocity (can be Quantity)

vz - vertical velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

nsigma - sets integration limits to $[-1, +1] * \text{nsigma} * \text{sigma}_R(R)$ for integration over v_R (default: 4)

OUTPUT:

p(vT,vz,R,z)

HISTORY:

2012-12-22 - Written - Bovy (IAS) 2018-01-12 - Added Gauss-Legendre integration prefactor
nsigma/2 - Trick (MPA)

galpy.df.quasiisothermaldf.pvz

quasiisothermaldf.pvz (vz, R, z, gl=True, ngl=20, nsigma=4.0, vTmax=1.5, _return_actions=False, _jr=None, _lz=None, _jz=None, _return_freqs=False, _rg=None, _kappa=None, _nu=None, _Omega=None, _sigmaR1=None)

NAME:

pvz

PURPOSE:

calculate the marginalized vz probability at this location (NOT normalized by the density)

INPUT:

vz - vertical velocity (can be Quantity)

R - radius (can be Quantity)

z - height (can be Quantity)

gl - use Gauss-Legendre integration (True, currently the only option)

ngl - order of Gauss-Legendre integration

nsigma - sets integration limits to $[-1, +1] * \text{nsigma} * \text{sigma_R}(R)$ for integration over vR (default: 4)

vTmax - sets integration limits to $[0, \text{vTmax}]$ for integration over vT (default: 1.5)

OUTPUT:

$p(\text{vz}, R, z)$

HISTORY:

2012-12-22 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.sampleV

`quasiisothermaldf.sampleV(R, z, n=1, **kwargs)`

NAME:

sampleV

PURPOSE:

sample a radial, azimuthal, and vertical velocity at R,z

INPUT:

R - Galactocentric distance (can be Quantity)

z - height (can be Quantity)

n= number of distances to sample

OUTPUT:

list of samples

HISTORY:

2012-12-17 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.sigmaR2

`quasiisothermaldf.sigmaR2(R, z, nsigma=None, mc=False, nmc=10000, gl=True, ngl=10, **kwargs)`

NAME:

sigmaR2

PURPOSE:

calculate σ_R^2 by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

σ_R^2

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.sigmaRz

quasiisothermaldf.**sigmaRz**(R, z, nsigma=None, mc=False, nmc=10000, gl=True, ngl=10,
**kwargs)

NAME:

sigmaRz

PURPOSE:

calculate σ_{RZ}^2 by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

scipy.integrate.tplquad kwargs epsabs and epsrel

mc= if True, calculate using Monte Carlo integration

nmc= if mc, use nmc samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

σ_{Rz}^2

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.sigmaT2

`quasiisothermaldf.sigmaT2(R, z, nsigma=None, mc=False, nmc=10000, gl=True, ngl=10,
 **kwargs)`

NAME:

sigmaT2

PURPOSE:

calculate σ_T^2 by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

`scipy.integrate.tplquad` *kwargs* *epsabs* and *epsrel*

mc= if True, calculate using Monte Carlo integration

nmc= if *mc*, use *nmc* samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if *gl*, use *ngl*-th order Gauss-Legendre integration for each dimension

OUTPUT:

sigma_T^2

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.sigmaz2

`quasiisothermaldf.sigmaz2(R, z, nsigma=None, mc=False, nmc=10000, gl=True, ngl=10,
 **kwargs)`

NAME:

sigmaz2

PURPOSE:

calculate σ_z^2 by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over

`scipy.integrate.tplquad` *kwargs* *epsabs* and *epsrel*

mc= if True, calculate using Monte Carlo integration

nmc= if *mc*, use *nmc* samples

gl= if True, calculate using Gauss-Legendre integration

ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

σ_z^2

HISTORY:

2012-07-30 - Written - Bovy (IAS@MPIA)

galpy.df.quasiisothermaldf.surfacemass_z

quasiisothermaldf.**surfacemass_z**(*R*, *nz*=7, *zmax*=1.0, *fixed_quad*=True, *fixed_order*=8,
***kwargs*)

NAME:

surfacemass_z

PURPOSE:

calculate the vertically-integrated surface density

INPUT:

R - Galactocentric radius (can be Quantity)

fixed_quad= if True (default), use Gauss-Legendre integration

fixed_order= (20), order of GL integration to use

nz= number of *zs* to use to estimate

zmax= maximum *z* to use (can be Quantity)

density *kwargs*

OUTPUT:

$\Sigma(R)$

HISTORY:

2012-08-30 - Written - Bovy (IAS)

galpy.df.quasiisothermaldf.tilt

quasiisothermaldf.**tilt**(*R*, *z*, *nsigma*=None, *mc*=False, *nmc*=10000, *gl*=True, *ngl*=10, ***kwargs*)

NAME:

tilt

PURPOSE:

calculate the tilt of the velocity ellipsoid by marginalizing over velocity

INPUT:

R - radius at which to calculate this (can be Quantity)

z - height at which to calculate this (can be Quantity)

OPTIONAL INPUT:

nsigma - number of sigma to integrate the velocities over
 scipy.integrate.tplquad kwargs epsabs and epsrel
 mc= if True, calculate using Monte Carlo integration
 nmc= if mc, use nmc samples
 gl= if True, calculate using Gauss-Legendre integration
 ngl= if gl, use ngl-th order Gauss-Legendre integration for each dimension

OUTPUT:

tilt in rad

HISTORY:

2012-12-23 - Written - Bovy (IAS)

2017-10-28 - Changed return unit to rad - Bovy (UofT)

galpy.df.quasiisothermaldf._vmomentdensity

```
quasiisothermaldf._vmomentdensity(R, z, n, m, o, nsigma=None, mc=False, nmc=10000,
                                   _returnnmc=False, _vrs=None, _vts=None, _vzs=None,
                                   _rawgausssamples=False, gl=False, ngl=10, _returngl=False,
                                   _glqeval=None, _return_actions=False, _jr=None,
                                   _lz=None, _jz=None, _return_freqs=False, _rg=None,
                                   _kappa=None, _nu=None, _Omega=None, _sigmaRl=None,
                                   _sigmazl=None, **kwargs)
```

Non-physical version of vmomentdensity, otherwise the same

Specific distribution functions

Quasi-isothermal DF

```
class galpy.df.quasiisothermaldf(hr, sr, sz, hsr, hsz, pot=None, aA=None, cutcounter=False,
                                  _precomputerg=True, _precomputergrmax=None, _precomputergnLz=51,
                                  refr=1.0, lo=0.005681818181818182, ro=None, vo=None)
```

Class that represents a ‘Binney’ quasi-isothermal DF

```
__init__(hr, sr, sz, hsr, hsz, pot=None, aA=None, cutcounter=False, _precomputerg=True,
          _precomputergrmax=None, _precomputergnLz=51, refr=1.0, lo=0.005681818181818182, ro=None, vo=None)
```

NAME:

```
__init__
```

PURPOSE:

Initialize a quasi-isothermal DF

INPUT:

hr - radial scale length (can be Quantity)

sr - radial velocity dispersion at the solar radius (can be Quantity)

sz - vertical velocity dispersion at the solar radius (can be Quantity)

hsr - radial-velocity-dispersion scale length (can be Quantity)

hsz - vertical-velocity-dispersion scale length (can be Quantity)

pot= Potential instance or list thereof

aA= actionAngle instance used to convert (x,v) to actions [must be an instance of an actionAngle class that computes (J, Omega, angle) for a given (x,v)]

cutcounter= if True, set counter-rotating stars' DF to zero

refr= reference radius for dispersions (can be different from ro) (can be Quantity)

lo= reference angular momentum below where there are significant numbers of retrograde stars (can be Quantity)

ro= distance from vantage point to GC (kpc; can be Quantity)

vo= circular velocity at ro (km/s; can be Quantity)

OTHER INPUTS:

_precomputerg= if True (default), pre-compute the rL(L)

_precomputergRmax= if set, this is the maximum R for which to pre-compute rg (default: 5*hr)

_precomputergnLz if set, number of Lz to pre-compute rg for (default: 51)

OUTPUT:

object

HISTORY:

2012-07-25 - Started - Bovy ([IAS@MPIA](#))

3.4.6 The distribution function of a tidal stream

From Bovy 2014; see *Dynamical modeling of tidal streams*.

General instance routines

galpy.df.streamdf.__call__

streamdf.__call__(*args, **kwargs)

NAME:

__call__

PURPOSE:

evaluate the DF

INPUT:

Either:

a) R,vR,vT,z,vz,phi ndarray [nobjects]

b) (Omegar,Omegaphi,Omegaz,angler,anglephi,anglez) tuple if aAInput

where:

Omegar - radial frequency
 Omegaphi - azimuthal frequency
 Omegaz - vertical frequency
 angler - radial angle
 anglephi - azimuthal angle
 anglez - vertical angle

c) Orbit instance or list thereof

log= if True, return the natural log

aaInput= (False) if True, option b above

OUTPUT:

value of DF

HISTORY:

2013-12-03 - Written - Bovy (IAS)

The stream DF

```
class galpy.df.streamdf(sigv, progenitor=None, pot=None, aA=None, useTM=False, tdis-
                        rupt=None, sigMeanOffset=6.0, leading=True, sigangle=None, deltaAn-
                        gleTrack=None, nTrackChunks=None, nTrackIterations=None, progIs-
                        Track=False, ro=None, vo=None, Vnorm=None, Rnorm=None, R0=8.0,
                        Zsun=0.025, vsun=[-11.1, 241.92, 7.25], multi=None, interpTrack=True,
                        useInterp=True, nosetup=False, nospreadsetup=False, approxConstTrack-
                        Freq=False, useTMHessian=False, custom_transform=None)
```

The DF of a tidal stream

```
__init__(sigv, progenitor=None, pot=None, aA=None, useTM=False, tdisrupt=None, sigMeanOff-
          set=6.0, leading=True, sigangle=None, deltaAngleTrack=None, nTrackChunks=None,
          nTrackIterations=None, progIsTrack=False, ro=None, vo=None, Vnorm=None,
          Rnorm=None, R0=8.0, Zsun=0.025, vsun=[-11.1, 241.92, 7.25], multi=None, interp-
          Track=True, useInterp=True, nosetup=False, nospreadsetup=False, approxConstTrack-
          Freq=False, useTMHessian=False, custom_transform=None)
```

NAME:

`__init__`

PURPOSE:

Initialize a quasi-isothermal DF

INPUT:

sigv - radial velocity dispersion of the progenitor (can be Quantity)

tdisrupt= (5 Gyr) time since start of disruption (can be Quantity)

leading= (True) if True, model the leading part of the stream if False, model the trailing part

progenitor= progenitor orbit as Orbit instance (will be re-integrated, so don't bother integrating the orbit before)

progIsTrack= (False) if True, then the progenitor (x,v) is actually the (x,v) of the stream track at zero angle separation; useful when initializing with an orbit fit; the progenitor's position will be calculated

pot= Potential instance or list thereof

aA= actionAngle instance used to convert (x,v) to actions

useTM= (False) if set to an actionAngleTorus instance, use this to speed up calculations

sigMeanOffset= (6.) offset between the mean of the frequencies and the progenitor, in units of the largest eigenvalue of the frequency covariance matrix (along the largest eigenvector), should be positive; to model the trailing part, set leading=False

sigangle= (sigv/122/[1km/s]=1.8sigv in natural coordinates) estimate of the angle spread of the debris initially (can be Quantity)

deltaAngleTrack= (None) angle to estimate the stream track over (rad; or can be Quantity)

nTrackChunks= (floor(deltaAngleTrack/0.15)+1) number of chunks to divide the progenitor track in

nTrackIterations= Number of iterations to perform when establishing the track; each iteration starts from a previous approximation to the track in (x,v) and calculates a new track based on the deviation between the previous track and the desired track in action-angle coordinates; if not set, an appropriate value is determined based on the magnitude of the misalignment between stream and orbit, with larger numbers of iterations for larger misalignments

interpTrack= (might change), interpolate the stream track while setting up the instance (can be done by hand by calling self._interpolate_stream_track() and self._interpolate_stream_track_aA())

useInterp= (might change), use interpolation by default when calculating approximated frequencies and angles

nosetup= (False) if True, don't setup the stream track and anything else that is expensive

nospreadsetup= (False) if True, don't setup the spread around the stream track (only for nosetup is False)

multi= (None) if set, use multi-processing

Coordinate transformation inputs:

vo= (220) circular velocity to normalize velocities with [used to be Vnorm; can be Quantity]

ro= (8) Galactocentric radius to normalize positions with [used to be Rnorm; can be Quantity]

R0= (8) Galactocentric radius of the Sun (kpc) [can be different from ro; can be Quantity]

Zsun= (0.025) Sun's height above the plane (kpc; can be Quantity)

vsun= ([-11.1,241.92,7.25]) Sun's motion in cylindrical coordinates (vR positive away from center) (can be Quantity)

custom_transform= (None) matrix implementing the rotation from (ra,dec) to a custom set of sky coordinates

approxConstTrackFreq= (False) if True, approximate the stream assuming that the frequency is constant along the stream (only works with useTM, for which this leads to a significant speed-up)

useTMHessian= (False) if True, compute the basic Hessian dO/dJ_{prog} using TM; otherwise use aA

OUTPUT:

object

HISTORY:

2013-09-16 - Started - Bovy (IAS)

2013-11-25 - Started over - Bovy (IAS)

galpy.df.streamdf.calc_stream_lb

streamdf.**calc_stream_lb** (*vo=None, ro=None, R0=None, Zsun=None, vsun=None*)

NAME:

calc_stream_lb

PURPOSE:

convert the stream track to observational coordinates and store

INPUT:

Coordinate transformation inputs (all default to the instance-wide values):

vo= circular velocity to normalize velocities with

ro= Galactocentric radius to normalize positions with

R0= Galactocentric radius of the Sun (kpc)

Zsun= Sun's height above the plane (kpc)

vsun= Sun's motion in cylindrical coordinates (vR positive away from center)

OUTPUT:

(none)

HISTORY:

2013-12-02 - Written - Bovy (IAS)

galpy.df.streamdf.callMarg

streamdf.**callMarg** (*xy, **kwargs*)

NAME:

callMarg

PURPOSE: evaluate the DF, marginalizing over some directions, in Galactocentric rectangular coordinates (or in observed $l, b, D, v_{\text{los}}, p_{\text{ml}}, p_{\text{mbb}}$ coordinates)

INPUT:

xy - phase-space point $[X, Y, Z, v_X, v_Y, v_Z]$; the distribution of the dimensions set to None is returned

interp= (object-wide interp default) if True, use the interpolated stream track

cindx= index of the closest point on the (interpolated) stream track if not given, determined from the dimensions given

nsigma= (3) number of sigma to marginalize the DF over (approximate sigma)

ngl= (5) order of Gauss-Legendre integration

lb= (False) if True, xy contains [l,b,D,vlos,pmll,pmbb] in [deg,deg,kpc,km/s,mas/yr,mas/yr] and the marginalized PDF in these coordinates is returned

vo= (220) circular velocity to normalize with when lb=True

ro= (8) Galactocentric radius to normalize with when lb=True

R0= (8) Galactocentric radius of the Sun (kpc)

Zsun= (0.025) Sun's height above the plane (kpc)

vsun= ([-11.1,241.92,7.25]) Sun's motion in cylindrical coordinates (vR positive away from center)

OUTPUT:

p(xy) marginalized over missing directions in xy

HISTORY:

2013-12-16 - Written - Bovy (IAS)

galpy.df.streamdf.density_par

`streamdf.density_par(dangle, coord='apar', tdisrupt=None, **kwargs)`

NAME:

density_par

PURPOSE:

calculate the density as a function of a parallel coordinate

INPUT:

dangle - parallel angle offset for this coordinate value

coord - coordinate to return the density in ('apar' [default], 'll', 'ra', 'customra', 'phi')

OUTPUT:

density(angle)

HISTORY:

2015-11-17 - Written - Bovy (UofT)

galpy.df.streamdf.estimateTdisrupt

`streamdf.estimateTdisrupt(deltaAngle)`

NAME:

estimateTdisrupt

PURPOSE:

estimate the time of disruption

INPUT:

deltaAngle- spread in angle since disruption

OUTPUT:

time in natural units

HISTORY:

2013-11-27 - Written - Bovy (IAS)

galpy.df.streamdf.find_closest_trackpoint

`streamdf.find_closest_trackpoint (R, vR, vT, z, vz, phi, interp=True, xy=False, usev=False)`

NAME:

find_closest_trackpoint

PURPOSE:

find the closest point on the stream track to a given point

INPUT:

R,vR,vT,z,vz,phi - phase-space coordinates of the given point

interp= (True), if True, return the index of the interpolated track

xy= (False) if True, input is X,Y,Z,vX,vY,vZ in Galactocentric rectangular coordinates; if xy, some coordinates may be missing (given as None) and they will not be used

usev= (False) if True, also use velocities to find the closest point

OUTPUT:

index into the track of the closest track point

HISTORY:

2013-12-04 - Written - Bovy (IAS)

galpy.df.streamdf.find_closest_trackpointLB

`streamdf.find_closest_trackpointLB (l, b, D, vlos, pmll, pmllb, interp=True, usev=False)`

NAME:

find_closest_trackpointLB

PURPOSE: find the closest point on the stream track to a given point in (l,b,...) coordinates

INPUT:

l,b,D,vlos,pmll,pmllb- coordinates in (deg,deg,kpc,km/s,mas/yr,mas/yr)

interp= (True) if True, return the closest index on the interpolated track

usev= (False) if True, also use the velocity components (default is to only use the positions)

OUTPUT:

index of closest track point on the interpolated or not-interpolated track

HISTORY:

2013-12-17- Written - Bovy (IAS)

galpy.df.streamdf.freqEigvalRatio

streamdf.**freqEigvalRatio** (*isotropic=False*)

NAME:

freqEigvalRatio

PURPOSE:

calculate the ratio between the largest and 2nd-to-largest (in abs) eigenvalue of $\sqrt{dO/dJ^T V_J dO/dJ}$ (if this is big, a 1D stream will form)

INPUT:

isotropic= (False), if True, return the ratio assuming an isotropic action distribution (i.e., just of dO/dJ)

OUTPUT:

ratio between eigenvalues of $\text{fabs}(dO / dJ)$

HISTORY:

2013-12-05 - Written - Bovy (IAS)

galpy.df.streamdf.gaussApprox

streamdf.**gaussApprox** (*xy, **kwargs*)

NAME:

gaussApprox

PURPOSE:

return the mean and variance of a Gaussian approximation to the stream DF at a given phase-space point in Galactocentric rectangular coordinates (distribution is over missing directions)

INPUT:

xy - phase-space point [X,Y,Z,vX,vY,vZ]; the distribution of the dimensions set to None is returned

interp= (object-wide interp default) if True, use the interpolated stream track

cindx= index of the closest point on the (interpolated) stream track if not given, determined from the dimensions given

lb= (False) if True, xy contains [l,b,D,vlos,pmll,pmbb] in [deg,deg,kpc,km/s,mas/yr,mas/yr] and the Gaussian approximation in these coordinates is returned

OUTPUT:

(mean,variance) of the approximate Gaussian DF for the missing directions in xy

HISTORY:

2013-12-12 - Written - Bovy (IAS)

galpy.df.streamdf.length

streamdf.**length** (*threshold=0.2, phys=False, ang=False, tdisrupt=None, **kwargs*)

NAME:

length

PURPOSE:

calculate the length of the stream

INPUT:

threshold - threshold down from the density near the progenitor at which to define the 'end' of the stream

phys= (False) if True, return the length in physical kpc

ang= (False) if True, return the length in sky angular arc length in degree

coord - coordinate to return the density in ('apar' [default], 'll', 'ra', 'customra', 'phi')

OUTPUT:

length (rad for parallel angle; kpc for physical length; deg for sky arc length)

HISTORY:

2015-12-22 - Written - Bovy (UofT)

galpy.df.streamdf.meanangledAngle

`streamdf.meanangledAngle` (*dangle*, *smallest=False*)

NAME:

meanangledAngle

PURPOSE:

calculate the mean perpendicular angle at a given angle

INPUT:

dangle - angle offset along the stream

smallest= (False) calculate for smallest eigenvalue direction rather than for middle

OUTPUT:

mean perpendicular angle

HISTORY:

2013-12-06 - Written - Bovy (IAS)

galpy.df.streamdf.meanOmega

`streamdf.meanOmega` (*dangle*, *oned=False*, *offset_sign=None*, *tdisrupt=None*)

NAME:

meanOmega

PURPOSE:

calculate the mean frequency as a function of angle, assuming a uniform time distribution up to a maximum time

INPUT:

dangle - angle offset

oned= (False) if True, return the 1D offset from the progenitor (along the direction of disruption)

offset_sign= sign of the frequency offset (shouldn't be set)

OUTPUT:

mean Omega

HISTORY:

2013-12-01 - Written - Bovy (IAS)

galpy.df.streamdf.meantdAngle

`streamdf.meantdAngle` (*dangle*)

NAME:

meantdAngle

PURPOSE:

calculate the mean stripping time at a given angle

INPUT:

dangle - angle offset along the stream

OUTPUT:

mean stripping time at this dangle

HISTORY:

2013-12-05 - Written - Bovy (IAS)

galpy.df.streamdf.misalignment

`streamdf.misalignment` (*isotropic=False, **kwargs*)

NAME:

misalignment

PURPOSE:

calculate the misalignment between the progenitor's frequency and the direction along which the stream disrupts

INPUT:

isotropic= (False), if True, return the misalignment assuming an isotropic action distribution

OUTPUT:

misalignment in rad

HISTORY:

2013-12-05 - Written - Bovy (IAS)

2017-10-28 - Changed output unit to rad - Bovy (UofT)

galpy.df.streamdf.pangledAngle

`streamdf.pangledAngle` (*angleperp*, *dangle*, *smallest=False*)

NAME:

`pangledAngle`

PURPOSE: return the probability of a given perpendicular angle at a given angle along the stream

INPUT:

angleperp - perpendicular angle

dangle - angle offset along the stream

smallest= (False) calculate for smallest eigenvalue direction rather than for middle

OUTPUT:

`p(angle_perp|dangle)`

HISTORY:

2013-12-06 - Written - Bovy (IAS)

galpy.df.streamdf.plotCompareTrackAAModel

`streamdf.plotCompareTrackAAModel` (***kwargs*)

NAME:

`plotCompareTrackAAModel`

PURPOSE:

plot the comparison between the underlying model's $d\Omega_{\text{perp}}$ vs. $d\text{angle}_r$ (line) and the track in (x,v) 's $d\Omega_{\text{perp}}$ vs. $d\text{angle}_r$ (dots; explicitly calculating the track's action-angle coordinates)

INPUT:

`bovy_plot.bovy_plot` kwargs

OUTPUT:

`plot`

HISTORY:

2014-08-27 - Written - Bovy (IAS)

galpy.df.streamdf.plotProgenitor

`streamdf.plotProgenitor` (*d1='x'*, *d2='z'*, **args*, ***kwargs*)

NAME:

`plotProgenitor`

PURPOSE:

plot the progenitor orbit

INPUT:

d1= plot this on the X axis ('x','y','z','R','phi','vx','vy','vz','vR','vt','ll','bb','dist','pml','pmbb','vlos')

d2= plot this on the Y axis (same list as for d1)

scaleToPhysical= (False), if True, plot positions in kpc and velocities in km/s

bovy_plot.bovy_plot args and kwargs

OUTPUT:

plot to output device

HISTORY:

2013-12-09 - Written - Bovy (IAS)

galpy.df.streamdf.plotTrack

streamdf.**plotTrack** (d1='x', d2='z', interp=True, spread=0, simple=True, *args, **kwargs)

NAME:

plotTrack

PURPOSE:

plot the stream track

INPUT:

d1= plot this on the X axis ('x','y','z','R','phi','vx','vy','vz','vR','vt','ll','bb','dist','pml','pmbb','vlos')

d2= plot this on the Y axis (same list as for d1)

interp= (True) if True, use the interpolated stream track

spread= (0) if int > 0, also plot the spread around the track as spread x sigma

scaleToPhysical= (False), if True, plot positions in kpc and velocities in km/s

simple= (False), if True, use a simple estimate for the spread in perpendicular angle

bovy_plot.bovy_plot args and kwargs

OUTPUT:

plot to output device

HISTORY:

2013-12-09 - Written - Bovy (IAS)

galpy.df.streamdf.pOparapar

streamdf.**pOparapar** (Opar, apar, tdisrupt=None)

NAME:

pOparapar

PURPOSE:

return the probability of a given parallel (frequency,angle) offset pair

INPUT:

Opar - parallel frequency offset (array) (can be Quantity)

apar - parallel angle offset along the stream (scalar) (can be Quantity)

OUTPUT:

p(Opar,apar)

HISTORY:

2015-12-07 - Written - Bovy (UofT)

galpy.df.streamdf.ptdAngle

streamdf.**ptdAngle**(*t, dangle*)

NAME:

ptdangle

PURPOSE:

return the probability of a given stripping time at a given angle along the stream

INPUT:

t - stripping time

dangle - angle offset along the stream

OUTPUT:

p(t,dangle)

HISTORY:

2013-12-05 - Written - Bovy (IAS)

galpy.df.streamdf.sample

streamdf.**sample**(*n, returnAdt=False, returndT=False, interp=None, xy=False, lb=False*)

NAME:

sample

PURPOSE:

sample from the DF

INPUT:

n - number of points to return

returnAdt= (False) if True, return (Omega,angle,dt)

returndT= (False) if True, also return the time since the star was stripped

interp= (object-wide default) use interpolation of the stream track

xy= (False) if True, return Galactocentric rectangular coordinates

lb= (False) if True, return Galactic l,b,d,vlos,pmll,pmdb coordinates

OUTPUT:

(R,vR,vT,z,vz,phi) of points on the stream in 6,N array

HISTORY:

2013-12-22 - Written - Bovy (IAS)

galpy.df.streamdf.sigangledAngle

`streamdf.sigangledAngle` (*dangle*, *assumeZeroMean=True*, *smallest=False*, *simple=False*)

NAME:

sigangledAngle

PURPOSE:

calculate the dispersion in the perpendicular angle at a given angle

INPUT:

dangle - angle offset along the stream

assumeZeroMean= (True) if True, assume that the mean is zero (should be)

smallest= (False) calculate for smallest eigenvalue direction rather than for middle

simple= (False), if True, return an even simpler estimate

OUTPUT:

dispersion in the perpendicular angle at this angle

HISTORY:

2013-12-06 - Written - Bovy (IAS)

galpy.df.streamdf.sigOmega

`streamdf.sigOmega` (*dangle*)

NAME:

sigmaOmega

PURPOSE:

calculate the 1D sigma in frequency as a function of angle, assuming a uniform time distribution up to a maximum time

INPUT:

dangle - angle offset

OUTPUT:

sigma Omega

HISTORY:

2013-12-05 - Written - Bovy (IAS)

galpy.df.streamdf.sigtdAngle

`streamdf.sigtdAngle` (*dangle*)

NAME:

sigtdAngle

PURPOSE:

calculate the dispersion in the stripping times at a given angle

INPUT:

dangle - angle offset along the stream

OUTPUT:

dispersion in the stripping times at this angle

HISTORY:

2013-12-05 - Written - Bovy (IAS)

galpy.df.streamdf.subhalo_encounters

`streamdf.subhalo_encounters` (*venc=inf, sigma=0.6818181818181818, nsubhalo=0.3, bmax=0.025, yoon=False*)

NAME:

subhalo_encounters

PURPOSE:

estimate the number of encounters with subhalos over the lifetime of this stream, using a formalism similar to that of Yoon et al. (2011)

INPUT:

venc= (numpy.inf) count encounters with (relative) speeds less than this (relative radial velocity in cylindrical stream frame, unless *yoon* is True) (can be Quantity)

sigma= (150/220) velocity dispersion of the DM subhalo population (can be Quantity)

nsubhalo= (0.3) spatial number density of subhalos (can be Quantity)

bmax= (0.025) maximum impact parameter (if larger than width of stream) (can be Quantity)

yoon= (False) if True, use erroneous Yoon et al. formula

OUTPUT:

number of encounters

HISTORY:

2016-01-19 - Written - Bovy (UofT)

3.4.7 The distribution function of a gap in a tidal stream

From Sanders, Bovy, & Erkal 2015; see *Modeling gaps in streams*. Implemented as a subclass of `streamdf`. No full implementation is available currently, but the model can be set up and sampled as in the above paper.

General instance routines

The stream gap DF

class galpy.df.streamgapdf(*args, **kwargs)

The DF of a gap in a tidal stream

__init__(*args, **kwargs)

NAME:

__init__

PURPOSE:

Initialize the DF of a gap in a stellar stream

INPUT:

streamdf args and kwargs

Subhalo and impact parameters:

 impactb= impact parameter (can be Quantity)

 subhalovel= velocity of the subhalo shape=(3) (can be Quantity)

 timpact time since impact (can be Quantity)

 impact_angle= angle offset from progenitor at which the impact occurred (rad) (can be Quantity)

 Subhalo: specify either 1(mass and size of Plummer sphere or 2(general spherical-potential object (kick is numerically computed)

 1(GM= mass of the subhalo (can be Quantity)

 rs= size parameter of the subhalo (can be Quantity)

 2(subhalopot= galpy potential object or list thereof (should be spherical)

 3(hernquist= (False) if True, use Hernquist kicks for GM/rs

 deltaAngleTrackImpact= (None) angle to estimate the stream track over to determine the effect of the impact [similar to deltaAngleTrack] (rad)

 nTrackChunksImpact= (floor(deltaAngleTrack/0.15)+1) number of chunks to divide the progenitor track in near the impact [similar to nTrackChunks]

 nKickPoints= (30xnTrackChunksImpact) number of points along the stream to compute the kicks at (kicks are then interpolated); '30' chosen such that higherorderTrack can be set to False and get calculations accurate to > 99%

 nokicksetup= (False) if True, only run as far as setting up the coordinate transformation at the time of impact (useful when using this in streampepperdf)

 spline_order= (3) order of the spline to interpolate the kicks with

 higherorderTrack= (False) if True, calculate the track using higher-order terms

OUTPUT:

 object

HISTORY:

 2015-06-02 - Started - Bovy (IAS)

Helper routines to compute kicks

galpy.df.impulse_deltav_plummer

`galpy.df.impulse_deltav_plummer` (*v*, *y*, *b*, *w*, *GM*, *rs*)

NAME:

impulse_deltav_plummer

PURPOSE:

calculate the delta velocity to due an encounter with a Plummer sphere in the impulse approximation;
allows for arbitrary velocity vectors, but *y* is input as the position along the stream

INPUT:

v - velocity of the stream (nstar,3)

y - position along the stream (nstar)

b - impact parameter

w - velocity of the Plummer sphere (3)

GM - mass of the Plummer sphere (in natural units)

rs - size of the Plummer sphere

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-04-30 - Written based on Erkal's expressions - Bovy (IAS)

galpy.df.impulse_deltav_plummer_curvedstream

`galpy.df.impulse_deltav_plummer_curvedstream` (*v*, *x*, *b*, *w*, *x0*, *v0*, *GM*, *rs*)

NAME:

impulse_deltav_plummer_curvedstream

PURPOSE:

calculate the delta velocity to due an encounter with a Plummer sphere in the impulse approximation;
allows for arbitrary velocity vectors, and arbitrary position along the stream

INPUT:

v - velocity of the stream (nstar,3)

x - position along the stream (nstar,3)

b - impact parameter

w - velocity of the Plummer sphere (3)

x0 - point of closest approach

v0 - velocity of point of closest approach

GM - mass of the Plummer sphere (in natural units)

rs - size of the Plummer sphere

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-05-04 - Written based on above - SANDERS

galpy.df.impulse_deltav_hernquist

galpy.df.**impulse_deltav_hernquist** (*v, y, b, w, GM, rs*)

NAME:

impulse_deltav_hernquist

PURPOSE:

calculate the delta velocity to due an encounter with a Hernquist sphere in the impulse approximation;
allows for arbitrary velocity vectors, but y is input as the position along the stream

INPUT:

v - velocity of the stream (nstar,3)

y - position along the stream (nstar)

b - impact parameter

w - velocity of the Hernquist sphere (3)

GM - mass of the Hernquist sphere (in natural units)

rs - size of the Hernquist sphere

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-08-13 SANDERS, using Wyn Evans calculation

galpy.df.impulse_deltav_hernquist_curvedstream

galpy.df.**impulse_deltav_hernquist_curvedstream** (*v, x, b, w, x0, v0, GM, rs*)

NAME:

impulse_deltav_plummer_hernquist

PURPOSE:

calculate the delta velocity to due an encounter with a Hernquist sphere in the impulse approximation;
allows for arbitrary velocity vectors, and arbitrary position along the stream

INPUT:

v - velocity of the stream (nstar,3)

x - position along the stream (nstar,3)

b - impact parameter

w - velocity of the Hernquist sphere (3)

x0 - point of closest approach

v0 - velocity of point of closest approach
GM - mass of the Hernquist sphere (in natural units)
rs - size of the Hernquist sphere

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-08-13 - SANDERS, using Wyn Evans calculation

galpy.df.impulse_deltav_general

`galpy.df.impulse_deltav_general` (*v*, *y*, *b*, *w*, *pot*)

NAME:

impulse_deltav_general

PURPOSE:

calculate the delta velocity to due an encounter with a general spherical potential in the impulse approximation; allows for arbitrary velocity vectors, but y is input as the position along the stream

INPUT:

v - velocity of the stream (nstar,3)
y - position along the stream (nstar)
b - impact parameter
w - velocity of the subhalo (3)
pot - Potential object or list thereof (should be spherical)

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-05-04 - SANDERS
2015-06-15 - Tweak to use galpy' potential objects - Bovy (IAS)

galpy.df.impulse_deltav_general_curvedstream

`galpy.df.impulse_deltav_general_curvedstream` (*v*, *x*, *b*, *w*, *x0*, *v0*, *pot*)

NAME:

impulse_deltav_general_curvedstream

PURPOSE:

calculate the delta velocity to due an encounter with a general spherical potential in the impulse approximation; allows for arbitrary velocity vectors and arbitrary shaped streams

INPUT:

v - velocity of the stream (nstar,3)
x - position along the stream (nstar,3)
b - impact parameter
w - velocity of the subhalo (3)
x0 - position of closest approach (3)
v0 - velocity of stream at closest approach (3)
pot - Potential object or list thereof (should be spherical)

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-05-04 - SANDERS
2015-06-15 - Tweak to use galpy's potential objects - Bovy (IAS)

galpy.df.impulse_deltav_general_orbitintegration

`galpy.df.impulse_deltav_general_orbitintegration(v, x, b, w, x0, v0, pot, tmax, galpot, tmaxfac=10.0, nsamp=1000, integrate_method='symplec4_c')`

NAME:

impulse_deltav_general_orbitintegration

PURPOSE:

calculate the delta velocity to due an encounter with a general spherical potential NOT in the impulse approximation by integrating each particle in the underlying galactic potential; allows for arbitrary velocity vectors and arbitrary shaped streams.

INPUT:

v - velocity of the stream (nstar,3)
x - position along the stream (nstar,3)
b - impact parameter
w - velocity of the subhalo (3)
x0 - position of closest approach (3)
v0 - velocity of stream at closest approach (3)
pot - Potential object or list thereof (should be spherical)
tmax - maximum integration time
galpot - galpy Potential object or list thereof
nsamp(1000) - number of forward integration points
integrate_method= ('symplec4_c') orbit integrator to use (see Orbit.integrate)

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-08-17 - SANDERS

galpy.df.impulse_deltav_general_fullplummerintegration

```
galpy.df.impulse_deltav_general_fullplummerintegration(v, x, b, w, x0, v0,
                                                         galpot, GM, rs, tmax-
                                                         fac=10.0, N=1000, inte-
                                                         grate_method='symplec4_c')
```

NAME:

impulse_deltav_general_fullplummerintegration

PURPOSE:

calculate the delta velocity to due an encounter with a moving Plummer sphere and galactic potential relative to just in galactic potential

INPUT:

v - velocity of the stream (nstar,3)

x - position along the stream (nstar,3)

b - impact parameter

w - velocity of the subhalo (3)

x0 - position of closest approach (3)

v0 - velocity of stream at closest approach (3)

galpot - Galaxy Potential object

GM - mass of Plummer

rs - scale of Plummer

tmaxfac(10) - multiple of rs/fabs(w - v0) to use for time integration interval

N(1000) - number of forward integration points

integrate_method('symplec4_c') - orbit integrator to use (see Orbit.integrate)

OUTPUT:

deltav (nstar,3)

HISTORY:

2015-08-18 - SANDERS

3.5 Utilities (galpy.util)

3.5.1 galpy.util.config

Configuration module

galpy.util.config.set_ro

galpy.util.config.set_ro(ro)

NAME: set_ro

PURPOSE: set the global configuration value of ro (distance scale)

INPUT: ro - scale in kpc or astropy Quantity

OUTPUT: (none)

HISTORY: 2016-01-05 - Written - Bovy (UofT)

galpy.util.config.set_vo

galpy.util.config.set_vo(vo)

NAME: set_vo

PURPOSE: set the global configuration value of vo (velocity scale)

INPUT: vo - scale in km/s or astropy Quantity

OUTPUT: (none)

HISTORY: 2016-01-05 - Written - Bovy (UofT)

3.5.2 galpy.util.bovy_plot

Warning: Importing `galpy.util.bovy_plot` (or having it be imported by other galpy routines) with seaborn installed may change the seaborn plot style. If you don't like this, set the configuration parameter `seaborn-plotting-defaults` to `False` in the *configuration file*

Various plotting routines:

galpy.util.bovy_plot.bovy_dens2d

galpy.util.bovy_plot.bovy_dens2d(X, **kwargs)

NAME:

bovy_dens2d

PURPOSE:

plot a 2d density with optional contours

INPUT:

first argument is the density

matplotlib.pyplot.imshow keywords (see http://matplotlib.sourceforge.net/api/axes_api.html#matplotlib.axes.Axes.imshow)

xlabel - (raw string!) x-axis label, LaTeX math mode, no \$s needed

ylabel - (raw string!) y-axis label, LaTeX math mode, no \$s needed

xrange

yrange

noaxes - don't plot any axes

overplot - if True, overplot

colorbar - if True, add colorbar

shrink= colorbar argument: shrink the colorbar by the factor (optional)

conditional - normalize each column separately (for probability densities, i.e., cntrmass=True)

gcf=True does not start a new figure (does change the ranges and labels)

Contours:

justcontours - if True, only draw contours

contours - if True, draw contours (10 by default)

levels - contour-levels

cntrmass - if True, the density is a probability and the levels are probability masses contained within the contour

cntrcolors - colors for contours (single color or array)

cntrlabel - label the contours

cntrlw, cntrls - linewidths and linestyle for contour

cntrlabelsize, cntrlabelcolors, cntrinline - contour arguments

cntrSmooth - use ndimage.gaussian_filter to smooth before contouring

onedhists - if True, make one-d histograms on the sides

onedhistcolor - histogram color

retAxes= return all Axes instances

retCont= return the contour instance

OUTPUT:

plot to output device, Axes instances depending on input

HISTORY:

2010-03-09 - Written - Bovy (NYU)

galpy.util.bovy_plot.bovy_end_print

`galpy.util.bovy_plot.bovy_end_print` (*filename*, ***kwargs*)

NAME:

bovy_end_print

PURPOSE:

saves the current figure(s) to filename

INPUT:

filename - filename for plot (with extension)

OPTIONAL INPUTS:

format - file-format

OUTPUT:

(none)

HISTORY:

2009-12-23 - Written - Bovy (NYU)

galpy.util.bovy_plot.bovy_hist

galpy.util.bovy_plot.**bovy_hist** (*x*, *xlabel=None*, *ylabel=None*, *overplot=False*, ***kwargs*)

NAME:

bovy_hist

PURPOSE:

wrapper around matplotlib's hist function

INPUT:

x - array to histogram

xlabel - (raw string!) x-axis label, LaTeX math mode, no \$s needed

ylabel - (raw string!) y-axis label, LaTeX math mode, no \$s needed

yrange - set the y-axis range

+all pyplot.hist keywords

OUTPUT: (from the matplotlib docs: http://matplotlib.sourceforge.net/api/pyplot_api.html#matplotlib.pyplot.hist)

The return value is a tuple (*n*, *bins*, *patches*) or (*[n0, n1, ...]*, *bins*, *[patches0, patches1, ...]*) if the input contains multiple data

HISTORY:

2009-12-23 - Written - Bovy (NYU)

galpy.util.bovy_plot.bovy_plot

galpy.util.bovy_plot.**bovy_plot** (**args*, ***kwargs*)

NAME:

bovy_plot

PURPOSE:

wrapper around matplotlib's plot function

INPUT:

see http://matplotlib.sourceforge.net/api/pyplot_api.html#matplotlib.pyplot.plot

xlabel - (raw string!) x-axis label, LaTeX math mode, no \$s needed

ylabel - (raw string!) y-axis label, LaTeX math mode, no \$s needed

xrange

yrange

scatter= if True, use `pyplot.scatter` and its options etc.
 colorbar= if True, and `scatter==True`, add colorbar
 crange - range for colorbar of `scatter==True`
 clabel= label for colorbar
 overplot=True does not start a new figure and does not change the ranges and labels
 gcf=True does not start a new figure (does change the ranges and labels)
 onedhists - if True, make one-d histograms on the sides
 onedhistcolor, onedhistfc, onedhistec
 onedhistxnormed, onedhistynormed - normed keyword for one-d histograms
 onedhistxweights, onedhistyweights - weights keyword for one-d histograms
 bins= number of bins for onedhists
 semilogx=, semilogy=, loglog= if True, plot logs

OUTPUT:

plot to output device, returns what `pyplot.plot` returns, or 3 Axes instances if `onedhists=True`

HISTORY:

2009-12-28 - Written - Bovy (NYU)

galpy.util.bovy_plot.bovy_print

`galpy.util.bovy_plot.bovy_print` (*fig_width=5, fig_height=5, axes_labelsize=16, text_fontsize=11, legend_fontsize=12, xtick_labelsize=10, ytick_labelsize=10, xtick_minor_size=2, ytick_minor_size=2, xtick_major_size=4, ytick_major_size=4*)

NAME:

`bovy_print`

PURPOSE:

setup a figure for plotting

INPUT:

`fig_width` - width in inches
`fig_height` - height in inches
`axes_labelsize` - size of the axis-labels
`text_fontsize` - font-size of the text (if any)
`legend_fontsize` - font-size of the legend (if any)
`xtick_labelsize` - size of the x-axis labels
`ytick_labelsize` - size of the y-axis labels
`xtick_minor_size` - size of the minor x-ticks
`ytick_minor_size` - size of the minor y-ticks

OUTPUT:

(none)

HISTORY:

2009-12-23 - Written - Bovy (NYU)

galpy.util.bovy_plot.bovy_text

galpy.util.bovy_plot.**bovy_text** (**args, **kwargs*)

NAME:

bovy_text

PURPOSE:

thin wrapper around matplotlib's text and annotate

use keywords:

 'bottom_left=True'

 'bottom_right=True'

 'top_left=True'

 'top_right=True'

 'title=True'

to place the text in one of the corners or use it as the title

INPUT:

see matplotlib's text (http://matplotlib.sourceforge.net/api/pyplot_api.html#matplotlib.pyplot.text)

OUTPUT:

prints text on the current figure

HISTORY:

2010-01-26 - Written - Bovy (NYU)

galpy.util.bovy_plot.scatterplot

galpy.util.bovy_plot.**scatterplot** (*x, y, *args, **kwargs*)

NAME:

scatterplot

PURPOSE:

make a 'smart' scatterplot that is a density plot in high-density regions and a regular scatterplot for outliers

INPUT:

x, y

xlabel - (raw string!) x-axis label, LaTeX math mode, no \$s needed

ylabel - (raw string!) y-axis label, LaTeX math mode, no \$s needed

xrange

yrange

bins - number of bins to use in each dimension

weights - data-weights

aspect - aspect ratio

conditional - normalize each column separately (for probability densities, i.e., `cntrmass=True`)

`gcf=True` does not start a new figure (does change the ranges and labels)

contours - if False, don't plot contours

justcontours - if True, only draw contours, no density

cntrcolors - color of contours (can be array as for `bovy_dens2d`)

cntrlw, cntrls - linewidths and linestyles for contour

cntrSmooth - use `ndimage.gaussian_filter` to smooth before contouring

levels - contour-levels; data points outside of the last level will be individually shown (so, e.g., if this list is descending, contours and data points will be overplotted)

onedhists - if True, make one-d histograms on the sides

onedhistx - if True, make one-d histograms on the side of the x distribution

onedhisty - if True, make one-d histograms on the side of the y distribution

onedhistcolor, onedhistfc, onedhistec

onedhistxnormed, onedhistynormed - normed keyword for one-d histograms

onedhistxweights, onedhistyweights - weights keyword for one-d histograms

cmap= cmap for density plot

hist= and edges= - you can supply the histogram of the data yourself, this can be useful if you want to censor the data, both need to be set and calculated using `scipy.histogramdd` with the given range

retAxes= return all Axes instances

OUTPUT:

plot to output device, Axes instance(s) or not, depending on input

HISTORY:

2010-04-15 - Written - Bovy (NYU)

`galpy` also contains a new matplotlib projection '`galpolar`' that can be used when working with older versions of matplotlib like '`polar`' to create a polar plot in which the azimuth increases clockwise (like when looking at the Milky Way from the north Galactic pole). In newer versions of matplotlib, this does not work, but the '`polar`' projection now supports clockwise azimuths by doing, e.g.,

```
>>> ax= pyplot.subplot(111,projection='polar')
>>> ax.set_theta_direction(-1)
```

3.5.3 galpy.util.bovy_conversion

Utility functions that provide conversions between `galpy`'s *natural* units and *physical* units. These can be used to translate `galpy` outputs in natural coordinates to physical units by multiplying with the appropriate function.

These could also be used to figure out the conversion between different units. For example, if you want to know how many GeV cm^{-3} correspond to $1 M_{\odot} \text{pc}^{-3}$, you can calculate

```
>>> from galpy.util import bovy_conversion
>>> bovy_conversion.dens_in_gevcc(1.,1.)/bovy_conversion.dens_in_msolpc3(1.,1.)
# 37.978342941703616
```

or $1 M_{\odot} \text{ pc}^{-3} \approx 40 \text{ GeV cm}^{-3}$.

Functions:

galpy.util.bovy_conversion.dens_in_criticaldens

galpy.util.bovy_conversion.**dens_in_criticaldens**(*vo*, *ro*, *H=70.0*)

NAME:

dens_in_criticaldens

PURPOSE:

convert density to units of the critical density

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

H= (default: 70) Hubble constant in km/s/Mpc

OUTPUT:

conversion from units where vo=1. at ro=1. to units of the critical density

HISTORY:

2014-01-28 - Written - Bovy (IAS)

galpy.util.bovy_conversion.dens_in_gevcc

galpy.util.bovy_conversion.**dens_in_gevcc**(*vo*, *ro*)

NAME:

dens_in_gevcc

PURPOSE:

convert density to GeV / cm³

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1. to GeV/cm³

HISTORY:

2014-06-16 - Written - Bovy (IAS)

galpy.util.bovy_conversion.dens_in_meanmatterdens

`galpy.util.bovy_conversion.dens_in_meanmatterdens` (*vo*, *ro*, *H*=70.0, *Om*=0.3)

NAME:

`dens_in_meanmatterdens`

PURPOSE:

convert density to units of the mean matter density

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

H= (default: 70) Hubble constant in km/s/Mpc

Om= (default: 0.3) Omega matter

OUTPUT:

conversion from units where *vo*=1. at *ro*=1. to units of the mean matter density

HISTORY:

2014-01-28 - Written - Bovy (IAS)

galpy.util.bovy_conversion.dens_in_msolpc3

`galpy.util.bovy_conversion.dens_in_msolpc3` (*vo*, *ro*)

NAME:

`dens_in_msolpc3`

PURPOSE:

convert density to Msolar / pc³

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where *vo*=1. at *ro*=1. to Msolar/pc³

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.force_in_2piGmsolpc2

`galpy.util.bovy_conversion.force_in_2piGmsolpc2` (*vo*, *ro*)

NAME:

`force_in_2piGmsolpc2`

PURPOSE:

convert a force or acceleration to 2piG x Msolar / pc²

INPUT:

vo - velocity unit in km/s
ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.force_in_pcMyr2

galpy.util.bovy_conversion.**force_in_pcMyr2**(vo, ro)

NAME:

force_in_pcMyr2

PURPOSE:

convert a force or acceleration to pc/Myr²

INPUT:

vo - velocity unit in km/s
ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.force_in_10m13kms2

galpy.util.bovy_conversion.**force_in_10m13kms2**(vo, ro)

NAME:

force_in_10m13kms2

PURPOSE:

convert a force or acceleration to 10⁻¹³ km/s²

INPUT:

vo - velocity unit in km/s
ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2014-01-22 - Written - Bovy (IAS)

galpy.util.bovy_conversion.force_in_kmsMyr

`galpy.util.bovy_conversion.force_in_kmsMyr(vo, ro)`

NAME:

force_in_kmsMyr

PURPOSE:

convert a force or acceleration to km/s/Myr

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.freq_in_Gyr

`galpy.util.bovy_conversion.freq_in_Gyr(vo, ro)`

NAME:

freq_in_Gyr

PURPOSE:

convert a frequency to 1/Gyr

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.freq_in_kmskpc

`galpy.util.bovy_conversion.freq_in_kmskpc(vo, ro)`

NAME:

freq_in_kmskpc

PURPOSE:

convert a frequency to km/s/kpc

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.surfdens_in_msolpc2

`galpy.util.bovy_conversion.surfdens_in_msolpc2(vo, ro)`

NAME:

surfdens_in_msolpc2

PURPOSE:

convert a surface density to Msolar / pc²

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.mass_in_msol

`galpy.util.bovy_conversion.mass_in_msol(vo, ro)`

NAME:

mass_in_msol

PURPOSE:

convert a mass to Msolar

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.mass_in_1010msol

`galpy.util.bovy_conversion.mass_in_1010msol` (*vo*, *ro*)

NAME:

mass_in_1010msol

PURPOSE:

convert a mass to 10^{10} x Msolar

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.time_in_Gyr

`galpy.util.bovy_conversion.time_in_Gyr` (*vo*, *ro*)

NAME:

time_in_Gyr

PURPOSE:

convert a time to Gyr

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2013-09-01 - Written - Bovy (IAS)

galpy.util.bovy_conversion.velocity_in_kpcGyr

`galpy.util.bovy_conversion.velocity_in_kpcGyr` (*vo*, *ro*)

NAME:

velocity_in_kpcGyr

PURPOSE:

convert a velocity to kpc/Gyr

INPUT:

vo - velocity unit in km/s

ro - length unit in kpc

OUTPUT:

conversion from units where vo=1. at ro=1.

HISTORY:

2014-12-19 - Written - Bovy (IAS)

3.5.4 galpy.util.bovy_coords

Warning: galpy uses a left-handed coordinate frame, as is common in studies of the kinematics of the Milky Way. Care should be taken when using the coordinate transformation routines below for coordinates in a right-handed frame, the routines do not always apply and are only tested for the standard galpy left-handed frame.

Various coordinate transformation routines with fairly self-explanatory names:

galpy.util.bovy_coords.cov_dvrpmlbb_to_vxyz

galpy.util.bovy_coords.cov_dvrpmlbb_to_vxyz(*d, e_d, e_vr, pmll, pmbb, cov_pmllbb, l, b,*
plx=False, degree=False)

NAME:

cov_dvrpmlbb_to_vxyz

PURPOSE:

propagate distance, radial velocity, and proper motion uncertainties to Galactic coordinates

INPUT:

d - distance [kpc, as/mas for plx]

e_d - distance uncertainty [kpc, [as/mas] for plx]

e_vr - low velocity uncertainty [km/s]

pmll - proper motion in l (*cos(b)) [[as/mas]/yr]

pmbb - proper motion in b [[as/mas]/yr]

cov_pmllbb - uncertainty covariance for proper motion [pmll is pmll x cos(b)]

l - Galactic longitude

b - Galactic latitude

KEYWORDS:

plx - if True, d is a parallax, and e_d is a parallax uncertainty

degree - if True, l and b are given in degree

OUTPUT:

cov(vx,vy,vz) [3,3] or[:,3,3]

HISTORY:

2010-04-12 - Written - Bovy (NYU)

galpy.util.bovy_coords.cov_pmrappmdec_to_pmlpmbb

`galpy.util.bovy_coords.cov_pmrappmdec_to_pmlpmbb` (*cov_pmrappmdec*, *ra*, *dec*, *degree=False*,
epoch=2000.0)

NAME:

`cov_pmrappmdec_to_pmlpmbb`

PURPOSE:

propagate the proper motions errors through the rotation from (ra,dec) to (l,b)

INPUT:

cov_pmrappmdec - uncertainty covariance matrix of the proper motion in ra (multiplied with cos(dec)) and dec [2,2] or [:,2,2]

ra - right ascension

dec - declination

degree - if True, ra and dec are given in degrees (default=False)

epoch - epoch of ra,dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

cov_pmlpmbb [2,2] or [:,2,2] [pml here is pml x cos(b)]

HISTORY:

2010-04-12 - Written - Bovy (NYU)

galpy.util.bovy_coords.cyl_to_rect

`galpy.util.bovy_coords.cyl_to_rect` (*R*, *phi*, *Z*)

NAME:

`cyl_to_rect`

PURPOSE:

convert from cylindrical to rectangular coordinates

INPUT:

R, *phi*, *Z* - cylindrical coordinates

OUTPUT:

X, *Y*, *Z*

HISTORY:

2011-02-23 - Written - Bovy (NYU)

galpy.util.bovy_coords.cyl_to_rect_vec

`galpy.util.bovy_coords.cyl_to_rect_vec` (*vr*, *vt*, *vz*, *phi*)

NAME:

`cyl_to_rect_vec`

PURPOSE:

transform vectors from cylindrical to rectangular coordinate vectors

INPUT:

vr - radial velocity

vt - tangential velocity

vz - vertical velocity

phi - azimuth

OUTPUT:

vx,vy,vz

HISTORY:

2011-02-24 - Written - Bovy (NYU)

galpy.util.bovy_coords.dl_to_rphi_2d

galpy.util.bovy_coords.dl_to_rphi_2d(*d, l, degree=False, ro=1.0, phio=0.0*)

NAME:

dl_to_rphi_2d

PURPOSE:

convert Galactic longitude and distance to Galactocentric radius and azimuth

INPUT:

d - distance

l - Galactic longitude [rad/deg if degree]

KEYWORDS:

degree= (False): l is in degrees rather than rad

ro= (1) Galactocentric radius of the observer

phio= (0) Galactocentric azimuth of the observer [rad/deg if degree]

OUTPUT:

(R,phi); phi in degree if degree

HISTORY:

2012-01-04 - Written - Bovy (IAS)

galpy.util.bovy_coords.galcencyl_to_XYZ

galpy.util.bovy_coords.galcencyl_to_XYZ(*R, phi, Z, Xsun=1.0, Zsun=0.0, _extra_rot=True*)

NAME:

galcencyl_to_XYZ

PURPOSE:

transform cylindrical Galactocentric coordinates to XYZ coordinates (wrt Sun)

INPUT:

R, phi, Z - Galactocentric cylindrical coordinates

Xsun - cylindrical distance to the GC (can be array of same length as R)

Zsun - Sun's height above the midplane (can be array of same length as R)

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

X,Y,Z

HISTORY:

2011-02-23 - Written - Bovy (NYU)

2017-10-24 - Allowed Xsun/Zsun to be arrays - Bovy (UofT)

galpy.util.bovy_coords.galcencyl_to_vxvyvz

`galpy.util.bovy_coords.galcencyl_to_vxvyvz` (*vR, vT, vZ, phi, vsun=[0.0, 1.0, 0.0], Xsun=1.0, Zsun=0.0, _extra_rot=True*)

NAME:

`galcencyl_to_vxvyvz`

PURPOSE:

transform cylindrical Galactocentric coordinates to XYZ (wrt Sun) coordinates for velocities

INPUT:

vR - Galactocentric radial velocity

vT - Galactocentric tangential velocity

vZ - Galactocentric vertical velocity

phi - Galactocentric azimuth

vsun - velocity of the sun in the GC frame ndarray[3] (can be array of same length as vRg; shape [3,N])

Xsun - cylindrical distance to the GC (can be array of same length as vRg)

Zsun - Sun's height above the midplane (can be array of same length as vRg)

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

vx,vy,vz

HISTORY:

2011-02-24 - Written - Bovy (NYU)

2017-10-24 - Allowed vsun/Xsun/Zsun to be arrays - Bovy (NYU)

galpy.util.bovy_coords.galcenrect_to_XYZ

`galpy.util.bovy_coords.galcenrect_to_XYZ` (*X*, *Y*, *Z*, *Xsun*=1.0, *Zsun*=0.0, *_extra_rot*=True)
NAME:

`galcenrect_to_XYZ`

PURPOSE:

transform rectangular Galactocentric to XYZ coordinates (wrt Sun) coordinates

INPUT:

X, *Y*, *Z* - Galactocentric rectangular coordinates

Xsun - cylindrical distance to the GC (can be array of same length as *X*)

Zsun - Sun's height above the midplane (can be array of same length as *X*)

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

(*X*, *Y*, *Z*)

HISTORY:

2011-02-23 - Written - Bovy (NYU)

2016-05-12 - Edited to properly take into account the Sun's vertical position; dropped *Ysun* keyword - Bovy (UofT)

2017-10-24 - Allowed *Xsun*/*Zsun* to be arrays - Bovy (UofT)

2018-04-18 - Tweaked to be consistent with astropy's Galactocentric frame - Bovy (UofT)

galpy.util.bovy_coords.galcenrect_to_vxvyvz

`galpy.util.bovy_coords.galcenrect_to_vxvyvz` (*vXg*, *vYg*, *vZg*, *vsun*=[0.0, 1.0, 0.0],
Xsun=1.0, *Zsun*=0.0, *_extra_rot*=True)

NAME:

`galcenrect_to_vxvyvz`

PURPOSE:

transform rectangular Galactocentric coordinates to XYZ coordinates (wrt Sun) for velocities

INPUT:

vXg - Galactocentric x-velocity

vYg - Galactocentric y-velocity

vZg - Galactocentric z-velocity

vsun - velocity of the sun in the GC frame ndarray[3] (can be array of same length as *vXg*; shape [3,N])

Xsun - cylindrical distance to the GC (can be array of same length as *vXg*)

Zsun - Sun's height above the midplane (can be array of same length as *vXg*)

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

[:,3]= vx, vy, vz

HISTORY:

2011-02-24 - Written - Bovy (NYU)

2016-05-12 - Edited to properly take into account the Sun's vertical position; dropped Ysun keyword - Bovy (UofT)

2017-10-24 - Allowed vsun/Xsun/Zsun to be arrays - Bovy (UofT)

2018-04-18 - Tweaked to be consistent with astropy's Galactocentric frame - Bovy (UofT)

galpy.util.bovy_coords.lb_to_radec

`galpy.util.bovy_coords.lb_to_radec(l, b, degree=False, epoch=2000.0)`

NAME:

lb_to_radec

PURPOSE:

transform from Galactic coordinates to equatorial coordinates

INPUT:

l - Galactic longitude

b - Galactic latitude

degree - (Bool) if True, l and b are given in degree and ra and dec will be as well

epoch - epoch of ra,dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

ra,dec

For vector inputs[:,2]

HISTORY:

2010-04-07 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

2016-05-13 - Added support for using astropy's coordinate transformations and for non-standard epochs - Bovy (UofT)

galpy.util.bovy_coords.lbd_to_xyz

`galpy.util.bovy_coords.lbd_to_xyz(l, b, d, degree=False)`

NAME:

lbd_to_XYZ

PURPOSE:

transform from spherical Galactic coordinates to rectangular Galactic coordinates (works with vector inputs)

INPUT:

l - Galactic longitude (rad)
b - Galactic latitude (rad)
d - distance (arbitrary units)
degree - (bool) if True, l and b are in degrees

OUTPUT:

[X,Y,Z] in whatever units d was in
For vector inputs[:,3]

HISTORY:

2009-10-24- Written - Bovy (NYU)
2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.pmlpmbb_to_pmrpmdec

```
galpy.util.bovy_coords.pmlpmbb_to_pmrpmdec(pml, pmbb, l, b, degree=False, epoch=2000.0)
```

NAME:

pmlpmbb_to_pmrpmdec

PURPOSE:

rotate proper motions in (l,b) into proper motions in (ra,dec)

INPUT:

pml - proper motion in l (multiplied with cos(b)) [mas/yr]
pmbb - proper motion in b [mas/yr]
l - Galactic longitude
b - Galactic latitude
degree - if True, l and b are given in degrees (default=False)
epoch - epoch of ra,dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

(pmra x cos(dec),pmdec), for vector inputs[:,2]

HISTORY:

2010-04-07 - Written - Bovy (NYU)
2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.pmrpmdec_to_pmlpmbb

```
galpy.util.bovy_coords.pmrpmdec_to_pmlpmbb(pmra, pmdec, ra, dec, degree=False, epoch=2000.0)
```

NAME:

pmrapmdec_to_pmllpmbb

PURPOSE:

rotate proper motions in (ra,dec) into proper motions in (l,b)

INPUT:

pmra - proper motion in ra (multiplied with cos(dec)) [mas/yr]

pmdec - proper motion in dec [mas/yr]

ra - right ascension

dec - declination

degree - if True, ra and dec are given in degrees (default=False)

epoch - epoch of ra,dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

(pmll x cos(b),pmbb) for vector inputs[:,2]

HISTORY:

2010-04-07 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.pmrpmdec_to_custom

galpy.util.bovy_coords.pmrpmdec_to_custom(*pmra, pmdec, ra, dec, T=None, degree=False, epoch=2000.0*)

NAME:

pmrpmdec_to_custom

PURPOSE:

rotate proper motions in (ra,dec) to proper motions in a custom set of sky coordinates (phi1,phi2)

INPUT:

pmra - proper motion in ra (multiplied with cos(dec)) [mas/yr]

pmdec - proper motion in dec [mas/yr]

ra - right ascension

dec - declination

T= matrix defining the transformation: new_rect= T dot old_rect, where old_rect = [cos(dec)cos(ra),cos(dec)sin(ra),sin(dec)] and similar for new_rect

degree= (False) if True, ra and dec are given in degrees (default=False)

epoch= (2000.) epoch of ra,dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

(pmphi1 x cos(phi2),pmphi2) for vector inputs[:,2]

HISTORY:

2016-10-24 - Written - Bovy (UofT/CCA)

galpy.util.bovy_coords.pupv_to_vRvz

`galpy.util.bovy_coords.pupv_to_vRvz` (*pu, pv, u, v, delta=1.0, oblate=False*)

NAME:

`pupv_to_vRvz`

PURPOSE:

calculate cylindrical vR and vz from momenta in prolate or oblate confocal u and v coordinates for a given focal length δ

INPUT:

pu - u momentum

pv - v momentum

u - u coordinate

v - v coordinate

δ = focus

$oblate$ = (False) if True, compute oblate confocal coordinates instead of prolate

OUTPUT:

(vR, vz)

HISTORY:

2017-12-04 - Written - Bovy (UofT)

galpy.util.bovy_coords.radec_to_lb

`galpy.util.bovy_coords.radec_to_lb` (*ra, dec, degree=False, epoch=2000.0*)

NAME:

`radec_to_lb`

PURPOSE:

transform from equatorial coordinates to Galactic coordinates

INPUT:

ra - right ascension

dec - declination

$degree$ - (Bool) if True, ra and dec are given in degree and l and b will be as well

$epoch$ - epoch of ra, dec (right now only 2000.0 and 1950.0 are supported when not using astropy's transformations internally; when internally using astropy's coordinate transformations, epoch can be None for ICRS, 'JXXXX' for FK5, and 'BXXXX' for FK4)

OUTPUT:

l, b

For vector inputs $[:,2]$

HISTORY:

2009-11-12 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

2016-05-13 - Added support for using astropy's coordinate transformations and for non-standard epochs - Bovy (UofT)

galpy.util.bovy_coords.radec_to_custom

`galpy.util.bovy_coords.radec_to_custom(ra, dec, T=None, degree=False, epoch=2000.0)`

NAME:

`radec_to_custom`

PURPOSE:

transform from equatorial coordinates to a custom set of sky coordinates

INPUT:

`ra` - right ascension

`dec` - declination

`T`= matrix defining the transformation: `new_rect= T dot old_rect`, where `old_rect = [cos(dec)cos(ra),cos(dec)sin(ra),sin(dec)]` and similar for `new_rect`

`degree` - (Bool) if True, `ra` and `dec` are given in degree and `l` and `b` will be as well

OUTPUT:

custom longitude, custom latitude (with longitude -180 to 180)

For vector inputs `[:,2]`

HISTORY:

2009-11-12 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.rectgal_to_sphergal

`galpy.util.bovy_coords.rectgal_to_sphergal(X, Y, Z, vx, vy, vz, degree=False)`

NAME:

`rectgal_to_sphergal`

PURPOSE:

transform phase-space coordinates in rectangular Galactic coordinates to spherical Galactic coordinates (can take vector inputs)

INPUT:

`X` - component towards the Galactic Center (kpc)

`Y` - component in the direction of Galactic rotation (kpc)

`Z` - component towards the North Galactic Pole (kpc)

`vx` - velocity towards the Galactic Center (km/s)

`vy` - velocity in the direction of Galactic rotation (km/s)

vz - velocity towards the North Galactic Pole (km/s)

degree - (Bool) if True, return l and b in degrees

OUTPUT:

(l,b,d,vr,pmll x cos(b),pmbb) in (rad,rad,kpc,km/s,mas/yr,mas/yr)

HISTORY:

2009-10-25 - Written - Bovy (NYU)

galpy.util.bovy_coords.rect_to_cyl

`galpy.util.bovy_coords.rect_to_cyl(X, Y, Z)`

NAME:

rect_to_cyl

PURPOSE:

convert from rectangular to cylindrical coordinates

INPUT:

X, Y, Z - rectangular coordinates

OUTPUT:

R,phi,z

HISTORY:

2010-09-24 - Written - Bovy (NYU)

galpy.util.bovy_coords.rect_to_cyl_vec

`galpy.util.bovy_coords.rect_to_cyl_vec(vx, vy, vz, X, Y, Z, cyl=False)`

NAME:

rect_to_cyl_vec

PURPOSE:

transform vectors from rectangular to cylindrical coordinates vectors

INPUT:

vx -

vy -

vz -

X - X

Y - Y

Z - Z

cyl - if True, X,Y,Z are already cylindrical

OUTPUT:

vR,vT,vz

HISTORY:

2010-09-24 - Written - Bovy (NYU)

galpy.util.bovy_coords.rphi_to_dl_2d

`galpy.util.bovy_coords.rphi_to_dl_2d(R, phi, degree=False, ro=1.0, phio=0.0)`

NAME:

`rphi_to_dl_2d`

PURPOSE:

convert Galactocentric radius and azimuth to distance and Galactic longitude

INPUT:

R - Galactocentric radius

phi - Galactocentric azimuth [rad/deg if degree]

KEYWORDS:

degree= (False): phi is in degrees rather than rad

ro= (1) Galactocentric radius of the observer

phio= (0) Galactocentric azimuth of the observer [rad/deg if degree]

OUTPUT:

(d,l); phi in degree if degree

HISTORY:

2012-01-04 - Written - Bovy (IAS)

galpy.util.bovy_coords.Rz_to_coshucosv

`galpy.util.bovy_coords.Rz_to_coshucosv(R, z, delta=1.0, oblate=False)`

NAME:

`Rz_to_coshucosv`

PURPOSE:

calculate prolate confocal cosh(u) and cos(v) coordinates from R,z, and delta

INPUT:

R - radius

z - height

delta= focus

oblate= (False) if True, compute oblate confocal coordinates instead of prolate

OUTPUT:

(cosh(u),cos(v))

HISTORY:

2012-11-27 - Written - Bovy (IAS)

2017-10-11 - Added oblate coordinates - Bovy (UofT)

galpy.util.bovy_coords.Rz_to_uv

galpy.util.bovy_coords.**Rz_to_uv** (*R*, *z*, *delta*=1.0, *oblate*=False)

NAME:

Rz_to_uv

PURPOSE:

calculate prolate or oblate confocal u and v coordinates from R,z, and delta

INPUT:

R - radius

z - height

delta= focus

oblate= (False) if True, compute oblate confocal coordinates instead of prolate

OUTPUT:

(u,v)

HISTORY:

2012-11-27 - Written - Bovy (IAS)

2017-10-11 - Added oblate coordinates - Bovy (UofT)

galpy.util.bovy_coords.sphergal_to_rectgal

galpy.util.bovy_coords.**sphergal_to_rectgal** (*l*, *b*, *d*, *vr*, *pmll*, *pmbb*, *degree*=False)

NAME:

sphergal_to_rectgal

PURPOSE:

transform phase-space coordinates in spherical Galactic coordinates to rectangular Galactic coordinates (can take vector inputs)

INPUT:

l - Galactic longitude (rad)

b - Galactic latitude (rad)

d - distance (kpc)

vr - line-of-sight velocity (km/s)

pmll - proper motion in the Galactic longitude direction ($\mu_l \cos(b)$) (mas/yr)

pmbb - proper motion in the Galactic latitude (mas/yr)

degree - (bool) if True, l and b are in degrees

OUTPUT:

(X,Y,Z,vx,vy,vz) in (kpc,kpc,kpc,km/s,km/s,km/s)

HISTORY:

2009-10-25 - Written - Bovy (NYU)

galpy.util.bovy_coords.uv_to_Rz

`galpy.util.bovy_coords.uv_to_Rz(u, v, delta=1.0, oblate=False)`

NAME:

`uv_to_Rz`

PURPOSE:

calculate R and z from prolate confocal u and v coordinates

INPUT:

u - confocal u

v - confocal v

delta= focus

oblate= (False) if True, compute oblate confocal coordinates instead of prolate

OUTPUT:

(R,z)

HISTORY:

2012-11-27 - Written - Bovy (IAS)

2017-10-11 - Added oblate coordinates - Bovy (UofT)

galpy.util.bovy_coords.vrpmllpmbb_to_vxvyvz

`galpy.util.bovy_coords.vrpmllpmbb_to_vxvyvz(vr, pmll, pmbb, l, b, d, XYZ=False, degree=False)`

NAME:

`vrpmllpmbb_to_vxvyvz`

PURPOSE:

Transform velocities in the spherical Galactic coordinate frame to the rectangular Galactic coordinate frame (can take vector inputs)

INPUT:

vr - line-of-sight velocity (km/s)

pmll - proper motion in the Galactic longitude ($\mu_l \cos(b)$)(mas/yr)

pmbb - proper motion in the Galactic latitude (mas/yr)

l - Galactic longitude

b - Galactic latitude

d - distance (kpc)

XYZ - (bool) If True, then l,b,d is actually X,Y,Z (rectangular Galactic coordinates)

degree - (bool) if True, l and b are in degrees

OUTPUT:

(vx,vy,vz) in (km/s,km/s,km/s)

For vector inputs [:,3]

HISTORY:

2009-10-24 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.vRvz_to_pupv

`galpy.util.bovy_coords.vRvz_to_pupv(vR, vz, R, z, delta=1.0, oblate=False, uv=False)`

NAME:

`vRvz_to_pupv`

PURPOSE:

calculate momenta in prolate or oblate confocal u and v coordinates from cylindrical velocities vR,vz for a given focal length delta

INPUT:

vR - radial velocity in cylindrical coordinates

vz - vertical velocity in cylindrical coordinates

R - radius

z - height

delta= focus

oblate= (False) if True, compute oblate confocal coordinates instead of prolate

uv= (False) if True, the given R,z are actually u,v

OUTPUT:

(pu,pv)

HISTORY:

2017-11-28 - Written - Bovy (UofT)

galpy.util.bovy_coords.vxvyvz_to_galcencyl

`galpy.util.bovy_coords.vxvyvz_to_galcencyl(vx, vy, vz, X, Y, Z, vsun=[0.0, 1.0, 0.0], Xsun=1.0, Zsun=0.0, galcen=False, _extra_rot=True)`

NAME:

`vxvyvz_to_galcencyl`

PURPOSE:

transform velocities in XYZ coordinates (wrt Sun) to cylindrical Galactocentric coordinates for velocities

INPUT:

vx - U

vy - V

vz - W

X - X in Galactocentric rectangular coordinates

Y - Y in Galactocentric rectangular coordinates

Z - Z in Galactocentric rectangular coordinates

vsun - velocity of the sun in the GC frame ndarray[3]

Xsun - cylindrical distance to the GC

Zsun - Sun's height above the midplane

galcen - if True, then X,Y,Z are in cylindrical Galactocentric coordinates rather than rectangular coordinates

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

vRg, vTg, vZg

HISTORY:

2010-09-24 - Written - Bovy (NYU)

galpy.util.bovy_coords.vxvyvz_to_galcenrect

galpy.util.bovy_coords.**vxvyvz_to_galcenrect** (vx, vy, vz, vsun=[0.0, 1.0, 0.0], Xsun=1.0, Zsun=0.0, _extra_rot=True)

NAME:

vxvyvz_to_galcenrect

PURPOSE:

transform velocities in XYZ coordinates (wrt Sun) to rectangular Galactocentric coordinates for velocities

INPUT:

vx - U

vy - V

vz - W

vsun - velocity of the sun in the GC frame ndarray[3]

Xsun - cylindrical distance to the GC

Zsun - Sun's height above the midplane

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

[:,3]= vXg, vYg, vZg

HISTORY:

2010-09-24 - Written - Bovy (NYU)

2016-05-12 - Edited to properly take into account the Sun's vertical position; dropped Ysun keyword - Bovy (UofT)

2018-04-18 - Tweaked to be consistent with astropy's Galactocentric frame - Bovy (UofT)

galpy.util.bovy_coords.vxvyvz_to_vrpmllpmbb

galpy.util.bovy_coords.**vxvyvz_to_vrpmllpmbb** (vx, vy, vz, l, b, d, XYZ=False, degree=False)
NAME:

vxvyvz_to_vrpmllpmbb

PURPOSE:

Transform velocities in the rectangular Galactic coordinate frame to the spherical Galactic coordinate frame (can take vector inputs)

INPUT:

vx - velocity towards the Galactic Center (km/s)

vy - velocity in the direction of Galactic rotation (km/s)

vz - velocity towards the North Galactic Pole (km/s)

l - Galactic longitude

b - Galactic latitude

d - distance (kpc)

XYZ - (bool) If True, then l,b,d is actually X,Y,Z (rectangular Galactic coordinates)

degree - (bool) if True, l and b are in degrees

OUTPUT:

(vr,pmll x cos(b),pmbb) in (km/s,mas/yr,mas/yr); pmll = mu_l * cos(b)

For vector inputs[:,3]

HISTORY:

2009-10-24 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

galpy.util.bovy_coords.XYZ_to_galcencyl

galpy.util.bovy_coords.**XYZ_to_galcencyl** (X, Y, Z, Xsun=1.0, Zsun=0.0, _extra_rot=True)
NAME:

XYZ_to_galcencyl

PURPOSE:

transform XYZ coordinates (wrt Sun) to cylindrical Galactocentric coordinates

INPUT:

X - X

Y - Y

Z - Z

Xsun - cylindrical distance to the GC

Zsun - Sun's height above the midplane

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

R,phi,z

HISTORY:

2010-09-24 - Written - Bovy (NYU)

galpy.util.bovy_coords.XYZ_to_galcenrect

galpy.util.bovy_coords.XYZ_to_galcenrect(*X*, *Y*, *Z*, *Xsun*=1.0, *Zsun*=0.0, *_extra_rot*=True)

NAME:

XYZ_to_galcenrect

PURPOSE:

transform XYZ coordinates (wrt Sun) to rectangular Galactocentric coordinates

INPUT:

X - X

Y - Y

Z - Z

Xsun - cylindrical distance to the GC

Zsun - Sun's height above the midplane

_extra_rot= (True) if True, perform an extra tiny rotation to align the Galactocentric coordinate frame with astropy's definition

OUTPUT:

(Xg, Yg, Zg)

HISTORY:

2010-09-24 - Written - Bovy (NYU)

2016-05-12 - Edited to properly take into account the Sun's vertical position; dropped Ysun keyword - Bovy (UofT)

2018-04-18 - Tweaked to be consistent with astropy's Galactocentric frame - Bovy (UofT)

galpy.util.bovy_coords.XYZ_to_lbd

galpy.util.bovy_coords.XYZ_to_lbd(*X*, *Y*, *Z*, *degree*=False)

NAME:

XYZ_to_lbd

PURPOSE:

transform from rectangular Galactic coordinates to spherical Galactic coordinates (works with vector inputs)

INPUT:

X - component towards the Galactic Center (in kpc; though this obviously does not matter))

Y - component in the direction of Galactic rotation (in kpc)

Z - component towards the North Galactic Pole (kpc)

degree - (Bool) if True, return l and b in degrees

OUTPUT:

[l,b,d] in (rad or degree,rad or degree,kpc)

For vector inputs[:,3]

HISTORY:

2009-10-24 - Written - Bovy (NYU)

2014-06-14 - Re-written w/ numpy functions for speed and w/ decorators for beauty - Bovy (IAS)

3.5.5 galpy.util.bovy_ars.bovy_ars

`galpy.util.bovy_ars.bovy_ars(domain, isDomainFinite, abcissae, hx, hpx, nsamples=1, hx-params=(), maxn=100)`

bovy_ars: Implementation of the Adaptive-Rejection Sampling algorithm by Gilks & Wild (1992): Adaptive Rejection Sampling for Gibbs Sampling, Applied Statistics, 41, 337 Based on Wild & Gilks (1993), Algorithm AS 287: Adaptive Rejection Sampling from Log-concave Density Functions, Applied Statistics, 42, 701

Input:

domain - [.,.] upper and lower limit to the domain

isDomainFinite - [.,.] is there a lower/upper limit to the domain?

abcissae - initial list of abcissae (must lie on either side of the peak in hx if the domain is unbounded)

hx - function that evaluates $h(x) = \ln g(x)$

hpx - function that evaluates $hp(x) = d h(x) / d x$

nsamples - (optional) number of desired samples (default=1)

hxparams - (optional) a tuple of parameters for $h(x)$ and $h'(x)$

maxn - (optional) maximum number of updates to the hull (default=100)

Output:

list with nsamples of samples from $\exp(h(x))$

External dependencies:

math scipy scipy.stats

History: 2009-05-21 - Written - Bovy (NYU)

Acknowledging galpy

If you use galpy in a publication, please cite the following paper

- *galpy: A Python Library for Galactic Dynamics*, Jo Bovy (2015), *Astrophys. J. Supp.*, **216**, 29 (arXiv/1412.3451).

and link to <http://github.com/jobovy/galpy>. Some of the code's functionality is introduced in separate papers:

- `galpy.actionAngle.EccZmaxRperiRap` and `galpy.orbit.Orbit` methods with `analytic=True`: Fast method for computing orbital parameters from *this section*: please cite [Mackey & Bovy \(2018\)](#).
- `galpy.actionAngle.actionAngleAdiabatic`: please cite [Binney \(2010\)](#).
- `galpy.actionAngle.actionAngleStaeckel`: please cite [Bovy & Rix \(2013\)](#) and [Binney \(2012\)](#).
- `galpy.actionAngle.actionAngleIsochroneApprox`: please cite [Bovy \(2014\)](#).
- `galpy.df.streamdf`: please cite [Bovy \(2014\)](#).
- `galpy.df.streamgapdf`: please cite [Sanders, Bovy, & Erkal \(2016\)](#).

Please also send me a reference to the paper or send a pull request including your paper in the list of galpy papers on this page (this page is at `doc/source/index.rst`). Thanks!

Papers using galpy

galpy is described in detail in this publication:

- *galpy: A Python Library for Galactic Dynamics*, Jo Bovy (2015), *Astrophys. J. Supp.*, **216**, 29 (2015ApJS..216...29B).

The following is a list of publications using galpy; please let me (bovy at astro dot utoronto dot ca) know if you make use of galpy in a publication.

1. ***Tracing the Hercules stream around the Galaxy***, Jo Bovy (2010), *Astrophys. J.* **725**, 1676 (2010ApJ...725.1676B):
Uses what later became the orbit integration routines and Dehnen and Shu disk distribution functions.
2. ***The spatial structure of mono-abundance sub-populations of the Milky Way disk***, Jo Bovy, Hans-Walter Rix, Chao Liu, et al.
Employs galpy orbit integration in `galpy.potential.MWPotential` to characterize the orbits in the SEGUE G dwarf sample.
3. ***On the local dark matter density***, Jo Bovy & Scott Tremaine (2012), *Astrophys. J.* **756**, 89 (2012ApJ...756...89B):
Uses `galpy.potential` force and density routines to characterize the difference between the vertical force and the surface density at large heights above the MW midplane.
4. ***The Milky Way's circular velocity curve between 4 and 14 kpc from APOGEE data***, Jo Bovy, Carlos Allende Prieto, Timothy
Utilizes the Dehnen distribution function to inform a simple model of the velocity distribution of APOGEE stars in the Milky Way disk and to create mock data.
5. ***A direct dynamical measurement of the Milky Way's disk surface density profile, disk scale length, and dark matter profile at 4***
Makes use of potential models, the adiabatic and Staackel actionAngle modules, and the quasiisothermal DF to model the dynamics of the SEGUE G dwarf sample in mono-abundance bins.
6. ***The peculiar pulsar population of the central parsec***, Jason Dexter & Ryan M. O'Leary (2013), *Astrophys. J. Lett.*, **783**, L7
Uses galpy for orbit integration of pulsars kicked out of the Galactic center.
7. ***Chemodynamics of the Milky Way. I. The first year of APOGEE data***, Friedrich Anders, Christina Chiappini, Basilio X. San
Employs galpy to perform orbit integrations in `galpy.potential.MWPotential` to characterize the orbits of stars in the APOGEE sample.
8. ***Dynamical modeling of tidal streams***, Jo Bovy (2014), *Astrophys. J.*, **795**, 95 (2014ApJ...795...95B):
Introduces `galpy.df.streamdf` and `galpy.actionAngle`.

`actionAngleIsochroneApprox` for modeling tidal streams using simple models formulated in action-angle space (see the tutorial above).

9. ***The Milky Way Tomography with SDSS. V. Mapping the Dark Matter Halo***, Sarah R. Loebman, Zeljko Ivezic Thomas R. Quinn (2015), *Mon. Not. Roy. Astron. Soc.*, **451**, 1005 (2015MNRAS...451.1005L).
Uses `galpy.potential` functions to calculate the acceleration field of the best-fit potential in Bovy & Rix (2013) above.
10. ***The Proper Motion of the Galactic Center Pulsar Relative to Sagittarius A****, Geoffrey C. Bower, Adam Deller, Paul Demores (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Utilizes `galpy.orbit` integration in Monte Carlo simulations of the possible origin of the pulsar PSR J1745-2900 near the black hole at the center of the Milky Way.
11. ***The power spectrum of the Milky Way: Velocity fluctuations in the Galactic disk***, Jo Bovy, Jonathan C. Bird, Ana E. Garcia (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Uses `galpy.df.evolveddiskdf` to calculate the mean non-axisymmetric velocity field due to different non-axisymmetric perturbations and compares it to APOGEE data.
12. ***The LMC geometry and outer stellar populations from early DES data***, Eduardo Balbinot, B. X. Santiago, L. Girardi, et al. (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Employs `galpy.potential.MWPotential` as a mass model for the Milky Way to constrain the mass of the LMC.
13. ***Generation of mock tidal streams***, Mark A. Fardal, Shuiyao Huang, & Martin D. Weinberg (2015), *Mon. Not. Roy. Astron. Soc.*, **451**, 1005 (2015MNRAS...451.1005L).
Uses `galpy.potential` and `galpy.orbit` for orbit integration in creating a *particle-spray* model for tidal streams.
14. ***The nature and orbit of the Ophiuchus stream***, Branimir Sesar, Jo Bovy, Edouard J. Bernard, et al. (2015), *Astrophys. J.*, **804**, 100 (2015ApJ...804..100B).
Uses the `Orbit.fit` routine in `galpy.orbit` to fit the orbit of the Ophiuchus stream to newly obtained observational data and the routines in `galpy.df.streamdf` to model the creation of the stream.
15. ***Young Pulsars and the Galactic Center GeV Gamma-ray Excess***, Ryan M. O’Leary, Matthew D. Kistler, Matthew Kerr, & J. H. Simon (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Uses `galpy.orbit` integration and `galpy.potential.MWPotential2014` as part of a Monte Carlo simulation of the Galactic young-pulsar population.
16. ***Phase Wrapping of Epicyclic Perturbations in the Wobbly Galaxy***, Alexander de la Vega, Alice C. Quillen, Jeffrey L. Carlin, et al. (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Employs `galpy.orbit` integration, `galpy.potential` functions, and `galpy.potential.MWPotential2014` to investigate epicyclic motions induced by the pericentric passage of a large dwarf galaxy and how these motions give rise to streaming motions in the vertical velocities of Milky Way disk stars.
17. ***Chemistry of the Most Metal-poor Stars in the Bulge and the $z \sim 10$ Universe***, Andrew R. Casey & Kevin C. Schlaufman (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
This paper employs `galpy.orbit` integration in `MWPotential` to characterize the orbits of three very metal-poor stars in the Galactic bulge.
18. ***The Phoenix stream: a cold stream in the Southern hemisphere***, E. Balbinot, B. Yanny, T. S. Li, et al. (2015), *Astrophys. J.*, **820**, 58 (2016ApJ...820...58B).
19. ***Discovery of a Stellar Overdensity in Eridanus-Phoenix in the Dark Energy Survey***, T. S. Li, E. Balbinot, N. Mondrik, et al. (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Both of these papers use `galpy.orbit` integration to integrate the orbit of NGC 1261 to investigate a possible association of this cluster with the newly discovered Phoenix stream and Eridanus-Phoenix overdensity.
20. ***The Proper Motion of Palomar 5***, T. K. Fritz & N. Kallivayalil (2015), *Astrophys. J.*, **811**, 123 (2015ApJ...811..123F).
This paper makes use of the `galpy.df.streamdf` model for tidal streams to constrain the Milky Way’s gravitational potential using the kinematics of the Palomar 5 cluster and stream.
21. ***Spiral- and bar-driven peculiar velocities in Milky Way-sized galaxy simulations***, Robert J. J. Grand, Jo Bovy, Daisuke Kawata, et al. (2015), *ApJ*, **804**, 100 (2015ApJ...804..100B).
Uses `galpy.df.evolveddiskdf` to calculate the mean non-axisymmetric velocity field due to the bar in different parts of the Milky Way.
22. ***Vertical kinematics of the thick disc at 4.5 $R \sim 9.5$ kpc***, Kohei Hattori & Gerard Gilmore (2015), *Mon. Not. Roy. Astron. Soc.*, **451**, 1005 (2015MNRAS...451.1005L).
This paper uses `galpy.potential` functions to set up a realistic Milky-Way potential for investigating the kinematics of stars in the thick disk.

23. *Local Stellar Kinematics from RAVE data - VI. Metallicity Gradients Based on the F-G Main-sequence Stars*, O. Plevne, T. A.
 This paper employs galpy orbit integration in `MWPotential2014` to calculate orbital parameters for a sample of RAVE F and G dwarfs to investigate the metallicity gradient in the Milky Way.
24. *Dynamics of stream-subhalo interactions*, Jason L. Sanders, Jo Bovy, & Denis Erkal (2015), *Mon. Not. Roy. Astron. Soc.*, 45
 Uses and extends `galpy.df.streamdf` to build a generative model of the dynamical effect of sub-halo impacts on tidal streams. This new functionality is contained in `galpy.df.streamgapdf`, a subclass of `galpy.df.streamdf`, and can be used to efficiently model the effect of impacts on the present-day structure of streams in position and velocity space.
25. *Extremely metal-poor stars from the cosmic dawn in the bulge of the Milky Way*, L. M. Howes, A. R. Casey, M. Asplund, et al.
 Employs galpy orbit integration in `MWPotential2014` to characterize the orbits of a sample of extremely metal-poor stars found in the bulge of the Milky Way. This analysis demonstrates that the orbits of these metal-poor stars are always close to the center of the Milky Way and that these stars are therefore true bulge stars rather than halo stars passing through the bulge.
26. *Detecting the disruption of dark-matter halos with stellar streams*, Jo Bovy (2016), *Phys. Rev. Lett.*, 116, 121301 (2016PhRvL
 Uses galpy functions in `galpy.df` to estimate the velocity kick imparted by a disrupting dark-matter halo on a stellar stream. Also employs `galpy.orbit` integration and `galpy.actionAngle` functions to analyze N -body simulations of such an interaction.
27. *Identification of Globular Cluster Stars in RAVE data II: Extended tidal debris around NGC 3201*, B. Anguiano, G. M. De S
 Employs `galpy.orbit` integration to study the orbits of potential tidal-debris members of NGC 3201.
28. *Young and Millisecond Pulsar GeV Gamma-ray Fluxes from the Galactic Center and Beyond*, Ryan M. O’Leary, Matthew D
 Uses `galpy.orbit` integration in `MWPotential2014` for orbit integration of pulsars kicked out of the central region of the Milky Way.
29. *Abundances and kinematics for ten anticentre open clusters*, T. Cantat-Gaudin, P. Donati, A. Vallenari, R. Sordo, A. Bragag
 Uses `galpy.orbit` integration in `MWPotential2014` to characterize the orbits of 10 open clusters located toward the Galactic anti-center, finding that the most distant clusters have high-eccentricity orbits.
30. *A Magellanic Origin of the DES Dwarfs*, P. Jethwa, D. Erkal, & V. Belokurov (2016), *Mon. Not. Roy. Astron. Soc.*, 461, 221
 Employs the C implementations of `galpy.potentials` to compute forces in orbit integrations of the LMC’s satellite-galaxy population.
31. *PSR J1024-0719: A Millisecond Pulsar in an Unusual Long-Period Orbit*, D. L. Kaplan, T. Kupfer, D. J. Nice, et al. (2016), *Astrophys. J.*, 826, 86 (arXiv/1604.00131):
32. *A millisecond pulsar in an extremely wide binary system*, C. G. Bassa, G. H. Janssen, B. W. Stappers, et al. (2016), *Mon. Not.
 Both of these papers use `galpy.orbit` integration in `MWPotential2014` to determine the orbit of the milli-second pulsar PSR J10240719, a pulsar in an unusual binary system.*
33. *The first low-mass black hole X-ray binary identified in quiescence outside of a globular cluster*, B. E. Tetarenko, A. Bahrami
 This paper employs `galpy.orbit` integration of orbits within the position-velocity uncertainty ellipse of the radio source VLA J213002.08+120904 to help characterize its nature (specifically, to rule out that it is a magnetar based on its birth location).
34. *Action-based Dynamical Modelling for the Milky Way Disk*, Wilma H. Trick, Jo Bovy, & Hans-Walter Rix (2016), *Astrophys
 Makes use of potential models, the Staeckel `actionAngle` modules, and the quasiisothermal DF to develop a robust dynamical modeling approach for recovering the Milky Way’s gravitational potential from kinematics of disk stars.*
35. *A Dipole on the Sky: Predictions for Hypervelocity Stars from the Large Magellanic Cloud*, Douglas Boubert & N. W. Evans
 Uses `galpy.orbit` integration to investigate the orbits of hyper-velocity stars that could be ejected from the Large Magellanic Cloud and their distribution on the sky.
36. *Linear perturbation theory for tidal streams and the small-scale CDM power spectrum*, Jo Bovy, Denis Erkal, & Jason L. Sar
 Uses and extends `galpy.df.streamdf` and `galpy.df.streamgapdf` to quickly compute the

- effect of impacts from dark-matter subhalos on stellar streams and investigates the structure of perturbed streams and how this structure relates to the CDM subhalo mass spectrum.
37. ***Local Stellar Kinematics from RAVE data - VII. Metallicity Gradients from Red Clump Stars***, O. Onal Tas, S. Bilir, G. M. Se...
Employs `galpy.orbit` integration in `MWPotential2014` to calculate orbital parameters for a sample of red clump stars in RAVE to investigate the metallicity gradient in the Milky Way.
 38. ***Study of Eclipsing Binary and Multiple Systems in OB Associations IV: Cas OB6 Member DN Cas***, V. Bakis, H. Bakis, S. Bili...
Uses `galpy.orbit` integration in `MWPotential2014` to calculate the orbit and orbital parameters of the spectroscopic binary DN Cas in the Milky Way.
 39. ***The shape of the inner Milky Way halo from observations of the Pal 5 and GD-1 stellar streams***, Jo Bovy, Anita Bahmanyar,
Makes use of the `galpy.df.streamdf` model for a tidal stream to constrain the shape and mass of the Milky Way's dark-matter halo. Introduced `galpy.potential.TriaxialNFWPotential`.
 40. ***The Rotation-Metallicity Relation for the Galactic Disk as Measured in the Gaia DR1 TGAS and APOGEE Data***, Carlos All...
Employs orbit integration in `MWPotential2014` to calculate the orbits of a sample of stars in common between Gaia DR1's TGAS and APOGEE to study the rotation-metallicity relation for the Galactic disk.
 41. ***Detection of a dearth of stars with zero angular momentum in the solar neighbourhood***, Jason A. S. Hunt, Jo Bovy, & Raym...
Uses `galpy.orbit` integration in `MWPotential2014` plus a hard Galactic core to calculate the orbits of stars in the solar neighborhood and predict how many of them should be lost to chaos.
 42. ***Differences in the rotational properties of multiple stellar populations in M 13: a faster rotation for the “extreme” chemical s...***
Employs `galpy.orbit` integration in `MWPotential2014` to investigate the orbit of the globular cluster M13 and in particular whether escaping stars from the cluster could contaminate the measurement of the rotation of different populations in the cluster.
 43. ***Using the Multi-Object Adaptive Optics demonstrator RAVEN to observe metal-poor stars in and towards the Galactic Centre***,
Uses `galpy.orbit` integration in `MWPotential2014` to characterize the orbits of three very metal-poor stars observed toward the Galactic center, to determine whether they are likely bulge members.
 44. ***The Radial Velocity Experiment (RAVE): Fifth Data Release***, Andrea Kunder, Georges Kordopatis, Matthias Steinmetz, et al...
Employs `galpy.orbit` integration to characterize the orbits of stars in the RAVE survey.
 45. ***The Proper Motion of Pyxis: the first use of Adaptive Optics in tandem with HST on a faint halo object***, Tobias K. Fritz, Sean...
Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbit of the globular cluster Pyxis using its newly measured proper motion and to search for potential streams associated with the cluster.
 46. ***The Galactic distribution of X-ray binaries and its implications for compact object formation and natal kicks***, Serena Repetto...
Uses `galpy.orbit` integration in `MWPotential2014` and that of [Paczynski \(1990\)](#) to study the orbits of X-ray binaries under different assumptions about their formation mechanism and natal velocity kicks.
 47. ***Kinematics of Subluminous O and B Stars by Surface Helium Abundance***, P. Martin, C. S. Jeffery, Naslim N., & V. M. Wool...
Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbits of different types of low-mass core-helium-burning stars.
 48. ***Is there a disk of satellites around the Milky Way?***, Moupiya Maji, Qirong Zhu, Federico Marinacci, & Yuexing Li (2017), s...
Employs `galpy.orbit` integration in `MWPotential2014` to predict the future paths of 11 classical Milky-Way satellites to investigate whether they remain in a disk configuration.
 49. ***The devil is in the tails: the role of globular cluster mass evolution on stream properties***, Eduardo Balbinot & Mark Gieles (2...
Uses `galpy.orbit` integration in `MWPotential2014` of globular clusters in the Milky-Way halo. These integrations are used to investigate the clusters' mass loss due to tidal stripping, taking the effects of collisional dynamics in the cluster into account, and to evaluate the visibility of their (potential) tidal tails.
 50. ***Absolute Ages and Distances of 22 GCs using Monte Carlo Main-Sequence Fitting***, Erin M. O'Malley, Christina Gilligan, &...
Employs `galpy.orbit` integration in `MWPotential2014` of globular clusters in the Milky Way, to study their orbits and classify them as disk or halo clusters.

51. **Siriusly, a newly identified intermediate-age Milky Way stellar cluster: A spectroscopic study of Gaia I**, J. D. Simpson, G. M. ...
Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbit in the Milky Way potential of a newly-confirmed stellar cluster found in the Gaia data.
52. **Action-based Dynamical Modeling for the Milky Way Disk: The Influence of Spiral Arms**, Wilma H. Trick, Jo Bovy, Elena D. ...
Uses various potential models, the `Staeckel` `actionAngle` modules, and the quasiisothermal DF to test a robust dynamical modeling approach for recovering the Milky Way's gravitational potential from kinematics of disk stars against numerical simulations with spiral arms.
53. **A spectroscopic study of the elusive globular cluster ESO452-SC11 and its surroundings**, Andreas Koch, Camilla Juul Hansen ...
Employs `galpy.orbit` integration in `MWPotential2014` to investigate the orbit in the Milky Way potential of two candidate cluster members of the bulge globular cluster ESO452-SC11.
54. **A Halo Substructure in Gaia Data Release 1**, G. C. Myeong, N. W. Evans, V. Belokurov, S. E. Koposov, & J. L. Sanders (2017) ...
Uses `galpy.actionAngle.actionAngleAdiabatic` routines to compute the actions using the adiabatic approximation for 268,588 stars in *Gaia* DR1 TGAS with line-of-sight velocities from spectroscopic surveys. Detects a co-moving group of 14 stars on strongly radial orbits and computes their orbits using `MWPotential2014`.
55. **An artificial neural network to discover Hypervelocity stars: Candidates in Gaia DR1/ TGAS**, T. Marchetti, E. M. Rossi, G. F. ...
Uses `galpy.orbit` integration in a custom Milky-Way-like potential built from `galpy.potential` models to investigate the orbits of hypervelocity-star candidates in *Gaia* DR1.
56. **GalRotpy: an educational tool to understand and parametrize the rotation curve and gravitational potential of disk-like galaxies** ...
These authors build an interactive tool to decompose observed rotation curves into bulge, disk (Miyamoto-Nagai or exponential), and NFW halo components on top of `galpy.potential` routines.
57. **The AMBRE Project: formation and evolution of the Milky Way disc**, V. Grisoni, E. Spitoni, F. Matteucci, A. Recio-Blanco, ...
Uses `galpy` to compute orbital parameters for stars in the AMBRE sample of high-resolution spectra and uses these orbital parameters to aid in the comparison between the data and chemical-evolution models.
58. **Low-mass X-ray binaries from black-hole retaining globular clusters**, Matthew Giesler, Drew Clausen, & Christian D. Ott (2017) ...
Employs `galpy.orbit` integration in `MWPotential2014` to compute the orbits of MW globular clusters and simulated ejected BH low-mass X-ray binaries from these globular clusters to determine where they end up today.
59. **ESO452-SC11: The lowest mass globular cluster with a potential chemical inhomogeneity**, Jeffrey D. Simpson, Gayandhi De Silva ...
Uses `galpy.orbit` in `MWPotential2014` to compute the orbit of the MW bulge globular cluster ESO452-SC11.
60. **The Hercules stream as seen by APOGEE-2 South**, Jason A. S. Hunt, Jo Bovy, Angeles Pérez-Villegas, et al. (2017) *Mon. Not. R. Astr. Soc.* ...
Uses `galpy.df.evolveddiskdf` to compute the effect of a fast-rotating Milky-Way bar model on the velocity distribution in the disk of the Milky Way near the Sun to compare to the observed line-of-sight velocity distribution in APOGEE-2 South.
61. **Detailed chemical abundance analysis of the thick disk star cluster Gaia I**, Andreas Koch, Terese T. Hansen, & Andrea Kunjathoor ...
Employs `galpy.orbit` integration to compute the orbits of four red-giant members of the *Gaia* I Milky Way star cluster, finding that the orbits of these stars are similar to those of the oldest stars in the Milky Way's disk.
62. **Proper motions in the VVV Survey: Results for more than 15 million stars across NGC 6544**, R. Contreras Ramos, M. Zoccali ...
Uses `galpy.orbit` integration in `MWPotential2014` to calculate the orbit of NGC 6544, a Milky-Way globular cluster, using a newly determined proper motion, finding that it is likely a halo globular cluster based on its orbit.
63. **How to make a mature accreting magnetar**, A. P. Igoshev & S. B. Popov (2017) *Mon. Not. Roy. Astron. Soc.*, in press ([arXiv:1705.08811](#)) ...
Employs `galpy.orbit` integration of the magnetar candidate 4U 0114+65 in the potential model from Irrgang et al. (2013) to aid in the determination of its likely age.

64. *iota Horologii is unlikely to be an evaporated Hyades star*, I. Ramirez, D. Yong, E. Gutierrez, M. Endl, D. L. Lambert, J.-D. Uses `galpy.orbit` integration in `MWPotential2014` to determine the approximate orbit of the star *iota Horologii*, a planet-hosting suspected former member of the Hyades cluster, to investigate whether it could have coincided with the Hyades cluster in the past.
65. *Confirming chemical clocks: asteroseismic age dissection of the Milky Way disk(s)*, V. Silva Aguirre, M. Bojsen-Hansen, D. S. Employs `galpy.orbit` integration in `MWPotential2014` to compute the orbits of a sample of 1989 red giants with spectroscopic and asteroseismic data from the APOKASC catalog, to shed light on the properties of stellar populations defined by age and metallicity.
66. *The universality of the rapid neutron-capture process revealed by a possible disrupted dwarf galaxy star*, Andrew R. Casey & Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbit and its uncertainty of 2MASS J151113.24–213003.0, an extremely metal-poor field star with measureable r-process abundances, and of other similar metal-poor stars. The authors find that all of these stars are on highly eccentric orbits, possibly indicating that they originated in dwarf galaxies.
67. *The Gaia-ESO Survey: Churning through the Milky Way*, M. R. Hayden, A. Recio-Blanco, P. de Laverny, et al. (2017) *Astro* Employs `galpy.orbit` integration in `MWPotential2014` to study the orbital characteristics (eccentricity, pericentric radius) of a sample of 2,364 stars observed in the Milky Way as part of the Gaia-ESO survey.
68. *The Evolution of the Galactic Thick Disk with the LAMOST Survey*, Chengdong Li & Gang Zhao (2017) *Astrophys. J.*, 850, Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbital characteristics (eccentricity, maximum height above the plane, angular momentum) of a sample of about 2,000 stars in the thicker-disk component of the Milky Way.
69. *The Orbit and Origin of the Ultra-faint Dwarf Galaxy Segue 1*, T. K. Fritz, M. Lokken, N. Kallivayalil, A. Wetzel, S. T. Lind Employs `galpy.orbit` integration in `MWPotential2014` and a version of this potential with a more massive dark-matter halo to investigate the orbit and origin of the dwarf-spheroidal galaxy Segue 1 using a newly measured proper motion with SDSS and LBC data.
70. *Prospects for detection of hypervelocity stars with Gaia*, T. Marchetti, O. Contigiani, E. M. Rossi, J. G. Albert, A. G. A. Brov Uses `galpy.orbit` integration in a custom Milky-Way-like potential built from `galpy.potential` models to create mock catalogs of hypervelocity stars in the Milky Way for different ejection mechanisms and study the prospects of their detection with *Gaia*.
71. *The AMBRE project: The thick thin disk and thin thick disk of the Milky Way*, Hayden, M. R., Recio-Blanco, A., de Laverny Employs `galpy.orbit` integration in `MWPotential2014` to characterize the orbits of 494 nearby stars analyzed as part of the AMBRE project to learn about their distribution within the Milky Way.
72. *KELT-21b: A Hot Jupiter Transiting the Rapidly-Rotating Metal-Poor Late-A Primary of a Likely Hierarchical Triple System*, Uses `galpy.orbit` integration in `MWPotential2014` to investigate the Galactic orbit of KELT-21b, a hot jupiter around a low-metallicity A-type star.
73. *GalDynPsr: A package to estimate dynamical contributions in the rate of change of the period of radio pulsars*, Dhruv Pathal Presents a python package to compute contributions to the GR spin-down of pulsars from the differential galactic acceleration between the Sun and the pulsar. The package uses `MWPotential2014` and `galpy.potential` functions to help compute this.
74. *Local Stellar Kinematics from RAVE data – VIII. Effects of the Galactic Disc Perturbations on Stellar Orbits of Red Clump St* Employs `galpy.orbit` integration in `MWPotential2014` and the non-axisymmetric `DehnenBarPotential` and `SteadyLogSpiralPotential` to study the orbits of Milky-Way red-clump stars.
75. *The VMC survey XXVIII. Improved measurements of the proper motion of the Galactic globular cluster 47 Tucanae*, F. Niede Uses `galpy.orbit` integration in `MWPotential2014` to investigate the orbit of the cluster 47 Tuc from a newly measured proper motion, finding that the orbit has an eccentricity of about 0.2 and reaches up to 3.6 kpc above the Galactic midplane.

76. *Characterising Open Clusters in the solar neighbourhood with the Tycho-Gaia Astrometric Solution*, T. Cantat-Gaudin, A. V. Employs `galpy.orbit` integration in `MWPotential2014` to investigate the orbits of a sample of 128 open clusters with astrometry from the Tycho-Gaia Astrometric Solution.
77. *Fast estimation of orbital parameters in Milky-Way-like potentials*, J. Ted Mackereth & Jo Bovy (2018) *Mon. Not. Roy. Astro* Introduces the `galpy.actionAngle.EccZmaxRperiRap` and related `galpy.orbit.Orbit` methods for the fast estimation of the orbital parameters (eccentricity, maximum vertical excursion, and peri/apocenter) in Milky-Way potentials. See [this section](#) of the documentation for more info.
78. *HI Kinematics Along The Minor Axis of M82*, Paul Martini, Adam K. Leroy, Jeffrey G. Mangum, Alberto Bolatto, Katie M. Use `galpy.potential` components to create a mass model for M82 that consists of a `HernquistPotential` bulge, `MN3ExponentialDiskPotential` disk, and `NFWPotential` dark-matter halo by matching photometric and rotation-curve data.
79. *Radial velocities of RR Lyrae stars in and around NGC 6441*, Andrea Kunder, Arthur Mills, Joseph Edgecomb, et al. (2018) Employs `galpy.orbit` integration in `MWPotential2014` to investigate a possible tidal stream for the globular cluster NGC 6441.
80. *The 4:1 Outer Lindblad Resonance of a long slow bar as a potential explanation for the Hercules stream*, Jason A. S. Hunt & Uses `galpy.potential` modeling tools (`CosmphiDiskPotential` and `wrap`pers to grow it [`DehnenSmoothWrapperPotential`] and make it rotate [`SolidBodyRotationWrapperPotential`]) and disk models in `galpy.df` to study the effect of the Milky Way bar's hexadecapole ($m=4$) moment on the velocity distribution of stars in the disk of the Milky Way, finding that it can have a big effect near the 4:1 outer Lindblad resonance, which may be close to the Sun.
81. *KELT-22Ab: A Massive Hot Jupiter Transiting a Near Solar Twin*, Jonathan Labadie-Bartz, Joseph E. Rodriguez, Keivan C. Employs `galpy.orbit` integration in `MWPotential2014` to explore the orbit of KELT-22A, the G2V host star of a transiting hot-jupiter exoplanet, in the Milky Way.
82. *Dissecting stellar chemical abundance space with t-SNE*, Friedrich Anders, Cristina Chiappini, Basílio X. Santiago, Gal Ma Uses the `galpy.actionAngle.EccZmaxRperiRap` and related `galpy.orbit.Orbit` methods introduced in [Mackereth & Bovy \(2018\)](#) to compute the eccentricity and maximum vertical excursion of a sample of solar-neighborhood stars with detailed chemical abundances, ages, and Galactic kinematics.
83. *Chemo-kinematics of the Milky Way from the SDSS-III MARVELS Survey*, Nolan Grieves, Jian Ge, Neil Thomas, et al. (2018) Employs `galpy.orbit` integration in a Milky-Way like potential (Miyamoto-Nagai disk, Hernquist bulge, and NFW dark-matter halo) to determine the eccentricity, maximum vertical excursion, and median Galactocentric radius of a sample of a few thousand solar neighborhood stars with spectra from the SDSS MARVELS survey.
84. *The demographics of neutron star - white dwarf mergers: rates, delay-time distributions and progenitors*, S. Toonen, H.B. Per Uses `galpy.potential` methods to build mass models for elliptical, dwarf elliptical, plus normal and bulgeless disk galaxies and `galpy.orbit` integration to integrate the orbits of neutron-star-white-dwarf binaries kicked by the supernova explosion that created the neutron star, to investigate the spatial distribution of such binaries at the time of their mergers.
85. *Absolute HST Proper Motion (HSTPROMO) Catalog of Distant Milky Way Globular Clusters: Three-dimensional Systemic V* Employs `galpy.orbit` integration in a custom Milky-Way-like potential (a scaled version of `MWPotential2014`) to study the orbits of 20 globular clusters in the Milky Way halo for which the authors measured new proper motions.
86. *Probing the nature of dark matter particles with stellar streams*, Nilanjan Banik, Gianfranco Bertone, Jo Bovy, & Nassim B. Uses `galpy.df.streamdf` and the [galpy extension](#) `galpy.df.streampepperdf` to compute the effect of impacts from dark-matter subhalos on stellar streams when dark matter is warm, when there are many fewer such subhalos, and compares this to the standard CDM case.
87. *Connecting the Milky Way potential profile to the orbital timescales and spatial structure of the Sagittarius Stream*, Mark A. I. Employs the `galpy.potential` module to build models for the Milky Way's gravitational potential

and uses `galpy.orbit` to integrate the orbits of members of the Sgr stream in a particle-spray model of the Sgr stream, to investigate how the Sgr stream is sensitive to the Milky Way's potential.

At this point, a flood of papers using `galpy` appeared, which we will not attempt to summarize:

88. *Evidence for accreted component in the Galactic discs*, Q. F. Xing. & G. Zhao (2018) *Mon. Not. Roy. Astron. Soc.*, in press ([2018MNRAS.476.5388X](#))
89. *The GALAH survey: An abundance, age, and kinematic inventory of the solar neighbourhood made with TGAS*, S. Buder, K. Lind, M. K. Ness, et al. (2018) *Astron. & Astrophys.*, submitted ([arXiv/1804.05869](#))
90. *The GALAH survey: Co-orbiting stars and chemical tagging*, Jeffrey D. Simpson, Sarah L. Martell, Gary Da Costa, et al. (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1804.05894](#))
91. *Correlations between age, kinematics, and chemistry as seen by the RAVE survey*, Jennifer Wojno, Georges Kordopatis, Matthias Steinmetz, et al. (2018) *Mon. Not. Roy. Astron. Soc.*, in press ([arXiv/1804.06379](#))
92. *On the kinematics of a runaway Be star population*, Douglas Boubert & N. Wyn Evans (2018) *Mon. Not. Roy. Astron. Soc.*, in press ([arXiv/1804.05849](#))
93. *Unbiased TGAS×LAMOST distances and the role of binarity*, Johanna Coronado, Hans-Walter Rix, & Wilma H. Trick (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1804.07760](#))
94. *The Origin of the 300 km/s Stream Near Segue 1*, Sal Wanying Fu, Joshua D. Simon, Matthew Shetrone, et al. (2018) *Astrophys. J.*, submitted ([arXiv/1804.08622](#))
95. *Anatomy of the hyper-runaway star LP 40-365 with Gaia*, R. Raddi, M. A. Hollands, B. T. Gaensicke, D. M. Townsley, J. J. Hermes, N. P. Gentile Fusillo, & D. Koester (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1804.09677](#))
96. *Revisiting hypervelocity stars after Gaia DR2*, Douglas Boubert, James Guillochon, Keith Hawkins, Idan Ginsburg, & N. Wyn Evans (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1804.10179](#))
97. *Gaia Proper Motions and Orbits of the Ultra-Faint Milky Way Satellites*, Joshua D. Simon (2018) *Astrophys. J.*, submitted ([arXiv/1804.10230](#))
98. *Three Hypervelocity White Dwarfs in Gaia DR2: Evidence for Dynamically Driven Double-Degenerate Double-Detonation Type Ia Supernovae*, Ken J. Shen, Douglas Boubert, Boris T. Gänsicke, et al. (2018) *Astrophys. J.*, submitted ([arXiv/1804.11163](#))
99. *Gaia DR2 Proper Motions of Dwarf Galaxies within 420 kpc: Orbits, Milky Way Mass, Tidal Influences, Planar Alignments, and Group Infall*, T. K. Fritz, G. Battaglia, M. S. Pawlowski, N. Kallivayalil, R. van der Marel, T. S. Sohn, C. Brook, & G. Besla (2018) *Astron. & Astrophys.*, submitted ([arXiv/1805.00908](#))
100. *The Lives of Stars: Insights From the TGAS-RAVE-LAMOST Dataset*, John J. Vickers & Martin C. Smith (2018) *Astrophys. J.*, in press ([arXiv/1805.02332](#))
101. *High precision pulsar timing and spin frequency second derivatives*, X. J. Liu, C. G. Bassa, & B. W. Stappers (2018) *Mon. Not. Roy. Astron. Soc.*, in press ([arXiv/1805.02892](#))
102. ***The Galactic Disc in Action Space as seen by Gaia DR2*, Wilma H. Trick, Johanna Coronado, Hans-Walter Rix (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1805.08481](#)):**
Uses `galpy.actionAngle.actionAngleStaeckel` to compute the actions in `galpy.potential.MWPotential2014` of stars in the extended solar neighborhood in *Gaia* DR2.
103. ***Tidal ribbons*, Walter Dehnen & Hasanuddin (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1805.08481](#)):**
Employs `galpy.actionAngle.actionAngleStaeckel` to compute the actions and frequencies of a disrupted satellite on an orbit close to the Galactic disk, to investigate the difference between the debris from such an orbit and from an orbit well within the halo.
104. *Apocenter Pile-Up: Origin of the Stellar Halo Density Break*, Alis J. Deason, Vasily Belokurov, Sergey E. Koposov, & Lachlan Lancaster (2018) *Astrophys. J. Lett.*, submitted ([arXiv/1805.10288](#))

105. *Bootes III is a disrupting dwarf galaxy associated with the Styx stellar stream*, Jeffrey L. Carlin & David J. Sand (2018) *Astrophys. J.*, submitted ([arXiv/1805.11624](#))
106. *Proper motions of Milky Way Ultra-Faint satellites with Gaia DR2 \times DES DR1*, Andrew B. Pace & Ting S. Li (2018) *Astrophys. J.*, submitted ([arXiv/1806.02345](#))
107. ***Transient spiral structure and the disc velocity substructure in Gaia DR2***, Jason A. S. Hunt, Jack Hong, Jo Bovy, Daisuke Kaibet, & David H. Weinbaum
Builds a simple model of co-rotating, transient spiral structure using the newly implemented `SpiralArmsPotential` and the `CorotatingRotationWrapperPotential` (and using `DehnenSmoothWrapperPotential` to make it transient). Then uses `galpy.df.evolveddiskdf` to compute the effect of such spiral structure on the kinematics in the extended solar neighborhood, finding qualitative agreement with the kinematics observed in *Gaia* DR2.
108. *Galactic Archeology with the AEGIS Survey: The Evolution of Carbon and Iron in the Galactic Halo*, Jinmi Yoon, Timothy C. Beers, Sarah Dietz, Young Sun Lee, Vinicius M. Placco, Gary Da Costa, Stefan Keller, Christopher I. Owen, & Mahavir Sharma (2018) *Astrophys. J.*, in press ([arXiv/1806.04738](#))
109. *The Formation and Evolution of Galactic Disks with APOGEE and the Gaia Survey*, Chengdong Li, Gang Zhao, Meng Zhai, & Yunpeng Jia (2018) *Astrophys. J.*, in press ([2018ApJ...860...53L](#))
110. *Spectroscopy of Dwarf Stars Around the North Celestial Pole*, Sarunas Mikolaitis, Grazina Tautvaisiene, Arnas Drazdauskas, & Renata Minkeviciute (2018) *Publ. Astron. Soc. Pacific*, in press ([2018PASP..130g4202M](#))
111. *The Study of Galactic Disk Kinematics with SCUSS and SDSS Data*, Xiyan Peng, Zhenyu Wu, Zhaoxiang Qi, Cuihua Du, Jun Ma, Xu Zhou, Yunpeng Jia, & Songhu Wang (2018) *Publ. Astron. Soc. Pacific*, in press ([2018PASP..130g4102P](#))
112. *Common origin for Hercules-Aquila and Virgo Clouds in Gaia DR2*, Iulia T. Simion, Vasily Belokurov, & Sergey E. Koposov (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1807.01335](#))
113. *The GALAH Survey and Gaia DR2: (Non)existence of five sparse high-latitude open clusters*, Janez Kos, Gayandhi de Silva, Joss Bland-Hawthorn, et al. (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1807.00822](#))
114. *Testing the universality of free fall by tracking a pulsar in a stellar triple system*, Anne M. Archibald, Nina V. Gusinskaia, Jason W. T. Hessels, Adam T. Deller, David L. Kaplan, Duncan R. Lorimer, Ryan S. Lynch, Scott M. Ransom, & Ingrid H. Stairs (2018) *Nature*, in press ([arXiv/1807.02059](#))
115. *On measuring the Galactic dark matter halo with hypervelocity stars*, O. Contigiani, E. M. Rossi, T. Marchetti (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1807.04468](#))
116. *Constraining the Milky Way Halo Potential with the GD-1 stellar stream*, Khyati Malhan & Rodrigo A. Ibata (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1807.05994](#))
117. *Interstellar magnetic cannon targeting the Galactic halo : A young bubble at the origin of the Ophiuchus and Lupus molecular complexes*, J.-F. Robitaille, A. M. M. Scaife, E. Carretti, M. Haverkorn, R. M. Crocker, M. J. Kesteven, S. Poppi, & L. Staveley-Smith (2018) *Astron. & Astrophys.*, submitted ([arXiv/1807.04054](#))
118. *Emergence of the Gaia Phase Space Spirals from Bending Waves*, Keir Darling & Lawrence M. Widrow (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1807.11516](#))
119. *Pristine Dwarf-Galaxy Survey I: A detailed photometric and spectroscopic study of the very metal-poor Draco II satellite*, Nicolas Longeard, Nicolas Martin, Else Starkenburg, et al. (2018) *Mon. Not. Roy. Astron. Soc.*, in press ([arXiv/1807.10655](#))
120. *The origin of accreted stellar halo populations in the Milky Way using APOGEE, Gaia, and the EAGLE simulations*, J. Ted Mackereth, Ricardo P. Schiavon, Joel Pfeffer, et al. (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1808.00968](#))
121. *The vertical motion history of disk stars throughout the Galaxy*, Yuan-Sen Ting & Hans-Walter Rix (2018) *Astrophys. J.*, submitted ([arXiv/1808.03278](#))

122. *A kinematical age for the interstellar object 1I/'Oumuamua*, F. Almeida-Fernandes & H. J. Rocha-Pinto (2018) *Mon. Not. Roy. Astron. Soc.*, submitted ([arXiv/1808.03637](#))
123. *Masses and ages for metal-poor stars: a pilot program combining asteroseismology and high-resolution spectroscopic follow-up of RAVE halo stars*, M. Valentini, C. Chiappini, D. Bossini, et al. (2018) *Astron. & Astrophys.*, submitted ([arXiv/1808.08569](#))
124. *On the radial metallicity gradient and radial migration effect of the Galactic disk*, Yunpeng Jia, Yuqin Chen, Gang Zhao, Xiangxiang Xue, Jingkun Zhao, Chengqun Yang, & Chengdong Li (2018) *Astrophys. J.*, in press ([arXiv/1808.05386](#))

CHAPTER 6

Indices and tables

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Symbols

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